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Interpreting the Fermi-LAT gamma ray excess in the simplified framework

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Abstract

We test the plausibility of the hypothesis that the annihilation of a Majorana fermion dark matter particle via a scalar mediator explains the gamma ray excess from the Galactic Center. Assuming that the mediator couples to all third generation fermions we calculate observables for dark matter abundance and scattering on nuclei, gamma, positron, and anti-proton cosmic ray fluxes, radio emission from dark matter annihilation, and the effect of dark matter annihilations on the CMB. After discarding the controversial radio observation, we show that the dark matter model simultaneously fits the observed excesses in the cosmic gamma ray, the positron, and anti-proton fluxes, while evading constraints from the CMB and direct detection. The experimental data are consistent with a dark matter (mediator) mass in the 10–100 (3–1000) GeV region and with weakly correlated couplings to bottom quarks and tau leptons with values of 10^{-3} –1 at the 68% credibility level.

I. INTRODUCTION

Since 2009 an increasingly significant deviation from background expectations has been identified in the data of the Large Area Telescope (LAT) on board the Fermi Gamma Ray Space Telescope satellite [1–9]. The deviation appears around 2 GeV in the energy spectrum of gamma ray flux originating from an extended region centered in the Galactic Center. The source of the excess photons presently is unclear. Their origin can be dark matter (DM) annihilation, a population of millisecond pulsars or supernova remnants [10–13], or cosmic rays injected in a burst-like or continuous event at the Galactic Center [14]. Based on their luminosity function it was debated whether the excess is explainable with millisecond pulsars [15–17]. Lately, however, the authors of Refs. [18] and [19] have provided observational and theoretical indication that millisecond pulsars are able to produce a gamma ray flux that can explain the observed excess from the galactic center.

Recently, several groups including Daylan et al. [7], Calore et al. [8], and the Fermi Collaboration [9] reanalyzed data from the Fermi-LAT [20] and concluded that the 1–3 GeV gamma ray signal is statistically significant and is in good agreement with that predicted by annihilating dark matter models. The peak in the energy distribution is broadly consistent with gamma rays originating from self-annihilation of dark matter particles [7, 21–26]. Since the signal extends to more than 10° from the Galactic Center, the possibility that it originates from millisecond pulsars is disfavored [7]. The intensity of the signal suggests a dark matter annihilation cross section of about $1\text{--}3 \times 10^{-26} \text{ cm}^3/\text{s}$, which would lead to the observed dark matter abundance at thermal freeze out [27–32]. The diffuse nature and morphology of the gamma ray excess is consistent with a Navarro-Frenk-White-like Galactic distribution of dark matter [8]. This gamma ray excess thus drew the attention of a number of particle model builders and phenomenologists [10, 14, 27, 33–35].

The conclusion that we have discovered dark matter particles, however, cannot be drawn yet. First, we have to be able to exclude the possibility of a standard astrophysical explanation. Second, we need to demonstrate that a dark matter particle that explains the gamma ray excess (with a given mass, spin, and interaction strength to the standard sector) is consistent with a large number of other observations. The latter concerns our paper. We aim to determine the microscopic properties of the dark matter particle from the gamma ray excess and check that these properties comply with limits from other experiments. We use dark

matter abundance and direct detection data, measurements of the gamma ray flux from the Galactic Center, near Earth positron and anti-proton flux data, Cosmic Microwave Background (CMB) observations, and measurements of galactic radio emission as experimental constraints.

Amongst the above listed experimental bounds the constraining role of radio emission has been debated in the literature. Bringmann et al. have shown that radio emission of dark matter annihilation products imposes severe constraints on dark matter annihilation in the Galactic Center [36]. The amount of energy emitted in the form of radio waves, however, is strongly affected by other energy loss (or gain) processes including synchrotron radiation, inverse Compton scattering, ionization, and bremsstrahlung. Most studies of the radio constraint on dark matter, including that of Bringmann et al., ignore energy loss processes other than synchrotron radiation. However, as pointed out by Cholis et al. in Ref. [37], there are several reasons why the other processes could be important. Cholis et al. have shown that after considering inverse Compton scattering, induced by high densities of radiation in the inner Milky Way, the radio constraint on dark matter is weakened by about three orders of magnitude [37]. As a result dark matter annihilating at the thermal rate remains compatible with the radio data. After considering the effect of diffusion the constraint will be further weakened. Due to this, we will exclude the radio data point from our combined fit.

As theoretical description of dark matter we use the simplified model framework. Within this ansatz we make minimal and general theoretical assumptions. We consider a single dark matter particle that couples to various standard fermions via a mediator. Our dark matter particle thus annihilates to several final states which all contribute to the observables mentioned above.

This paper is organized as follows. In Sec. II we introduce the simplified dark matter model we use. In Sec. III, we describe the observables of dark matter abundance and scattering on nuclei, gamma, e^+ , and \bar{p} cosmic ray fluxes, and the effect of dark matter annihilations on the CMB. Our numerical results are given in Sec. IV. Finally in Sec. V we summarize our main results. We collect the formulae of Bayesian inference and likelihood functions in the Appendix.

II. THEORETICAL HYPOTHESIS

In this section we motivate and describe the theoretical hypothesis we test. Our analysis is a follow-up of Ref. [38], which is based on Ref. [33]. The authors of Ref. [33] surveyed all possible simplified model operators with the combination of a spin 0, 1/2, and 1 dark matter particle and mediator. They found that out of the 16 possible combinations 8 types of operators are able to produce the observed gamma ray flux from the galactic center. Out of these 8 promising operators 4 leads to momentum suppressed SI direct detection cross section. In Ref. [38] we compared Bayesian evidences for 3 of the remaining 4 types of operators and found that the experimental data overwhelmingly supports a Majorana dark matter candidate coupled to Standard Model (SM) fermions via a spin-0 mediator. This is the case we analyze further in the present work.

Consequently, in this work we assume that the dark matter particle is a Majorana fermion, which we denote with χ . Inspired by the Higgs portal mechanism [39], we use a simplified model to describe interactions between χ and SM matter. We assume that the dark-standard mediator is a spin-0 field, S , and the form of the dark matter to mediator coupling is

$$\mathcal{L}_\chi \supset \frac{i\lambda_\chi}{2} \bar{\chi} \gamma_5 \chi S. \quad (1)$$

The presence of γ_5 is essential since it lifts the velocity suppression that one otherwise encounters in the indirect detection cross section, thus making this operator capable of explaining the gamma ray excess.

The interaction between the mediator and SM fermions f is assumed to be

$$\mathcal{L}_S \supset \lambda_f \bar{f} f S. \quad (2)$$

Together with the coupling in Eq. (1) it leads to a momentum suppressed spin independent (SI) nucleon-dark matter cross section. Without this suppression direct detection rates could be dangerously enhanced by loop contributions [40]. In line with minimal flavor violation [41], we only consider the third generation fermions, i.e. $f = b, t, \tau$.

For simplicity we assume that mediator pair final states are not present in the dark matter annihilation and only consider s-channel annihilation diagrams. According to power counting of the dark matter transfer momentum or velocity [42], with the bi-linears in Eqs. (1) and (2) the annihilation cross section of the fermionic dark matter candidate is not velocity

suppressed. The dark matter-nucleon elastic scattering cross section is spin-independent (SI) and momentum suppressed.

III. OBSERVABLES

In this section we describe the calculation of the observables that we use to constrain the parameter space of our hypothesis. TABLE I summarizes these observables.

observable name	expression	experiment	data points	data source
dark matter abundance	$\Omega_{\text{DM}}h^2$	PLANCK	1	Ref. [43]
γ -ray flux	$\frac{d^2\Phi_\gamma}{dEd\Omega}$	Fermi-LAT	24	Ref. [8]
cosmic e^+ flux	$\frac{d\Phi_{e^+}}{dE}$	AMS-02	72	Ref. [44]
cosmic \bar{p} flux	$\frac{d\Phi_{\bar{p}}}{dE}$	PAMELA	23	Ref. [45]
Cosmic Microwave Background	f_{eff}	PLANCK	3	Ref. [46]
dark matter direct detection	s	LUX	1	Ref. [47]
radio emission	S_ν	Jodrell Bank	1	Ref. [48]

TABLE I: Summary of observables we use to constrain our dark matter scenario. The expressions in the second column are defined in the text of this section.

A. Dark matter abundance

We assume that dark matter particles, as standard thermal relics, have frozen out in the early universe acquiring their present abundance. We calculate this abundance using micrOmegas version 3.6.9 [49]. We imagine that χ is the only dark matter candidate, that is we use a Gaussian likelihood function with a mean and width determined by PLANCK [43]

$$\Omega_{\text{DM}}h^2 = 0.1199 \pm 0.0027. \quad (3)$$

It is challenging to estimate the theoretical uncertainty of the abundance calculation in a simplified model and the task is the subject of a separate paper. In supersymmetric models,

for example, the theoretical uncertainty is comparable to the experimental one over the bulk of the parameter space. Based on this, we assume an extra theoretical uncertainty of the same size as the experimental error.

B. Gamma ray flux from the Galactic center

In the theoretical scenario under scrutiny the excess gamma ray flux observed by Fermi-LAT is generated by the self-annihilation of χ particles. The differential flux of photons as the function of energy E and observation region Ω is given by

$$\frac{d^2\Phi_\gamma}{dEd\Omega} = \frac{\langle\sigma v\rangle}{8\pi m_\chi^2} J(\psi) \sum_f B_f \frac{dN_\gamma^f}{dE}. \quad (4)$$

Here $\langle\sigma v\rangle$ is the velocity averaged dark matter annihilation cross section at the Galactic center, $B_f = \langle\sigma v\rangle_f/\langle\sigma v\rangle$ is the annihilation fraction into the $f\bar{f}$ final state, and dN_γ^f/dE is the energy distribution of photons produced in the annihilation channel with final state $f\bar{f}$. The J factor in Eq. (4) is a function of the direction of observation ψ

$$J(\psi) = \int_{los} \rho_\chi^2(r) dl, \quad (5)$$

with

$$r = \sqrt{l^2 + r_\odot^2 - 2lr_\odot \cos\psi}. \quad (6)$$

The dark matter distribution in the Galaxy is described by a generalized Navarro-Frenk-White (NFW) dark matter profile [50]

$$\rho_\chi(r) = \rho_0 \frac{(r/r_s)^{-\gamma}}{(1 + r/r_s)^{3-\gamma}}. \quad (7)$$

Here $r_s = 20$ kpc is the radius of the galactic diffusion disk, $r_\odot = 8.5$ kpc is the solar distance from the Galactic center, and ρ_0 is set to reproduce the local dark matter density $\rho_\chi(r_\odot) = 0.3$ GeV/cm³. Following Refs. [7, 8] we fix the inner slope of the NFW halo profile to $\gamma = 1.26$ and set $\psi = 5^\circ$ [27]. Within 1–2° of the Galactic Plane bremsstrahlung can significantly contribute to the gamma-ray spectrum from annihilating dark matter [7].

The differential yield dN_γ^f/dE is different for the three final states we consider. As seen from Eq. (4), the total differential yield determining the gamma ray flux is the annihilation-fraction-weighted sum of the differential yields into specific final states. We sum over the

contributions of the three individual SM fermions (b, t, τ) . As B_f depends on the model parameters, the gamma ray data plays an important role in constraining the coupling of the mediator to SM fermions.

We use micrOmegas version 3.6.9 to evaluate the theoretical prediction for the differential gamma ray flux [49]. The gamma ray spectral data points that we input into our Gaussian likelihood function are taken from Ref. [8], including both statistical errors and empirical model systematics.

C. Cosmic positron flux near Earth

The third generation fermion states produced by dark matter annihilation in our model can produce stable leptons in a variety of ways, including production via the decay of top quarks or tau leptons, or secondary production from hadron decays. These charged particles provide extra sources of cosmic flux in addition to the expected astrophysical backgrounds. Consequently the measurement of the electron and positron flux allows us to set constraints on the dark matter properties. Since the prediction of the electron flux poses an additional challenge and it is the source of considerable uncertainties, we only focus on the positron flux and do not consider the electron flux or the positron fraction in this paper.

The propagation of positrons within the Galaxy is well-described by the following simplified transport equation

$$\frac{\partial f_{e^+}}{\partial t} - \nabla(K(E, r)\nabla f_{e^+}) - \frac{\partial}{\partial E}(b(E, r)f_{e^+}) = Q_{e^+}(E, r), \quad (8)$$

in the diffusion zone approximated by a cylinder with thickness $2L$. In the above equation $f_{e^+}(r, t, E)$ is the number density of positrons, $K(E, r)$ is the diffusion coefficient which is parameterized as $K(E, r) = K_0(E/\text{GeV})^\delta$, and $b(E, r)$ is the rate of energy loss. The source term reads as

$$Q_{e^+}(E, r) = \frac{\rho_\chi^2(r)\langle\sigma v\rangle}{2m_\chi^2} \left(\sum_f B_f \frac{dN_{e^+}^f}{dE} \right), \quad (9)$$

with $dN_{e^+}^f/dE$ being the energy spectrum of positrons produced in the annihilation channel into $f\bar{f}$. The differential positron flux is given by

$$\frac{d\Phi_{e^+}}{dE} = \frac{v_{e^+}}{4\pi} f_{e^+}, \quad (10)$$

with v_{e^+} being the positron velocity.

For the dark matter induced positron flux calculation in micrOmegas, we fix the values of the propagation parameters to reduce the number of free degrees of freedom in the problem. Among the three defaults (MIN, MED, MAX) we choose the MED model. (We justify this choice in the next section.) The MED diffusion parameters are: the index of the diffusion coefficient $\delta = 0.7$, the normalization factor $K_0 = 0.0112 \text{ kpc}^2/\text{Myr}$, and the thickness of the diffusive cylinder $L = 4 \text{ kpc}$ [51].

For the astrophysical backgrounds, we adopt the following parametrization for the interstellar positron flux and the flux at the top of the atmosphere (TOA) [52]

$$\Phi_{e^+}^{\text{bkg}}(E) = C_s E^{-\gamma_s} \exp(-E/E_s) + C_{e^+} E^{-\gamma_{e^+}}, \quad (11)$$

$$\Phi_{e^+}^{\text{TOA}}(E) = \frac{E^2}{(E + \phi_{e^+})^2} \Phi_{e^+}^{\text{bkg}}(E + \phi_{e^+}), \quad (12)$$

with best fit parameters $C_{e^+} = 72 \text{ s}^{-1} \text{ sr}^{-1} \text{ m}^{-2} \text{ GeV}^{-1}$, $C_s = 1.6 \text{ s}^{-1} \text{ sr}^{-1} \text{ m}^{-2} \text{ GeV}^{-1}$, $\gamma_{e^+} = 3.7$, $\gamma_s = 2.51$, $E_s = 1 \text{ TeV}$, and solar modulation parameter $\phi_{e^+} = 0.93 \text{ GV}$ obtained in Ref. [52]. In Eq. (11) the first term on the right hand side describes the primary positron component arising from, among other sources, gamma-ray burst, hadronic interactions inside supernova remnants, or interactions of gamma rays with strong magnetic fields. The second term describes the secondary component produced in collisions of various cosmic rays in the interstellar medium.

The parametrization of the background given in Eq. (11) and Eq. (12) is important since when subtracted from the data it indicates the size of a signal potential arising from dark matter annihilation. The overall effect of the AMS-02 data on the combined result, however, is relatively weak. This is because the AMS-02 positron flux in the low energy region (below 10 GeV) is in fairly good agreement with the background expectations and this leaves small room for dark matter contribution. The dark matter model we analyze contributes only negligibly to the positron flux. Thus the parametrization of the AMS-02 background only weakly affects our final results.

The treatment of solar modulation of charged cosmic rays has been shown to be a considerable source of uncertainty in the prediction of these fluxes near Earth [53]. These uncertainties arising from solar modulation, however, are comparable to the experimental error and the theoretical uncertainties in cosmic ray propagation. Since the dark matter model under scrutiny contributes very moderately to the AMS-02 positron spectrum the

treatment of solar modulation only mildly affects our final results.

As experimental input for the positron flux we use the new release of AMS-02 data [44]. We assume the theoretical uncertainty is the same as the AMS-02 experimental error and the form of the likelihood is a composite Gaussian [54, 55].

D. Cosmic anti-proton flux near Earth

The propagation of anti-protons through the Galactic cylinder follows a similar diffusion equation as Eq. (8) but there is an additional effect from the galactic wind and the source term includes the annihilation of anti-protons in the interstellar medium as well as the annihilation of dark matter. The energy loss of anti-protons, however, is negligible compared with that of the positrons.

The authors of Ref. [36] found that the January 10 release of the PAMELA anti-proton flux is in tension with the dark matter annihilation interpretation of the gamma ray excess. They, however, only consider annihilation to arbitrary combination of quark final states. In this work, additionally to quark final states, we consider simultaneous annihilation into taus, which, as we demonstrate considerably eases the PAMELA constraint. Furthermore, Ref. [56] notes that uncertainties in the cosmic ray propagation model markedly ease the PAMELA constraint even in the case of quark only final states. They conclude that the PAMELA data allow for an approximately 35 GeV dark matter particle annihilating into $b\bar{b}$ final state with the rate of $\langle\sigma v\rangle \sim 10^{-26} \text{ cm}^3/\text{s}$.

For the anti-proton flux we calculate the astrophysical background by adopting the KRA model in Ref. [53], which reflects the standard assumptions on the Galactic propagation parameters. The authors of Ref. [53] found that under these standard propagation assumptions the anti-proton flux severely constrains the dark matter interpretation of the gamma ray excess. Our study, however, evades the concerns raised in Ref. [53] because we also allow leptonic annihilation final states which, as we show later, play a significant role in generating the gamma ray excess.

To calculate the anti-proton flux from dark matter annihilation, i.e. $d\Phi_{\bar{p}}/dE$, we use the same MED propagation model. As it is shown in Ref. [57] the shape of the anti-proton spectrum is relatively insensitive to this choice. Additionally, the shape difference due to reasonable choices of propagation parameters is only significant below about 10 GeV (for

anti-protons) and its extent is comparable to the experimental error [53]. This uncertainty arising from the choice of the propagation parameters is folded into our analysis as a part of the theoretical uncertainty. We assume the velocity of the convective wind to be $V_{\text{conv}} = 12$ km/s.

The anti-proton flux at low energies is also altered by solar modulation effects. We use the solar modulation potential $\phi_{\bar{p}} = 0.65$ GV for anti-proton, which relates the local interstellar anti-proton flux to the one measured at the top of the atmosphere, as described in the KRA model. The application of the solar modulation effect also follows Eq. (12).

We use the latest release of PAMELA data as experimental input for the anti-proton flux [45]. Note that the error bars in this data release are only statistical. Systematic error bars are expected to be of the same order of magnitude as in the first release of PAMELA data [58]. We combine the uncertainties in quadrature and assume that the theoretical uncertainty is the same as the experimental error in the composite Gaussian likelihood function.

E. Cosmic Microwave Background

Dark matter annihilation in the early universe affects the CMB temperature and polarization fluctuations. Thus the CMB power spectrum measurement from PLANCK provides constraints on dark matter properties. A key quantity for determining the constraint on a given dark matter model is the efficiency for producing ionizing radiation. The authors of Ref. [46] provide values of the effective efficiency f_{eff} for different annihilation channels and dark matter masses that can be easily interpolated. We quantify the CMB constraints using the following likelihood function

$$\mathcal{L}_{\text{CMB}} = \exp \left[-\frac{1}{2} f_{\text{eff}}^2 \lambda_1 c_1^2 \left(\frac{\langle \sigma v \rangle}{2 \times 10^{-27} \text{cm}^3 \text{s}^{-1}} \right)^2 \left(\frac{\text{GeV}}{m_\chi} \right)^2 \right], \quad (13)$$

$$f_{\text{eff}} = \sum_{i=\tau, b, t} f_{\text{eff}, i} B_i, \quad (14)$$

with $\lambda_1 = 3.16$ and $c_1 = 4.64$ for the PLANCK data. Here B_i is the annihilation branching fraction defined earlier.

F. Dark matter direct detection

Direct detection utilizes dark matter particles scattering on nuclei of a target material in a well shielded detector. The differential recoil rate of dark matter on nuclei, as a function of the recoil energy E_R , is

$$\frac{dR}{dE_R} = \frac{\rho_\chi}{m_\chi m_A} \int dv v f(v) \frac{d\sigma_A^{\text{SI}}}{dE_R}, \quad (15)$$

where m_A is the nucleus mass, $f(v)$ is the dark matter velocity distribution function and

$$\frac{d\sigma_A^{\text{SI}}}{dE_R} = G^\chi(q^2) \frac{1}{E_{\text{max}}} \frac{4\mu_A^2}{\pi} [Z f_p^\chi + (A - Z) f_n^\chi]^2 F_A^2(q), \quad (16)$$

with $E_{\text{max}} = 2\mu_A^2 v^2 / m_A$, $G^\chi(q^2) = \frac{q^2}{4m_\chi^2}$ [33], and $f_N^\chi = \frac{\lambda_\chi}{2m_S^2} g_{SNN}$ ($N = p, n$). $F_A(q)$ is the nucleus form factor and $\mu_A = m_\chi m_A / (m_\chi + m_A)$ is the reduced dark matter-nucleon mass. We assume that the local disk rotation speed is 220 km/s with the same value for the most probable speed of the dark matter's Maxwell-Boltzmann velocity distribution. The Galactic escape speed is 544 km/s [47].

As we only consider the interaction mediated by the scalar between the dark matter particles and the third generation quarks, the strength of the mediator-nucleon (N) interaction reads

$$g_{SNN} = \frac{2}{27} m_N f_{TG} \sum_{f=b,t} \frac{\lambda_f}{m_f}. \quad (17)$$

Above $f_{TG} = 1 - f_{T_u}^N - f_{T_d}^N - f_{T_s}$ and we adopt $f_{T_u}^p = f_{T_d}^n = 0.02$, $f_{T_d}^p = f_{T_u}^n = 0.026$, $f_{T_s} = 0.043$ [33, 59, 60].

For the LUX likelihood function, we use a Poisson distribution in the observed number of events N ,

$$\mathcal{L}(s|N) = P(N|s) = \frac{(b+s)^N e^{-(b+s)}}{N!}, \quad (18)$$

where b is the expected number of background events,

$$s = MT \int_0^\infty dE \phi(E) \frac{dR}{dE_R}(E) \quad (19)$$

is the expected signal, MT is the detector mass \times time exposure, and $\phi(E)$ is a global efficiency factor that takes into account trigger efficiencies, energy resolution, and analysis cuts. Likelihood calculations are performed using a version of `LUXCalc` [61] modified to include the additional momentum dependence in Eqn. (16). For the LUX analysis region used by `LUXCalc`, $N = 1$ and $b = 0.64$; the efficiency curve $\phi(E)$ was generated by `TPCMC` [62] using the `NEST` model [63, 64]. See Ref. [61] for further details.

G. Radio Emission

Electrons and positrons from dark matter annihilation are expected to lose energy through synchrotron radiation in the presence of large scale magnetic fields. Thus the radio emission in galaxies and galaxy clusters can also be used to place constraints on the dark matter properties. The synchrotron flux density is given by

$$S_\nu = \frac{1}{4\pi} \frac{J}{\rho_\chi^2} \int 2 \frac{d\Phi_{e^+}}{dE} \frac{dW_{\text{syn}}}{d\nu} dE_e, \quad (20)$$

where $d\Phi_{e^+}/dE$ is the positron flux in units of $(\text{GeV cm}^2 \text{ s sr})^{-1}$. The synchrotron power per frequency reads

$$\frac{dW_{\text{syn}}}{d\nu} = \frac{\sqrt{3}}{6\pi} \frac{e^3 B}{m_e} F\left(\frac{\nu}{\nu_{\text{syn}}}\right), \quad (21)$$

with

$$F(x) = x \int_x^\infty K_{5/3}(\xi) d\xi \approx \frac{8\pi}{9\sqrt{3}} \delta(x - 1/3). \quad (22)$$

The δ -function implies

$$\nu_{\text{syn}} = 3\nu = \frac{3eBp^2}{4\pi m_e^3} \approx \frac{3eBE_e^2}{4\pi m_e^3}. \quad (23)$$

For simplicity, we fix the magnetic field strength at a conservative lower limit [36]

$$B = 50 \mu\text{G}. \quad (24)$$

The integration cone in the J factor corresponds to a $4''$ region around the Galactic center.

IV. RESULTS

We coded the Lagrangian of the relevant simplified dark matter model in FeynRules [65]. Calculation of observables, including the dark matter relic density and nucleon scattering interactions, differential gamma ray, e^+ and \bar{p} fluxes, and radio signal were performed using a modified version of micrOmegas 3.6.9 [49]. Nested sampling and posterior distribution calculations were performed by MultiNest [66]. The nested sampling algorithm was developed to calculate marginalized posterior probability distributions and it is a Bayesian's way to numerically implement Lebesgue integration [67]. Since the relevant part of the likelihood

parameter (unit)	m_χ (GeV)	m_S (GeV)	λ_b	λ_t	λ_τ
scan range	$1 - 10^3$	$1 - 10^3$	$10^{-5} - 10$	$10^{-5} - 10$	$10^{-5} - 10$
prior type	log	log	log	log	log

TABLE II: Scan ranges and prior types used for the scanned parameters.

distribution spans multiple orders of magnitude, we use log priors for all parameters. We present further details of our statistical analysis in the Appendix.

In our numerical calculation we fix the dark matter to mediator coupling as $\lambda_\chi = 1$, and we scan the following free parameters:

$$P = \{m_\chi, m_S, \lambda_b, \lambda_t, \lambda_\tau\}. \quad (25)$$

Here m_χ is the mass of the dark matter particle, m_S is the mass of the scalar mediator, and λ_f ($f = \tau, b, t$) is the coupling of the mediator to the SM fermion pair $f\bar{f}$ as defined in Eqs. (1) and (2). The range of our scan over the above parameters and the type of prior we use is given in TABLE II.

To build some intuition, first we examine the constraining effect of each observable one by one. To this end we plot the posterior probability distributions marginalized to the scanned model parameters such that the likelihood function only contains the dark matter abundance and one of the other observables. In Fig. 1 we show marginalized posterior probability distributions taking into account the dark matter abundance and the gamma ray data. The first frame of Fig. 1 confirms that the gamma ray data restrict the range of the dark matter mass close to 35–60 GeV [7, 8]. It is less appreciated, however, that uncertainties still allow a 40–100 (25–160) GeV dark matter mass range at the 68% (95%) credibility level. The gamma ray data, coupled with the dark matter relic density, allows the whole mass range of the scalar mediator. It is also interesting to note that the preferred dark matter mass region is dissected by a diagonal band with a lower posterior around the on-shell resonance region $m_S = 2m_\chi$. In this valley, dark matter resonantly annihilates via the s -channel mediator, depleting its abundance. Thus, it is harder for the model to match the PLANCK constraint.

The relevant interaction strengths also remain virtually unconstrained as shown by the

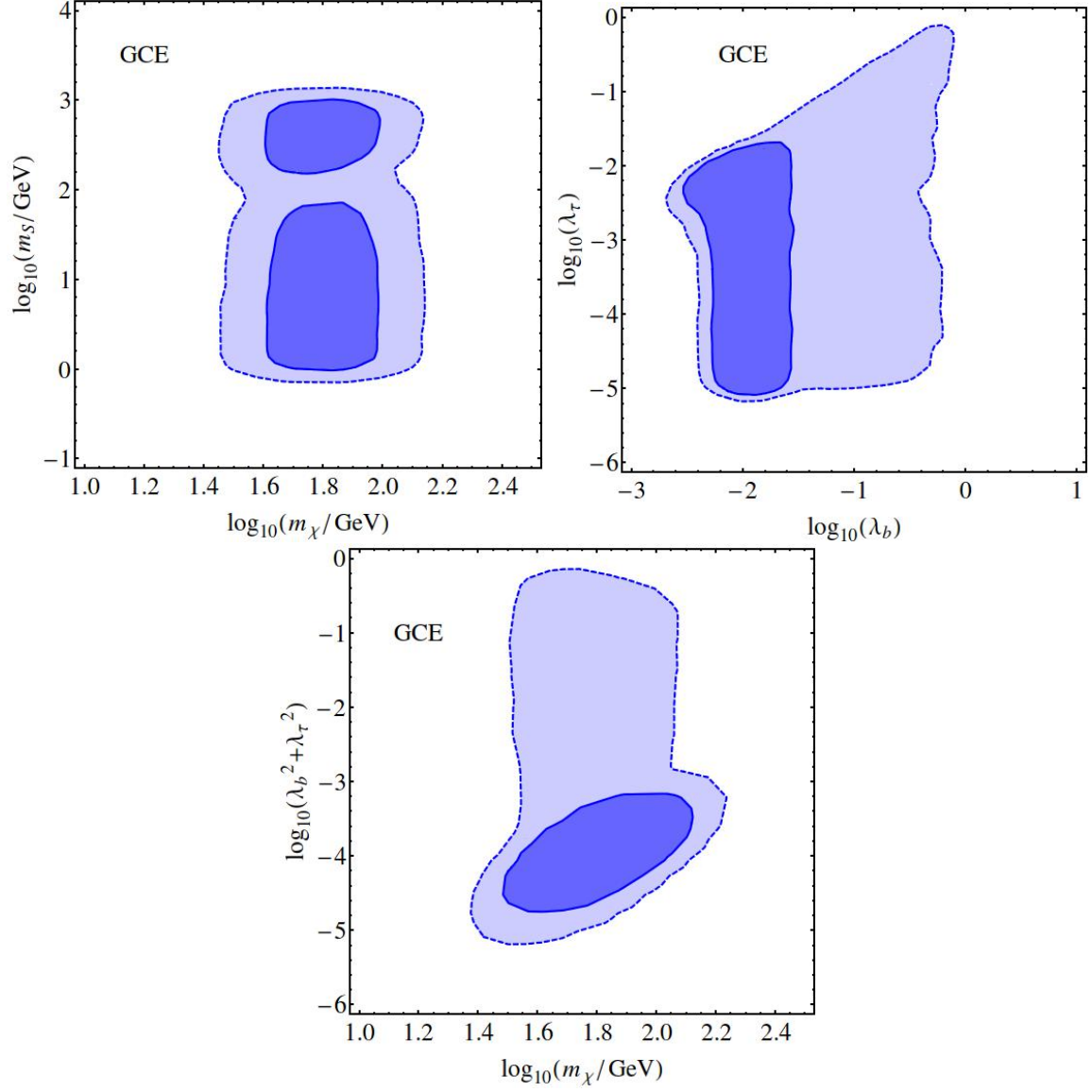


FIG. 1: Posterior probability distributions marginalized to the scanned model parameters. The likelihood function for these plots only contains the dark matter abundance and the anomalous Fermi-LAT gamma ray data. The dark and light regions hereinafter correspond to 68% and 95% credible regions, respectively.

right frame of Fig. 1. PLANCK and the anomalous Fermi-LAT gamma ray data only restrict these coupling in the $\lambda_\tau = 1 \times 10^{-5} - 2.5 \times 10^{-2}$ ($6 \times 10^{-6} - 1$) and $\lambda_b = 2.5 \times 10^{-3} - 2.5 \times 10^{-2}$ ($1.6 \times 10^{-3} - 1$) ranges at the 68% (95%) credibility level. Simultaneous order 1 couplings are marginally allowed and appear in the part of the parameter space where the annihilation cross section is suppressed by sizable m_S .

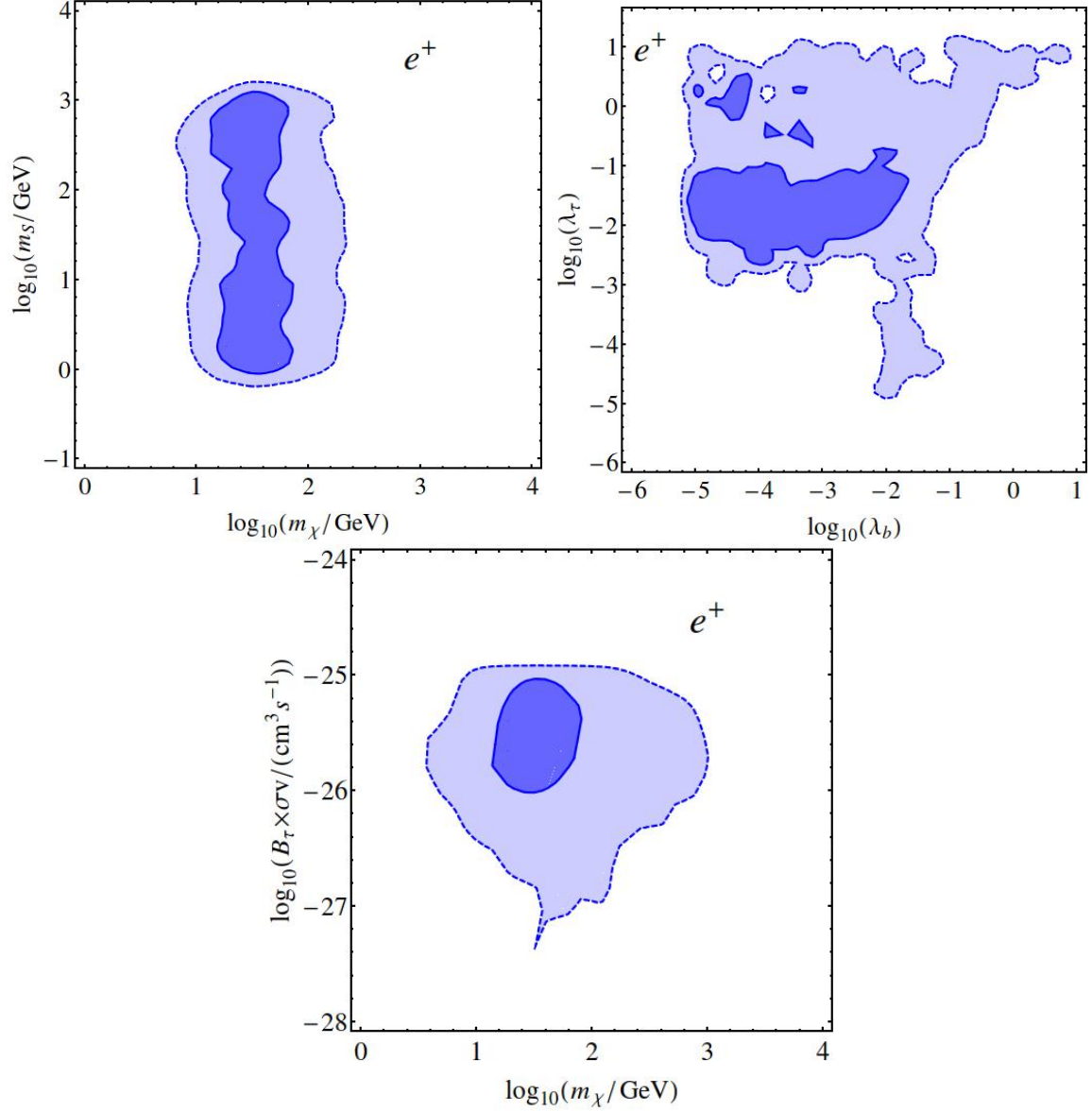


FIG. 2: Posterior probability distributions marginalized to the scanned model parameters. The likelihood function for these plots only contains the dark matter abundance and the AMS-02 positron flux data.

The last frame of Fig. 1 the coupling of the mediator to the light third generation fermions (b quarks plus τ leptons) versus the mass of the dark matter particle. The 68% confidence region clearly shows the presence of a marked correlation while this correlation is weaker at the 95% confidence.

In Fig. 2 we show marginalized posterior probability distributions with the likelihood

function containing only the dark matter abundance and the AMS-02 positron flux data. The AMS-02 measurement of the positron flux features a small upward kink, a sudden change of slope, around 35 GeV. A smooth background prediction has a hard time to reproduce this kink and systematically falls below the experimental points in the 35–50 GeV region. Positrons originating from the annihilation of a 35–50 GeV dark matter particle can fill the gap between the background and the data. Hence the AMS-02 data show a mild preference toward a dark matter candidate with 16–65 (10–160) GeV mass at 68% (95%) credibility level. AMS-02 also restricts the dominant decay to τ leptons with a λ_τ coupling around $2.5 \times 10^{-3} - 0.1$ and λ_b coupling below about 10^{-2} at the 68% credibility level.

The last frame of Fig. 2 shows the correlation between $B_\tau \times \sigma v$ and the dark matter mass. It is apparent from the plot that the AMS-02 positron data does not significantly modify the preference of the relic abundance constraint: a 10–90 GeV dark matter particle annihilating near the standard thermal rate of $3 \times 10^{-26} \text{ cm}^3/\text{s}$.

Fig. 3 shows marginalized posterior probability distributions with the likelihood function including the PLANCK and PAMELA anti-proton cosmic ray flux data only. The PAMELA data in itself does not prefer any particular parameter region. Dark matter and mediator masses are both allowed in the full scanned range at the 95% credibility level. This happens with the exception of a small island around $m_\chi \sim m_S \sim 10 \text{ GeV}$ where the combined PLANCK and PAMELA constraints are harder to satisfy. The reasons for this are that this island falls on the $m_S = 2m_\chi$ resonant annihilation corridor and the PAMELA data around 10 GeV leave very little room for dark matter. This situation improves for lower dark matter masses.

Since dark matter masses above the top quark mass are allowed by the combination of PLANCK and PAMELA, the λ_t coupling comes into play. These data, however, are not sufficient to constrain λ_t . It is interesting to note that PLANCK and PAMELA allow fairly large values of λ_t , λ_b , and λ_τ for heavier m_χ and m_S in order to accommodate the correct relic abundance.

In Fig. 4 we show marginalized posterior distributions for the dark matter abundance and CMB likelihood function, defined in Eq. (14). The diagonal depletion of the likelihood function due to dark matter resonant annihilation is apparent in the m_χ vs m_S frame. The CMB likelihood function suppresses the posterior around $m_\chi = 10 \text{ GeV}$ providing more constraint on low mass dark matter.

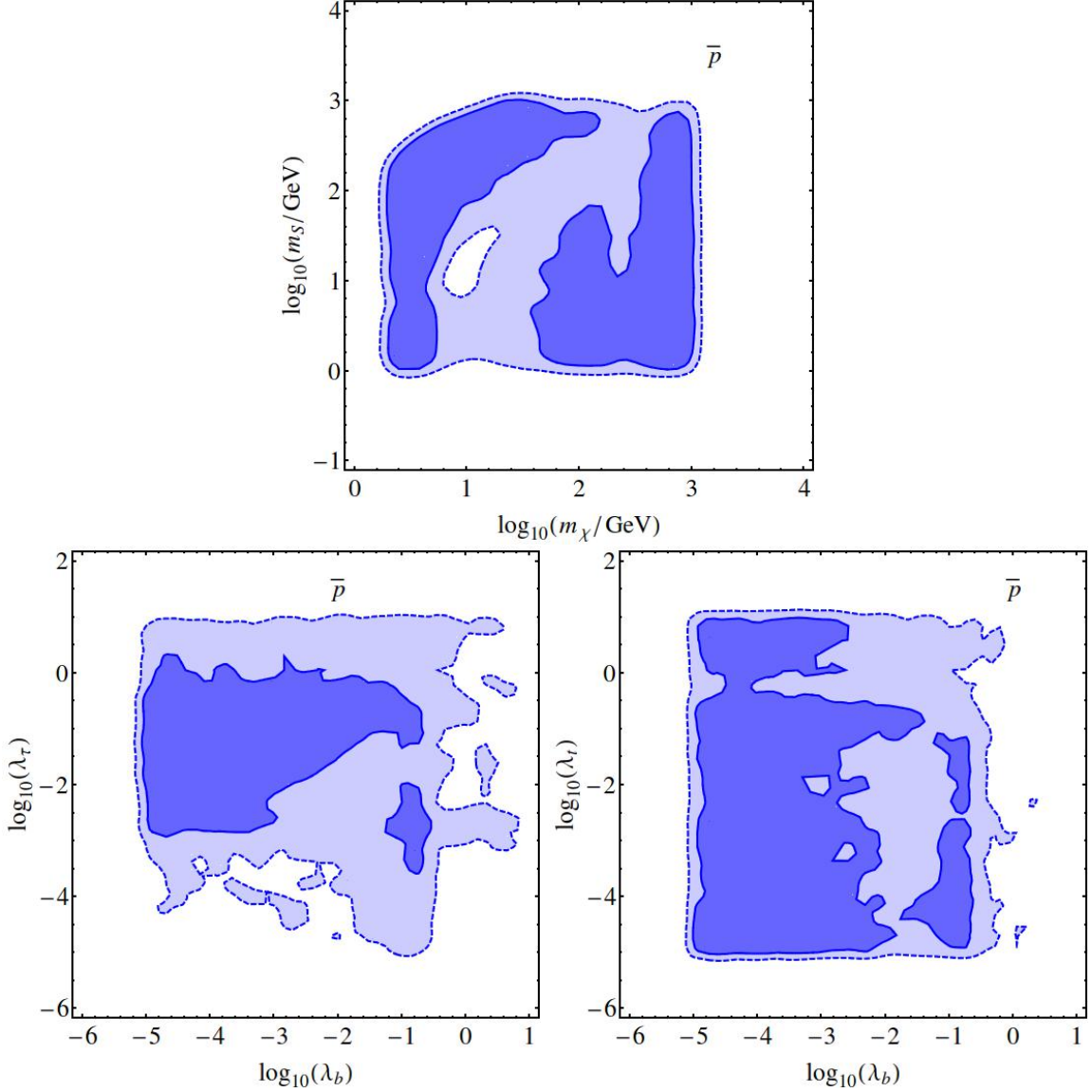


FIG. 3: Posterior probability distributions marginalized to the scanned model parameters. The likelihood function for these plots only contains the dark matter abundance and the PAMELA anti-proton cosmic ray flux data.

The posterior probability distribution projected to the λ_b vs. λ_τ couplings shows a peculiar pattern. This pattern is the combined result of two relatively simple sets of constraints. Dark matter abundance is responsible for the low likelihood values at low λ_b and λ_τ . It turns out that the PLANCK constraint on the amount of relic dark matter is hard to respect unless one of these couplings is sizable, that is λ_b or $\lambda_\tau \gtrsim 10^{-2}$ at the 68% credibility level. If both of these couplings are small then annihilation is slow and dark matter is overproduced

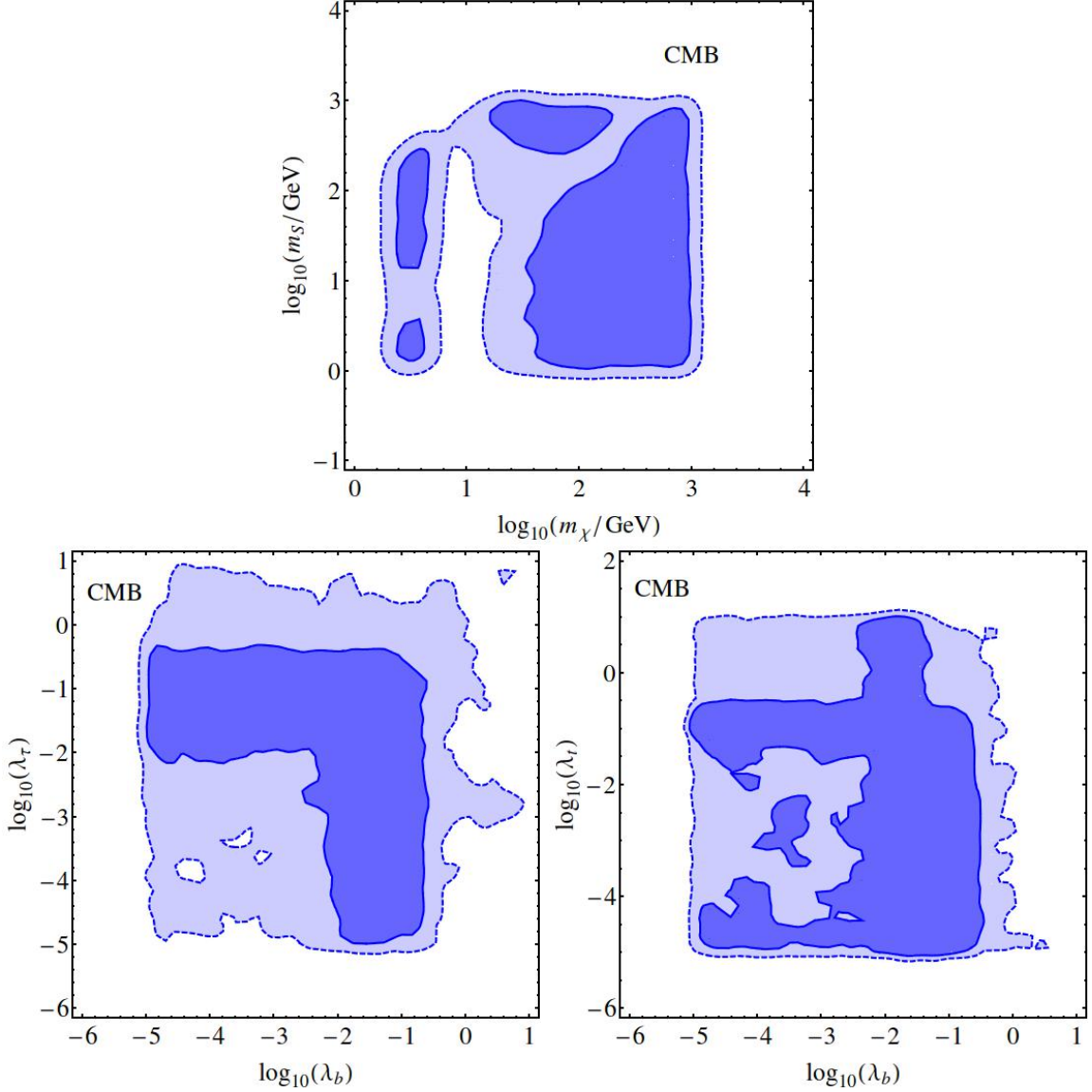


FIG. 4: Posterior probability distributions marginalized to the scanned model parameters. The likelihood function for these plots only contains the dark matter abundance and CMB terms, as defined in Eq. (14).

in the early universe. In the large coupling region, on the other hand, the CMB constrains λ_b and λ_τ from above. If any of these couplings are larger than about 0.1 then dark matter tends to become under-produced and the CMB receives too much modification from dark matter. The λ_t coupling is hardly constrained by the CMB at the 95% credibility level.

In Fig. 5 we show the marginalized posterior distribution with the likelihood function containing only the PLANCK and LUX data. Due to momentum suppression of the nucleon-

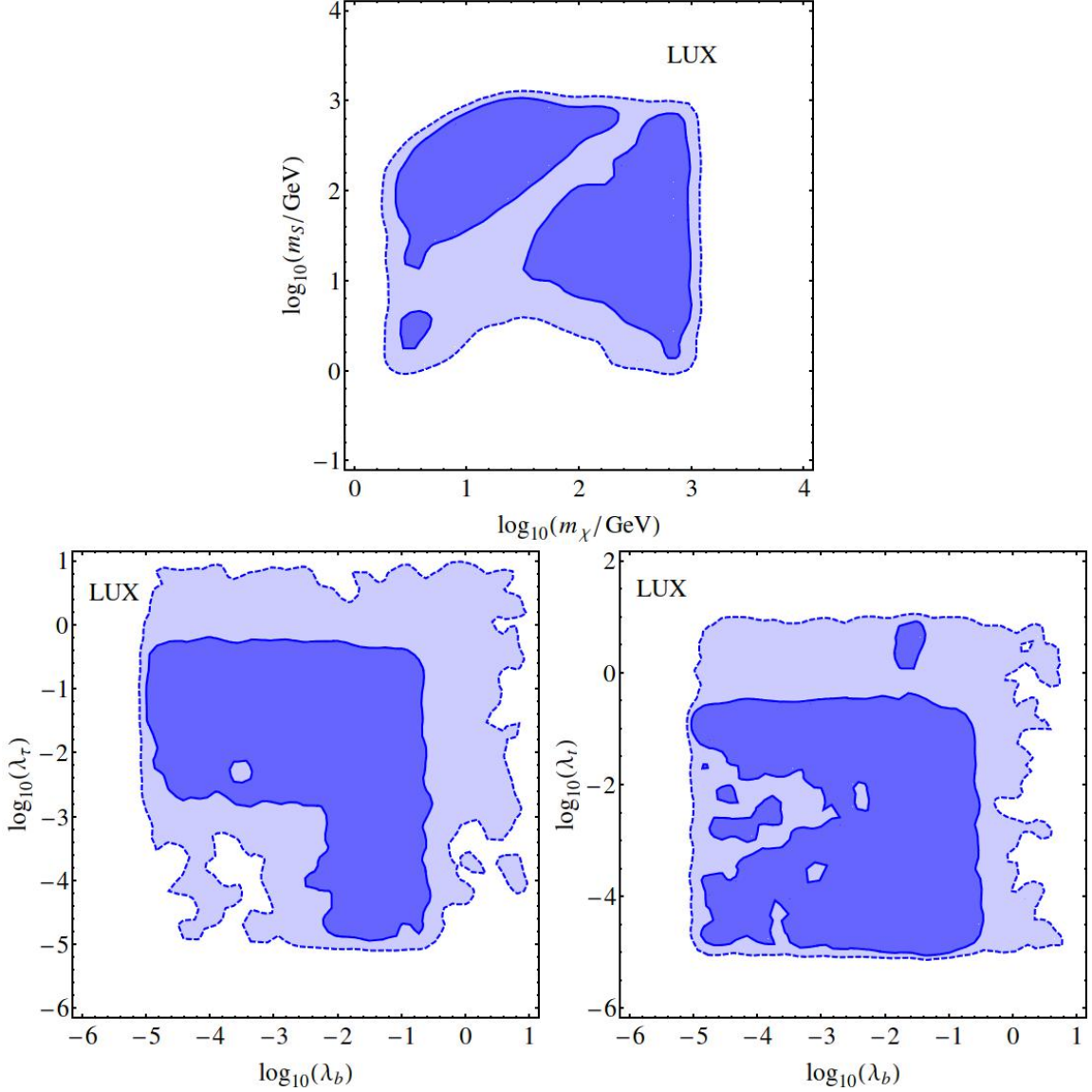


FIG. 5: Posterior probability distributions marginalized to the scanned model parameters. The likelihood function for these plots only contains the dark matter abundance and LUX data.

χ elastic scattering, the LUX data very weakly constrain the dark matter or mediator mass. In the m_χ vs. m_S plane the diagonal resonant annihilation valley is visible, but no other structure is present. The posterior probability distribution for the couplings is very similar to that in Fig. 4. Similarly to the case of the CMB, PLANCK and LUX only impose a constraint on the order 1 couplings.

As discussed in the Introduction, the radio signal potentially very strictly constrains dark matter [36]. Assuming that dark matter contributes to the radio signal only by synchrotron

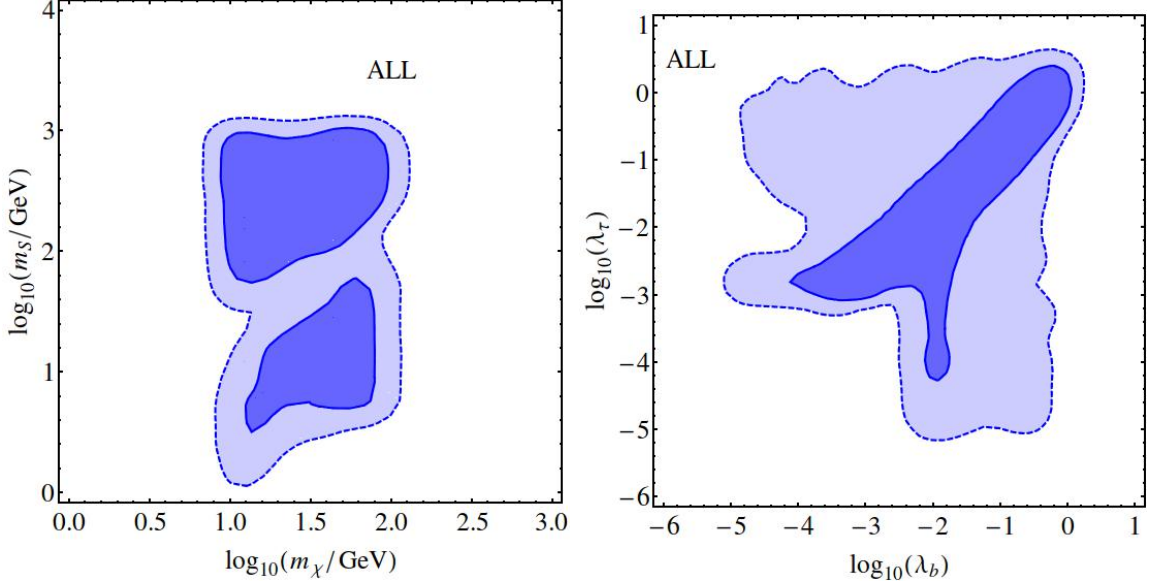


FIG. 6: Posterior probability distributions marginalized to the scanned model parameters. The likelihood function for these plots contains all of dark matter abundance, Fermi-LAT gamma ray data, AMS-02 positron flux data, PAMELA anti-proton ray flux data, CMB and LUX data.

radiation we find the radio flux upper limit of Jodrell Bank at 408 MHz [48] excludes the dark matter hypothesis we consider by two orders of magnitude. Our finding fully confirms that of Ref. [36]. This exclusion, on the other hand, is lifted if inverse Compton scattering, ionization, and bremsstrahlung are also considered as energy loss mechanisms of dark matter annihilation products, altering radio emission [37, 68]. As shown by Ref. [37] the bound from the radio data is weakened by about three orders of magnitude if inverse Compton scattering is considered and is expected to pose no constraint after including Galactic diffusion effects. Due to this, we do not include the radio emission data point in our combined likelihood function.

The summary of all constraints is presented in Fig. 6. The constraint on the dark matter mass is dominated by the gamma ray data and the final combination restricts m_χ to the 10–100 (7–125) GeV region with 68% (95%) credibility. Less of the low mediator mass region survives the scrutiny of the combined constraints, leaving the 3–1000 GeV m_S region preferred at the 68% credibility level. The combined constraints provide enhanced information on the coupling between the dark and standard sectors. They prefer a correlated pair of λ_τ and λ_b couplings in the intermediate 10^{-3} –1 region at the 68% credibility level. Small

($\lambda < 10^{-3}$) and large ($\lambda > 1$) values of couplings are disfavored mostly by PLANCK at the 95% credibility. Comparing the second frame of Fig. 6 to those showing the individual constraints it is clear that simultaneous order 1 couplings are mildly under stress from almost all the data.

Figure 7 shows best fit predictions to the observed gamma ray, positron and anti-proton fluxes. The lines show theoretical predictions including the astrophysical background and dark matter contributions. The solid red line is the prediction with dark matter parameter values, listed in Eq. (25), that give the best fit to all the experimental data listed in TABLE I. The prediction shown by the solid blue line uses the posterior mean parameter values preferred by all observables. The blue dashed lines show the 1 credibility region deviation from the mean. Black marks indicate experimental data as indicated in TABLE I. It is intriguing to see that the dark matter model described in Section II, together with the astrophysical assumptions spelled out in the later sections, can simultaneously fit the gamma ray excess and the charged cosmic ray observations while satisfying the CMB and direct detection constraints. This goodness of fit ensures that our parameter extraction exercise, summarized by the posterior plots, is physically meaningful.

V. CONCLUSIONS

In this work we perform a comprehensive statistical analysis of the gamma ray excess from the Galactic Center in a simplified dark matter model framework. According to our previous study, Majorana fermion dark matter interacting with standard model fermions via a scalar mediator is the most favoured explanation of the galactic center excess when characterised by Bayesian evidence. We locate the most plausible parameter regions of this theoretical hypothesis using experimental data on the dark matter abundance and direct detection interactions, the gamma ray flux from the Galactic center, near Earth positron and anti-proton fluxes, the Cosmic Microwave Background, and galactic radio emission.

We find that the radio data excludes the model if we include synchrotron radiation as the only energy loss channel for dark matter annihilation products. Since it was shown that inclusion of other types of energy losses lifts this exclusion we discard the single radio data point from our combined likelihood [37]. Excluding the radio observation, we find that the Majorana dark matter particle with a scalar mediator simultaneously fits the observed

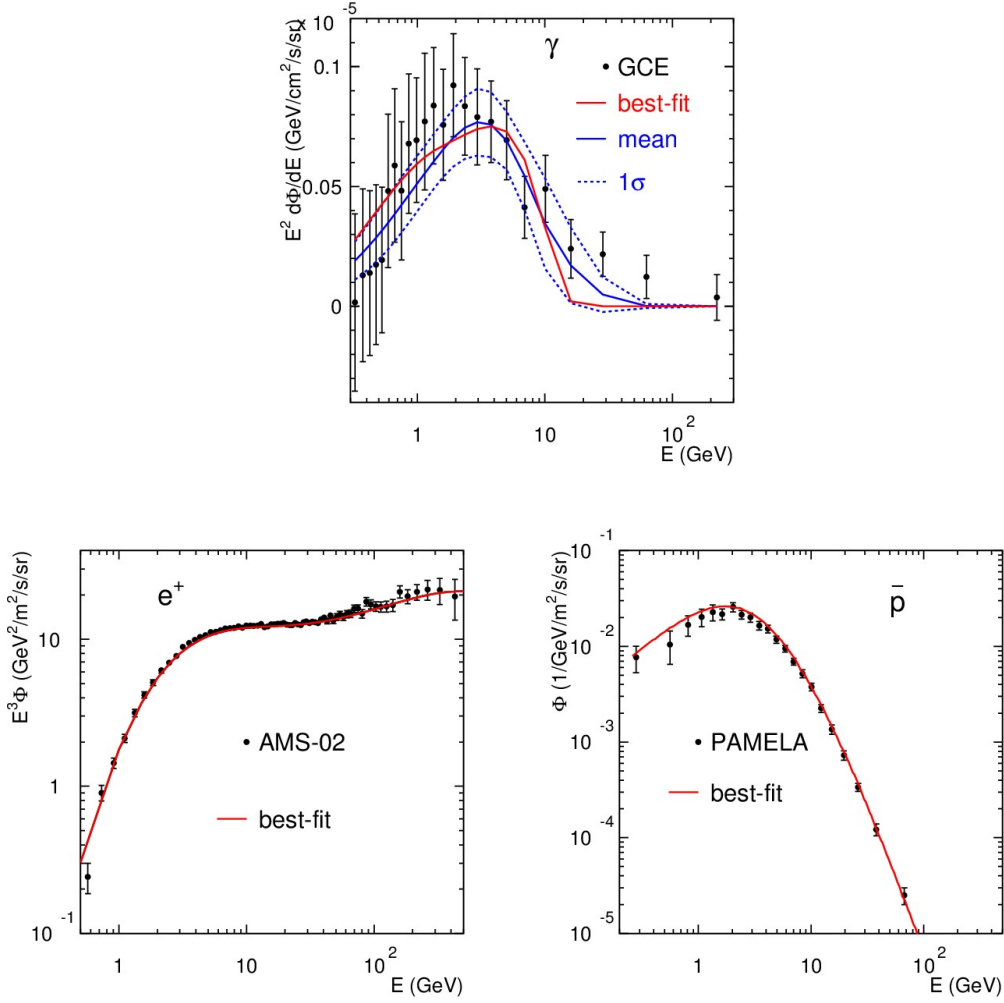


FIG. 7: Best fit predictions to the observed gamma ray, positron and anti-proton fluxes. The lines show theoretical predictions including the astrophysical background and dark matter contributions. The solid red line is the prediction with dark matter parameter values, listed in Eq. (25), that give the best fit to all the experimental data listed in TABLE I. The prediction shown by the solid blue line uses the posterior mean parameter values preferred by all observables. The blue dashed lines show the 1 credibility region deviation from the mean. Black marks indicate experimental data as indicated in TABLE I.

excesses in the cosmic gamma ray, the positron, and anti-proton fluxes, while evading constraints from the CMB and direct detection. The combination of the data prefers a dark matter (mediator) mass in the 10–100 (3–1000) GeV region and weakly correlated couplings

to bottom quarks and tau leptons with values of 10^{-3} –1 at the 68% credibility level.

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Appendix A: Bayesian Inference

In this section we summarize the statistical background of our analysis. Let $P(A|I)$ and $P(B|I)$ denote the plausibility of two non-exclusive propositions, A and B , in light of some prior information, I . The probability that both A and B are correct is given by the conditional expression

$$P(AB|I) = P(A|BI)P(B|I). \quad (\text{A1})$$

Bayes theorem follows from the symmetry of the conditional probability under the exchange of A and B :

$$P(A|BI) = \frac{P(B|AI)P(A|I)}{P(B|I)}. \quad (\text{A2})$$

In this context $P(A|I)$ is typically called the prior probability and represents the plausibility of our hypothesis given information prior the observation B . The likelihood function $P(B|AI)$ indicates how accurately the hypothesis can replicate the data. The posterior probability $P(A|BI)$ quantifies the plausibility of the hypothesis A given the data B . The evidence $P(B|I)$ serves to normalize the posterior.

For theoretical models with a continuous parameter θ Bayes’ theorem can be recast in the form

$$\mathcal{P}(\theta|B, I) = \frac{\mathcal{L}(B|\theta, I)\pi(\theta, I)}{\epsilon(B, I)}. \quad (\text{A3})$$

The posterior distribution can be used to estimate the most likely region of θ . The evidence is calculated via an integral over the full parameter space

$$\epsilon(B, I) = \int_{\theta} \mathcal{L}(B|\theta, I) \pi(\theta, I) d\theta. \quad (\text{A4})$$

For more than one continuous parameters, θ_i , marginalization is performed by integrating the posterior over various parameters in the higher dimensional parameter space

$$\mathcal{P}(\theta_j) = \int \prod_{i \neq j} d\theta_i \mathcal{P}(\theta_i). \quad (\text{A5})$$

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- [1] L. Goodenough and D. Hooper, arXiv:0910.2998 [hep-ph].
 - [2] D. Hooper and L. Goodenough, Phys. Lett. B **697**, 412 (2011) [arXiv:1010.2752 [hep-ph]].
 - [3] A. Boyarsky, D. Malyshev and O. Ruchayskiy, Phys. Lett. B **705**, 165 (2011) [arXiv:1012.5839 [hep-ph]].
 - [4] K. N. Abazajian and M. Kaplinghat, Phys. Rev. D **86**, 083511 (2012) [arXiv:1207.6047 [astro-ph.HE]].
 - [5] D. Hooper and T. R. Slatyer, Phys. Dark Univ. **2**, 118 (2013) [arXiv:1302.6589 [astro-ph.HE]].
 - [6] C. Gordon and O. Macias, Phys. Rev. D **88**, 083521 (2013) [arXiv:1306.5725 [astro-ph.HE]].
 - [7] T. Daylan, D. P. Finkbeiner, D. Hooper, T. Linden, S. K. N. Portillo, N. L. Rodd and T. R. Slatyer, arXiv:1402.6703 [astro-ph.HE].
 - [8] F. Calore, I. Cholis and C. Weniger, JCAP **1503**, 038 (2015) [arXiv:1409.0042 [astro-ph.CO]].
 - [9] S. Murgia, presentation at Fifth Fermi Symposium, 20-24 Oct. 2014.
 - [10] E. Carlson and S. Profumo, Phys. Rev. D **90**, 023015 (2014) [arXiv:1405.7685 [astro-ph.HE]].
 - [11] R. M. O’Leary, M. D. Kistler, M. Kerr and J. Dexter, arXiv:1504.02477 [astro-ph.HE].
 - [12] J. Petrovic, P. D. Serpico and G. Zaharijas, JCAP **1502**, no. 02, 023 (2015) [arXiv:1411.2980 [astro-ph.HE]].
 - [13] Q. Yuan and K. Ioka, Astrophys. J. **802**, no. 2, 124 (2015) [arXiv:1411.4363 [astro-ph.HE]].
 - [14] J. Petrovic, P. D. Serpico and G. Zaharijas, arXiv:1405.7928 [astro-ph.HE].
 - [15] I. Cholis, D. Hooper and T. Linden, arXiv:1407.5625 [astro-ph.HE].
 - [16] I. Cholis, D. Hooper and T. Linden, arXiv:1407.5583 [astro-ph.HE].
 - [17] R. Bartels, S. Krishnamurthy and C. Weniger, arXiv:1506.05104 [astro-ph.HE].

- [18] S. K. Lee, M. Lisanti, B. R. Safdi, T. R. Slatyer and W. Xue, arXiv:1506.05124 [astro-ph.HE].
- [19] T. D. Brandt and B. Kocsis, *Astrophys. J.* **812**, no. 1, 15 (2015) [arXiv:1507.05616 [astro-ph.HE]].
- [20] W. B. Atwood *et al.* [LAT Collaboration], *Astrophys. J.* **697**, 1071 (2009) [arXiv:0902.1089 [astro-ph.IM]].
- [21] G. A. Gómez-Vargas, M. A. Sánchez-Conde, J. H. Huh, M. Peiró, F. Prada, A. Morselli, A. Klypin and D. G. Cerdeño *et al.*, *JCAP* **1310**, 029 (2013) [arXiv:1308.3515 [astro-ph.HE]].
- [22] K. N. Abazajian, N. Canac, S. Horiuchi and M. Kaplinghat, *Phys. Rev. D* **90**, 023526 (2014) [arXiv:1402.4090 [astro-ph.HE]].
- [23] S. Ipek, D. McKeen and A. E. Nelson, arXiv:1404.3716 [hep-ph].
- [24] P. Ko, W. I. Park and Y. Tang, arXiv:1404.5257 [hep-ph].
- [25] J. M. Cline, G. Dupuis, Z. Liu and W. Xue, *JHEP* **1408**, 131 (2014) [arXiv:1405.7691 [hep-ph]].
- [26] P. Ko and Y. Tang, arXiv:1407.5492 [hep-ph].
- [27] A. Martin, J. Shelton and J. Unwin, arXiv:1405.0272 [hep-ph];
- [28] K. Kong and J. C. Park, arXiv:1404.3741 [hep-ph].
- [29] C. Boehm, M. J. Dolan and C. McCabe, *Phys. Rev. D* **90**, 023531 (2014) [arXiv:1404.4977 [hep-ph]].
- [30] D. K. Ghosh, S. Mondal and I. Saha, arXiv:1405.0206 [hep-ph].
- [31] L. Wang, arXiv:1406.3598 [hep-ph].
- [32] B. D. Fields, S. L. Shapiro and J. Shelton, *Phys. Rev. Lett.* **113**, 151302 (2014) [arXiv:1406.4856 [astro-ph.HE]].
- [33] A. Berlin, D. Hooper and S. D. McDermott, arXiv:1404.0022 [hep-ph].
- [34] K. P. Modak, D. Majumdar and S. Rakshit, *JCAP* **1503**, 011 (2015) [arXiv:1312.7488 [hep-ph]]; K. N. Abazajian, N. Canac, S. Horiuchi and M. Kaplinghat, arXiv:1402.4090 [astro-ph.HE]; E. Hardy, R. Lasenby and J. Unwin, arXiv:1402.4500 [hep-ph]; D. P. Finkbeiner and N. Weiner, arXiv:1402.6671 [hep-ph]; T. Lacroix, C. Boehm and J. Silk, arXiv:1403.1987 [astro-ph.HE]; A. Hektor and L. Marzola, arXiv:1403.3401 [hep-ph]; A. Alves, S. Profumo, F. S. Queiroz and W. Shepherd, arXiv:1403.5027 [hep-ph]; P. Agrawal, B. Batell, D. Hooper and T. Lin, arXiv:1404.1373 [hep-ph]; E. Izaguirre, G. Krnjaic and B. Shuve, arXiv:1404.2018 [hep-ph]; V. Gammaldi, J. A. R. Cembranos, A. de la Cruz-Dombriz, R. A. Lineros and

- A. L. Maroto, arXiv:1404.2067 [hep-ph]; Q. Yuan and B. Zhang, arXiv:1404.2318 [astro-ph.HE]; D. G. Cerdeno, M. Peiro and S. Robles, arXiv:1404.2572 [hep-ph]; M. Demianski and A. Doroshkevich, arXiv:1404.3362 [astro-ph.CO]; S. Ipek, D. McKeen and A. E. Nelson, arXiv:1404.3716 [hep-ph]; K. Kong and J. -C. Park, arXiv:1404.3741 [hep-ph]; C. Boehm, M. J. Dolan and C. McCabe, arXiv:1404.4977 [hep-ph]; P. Ko, W. -I. Park and Y. Tang, arXiv:1404.5257 [hep-ph]; M. Shirasaki, S. Horiuchi and N. Yoshida, arXiv:1404.5503 [astro-ph.CO]; M. Abdullah, A. DiFranzo, A. Rajaraman, T. M. P. Tait, P. Tanedo and A. M. Wijangco, arXiv:1404.6528 [hep-ph]; D. K. Ghosh, S. Mondal and I. Saha, arXiv:1405.0206 [hep-ph]; A. Drlica-Wagner, G. A. Gomez-Vargas, J. W. Hewitt, T. Linden and L. Tibaldo, arXiv:1405.1030 [astro-ph.HE]; J. Bramante and T. Linden, arXiv:1405.1031 [astro-ph.HE]; A. Berlin, P. Gratia, D. Hooper and S. D. McDermott, arXiv:1405.5204 [hep-ph]; T. Basak and T. Mondal, arXiv:1405.4877 [hep-ph]; N. Bernal, J. E. Forero-Romero, R. Garani and S. Palomares-Ruiz, arXiv:1405.6240 [astro-ph.CO]; P. Agrawal, M. Blanke and K. Gemmeller, arXiv:1405.6709 [hep-ph]; T. M. Yoast-Hull, J. S. Gallagher and E. G. Zweibel, arXiv:1405.7059 [astro-ph.HE]; K. Agashe, Y. Cui, L. Necib and J. Thaler, arXiv:1405.7370 [hep-ph]; E. Carlson and S. Profumo, arXiv:1405.7685 [astro-ph.HE]; J. M. Cline, G. Dupuis, Z. Liu and W. Xue, arXiv:1405.7691 [hep-ph]; J. Petrovic, P. D. Serpico and G. Zaharijas, arXiv:1405.7928 [astro-ph.HE]; S. K. N. Portillo and D. P. Finkbeiner, arXiv:1406.0507 [astro-ph.IM]; T. Han, Z. Liu and S. Su, arXiv:1406.1181 [hep-ph]; W. Detmold, M. McCullough and A. Pochinsky, arXiv:1406.2276 [hep-ph]; L. Wang, arXiv:1406.3598 [hep-ph]; C. Boehm, P. Gondolo, P. Jean, T. Lacroix, C. Norman and J. Silk, arXiv:1406.4683 [astro-ph.HE]; B. D. Fields, S. L. Shapiro and J. Shelton, arXiv:1406.4856 [astro-ph.HE]; A. Askew, S. Chauhan, B. Penning, W. Shepherd and M. Tripathi, International Journal of Modern Physics A, Vol. 29 (2014) 1430041 [arXiv:1406.5662 [hep-ph]]; C. Cheung, M. Papucci, D. Sanford, N. R. Shah and K. M. Zurek, arXiv:1406.6372 [hep-ph]; S. D. McDermott, arXiv:1406.6408 [hep-ph].
- [35] P. Ko and Y. Tang, JCAP **1501**, 023 (2015) [arXiv:1407.5492 [hep-ph]]; K. Ghorbani, JCAP **1501**, 015 (2015) [arXiv:1408.4929 [hep-ph]]; A. D. Banik and D. Majumdar, Phys. Lett. B **743**, 420 (2015) [arXiv:1408.5795 [hep-ph]]; D. Borah and A. Dasgupta, Phys. Lett. B **741**, 103 (2015) [arXiv:1409.1406 [hep-ph]]; J. Guo, J. Li, T. Li and A. G. Williams, Phys. Rev. D **91**, no. 9, 095003 (2015) [arXiv:1409.7864 [hep-ph]]; J. Cao, L. Shang, P. Wu,

- J. M. Yang and Y. Zhang, Phys. Rev. D **91**, no. 5, 055005 (2015) [arXiv:1410.3239 [hep-ph]]; M. Heikinheimo and C. Spethmann, JHEP **1412**, 084 (2014) [arXiv:1410.4842 [hep-ph]]; P. Agrawal, B. Batell, P. J. Fox and R. Harnik, JCAP **1505**, no. 05, 011 (2015) [arXiv:1411.2592 [hep-ph]]; K. Cheung, W. C. Huang and Y. L. S. Tsai, arXiv:1411.2619 [hep-ph]; F. Calore, I. Cholis, C. McCabe and C. Weniger, Phys. Rev. D **91**, no. 6, 063003 (2015) [arXiv:1411.4647 [hep-ph]]; A. Biswas, arXiv:1412.1663 [hep-ph]; M. J. Dolan, C. McCabe, F. Kahlhoefer and K. Schmidt-Hoberg, JHEP **1503**, 171 (2015) [arXiv:1412.5174 [hep-ph]]; K. Ghorbani and H. Ghorbani, arXiv:1501.00206 [hep-ph]; J. Kozaczuk and T. A. W. Martin, JHEP **1504**, 046 (2015) [arXiv:1501.07275 [hep-ph]]; C. H. Chen and T. Nomura, arXiv:1501.07413 [hep-ph]; K. P. Modak and D. Majumdar, arXiv:1502.05682 [hep-ph]; A. Achterberg, S. Caron, L. Hendriks, R. Ruiz de Austri and C. Weniger, arXiv:1502.05703 [hep-ph]; J. Conrad, J. Cohen-Tanugi and L. E. Strigari, arXiv:1503.06348 [astro-ph.CO]; J. M. Cline, G. Dupuis, Z. Liu and W. Xue, arXiv:1503.08213 [hep-ph]; E. C. F. S. Fortes, V. Pleitez and F. W. Stecker, arXiv:1503.08220 [hep-ph]; K. Ghorbani and H. Ghorbani, Phys. Rev. D **91**, no. 12, 123541 (2015) [arXiv:1504.03610 [hep-ph]]; P. Ko and Y. Tang, arXiv:1504.03908 [hep-ph]; J. Kim, J. C. Park and S. C. Park, arXiv:1505.04620 [hep-ph]; S. Dado and A. Dar, arXiv:1505.04988 [astro-ph.HE]; O. Buchmueller, S. A. Malik, C. McCabe and B. Penning, arXiv:1505.07826 [hep-ph]; I. Cholis, C. Evoli, F. Calore, T. Linden, C. Weniger and D. Hooper, arXiv:1506.05119 [astro-ph.HE]; T. Mondal and T. Basak, arXiv:1507.01793 [hep-ph]; A. Butter, T. Plehn, M. Rauch, D. Zerwas, S. Henrot-Versill and R. Lafaye, arXiv:1507.02288 [hep-ph]; A. Achterberg, M. van Beekveld, W. Beenakker, S. Caron and L. Hendriks, arXiv:1507.04644 [hep-ph]; G. Bertone, F. Calore, S. Caron, R. R. de Austri, J. S. Kim, R. Trotta and C. Weniger, arXiv:1507.07008 [hep-ph]; D. Kim and J. C. Park, arXiv:1507.07922 [hep-ph]; N. Fonseca, L. Necib and J. Thaler, arXiv:1507.08295 [hep-ph].
- [36] T. Bringmann, M. Vollmann and C. Weniger, Phys. Rev. D **90**, no. 12, 123001 (2014) [arXiv:1406.6027 [astro-ph.HE]].
- [37] I. Cholis, D. Hooper and T. Linden, Phys. Rev. D **91**, no. 8, 083507 (2015) [arXiv:1408.6224 [astro-ph.HE]].
- [38] C. Balázs and T. Li, Phys. Rev. D **90**, no. 5, 055026 (2014) [arXiv:1407.0174 [hep-ph]].
- [39] For recent discussions, see J. M. Cline, K. Kainulainen, P. Scott and C. Weniger, Phys. Rev. D **88**, 055025 (2013) [arXiv:1306.4710 [hep-ph]] and references therein.

- [40] F. D’Eramo and M. Procura, JHEP **1504** (2015) 054 [arXiv:1411.3342 [hep-ph]].
- [41] G. D’Ambrosio, G. F. Giudice, G. Isidori and A. Strumia, Nucl. Phys. B **645**, 155 (2002) [hep-ph/0207036].
- [42] J. Kumar and D. Marfatia, Phys. Rev. D **88**, no. 1, 014035 (2013) [arXiv:1305.1611 [hep-ph]].
- [43] P. A. R. Ade *et al.* [Planck Collaboration], arXiv:1303.5076 [astro-ph.CO].
- [44] M. Aguilar [AMS Collaboration], Phys. Rev. Lett. **113**, 121102 (2014).
- [45] O. Adriani, G. A. Bazilevskaya, G. C. Barbarino, R. Bellotti, M. Boezio, E. A. Bogomolov, V. Bonvicini and M. Bongi *et al.*, JETP Lett. **96**, 621 (2013) [Pisma Zh. Eksp. Teor. Fiz. **96**, 693 (2012)].
- [46] J. M. Cline and P. Scott, JCAP **1303**, 044 (2013) [Erratum-ibid. **1305**, E01 (2013)] [arXiv:1301.5908 [astro-ph.CO]].
- [47] D. S. Akerib *et al.* [LUX Collaboration], Phys. Rev. Lett. **112**, 091303 (2014) [arXiv:1310.8214 [astro-ph.CO]].
- [48] R. D. Davies, D. Walsh, and R. S. Booth, Mon. Not. R. Astron. Soc. **177**, 319 (1976).
- [49] G. Belanger, F. Boudjema, A. Pukhov and A. Semenov, Comput. Phys. Commun. **185**, 960 (2014) [arXiv:1305.0237 [hep-ph]].
- [50] J. F. Navarro, C. S. Frenk and S. D. M. White, Astrophys. J. **462**, 563 (1996) [astro-ph/9508025]; J. F. Navarro, C. S. Frenk and S. D. M. White, Astrophys. J. **490**, 493 (1997) [astro-ph/9611107].
- [51] G. Belanger, F. Boudjema, P. Brun, A. Pukhov, S. Rosier-Lees, P. Salati and A. Semenov, Comput. Phys. Commun. **182**, 842 (2011) [arXiv:1004.1092 [hep-ph]].
- [52] A. Ibarra, A. S. Lamperstorfer and J. Silk, Phys. Rev. D **89**, 063539 (2014) [arXiv:1309.2570 [hep-ph]].
- [53] M. Cirelli, D. Gaggero, G. Giesen, M. Taoso and A. Urbano, arXiv:1407.2173 [hep-ph].
- [54] H. B. Jin, Y. L. Wu and Y. F. Zhou, arXiv:1410.0171 [hep-ph].
- [55] H. B. Jin, Y. L. Wu and Y. F. Zhou, arXiv:1504.04604 [hep-ph].
- [56] D. Hooper, T. Linden and P. Mertsch, JCAP **1503**, no. 03, 021 (2015) [arXiv:1410.1527 [astro-ph.HE]].
- [57] G. Belanger, F. Boudjema and A. Pukhov, arXiv:1402.0787 [hep-ph].
- [58] O. Adriani *et al.* [PAMELA Collaboration], Phys. Rev. Lett. **105**, 121101 (2010) [arXiv:1007.0821 [astro-ph.HE]].

- [59] P. Junnarkar and A. Walker-Loud, Phys. Rev. D **87**, no. 11, 114510 (2013) [arXiv:1301.1114 [hep-lat]].
- [60] A. Crivellin, M. Hoferichter and M. Procura, Phys. Rev. D **89**, 054021 (2014) [arXiv:1312.4951 [hep-ph]].
- [61] C. Savage, A. Scaffidi, M. White and A. G. Williams, arXiv:1502.02667 [hep-ph].
- [62] C. Savage, *in preparation*.
- [63] M. Szydagis, N. Barry, K. Kazkaz, J. Mock, D. Stolp, M. Sweany, M. Tripathi and S. Uvarov *et al.*, JINST **6**, P10002 (2011) [arXiv:1106.1613 [physics.ins-det]].
- [64] M. Szydagis, A. Fyhrie, D. Thorngren and M. Tripathi, JINST **8**, C10003 (2013) [arXiv:1307.6601 [physics.ins-det]].
- [65] A. Alloul, N. D. Christensen, C. Degrande, C. Duhr and B. Fuks, Comput. Phys. Commun. **185**, 2250 (2014) [arXiv:1310.1921 [hep-ph]].
- [66] F. Feroz, M. P. Hobson and M. Bridges, Mon. Not. Roy. Astron. Soc. **398**, 1601 (2009) [arXiv:0809.3437 [astro-ph]].
- [67] J. Skilling, Bayesian Analysis 1 (4) 833860 (2006).
- [68] Y. Mambrini, M. H. G. Tytgat, G. Zaharijas and B. Zaldivar, JCAP **1211**, 038 (2012) [arXiv:1206.2352 [hep-ph]].