



This is the accepted manuscript made available via CHORUS. The article has been published as:

# Reactivity of Xenon with Ice at Planetary Conditions

Chrystèle Sanloup, Stanimir A. Bonev, Majdi Hochlaf, and Helen E. Maynard-Casely

Phys. Rev. Lett. **110**, 265501 — Published 24 June 2013

DOI: 10.1103/PhysRevLett.110.265501

# Reactivity of Xenon with Ice at planetary conditions

### Chrystèle Sanloup

UPMC Univ Paris 06, UMR 7193, ISTEP, 75005 Paris, France; now at School of Physics and Astronomy and Center for Science at Extreme Conditions, University of Edinburgh, Edinburgh, EH9 3JZ, UK.

#### Stanimir A. Bonev

Lawrence Livermore National Laboratory, Livermore, CA 94550, US and Department of Physics, Dalhousie University, Halifax, NS B3H 3J5, Canada

#### Majdi Hochlaf

Université Paris-Est, Laboratoire Modélisation et Simulation MultiEchelle, MSME UMR 8208 CNRS, 77454 Marne-La-Vallée, France

## Helen E. Maynard-Casely

School of Chemistry, University of Edinburgh, UK; now at Australian Synchrotron (Dated: May 30, 2013)

We report results from high pressure and temperature experiments that provide evidence for the reactivity of xenon with water ice at pressure above 50 GPa and temperature of 1,500 K – conditions that are found in the interiors of Uranus and Neptune. The x-ray data are sufficient to determine a hexagonal lattice with four Xe atoms per unit cell and several possible distributions of O atoms. The measurements are supplemented with ab initio calculations, on the basis of which a crystallographic structure with a  $Xe_4O_{12}H_{12}$  primitive cell is proposed. The newly discovered compound is formed in the stability fields of superionic ice and  $\eta$ -O<sub>2</sub>, and has the same oxygen sub-network as the latter. Furthermore, it has a weakly metallic character and likely undergoes sub-lattice melting of the H subsystem. Our findings indicate that Xe is expected to be depleted in the atmospheres of the giant planets as a result of sequestration at depth.

Knowledge of the chemistry of xenon (Xe) with planetary materials under high pressure (P) and temperature (T) conditions is a pre-requisite for understanding the abundance of Xe and its isotopic ratios in the atmospheres of giant planets. The noble gases Ne, Ar, Kr, Xe are indeed among the most critical heavy elements [1], as their abundances constrain the models for giant planet formation and the origin of their atmospheres. The field of Xe chemistry at ambient P was initiated with the synthesis of fluoride salts [2] and later extended to H, C, O, N, S, other halogens, and even metals (Au, Hg), so that close to a hundred Xe compounds are currently known [3, 4]. However, the use of pressure as a means to enforce Xe to bond to other elements has seldom been investigated [5]. Geophysicists first started to explore this track as a potential explanation for the observed Xe deficiency in the terrestrial and martian atmospheres [6]. Xe-Fe compounds were shown to be unstable up to 70 GPa [7], although up to 0.8 mol% Xe could be alloyed to Fe at 300 GPa [8]. It was later proposed that Xe depletion from the Earth's atmosphere could occur by Xe substitution for Si in the silicates network under the P-Tconditions of the deep crust and mantle, therefore forming Xe-doped oxides [9–11]; a synthesis shortly followed by that of pure  $XeO_2$  at ambient P [12].

Water ice is one of the most prevalent substances in the solar system, with the majority of it existing at high P and T conditions in the interiors of giant planets. Xe is among the gases that stabilizes clathrate hydrates structure through van der Waals interactions. However, Xe hydrates are stable only up to 2.5 GPa, before dissociating into Xe plus ice VII [13]. The chemistry of Xe with water has nonetheless been successfully explored by UV photolysis of water in solid Xe at ambient pressure. Among the compounds hitherto obtained are HXeOH [14] and, the latest one to date, HXeOXeH [15], which results from the insertion of Xe atoms in the water molecule. These results reopened the perspective to synthesize covalent compounds in the Xe-H<sub>2</sub>O system by applying extreme P-T conditions.

In this Letter we report the first experimental evidence for the reactivity of Xe with water ice at conditions found in the interior of giant planets. The resulting crystalline structure is resolved from x-ray diffraction data with additional input from ab-initio calculations. Our analysis indicates the participation of H in the structure formation, and brackets for the H content are given with an emphasis on the  $\rm Xe_4O_{12}H_{12}$  structure.

High P-high T conditions were produced in laserheated diamond-anvil cells and x-ray diffraction experiments were carried in situ on ID27 beamline of the ESRF. A ring was laser-cut in a 2.5  $\mu$ m thick platinum (Pt) foil and inserted in a 80  $\mu$ m hole drilled in a rhenium gasket on top of two ruby spheres, Pt acting as a coupler with the IR laser. Milli-Q deionised water was then loaded in the sample chamber, and Xe added cryogenically in a N<sub>2</sub> atmosphere. The sample did not contain N<sub>2</sub> after closing the cell as attested by the absence of the N<sub>2</sub>

vibron on Raman spectra. The pressure was measured from the ruby fluorescence signal [16], and the temperature was calculated from the thermal equation of state of Pt [17]. Temperatures were also measured from the sample thermal emission spectra. However, volumes measured from the x-ray diffraction patterns show that the sample T was lower than the surface T as previously shown when a metallic laser coupler is used [18]. Data were collected on a MAR-CCD using a monochromatic x-ray beam ( $\lambda$ =0.3738 Å) focussed to a 3  $\mu$ m×4  $\mu$ m area on the sample. We carried out experiments at P ranging from 17 GPa to 80 GPa and T below the melting point of water ice.

During laser heating at 50 GPa, we observed two phases above 1.500 K in addition to Xe. Pt and bcc ice and will refer to them as phase-1 and phase-2 (Fig.1). X-ray diffraction peaks from phase-1 can be assigned to a face-centered cubic structure. Phase-1 is not quenchable back to room T and is equally observed in blank experiments on Pt-water mixtures. No blank experiment was conducted on the Xe-Pt system as it has been experimentally explored at P-T conditions covering those of the present work and no reaction was observed [7]. Phase-2 is stable back to room T and was systematically synthesized upon heating up to 80 GPa, the maximum P investigated here. The x-ray diffraction signal from phase-2 is indexed by a hexagonal cell with  $a=b=5.0539\pm0.0003$  Å and  $c = 8.210 \pm 0.001 \text{ Å}$  at 58 GPa and 1,500 K. Its compressibility as obtained from a second-order Birch-Murnaghan equation of state is  $77 \pm 5$  GPa at room T (fitted with zero pressure volume parameter  $V_0=255.8$ Å<sup>3</sup>), and  $67\pm5$  GPa at 1,500 K (V<sub>0</sub>=269.8 Å<sup>3</sup>). The thermal expansion coefficient at high P is  $1.8\pm0.1\times10^{-5}$  K<sup>-1</sup> (Fig.2).

The x-ray diffraction pattern at 58 GPa and 1,500 K has a sufficiently high powder-quality pattern for the atomic positions of the Xe and oxygen (O) to be refined; P 63/mmc is the highest-symmetry space-group allowed. From density considerations, four atoms of Xe can be accommodated in the unit cell. The solution is however not unique for the O atoms, with three possibilities equallywell satisfying the x-ray pattern intensities. The three possible structures are with either 12, 14 or 16 O atoms per unit cell. In order to determine the correct solution for O atoms and to investigate the presence of hydrogen in the structure, we have carried out a theoretical analysis.

First principles density functional theory (DFT) calculations were carried with VASP [19] using projector augmented wave pseudopotentials with 8 and 6 valence electrons for the Xe and O atoms, respectively, and a 500 eV plane wave cut off. Full structural optimizations were performed with a uniform  $6\times6\times6$  k-point grid sampling of the Brillouin zone, which ensures the desired degree of convergence (meV/atom). The choice of exchange-correlation functional can lead to significant variations in the specific volume of the Xe-O-H systems being studied. In order to provide an estimate for the theoretical

uncertainties originating from the exchange-correlation approximation, we have carried out calculations within the local density approximation (LDA), the Perdew-Burke-Ernzerhof [20] generalized gradient approximation (PBE), and the PBE revised for solids [21] (PBEsol). Additionally, for selected structures van der Waals interactions were included using the vdW-D2 method [22] as implemented in VASP. However, this did not change the results.

Given the fact the newly synthesized structure is thermodynamically stable only at high T, as well as its complexity, we have not attempted to perform a complete phase space crystalline search. Instead, we have limited the theoretical analysis to within the constraints imposed by the experimental data. Starting from the experimental lattice parameters and refined atomic positions, structural relaxations were initially performed on systems with Xe and O atoms only. The optimizations with 16, 14 and 12 O atoms resulted in theoretically stable structures that differ significantly from the experimental fits – respectively +1%, -7% and -5% in the a lattice parameter, +6%, +8% and +9% in the c lattice parameter (within PBE). The internal O coordinates also change significantly. These structural differences, which are accompanied by large energy differences on the order of tens of eV/atom, are beyond the uncertainties of both the theoretical and experimental methods. The fact that no H-free structure could be identified is in agreement with recent theoretical investigation [23] where no thermodynamically stable xenon oxides were found below 83 GPa.

Analysis for the effect of hydrogen was carried out by initially placing H atoms into the system both at random positions and at selected high-symmetry points. This was followed by full structural optimization at 58 GPa. To ensure that the entire phase space for H positions had been sampled, we also performed selective molecular dynamics where the H subsystem was heated while Xe and O atoms were held fixed. This procedure was repeated with 4, 12 and 24 H atoms. For the structures with 16 and 14 O atoms per unit cell, adding H atoms increases even further the discrepancy with the experimental fits. Therefore, the structures with 16 and 14 O atoms can be ruled out conclusively. On the other hand, introducing H in calculations with 12 O atoms per unit cell yields structures that match the x-ray data. The best fitting result is for 12 H atoms, for which the finite temperature lattice parameters, optimized to give an isotropic stress tensor at each P-300 K point, are compared with the measured ones (Fig.2). Using 4 H atoms, the calculated structure has an excellent agreement with the measured volume (Fig.2); however, the H coordinates cannot be reproduced within any hexagonal space-group. The structures with 24 H atoms do not give satisfactory agreement with the experimental cell parameters. It is nonetheless difficult to determine theoretically the exact H content - between 4 and 12 atoms per unit cell. In fact, it may be varied due the diffusive character of H at high temperature as described below. To answer these questions requires a full free energy thermodynamic stability analysis at 1,500 K, which goes beyond the scope of this paper.

The optimized Xe and O atomic positions further refined for the  $\rm Xe_4O_{12}$   $\rm H_{12}$  structure by a Rietveld fit to the data (Fig.3) are a 12(k) site (x, y, z) for O with  $\rm x=0.157\pm0.002$ ,  $\rm y=0.314\pm0.003$ ,  $\rm z=0.622\pm0.002$ , and a  $\rm 4(\it f)$  site  $(1/\it 3, 2/\it 3, z)$  for Xe with  $\rm z=0.071\pm0.001$ . The H atoms, according to the theoretical calculations, are located on a  $12(\it j)$  site  $(x, y, 1/\it 4)$  with  $\rm x=0.9427\pm0.002$ ,  $\rm y=0.2886\pm0.004$ . The structure consists of two  $\rm Xe_2O_6H_6$  units per unit cell (Fig.3 inset), and the Xe-O distance is  $\rm 2.21\pm0.01$  Å . This distance is similar to that observed in the HXeOH and HXeOXeH molecules, and equal to  $\rm 2.208$  Å and  $\rm 2.15$  Å respectively [14, 15].

In order to examine the finite temperature properties of the Xe<sub>4</sub>O<sub>12</sub>H<sub>12</sub> structure, we performed DFT molecular dynamics simulations on supercells containing 224 and 504 atoms and Γ-point sampling of the Brillouin zone. While the simulations with the two cell sizes give very close results, we determined that a denser k-point grid is required for well converged results. Nevertheless, it is worth noting the observed tendency of the H subsystem to become diffusive at finite T. Its self-diffusion coefficient is estimated to be 5  $\text{Å}^2 \cdot \text{ps}^{-1}$  at 60 GPa and 1,500 K. The diffusivity of H is similar to what is observed in H<sub>2</sub>O at these conditions, except that while water is superionic Xe<sub>4</sub>O<sub>12</sub>H<sub>12</sub> is metallic. The electronic properties of the latter, computed within PBEsol, are shown in Fig. 4. Notice that the electronic states near the Fermi level are due to the Xe and O atoms, so that the metallic character would be preserved even if H undergoes subatomic melting at high T.

No oxygen phase was detected in the x-ray diffraction patterns. However, the  $Xe_4O_{12}H_{12}$  structure we propose may remarkably be obtained by the interpenetration of the high-P Xe and  $\eta$ -O<sub>2</sub> hcp structures [24]. The O atoms in the unit-cell of the newly discovered phase can be matched by summing four  $\eta$ -O<sub>2</sub> unit-cells expanded by 25% along the c axis, i.e. along the intramolecular bond, resulting within 1% in an identical atomic volume for O between both phases.  $\eta$ -O<sub>2</sub> has been described to be stable below the O<sub>2</sub> melting line between 15 GPa and 20 GPa [25], but recent experiments carried out at much higher P evidenced another domain of stability just below the melting line above 44 GPa [26]. The stability field of the new Xe compound as mapped in this study thus coincides with that of high  $P \eta$ -O<sub>2</sub>. In contrast to other  $O_2$  solid phases,  $\eta$ - $O_2$  is characterized by its orientational disorder and a high degree of charge transfers that provides an explanation for the relatively short Xe-O distance in our present phase. It is also interesting to note that the Xe<sub>4</sub>O<sub>12</sub>H<sub>12</sub> structure bears similarities to the hypothetical clathrate V structure proposed [27] for 4X-8Y-68H<sub>2</sub>O, where X=CCl<sub>4</sub> and Y=Xe, with partial occupancy for the latter. Both are hexagonal with the P63/mmc space-group, a ...ABAB... stacking sequence, and c/a ratio of 1.61.

Although theoretical calculations reproduced the experimental results without considering Pt, its role in the energetics of the reaction has to be considered as its hydrogenation is an exothermic reaction [28]. PtH x-ray diffraction signal [29] was observed locally upon heating at 40 GPa, a P value lower than that of Xe reactivity. It is therefore not possible to decipher between newly formed and inherited PtH from previous heating cycles at lower P. PtH was never observed simultaneously with the Xe compound (Fig.1), but that could be due to the very small x-ray beam (3×4  $\mu$ m) and high diffusion rate of hydrogen. In addition, Raman spectra collected after heating showed no sign of H<sub>2</sub> vibron. We therefore propose the following global reaction:

$$2Xe + 6H_2O + (12-x/2)Pt \rightarrow Xe_2O_6H_{x/2} + (12-x/2)PtH$$
(1)

The reaction must have a negative  $\Delta V/V$  value to be pressure induced, which further brackets the H content between 8 ( $\Delta V/V$ =-0.2% per mole) and 12 ( $\Delta V/V$ =-2.2% per mole) atoms per cell.

Water ice and Xe are consequently expected to react at P-T conditions higher than 50 GPa and 1,500 K. These conditions are gathered inside giant planets, and particularly within Uranus and Neptune that contain a larger proportion of planetary ices. These conditions are also gathered within subducted slabs in the deep Earth, provided that sea-water carries Xe within the deep mantle [30]. The observed enrichment of the jovian atmosphere in noble gases by a factor of 2 compared to the solar abundances has been explained by different scenarios including their clathration and/or absorption on water ice in pre-planetary objects and their further release into the planet's atmosphere [31]. However, Xe abundance is overestimated in these scenarios. Any reactivity with Xe at depth would explain this mismatch and future probe missions to Saturn [32], Uranus and Neptune could provide a test of the present results. These planetary implications also fit into the wider context of noble gases sequestration at depth inside giant planets as demonstrated for Ne in the interiors of Jupiter [33].

We acknowledge the ESRF for provision of beamtime on ID27 and LLNL for computational resources. C.S. acknowledges useful discussions with Y. Noel and M. Marques. S.A.B. performed work at LLNL under the auspices of the US Department of Energy under contract No. DE-AC52-07NA27344.

- [1] T. Guillot, Annu. Rev. Earth Planet. Sci. 33, 493 (2005).
- [2] N. Bartlett, Proc. Chem. Soc. 27, 218 (1962).
- [3] W. Grochala, Chem. Soc. Rev. 36, 1632 (2007).
- [4] R. B. Gerber, Annual Review of Physical Chemistry 55, 55 (2004).
- [5] M. Somayazulu, P. Dera, A. F. Goncharov, S. A. Gramsch, P. Liermann, W. Yang, Z. Liu, H. k. Mao, and R. J. Hemley, Nature Chemistry 427, 530 (2009).
- [6] E. Anders and T. Owen, Science **198**, 453 (1977).
- [7] W. A. Caldwell, J. H. Nguyen, B. G. Pfrommer, F. Mauri, S. G. Louie, and R. Jeanloz, Science 277, 930 (1997).
- [8] K. M. Lee and G. Steinle-Neumann, J. Geophys. Res. 111, B02202 (2006).
- [9] C. Sanloup, B. C. Schmidt, E. C. Perez, A. Jambon, E. Gregoryanz, and M. Mezouar, Science 310, 1174 (2005).
- [10] M. I. J. Probert, J. Phys.:Condens. Matter 22, 025501 (2010).
- [11] C. Sanloup, B. C. Schmidt, G. Gudfinnsson, A. Dewaele, and M. Mezouar, Geochim. Cosmochim. Acta 75, 6271 (2011).
- [12] D. S. Brock and G. J. Schrobilgen, J. Amer. Chem. Soc. 133, 6265 (2011).
- [13] C. Sanloup, H.-K. Mao, and R. J. Hemley, Proc. Nat. Acad. Sci. 99, 25 (2002).
- [14] M. Pettersson, J. Lundell, and M. Räsänen, Eur. J. Inorg. Chem. 505, 729 (1999).
- [15] L. Khriachtchev, K. Isokoski, A. Cohen, M. Räsänen, and R. B. Gerber, J. Am. Chem. Soc. 130, 6114 (2008).
- [16] H. K. Mao, J. Xu, and P. M. Bell, J. Geophys. Res. 91, 4673 (1986).
- [17] C.-S. Zha, K. Mibe, W. A. Bassett, O. Tschauner, H.-K. Mao, and R. J. Hemley, J. App. Phys. 103, 054908 (2008).
- [18] J. F. Lin, M. Santoro, V. V. Struzhkin, H. K. Mao, and R. J. Hemley, Rev. Sci. Instrum. 75, 3302 (2004).
- [19] G. Kresse and J. Hafner, Phys. Rev. B 47, 558 (1993);
   G. Kresse and X. Furthmuller, Comp. Mat. Sci. 6, 15 (1996).
- [20] J. P. Perdew, K. Burke, and M. Ernzerhof, Phys. Rev. Lett. 77, 3865 (1996).
- [21] J. P. Perdew, A. Ruzsinszky, G. I. Csonka, O. A. Vydrov, G. E. Scuseria, L. A. Constantin, X. Zhou, and K. Burke, Phys. Rev. Lett. 100, 136406 (2008).
- [22] S. Grimme, J. Comp. Chem. 27, 1787 (2006).
- [23] Q. Zhu, D. Y. Jung, A. R. Oganov, C. W. Glass, C. Gatti, and A. O. Lyakhov, Nat. Chem. 5, 61 (2013).
- [24] L. F. Lundegaard, C. Guillaume, M. I. McMahon, E. Gregoryanz, and M. Merlini, J. Chem. Phys. 130, 164516 (2009).
- [25] M. Santoro, E. Gregoryanz, H.-K. Mao, and R. J. Hemley, Phys. Rev. Lett. 93, 265701 (2004).
- [26] A. F. Goncharov, T. R. R. N. Subramanian, M. So-mayazulu, V. B. Prakapenka, and R. J. Hemley, J. Chem. Phys. 135, 084512 (2011).
- [27] E. A. Smelik and H. King, Z. Kristallogr. 211, 84 (1996).
- [28] L. Visscher, T. Saue, W. C. Nieuwpoort, K. Faegri, and O. Gropen, J. Chem. Phys. 99, 6704 (1993).
- [29] N. Hirai, H. Fujihisa, Y. Ohishi, K. Takemura, and T. Kikegawa, Acta Cryst. A 64, C609 (2008).

- [30] G. Holland and C. J. Ballentine, Nature 441, 186 (2006).
- [31] F. Hersant, D. Gautier, and J. I. Lunine, Planet. Space Sci. 52, 623 (2004).
- [32] B. Marty, T. Guillot, A. Coustenis, N. Achilleos, Y. Alibert, S. Asmar, D. Atkinson, S. Atreya, G. Babasides, K. Baines, T. Balint, D. Banfield, S. Barber, B. Bezard, G. L. Bjoraker, M. Blanc, S. Bolton, N. Chanover, S. Charnoz, E. Chassefiere, J. E. Colwell, E. Deangelis, M. Dougherty, P. Drossart, F. M. Flasar, T. Fouchet, R. Frampton, I. Franchi, D. Gautier, L. Gurvits, R. Hueso, B. Kazeminejad, T. Krimigis, A. Jambon, G. Jones, Y. Langevin, M. Leese, E. Lellouch, J. Lunine, A. Milillo, P. Mahaffy, B. Mauk, A. Morse, M. Moreira, X. Moussas, C. Murray, I. Mueller-Wodarg, T. C. Owen, S. Pogrebenko, R. Prange, P. Read, A. Sanchez-Lavega, P. Sarda, D. Stam, G. Tinetti, P. Zarka, J. Zarnecki, and Kronos Consortium, Experimental astronomy 33, 947 (2009).
- [33] H. F. Wilson and B. Militzer, Phys. Rev. Lett. 104, 121101 (2010).

FIG. 1. X-ray diffraction patterns. Spectra obtained upon and after heating a  $Xe+H_2O+Pt$  sample (black lines) and upon and after heating a  $H_2O+Pt$  sample (grey lines). Circles indicate peaks from phase-1 and asterisks indicate peaks from phase-2.

FIG. 3. Rietveld fit of the  $\rm Xe_4O_{12}H_{12}$  structure at 58 GPa and 1,500 K. Vertical ticks indicate diffraction peaks from phase-2 (blue) and from Xe (black). The pattern has been cut between 9.4 and 12.3 degrees, as the signal in this region is overloaded by Pt and water ice diffraction peaks which would prevent an accurate refinement of the less intense  $\rm Xe_4O_{12}H_{12}$  peaks. Insets: structure of the  $\rm Xe_2O_6H_6$  unit (2 per cell), inner half of the corresponding 2-D diffraction pattern (the ring at 10 degrees is the (111) Pt peak, all rings inwards are from either phase-2 or Xe).

FIG. 2. Pressure dependence of the unit cell volume (left) and cell parameters (right). Solid curves are a  $2^{nd}$  order Birch-Murnaghan fit to the data. Theoretical calculations: squares are results obtained with the PBEsol functional, the error bars extend from results obtained using LDA (lower bar) to results within PBE (upper bar).

FIG. 4. Electronic density of states (EDOS) of  $Xe_4O_{12}H_{12}$  computed at 60 GPa within PBEsol. The total EDOS are shown with a solid line. The dashed lines indicate partial EDOS projected on different atomic species (as indicated in the figure legend). The energy scale is shifted so that the Fermi level is at 0.

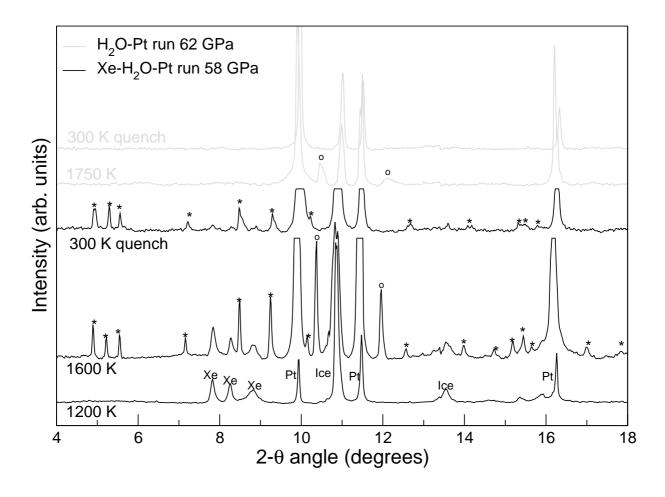


Figure 1 LZ13568 30May2013

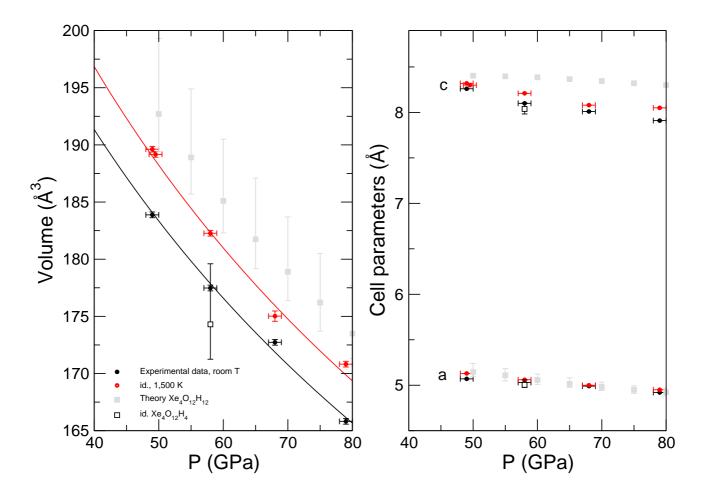


Figure 2 LZ13568 30May2013

