



This is the accepted manuscript made available via CHORUS. The article has been published as:

## Superconformal symmetry, NMSSM, and inflation

Sergio Ferrara, Renata Kallosh, Andrei Linde, Alessio Marrani, and Antoine Van Proeyen

Phys. Rev. D **83**, 025008 — Published 10 January 2011

DOI: [10.1103/PhysRevD.83.025008](https://doi.org/10.1103/PhysRevD.83.025008)

# Superconformal Symmetry, NMSSM, and Inflation

Sergio Ferrara<sup>1,2</sup>, Renata Kallosh<sup>3</sup>, Andrei Linde<sup>3</sup>,  
Alessio Marrani<sup>3</sup>, and Antoine Van Proeyen<sup>4</sup>

<sup>1</sup>*Physics Department, Theory Unit, CERN, CH 1211, Geneva 23, Switzerland*

<sup>2</sup>*INFN - Laboratori Nazionali di Frascati, Via Enrico Fermi 40, 00044 Frascati, Italy*

<sup>3</sup>*Department of Physics, Stanford University, Stanford, California 94305 USA*

<sup>4</sup>*Instituut voor Theoretische Fysica, Katholieke Universiteit Leuven,  
Celestijnenlaan 200D, B-3001 Leuven, Belgium*

## Abstract

We identify a particularly simple class of supergravity models describing superconformal coupling of matter to supergravity. In these models, which we call the canonical superconformal supergravity (CSS) models, the kinetic terms in the Jordan frame are canonical, and the scalar potential is the same as in the global theory. The pure supergravity part of the total action has a local Poincaré supersymmetry, whereas the chiral and vector multiplets coupled to supergravity have a larger local superconformal symmetry.

The scale-free globally supersymmetric theories, such as the NMSSM with a scale-invariant superpotential, can be naturally embedded into this class of theories. After the supergravity embedding, the Jordan frame scalar potential of such theories remains scale free; it is quartic, it contains no mass terms, no nonrenormalizable terms, no cosmological constant.

The local superconformal symmetry can be broken by additional terms, which, in the small field limit, are suppressed by the gravitational coupling. This can be achieved by introducing the nonminimal scalar-curvature coupling, and by taking into account interactions with a hidden sector. In this approach, the smallness of the mass parameters in the NMSSM may be traced back to the original superconformal invariance. This allows to address the  $\mu$ -problem and the cosmological domain wall problem in this model, and to implement chaotic inflation in the NMSSM. We discuss the gravitino problem in the NMSSM inflation, as well as the possibility to obtain a broad class of new versions of chaotic inflation in supergravity.

# 1 Introduction

This work is a continuation of our previous paper [1] where we studied generic supergravity in the Jordan frame and the possibility to implement the Higgs-type inflation [2] in the context of the next-to-minimal supersymmetric standard model (NMSSM) [1, 3, 4].

These recent developments were based on a combination of long efforts of many authors in several seemingly unrelated directions.

1) Many decades ago, one of the most popular ways to describe gravitational interactions of a scalar field  $\varphi$  was to assume that it is conformally coupled to gravity, which means that its Lagrangian contains a term  $-\sqrt{-g} \frac{\varphi^2}{12} R$ , see e.g. [5]. With the invention of inflation, which was difficult to achieve for conformally coupled scalars, the concept of scalar fields conformally coupled to gravity gradually lost part of its appeal. On the other hand, several authors emphasized that inflation may occur in a very natural way if a scalar field nonminimally couples to gravity, with a sign opposite to that of the conformal coupling, see e.g. [6]. Recently there was a revival of interest in this possibility after it was realized that it may allow inflation in the standard model, with the Higgs field playing the role of the inflaton [2]. However, for a while it was not clear whether one can implement this idea in supersymmetric generalizations of the standard model. Some progress in this direction was reached only very recently [1, 3, 4]. In this paper we will develop a more systematic approach to this issue.

2) Conformal invariance plays an important role in the formulation of supergravity. The general formulation of supergravity starts with the superconformal theory. Then, after gauge fixing, which, in particular, makes the conformal compensator field proportional to the Planck mass, one derives the standard textbook formulation of supergravity [7–13]. Once this step is made, the theory is formulated in the Einstein frame, all scalars have minimal coupling to gravity, and the superconformal origin of supergravity becomes well hidden.

In [1] we performed an alternative gauge-fixing of the version of the superconformal theory, which allows to derive the supergravity action in an arbitrary Jordan frame. This provides a complete locally supersymmetric theory for scalars with a nonminimal coupling to gravity.

3) Prior to the discovery of supergravity, the development of particle physics was successfully guided by the principle of gauge invariance and renormalizability. However, supergravity is nonrenormalizable. In general, one can write any kind of superpotential which may lead to nonrenormalizable interactions which become important even at low energy. It would be nice to have a formulation of supergravity

where the low-energy renormalizability appears as a result of some general principle, similar to the principle of spontaneously broken gauge invariance.

The extraordinary smallness of the Higgs mass, as compared to the Planck mass, can be protected by supersymmetry, but only if the Higgs mass was extremely small to start with. The minimal supersymmetric standard model (MSSM) and the general next-to-minimal supersymmetric standard model (NMSSM) include several other dimensional parameters which are required to be extraordinary small ( $\mu$ -problem, tadpole problem). These issues can be addressed in the context of the  $\mathbb{Z}_3$ -invariant NMSSM, which requires that the superpotential describing the standard model is scale-invariant [14]. However, it would be important to find some fundamental underpinnings of this requirement. Moreover,  $\mathbb{Z}_3$  symmetry of the scale-invariant superpotential leads to the cosmological domain wall problem.

All of these problems have been discussed extensively in the existing literature, but recent developments stimulated us to look at these issues again in [1], returning back to the superconformal origin of supergravity. As we will see, many of these problems become much easier to address in a class of models where the original superconformal invariance remains at least partially preserved, being broken only by gravitational effects, or by anomalies. This symmetry may naturally explain renormalizability and smallness of the mass parameters in the standard model. It leads to a formulation of supergravity in the Jordan frame, where in certain cases the potentials and kinetic terms look as simple as in the global SUSY. In this context, one can achieve inflation and simultaneously solve the domain wall problem with the help of the terms describing scalars nonminimally coupled to gravity.

To explain our main idea, let us consider a nonsupersymmetric conformally invariant toy model describing gravity and two real scalar fields,  $\phi$  and  $h$ :

$$\mathcal{L} = \sqrt{-g} \left[ \frac{1}{2} \partial_\mu \phi \partial_\nu \phi g^{\mu\nu} + \frac{\phi^2}{12} R(g) - \frac{1}{2} \partial_\mu h \partial_\nu h g^{\mu\nu} - \frac{h^2}{12} R(g) - \frac{\lambda}{4} h^4 \right]. \quad (1.1)$$

The field  $\phi(x)$  is referred to as a conformal compensator. This theory is locally conformal invariant under the following transformations:

$$g'_{\mu\nu} = e^{-2\sigma(x)} g_{\mu\nu}, \quad \phi' = e^{\sigma(x)} \phi, \quad h' = e^{\sigma(x)} h. \quad (1.2)$$

Note that the kinetic term of the conformal compensator  $\phi$  has a wrong sign. This is not a problem because there are no physical degrees of freedom associated with it; the field  $\phi$  can be removed from the theory by fixing the gauge symmetry (1.2). If we choose the gauge  $\phi(x) = \sqrt{6} M_P$ , the  $\phi$ -terms in (1.1) reduce to the Einstein action. The full Lagrangian in the Jordan frame is

$$\mathcal{L}_{\text{total}} = \mathcal{L}_E + \mathcal{L}_{\text{conf}} = \sqrt{-g} \frac{M_P^2}{2} R(g) - \sqrt{-g} \left[ \frac{1}{2} \partial_\mu h \partial_\nu h g^{\mu\nu} + \frac{h^2}{12} R(g) + \frac{\lambda}{4} h^4 \right]. \quad (1.3)$$

It consists of two parts, the Einstein Lagrangian  $\sqrt{-g} \frac{M_P^2}{2} R(g)$ , which is not conformally invariant, and the conformally invariant theory of the canonically normalized scalar field  $h$ ,

$$\mathcal{L}_{\text{conf}} = -\sqrt{-g} \left[ \frac{1}{2} \partial_\mu h \partial_\nu h g^{\mu\nu} + \frac{h^2}{12} R(g) + \frac{\lambda}{4} h^4 \right]. \quad (1.4)$$

As we already mentioned, theories of this type played a very important role in the development of particle physics and cosmology many decades ago, see e.g. [5]. One of the main reasons is that the Friedmann universe is conformally flat. By making a conformal transformation, one could represent equations of motion of the scalar field in the Friedmann universe in terms of equations of motion of a conformally transformed field in Minkowski space, which is a tremendous simplification.

The theory (1.1) is unique if we require that the local conformal symmetry of the  $h$  part of the action, which has canonical kinetic terms, should be preserved after the gauge-fixing. It is determined by the condition that the conformal compensator  $\phi(x)$  is decoupled from the field  $h(x)$ .

The conformal symmetry of the matter action in Eq. (1.4) is manifest in the Jordan frame (1.3). One can make a certain field and metric transformation and switch to the Einstein frame, where the term  $-\sqrt{-g} \frac{h^2}{12} R(g)$  is absorbed into the Einstein action. This allows one to use the standard Einstein equations. However, after this transformation both the gravity part as well as the matter part of the action have conformal symmetry broken.

Similarly, the standard formulation of supergravity interacting with matter brings us directly to the Einstein frame, where the original superconformal symmetry is lost even in the special class of models where matter fields are decoupled from the conformal compensator. That is why it was hard to see any advantages of this class of models in the standard textbook formulation of supergravity. Meanwhile, as we will see shortly, in the class of models with conformal coupling of scalars, the matter Lagrangian in the Jordan frame looks exceptionally simple: all kinetic terms are canonical in the simplest case, the superpotential contains only cubic terms, and the scalar potential is quartic with respect to the scalar fields, just as in our toy model (1.3). The theories of this class provide a very natural supergravity embedding of the  $\mathbb{Z}_3$ -invariant NMSSM with a scale-free superpotential.

Of course, in the end of the day we want to make most of the particles massive. Thus, we would need to break superconformal invariance, but we would like to do it in a way that preserves some of the most attractive features of the original superconformal theory.

Superconformal symmetry may be broken by anomalies, by interaction with a hidden sector, or by gravitational interactions suppressed by inverse powers of the Planck mass. However, one should try

to avoid introducing into the original theory any terms proportional to the compensator field. For example, one could add to (1.1) a term  $c^2 h^2 \phi^2$  without breaking the conformal invariance of the model (1.1). However, in the gauge  $\phi = \sqrt{6} M_P$ , the matter part of the theory (1.3) would acquire the term  $6 c^2 M_P^2 h^2$ . This term strongly breaks the conformal invariance by giving the field  $h$  a mass squared  $12 c^2 M_P^2$ , which is enormously large unless the dimensionless constant  $c$  is extremely small. This would lead to the hierarchy problem if one tries to use the field  $h$  for the description of the low-energy physics. Similarly, the term  $c \phi^3 h$  would introduce a huge tadpole  $\sim M_P^3 h$ , and the term  $\phi^4$  would introduce an enormously large cosmological constant  $\sim M_P^4$ . We would like to use the original (super)conformal symmetry to protect us against such problems.

On the other hand, the terms  $h^{4+n}/\phi^n$ , which are inversely proportional to  $\phi^n$ , would lead to nonrenormalizable interaction terms  $\sim h^{4+n}/M_P^n$ . While such terms are unpleasant, they usually appear in the Einstein frame anyway, and they are not expected to affect particle physics at energies and fields much smaller than  $M_P$ .

In this paper, we will break the local superconformal symmetry of the matter coupling via the real part of the quadratic holomorphic function in scalar-curvature coupling, by the terms in the Kähler potential which are suppressed by the inverse Planck mass, and also by the interaction to the hidden sector. This produces additional terms in the action. At small fields and low energies, the new terms are suppressed by inverse powers of the Planck mass. In other words, at small fields and low energies, the original superconformal symmetry is broken only by effects suppressed by the small gravitational coupling. This can be very helpful in particle phenomenology. The smallness of the new terms helps to explain the smallness of the Higgs mass and of the  $\mu$ -term, which appear only because of the breaking of the superconformal invariance. On the other hand, these terms can be large enough to address the domain wall problem in the NMSSM. Moreover, in the large field limit, some of the new terms become dominant and allow to implement inflation in supergravity along the lines of [2, 6] and [1, 3, 4].

The paper is organized as follows. In Sec. 2 we give a short summary of the results of Ref. [1] on the supergravity action in an arbitrary Jordan frame defined by the function of scalars  $\Phi(z, \bar{z})$ . We then focus on a special case of  $\Phi(z, \bar{z}) = -3M_P^2 e^{-\frac{\mathcal{K}(z, \bar{z})}{3M_P^2}} = -3M_P^2 + \delta_{\alpha\bar{\beta}} z^\alpha \bar{z}^\beta + J(z) + \bar{J}(\bar{z})$ , where  $J(z)$  is a holomorphic function quadratic in  $z$ . It has been found in [1] that in this case the kinetic term for scalars is canonical since it is defined by  $\Phi_{\alpha\bar{\beta}} = \delta_{\alpha\bar{\beta}}$ . We explain the role of the auxiliary supergravity vector fields  $A_\mu$ .

Sec. 3 presents the simplest possible embedding of the globally superconformal theory into su-

pergravity. We start with the analog of (1.1), the theory with the  $SU(2,2|1)$  local superconformal symmetry and no dimensional parameters. It contains an extra chiral multiplet, the compensator one. The theory is based on the results obtained in [8,9], and more recent results of Refs. [13] and [1]. We then specify the superconformal action for the case when a) all kinetic terms are canonical and b) the matter multiplets decouple from the compensator. We find that in these models the total supergravity action consists of the pure supergravity part, which breaks superconformal symmetry, and the matter part, which remains superconformal after the gauge-fixing. The scalars are conformally coupled to gravity, kinetic terms are canonical, and the supergravity potential coincides with the global theory potential. We call these theories *Canonical Superconformal Supergravity* (CSS) models.

In this sense, the embedding of the globally superconformal theory into supergravity in the Jordan frame becomes a simple additive operation: One adds the action of the global SUSY, interacting with gravity with conformal scalar-curvature coupling, to the action of supergravity; that is it. However, this simple operation looks much more complicated in the Einstein frame.

We further develop a geometric way to break the superconformal coupling of matter to supergravity. The flat Kähler geometry of the chiral multiplets, including the compensator field, is replaced by a non-flat geometry without introducing any new dimensional parameters. Specifically, we study CSS models with superconformal symmetry broken by the  $\chi$ -term: the real part of the holomorphic function defining the Jordan frame. This leads to useful applications both in particle physics and cosmology.

In Sec. 4 we apply the method of embedding globally supersymmetric theories into supergravity described in Sec. 3 to the scale-invariant version of the NMSSM. Sec. 5 has a short discussion of the issues of the NMSSM phenomenology, including the  $\mu$ -problem and domain wall problem. We argue that using the superconformal matter action with the  $\chi$ -term in combination with a hidden sector may resolve both of these problems.

Sections. 6 and 7 are devoted to inflation. We first review the Higgs type inflation in the standard model, following [2]. We argue that this inflationary model, as well as its NMSSM generalization proposed in [3] and developed in [1,4], does not suffer from the problems related to the unitarity bound discussed in [15–18]. We describe observational implications of inflation in the NMSSM and find that these implications are invariant with respect to a certain rescaling of the parameters of this model. We describe a mechanism of stabilization of the inflationary trajectory, which is necessary for consistency of this scenario. It includes a requirement of special corrections which stabilize some moduli at the origin of the moduli space. We discuss the gravitino problem, which may appear in

this scenario, and point out the existence of a broad class of new inflationary models based on the ideas described in our paper, where the inflaton is not necessarily related to the Higgs field and the gravitino problem may not arise [21].

Our results are briefly summarized in the Conclusions. Appendix A presents a complete action of superconformal matter coupled to gravity, including vector multiplets and fermions with canonical kinetic terms. It is given in equations (A.10)-(A.14) which provide the generalization of the scalar-gravity action in equations (3.25), (3.26), when all fermions and vectors are included. Appendix B shows how one can derive the simple CSS potential, which is the same as in global SUSY theories, starting from the generic Einstein frame supergravity potential. Appendix C provides the metric of the moduli space for CSS models with superconformal symmetry broken by the real part of the holomorphic function. Appendix D presents a detailed expression for the potential of the scalar field  $s$  and the inflaton field  $h$  in the Jordan frame and in the Einstein frame.

Thus, the paper essentially consists of two parts. Those who are interested mainly in phenomenological and cosmological implications of our construction, may take a quick look at Sections 2-4 and then proceed directly to Sections 5-7. However, we believe that the superconformal approach to supergravity and the CSS models described in Sections 2-4 deserve further investigation quite independently of their immediate implications for inflation and the NMSSM.

## 2 Supergravity in the Jordan frame

The general theory of supergravity in an arbitrary Jordan frame was derived in [1] by a gauge-fixing of the  $SU(2, 2|1)$  superconformal theory [13]. This approach is based on earlier work on superconformal origin of the supergravity theory in [22]. The extra gauge symmetries of the superconformal theory, including a local conformal symmetry, which rescales the metric, allow a possibility to derive the supergravity action either in the Einstein frame or in an arbitrary Jordan frame. The Einstein frame Lagrangian, in units  $M_P = 1$ , is  $\mathcal{L}_E = \sqrt{-g_E} \frac{1}{2} R(g_E) + \dots$ , there is no direct scalar-curvature coupling. The Jordan frame Lagrangian is  $\mathcal{L}_J = -\sqrt{-g_J} \frac{\Phi(z, \bar{z})}{6} R(g_J) + \dots$ , where  $\Phi(z, \bar{z})$  is an arbitrary function of complex scalar fields  $z, \bar{z}$ . Therefore, in general, there is a scalar-curvature coupling in the Jordan frame. A local conformal symmetry allows to make a choice of  $\Phi(z, \bar{z}) = -3$  to get the Einstein frame supergravity. Otherwise with the frame function depending on scalars we get the Jordan frame supergravity. The relation between the space-time metrics is given by  $g_E^{\mu\nu} = \Omega^2(z, \bar{z}) g_J^{\mu\nu}$ , where



$$\Omega^2(z, \bar{z}) = -\frac{1}{3} \Phi(z, \bar{z}).$$

The scalar-gravity part of the  $\mathcal{N} = 1$ ,  $d = 4$  supergravity in a generic Jordan frame with frame function  $\Phi(z, \bar{z})$ , a Kähler potential  $\mathcal{K}(z, \bar{z})$  independent on the frame function, and superpotential  $W(z)$  is, according to [1],

$$\mathcal{L}_J^{\text{scalar-grav}} = \sqrt{-g_J} \left[ \Phi \left( -\frac{1}{6} R(g_J) + \mathcal{A}_\mu^2(z, \bar{z}) \right) + \left( \frac{1}{3} \Phi g_{\alpha\bar{\beta}} - \frac{\Phi_\alpha \Phi_{\bar{\beta}}}{\Phi} \right) \hat{\partial}_\mu z^\alpha \hat{\partial}^\mu \bar{z}^{\bar{\beta}} - V_J \right]. \quad (2.1)$$

Here

$$\begin{aligned} \Phi_\alpha &\equiv \frac{\partial}{\partial z^\alpha} \Phi(z, \bar{z}), & \Phi_{\bar{\beta}} &\equiv \frac{\partial}{\partial \bar{z}^{\bar{\beta}}} \Phi(z, \bar{z}) = \bar{\Phi}_{\bar{\beta}}, \\ g_{\alpha\bar{\beta}} &= \frac{\partial^2 \mathcal{K}(z, \bar{z})}{\partial z^\alpha \partial \bar{z}^{\bar{\beta}}} \equiv \mathcal{K}_{\alpha\bar{\beta}}(z, \bar{z}), \end{aligned} \quad (2.2)$$

and  $\mathcal{A}_\mu$  is the purely bosonic part of the on-shell value of the auxiliary field  $A_\mu$ . On shell it depends on scalar fields as follows:

$$\mathcal{A}_\mu(z, \bar{z}) \equiv -\frac{i}{2\Phi} \left( \hat{\partial}_\mu z^\alpha \partial_\alpha \Phi - \hat{\partial}_\mu \bar{z}^{\bar{\alpha}} \partial_{\bar{\alpha}} \Phi \right). \quad (2.3)$$

The gauge covariant derivative  $\hat{\partial}_\mu z^\alpha$  in Eqs. (2.1), (2.3) is

$$\hat{\partial}_\mu z^\alpha \equiv \partial_\mu z^\alpha - A_\mu^A k_A^\alpha, \quad (2.4)$$

where  $A_\mu^A$  is the vector gauge field and  $k_A^\alpha$  is the Killing vector, defining the gauge transformations of scalars,  $\delta z^\alpha = \theta^A k_A^\alpha$ . The Jordan frame potential

$$V_J = \frac{\Phi^2}{9} V_E \quad (2.5)$$

is defined via the Einstein frame potential

$$V_E = V_E^F + V_E^D = e^{\mathcal{K}} \left( -3W\bar{W} + \nabla_\alpha W g^{\alpha\bar{\beta}} \nabla_{\bar{\beta}} \bar{W} \right) + \frac{1}{2} (\text{Re } f)^{-1AB} P_A P_B, \quad (2.6)$$

where  $\nabla_\alpha W$  denotes the Kähler-covariant derivative of the superpotential and  $P_A$  is a momentum map. A special important class of the superconformal models with

$$\Phi(z, \bar{z}) = -3e^{-\frac{1}{3}\mathcal{K}(z, \bar{z})} \quad (2.7)$$

and the corresponding actions in the Jordan frame were derived in components in [7, 9], and in superspace in [10, 11]. In this case the simpler form of  $\mathcal{L}_J$  given by (2.1) was found in [1]:

$$\mathcal{L}_J = \sqrt{-g_J} \left[ \Phi \left( -\frac{1}{6} R(g_J) + \mathcal{A}_\mu^2(z, \bar{z}) \right) - \Phi_{\alpha\bar{\beta}} \hat{\partial}_\mu z^\alpha \hat{\partial}^\mu \bar{z}^{\bar{\beta}} - \frac{\Phi^2}{9} V_E \right]. \quad (2.8)$$

Also an interesting observation about the Jordan frame kinetic terms for scalars was made: *For a particular choice of the frame function the kinetic scalar terms are canonical* when the on-shell auxiliary axial-vector field  $\mathcal{A}_\mu$  vanishes. This requires that

$$\Phi(z, \bar{z}) = -3e^{-\frac{1}{3}\mathcal{K}(z, \bar{z})} = -3 + \delta_{\alpha\bar{\beta}} z^\alpha \bar{z}^{\bar{\beta}} + J(z) + \bar{J}(\bar{z}) , \quad (2.9)$$

and it follows that

$$\Phi_{\alpha\bar{\beta}} \equiv \frac{\partial^2 \Phi(z, \bar{z})}{\partial z^\alpha \partial \bar{z}^{\bar{\beta}}} = \delta_{\alpha\bar{\beta}} , \quad (2.10)$$

where  $\mathcal{K}(z, \bar{z})$  is the Kähler potential and  $J(z)$  is holomorphic. For the choice (2.9) the action in the Jordan frame is

$$\frac{\mathcal{L}_J}{\sqrt{-g_J}} = \Phi(z, \bar{z}) \left( -\frac{1}{6}R(g_J) + \mathcal{A}_\mu^2(z, \bar{z}) \right) - \delta_{\alpha\bar{\beta}} \hat{\partial}_\mu z^\alpha \hat{\partial}^\mu \bar{z}^{\bar{\beta}} - V_J(z, \bar{z}) , \quad (2.11)$$

where  $\mathcal{A}_\mu(z, \bar{z})$  is defined in Eq. (2.3). It vanishes in many cosmological applications with either real or imaginary scalar fields. For such configurations with  $\mathcal{A}_\mu = 0$  the second term in Eq. (2.11) is a canonical kinetic term for scalars. This simplification of the supergravity theory in the Jordan frame with regard to kinetic terms of scalars is, as we will see below, a particular property of the class of supergravity theories which have a superconformal matter-supergravity coupling.

### 3 Superconformal matter coupling in the Jordan frame supergravity

#### 3.1 Locally superconformal theory

Superconformal theory is the starting point to derive supergravity. We will consider here a class of models where the chiral and vector multiplets do not interact with the superconformal compensator field. This will provide a simple embedding of globally supersymmetric models into supergravity in the Jordan frame. We will further introduce a geometric mechanism of breaking of the superconformal symmetry, which is suitable for phenomenology and cosmology. The superconformal symmetry will also be broken by radiative corrections and by terms suppressed by inverse powers of  $M_P^2$ .

To embed a given globally supersymmetric model into supergravity, a particular Kähler potential has to be chosen, which could be any real function of all scalars (with positive definite metric of moduli space). In the Einstein frame, where there is no direct coupling of curvature to scalars, the kinetic term for scalars is  $\mathcal{L}_{\text{kin}} = \mathcal{K}_{\alpha\bar{\beta}}(z, \bar{z}) \partial z^\alpha \partial \bar{z}^{\bar{\beta}}$  and the F-term potential in the Einstein frame is  $V_E^F = e^{\mathcal{K}} \left( \nabla_\alpha W \mathcal{K}^{\alpha\bar{\beta}} \nabla_{\bar{\beta}} \bar{W} - 3W\bar{W} \right)$  where  $\nabla_\alpha W$  denotes the Kähler-covariant derivative of the

superpotential. If one would like to preserve the property of the globally supersymmetric theory to have canonical kinetic terms, one has to take  $\mathcal{K} = \delta_{\alpha\bar{\beta}} z^\alpha \bar{z}^\beta$ , up to Kähler transformation. But in such case the F-term potential is quite different from the global supersymmetry case  $V_{\text{global}} = |\partial W|^2$ .

Here we would like to present a simple case of the embedding of a class of scale-invariant globally supersymmetric models into supergravity. Embedding, in general, means that the total action of supergravity and chiral and vector  $\mathcal{N}=1$  multiplets has a local Poincaré supersymmetry. A special class that we will present here has the property that the part of the action describing chiral and vector multiplets coupled to supergravity has a much larger local superconformal symmetry. This symmetry is broken down to the local Poincaré supersymmetry only by the part of the action describing the self-interacting supergravity multiplet. First we start with the analog of (1.1), the theory with the  $SU(2, 2|1)$  superconformal symmetry and no dimensional parameters: it contains an extra chiral multiplet, the compensator one. The theory is based on Refs. [8, 9] and [1, 13]. We then specify the superconformal action for the case when the matter multiplets decouple from the compensator. As a result, we find in this class of theories that the total supergravity action consists of the pure supergravity part, which breaks superconformal symmetry, and the matter part which remains superconformal after the gauge-fixing.

The supergravity Weyl multiplet consists of the vierbein, gravitino and the vector gauge field of the  $U(1) \mathcal{R}$  symmetry:  $e_\mu^a, \psi_\mu, A_\mu$ . The chiral multiplet has scalars and spinors, the vector multiplet has gauge fields and gauginos.

We start here with the superconformal action described in details in [13] in eqs. (3.3)-(3.8) and more recently <sup>1</sup> in [1] in Sec. 5.1. This action has a local  $SU(2, 2|1)$  superconformal symmetry and no dimensional parameters. The symmetries include local dilatation, special conformal symmetry, special supersymmetry and local  $U(1) \mathcal{R}$  symmetry, in addition to all local symmetries of supergravity. The special conformal symmetry has an independent field  $b_\mu$  as a gauge field, and the local  $U(1) \mathcal{R}$  symmetry has  $A_\mu$  as a gauge field. When the local dilatation, special conformal symmetry, special supersymmetry and local  $U(1) \mathcal{R}$  symmetry are gauge-fixed, one finds a generic supergravity theory without extra symmetries. In notation of [1] the action contains 3 superconformal-invariant terms

$$\mathcal{L}_{\text{sc}} = [\mathcal{N}(X, \bar{X})]_D + [\mathcal{W}(X)]_F + [f_{AB} \bar{\lambda}^A P_L \lambda^B]_F. \quad (3.1)$$

---

<sup>1</sup>The action in eqs. (3.3)-(3.8) of [13] has an auxiliary vector field  $A_\mu$  as an independent field, before the equations of motion have been used. In Sec. 5.1 of [1] the superconformal action has  $A_\mu$  already on-shell, with purely bosonic and fermionic part respectively given by the first and second expression of Eq. (5.13) of [1]. Here it is important for us to keep  $A_\mu$  as an off-shell independent field.

The first one codifies the Kähler potential for the  $n + 1$  superconformal fields  $X^I, \bar{X}^{\bar{J}}$ , the second introduces a superpotential, and the third involves the chiral kinetic matrix  $f_{AB}(X)$  (where  $A, B$  are the gauge indices), and gauginos  $\lambda^A$ ,  $P_L$  projects on the left-handed fermions. The superconformal chiral multiplets contain the bosonic fields  $X^I$ , fermions  $\Omega^I$ , and auxiliary fields  $F^I$ ,  $I = 0, 1, \dots, n$ . One of the multiplets,  $X^0, \Omega^0, F^0$ , can be viewed as a compensator multiplet. Its purpose is to provide the part of the local superconformal symmetries which are absent in supergravity.

The dilatation symmetry implies  $\mathcal{N}(X, \bar{X})$  to be homogeneous of first degree in both  $X$  and  $\bar{X}$ ,  $\mathcal{W}(X)$  to be homogeneous of third degree in  $X$ , and  $f_{AB}(X)$  to be homogeneous of zeroth degree in  $X$ . Under chiral  $U(1) \mathcal{R}$  symmetry  $\mathcal{N}(X, \bar{X})$  and  $f_{AB}(X)$  are neutral,  $\mathcal{W}(X)$  has chiral weight 3,  $\bar{\mathcal{W}}(\bar{X})$  has chiral weight -3.

The  $n + 1$  scalars including the compensator multiplet form a Kähler manifold with metric, connection and curvature given, respectively, by

$$G_{I\bar{J}} = \partial_I \partial_{\bar{J}} \mathcal{N} \equiv \frac{\partial \mathcal{N}(X, \bar{X})}{\partial X^I \partial \bar{X}^{\bar{J}}} \quad (3.2)$$

and  $\Gamma_{JK}^I = G^{I\bar{L}} \mathcal{N}_{JK\bar{L}}$ ,  $R_{I\bar{K}J\bar{L}} = \mathcal{N}_{IJ\bar{K}\bar{L}} - \mathcal{N}_{IJ\bar{M}} G^{M\bar{M}} \mathcal{N}_{M\bar{K}\bar{L}}$ . For example, the complete gravity-scalar part of the  $SU(2, 2|1)$  invariant superconformal action has a gravity part, kinetic terms for scalars and a potential:

$$\frac{1}{\sqrt{-g}} \mathcal{L}_{\text{sc}}^{\text{scalar-grav}} = -\frac{1}{6} \mathcal{N}(X, \bar{X}) R - G_{I\bar{J}} D^\mu X^I D_\mu \bar{X}^{\bar{J}} - V_{\text{sc}}, \quad (3.3)$$

where

$$V_{\text{sc}} = V_F + V_D = G^{I\bar{J}} \mathcal{W}_I \bar{\mathcal{W}}_{\bar{J}} + \frac{1}{2} (\text{Re } f)^{-1AB} \mathcal{P}_A \mathcal{P}_B. \quad (3.4)$$

Here

$$\mathcal{W}_I \equiv \frac{\partial \mathcal{W}}{\partial X^I}, \quad \bar{\mathcal{W}}_{\bar{J}} \equiv \frac{\partial \bar{\mathcal{W}}}{\partial \bar{X}^{\bar{J}}}. \quad (3.5)$$

The F-term potential originates from the solution for the auxiliary field for the chiral multiplet,  $F^I = G^{I\bar{J}} \bar{\mathcal{W}}_{\bar{J}}$ . The D-term potential originates from the solution for the auxiliary field for the vector multiplet,  $D^A = (\text{Re } f)^{-1AB} \mathcal{P}_B$ , where  $\mathcal{P}_A$  is the momentum map defining the D-term potential. The covariant derivative  $D_\mu$  in Eq. (3.3) is

$$D_\mu X^I = \partial_\mu X^I - b_\mu X^I - i A_\mu X^I - A_\mu^A k_A^I, \quad (3.6)$$

where  $A_\mu^A$  is the vector gauge field and  $k_A^I$  is the Killing vector, defining the gauge transformations of scalars  $\delta X^I = \theta^A k_A^I$ . We have included the gauge vectors  $A_\mu^A$  into the covariant derivatives (3.6) to make clear the relation between gauge symmetries and D-term potential where

$$\mathcal{P}_A = i \mathcal{N}_I k_A^I = \mathcal{P}_A^\dagger. \quad (3.7)$$

### 3.2 Canonical superconformal supergravity (CSS) models

Here we introduce a set of canonical superconformal supergravity (CSS) models, starting from the  $SU(2, 2|1)$  superconformal action. In these models, after the gauge-fixing of local dilatation, special conformal symmetry, special supersymmetry and local  $U(1)\mathcal{R}$  symmetry *the resulting action for the  $n$  chiral multiplets and all vector multiplets remains superconformal*. The reason for this is that the original superconformal theory before gauge-fixing has all  $n$  chiral multiplets and vectors multiplets *decoupled from the compensator multiplet  $X^0, \Omega^0, F^0$* . We focus here on the *simplest version of such superconformal matter coupling models when kinetic terms are canonical*. We therefore define CSS class of models by the following conditions:

1. We choose<sup>2</sup> flat  $SU(1, n)$  Kähler manifold for all  $n + 1$  chiral multiplets  $X^I$ , including the compensator field  $X^0$

$$\mathcal{N}(X, \bar{X}) = -|X^0|^2 + |X^\alpha|^2, \quad \alpha = 1, \dots, n. \quad (3.8)$$

This means that  $\eta_{0\bar{0}} = -1$ ,  $\eta_{\alpha\bar{\beta}} = \delta_{\alpha\bar{\beta}}$ , where

$$G_{I\bar{J}} = \mathcal{N}_{I\bar{J}} = \eta_{I\bar{J}}, \quad G^{I\bar{J}} = \eta^{I\bar{J}}, \quad \Gamma_{JK}^I = 0, \quad R_{I\bar{K}J\bar{L}} = 0. \quad (3.9)$$

2. We choose a cubic,  $X^0$  independent superpotential, which breaks the  $SU(1, n)$  symmetry:

$$\mathcal{W}(X) = \frac{1}{3}d_{\alpha\beta\gamma}X^\alpha X^\beta X^\gamma \quad \Rightarrow \quad \mathcal{W}_0 \equiv \frac{\partial \mathcal{W}}{\partial X^0} = 0. \quad (3.10)$$

3. We choose a constant complex vector kinetic matrix and  $\text{Re } f_{AB}$  is a constant positive definite matrix.

4. We choose an  $X^0$  independent momentum map. The conformal requirements respecting (3.8) implies that the transformations are of the form

$$k_A^\alpha = (m_A)^\alpha{}_\beta X^\beta, \quad k_A^{\bar{\alpha}} = (\bar{m}_A)^{\bar{\alpha}}{}_{\bar{\beta}} X^{\bar{\beta}}, \quad \mathcal{P}_A = i\delta_{\alpha\bar{\beta}} X^{\bar{\beta}} (m_A)^\alpha{}_\gamma X^\gamma = -i\delta_{\alpha\bar{\beta}} X^\alpha (\bar{m}_A)^{\bar{\beta}}{}_{\bar{\gamma}} \bar{X}^{\bar{\gamma}}, \quad (3.11)$$

where  $m_A$  are anti-hermitian matrices:  $\delta_{\alpha\bar{\beta}}(m_A)^\alpha{}_\beta = -\delta_{\beta\bar{\alpha}}(\bar{m}_A)^{\bar{\alpha}}{}_{\bar{\beta}}$ . Hence the gauge group is part of  $U(n)$ . This means that the compensator  $X^0$  does not participate in Yang-Mills transformations, the zero component of the Killing vector vanishes,  $k_A^0 = 0$ , and other components do not depend on  $X^0$ .

For the decoupling of matter from compensator a more general class of the Kähler manifold for all  $n + 1$  chiral multiplets is possible. However, more general choices after gauge-fixing of the local confor-

---

<sup>2</sup>After the compensator field is gauge-fixed, the manifold of physical  $n$  complex scalars becomes that of  $\frac{SU(1, n)}{U(n)}$  non-compact space.

mal symmetry will lead to non-canonical kinetic terms for scalars. Our choice of scalar-independent  $\text{Re } f_{AB}$  was made to have canonical kinetic terms for vectors. Our choice of  $\text{Im } f_{AB}$  and of cubic superpotentials could have included a dependence on some ratios of homogeneous scalars  $\frac{X^\alpha}{X^\beta}$ . We do not consider such theories here for simplicity, and also since they would blow up at  $\langle X^\beta \rangle = 0$ .

We impose our 4 conditions above on the superconformal action (3.1) and find a superconformal action of this special kind. The scalar-gravity part of the superconformal action (neglecting fermions and gauge vector fields)<sup>3</sup> becomes

$$\frac{1}{\sqrt{-g}}\hat{\mathcal{L}}_{\text{sc}} = \frac{1}{6}(|X^0|^2 - |X^\alpha|^2)R - \eta_{I\bar{J}}D^\mu X^I D_\mu \bar{X}^{\bar{J}} - \delta^{\alpha\bar{\beta}}\mathcal{W}_\alpha \overline{\mathcal{W}}_{\bar{\beta}} - \frac{1}{2}(\text{Re } f)^{-1AB}\mathcal{P}_A\mathcal{P}_B. \quad (3.12)$$

We may split the total superconformal action into parts depending on the compensator multiplet  $X^0, \Omega^0, F^0$  and the part not depending on it. In our class of models we get

$$\frac{1}{\sqrt{-g}}\hat{\mathcal{L}}_{\text{sc}}^0 = \frac{1}{6}|X^0|^2 R + D^\mu X^0 D_\mu \bar{X}^{\bar{0}}, \quad (3.13)$$

$$\frac{1}{\sqrt{-g}}\hat{\mathcal{L}}_{\text{sc}}^{\text{m}} = -\frac{1}{6}|X^\alpha|^2 R - \delta_{\alpha\bar{\beta}}D^\mu X^\alpha D_\mu \bar{X}^{\bar{\beta}} - \delta^{\alpha\bar{\beta}}\mathcal{W}_\alpha \overline{\mathcal{W}}_{\bar{\beta}} - \frac{1}{2}(\text{Re } f)^{-1AB}\mathcal{P}_A\mathcal{P}_B. \quad (3.14)$$

Each of these two actions is separately superconformal (when fermions and vectors are added). In the absence of fermions and vectors they have local conformal and local  $U(1)$  chiral symmetry. The matter part of the action  $\hat{\mathcal{L}}_{\text{sc}}^{\text{m}}$  does not depend on  $X^0$  and therefore it remains superconformal after the gauge-fixing.

### 3.3 Gauge-fixing

Now we proceed with the gauge-fixing of local symmetries that are absent in supergravity. We change variables from the basis  $\{X^I\}$  to a basis  $\{y, z^\alpha\}$ , where  $\alpha = 1, \dots, n$  using  $X^I = y Z^I(z)$ . We now fix *the special conformal symmetry*:

$$b_\mu = 0. \quad (3.15)$$

The *dilatational and  $U(1)$  symmetries* are fixed by a choice  $\mathcal{N}(X, \bar{X}) = -|X^0|^2 + |X^\alpha|^2 = \Phi(z, \bar{z})$  and<sup>4</sup>

$$X^0 = \bar{X}^{\bar{0}} = \sqrt{3} M_P, \quad y = \bar{y} = 1, \quad X^\alpha = z^\alpha. \quad (3.16)$$

---

<sup>3</sup>The complete action for this class of models is presented in the Appendix A.

<sup>4</sup>Here we restore the value of  $M_P$  to stress that after the gauge-fixing of the superconformal action only one dimensional parameter,  $M_P$ , appears in the supergravity action. Moreover, in the class of models described above, in the Jordan frame, the matter part of the action does not depend on  $M_P$ , since it was independent on  $X^0$ .

The special supersymmetry is fixed<sup>5</sup> by the matching requirement on fermions that

$$\Omega^0 = 0, \quad \Omega^\alpha = \chi^\alpha. \quad (3.17)$$

This choice of the gauge-fixing provides a decoupling of the matter multiplets ( $X^\alpha = z^\alpha, \Omega^\alpha = \chi^\alpha, F^\alpha = \delta^{\alpha\bar{\beta}} \frac{\partial \bar{W}(\bar{z})}{\partial \bar{z}^{\bar{\beta}}}$ ) from the compensator multiplet ( $X^0, \Omega^0, F^0$ ). This leads to

$$\hat{\Phi}(z, \bar{z}) = -3M_P^2 + \delta_{\alpha\bar{\beta}} z^\alpha \bar{z}^{\bar{\beta}}, \quad \mathcal{W}(X) = W(z) = \frac{1}{3} d_{\alpha\beta\gamma} z^\alpha z^\beta z^\gamma. \quad (3.18)$$

After the gauge fixing the scalar-gravity part of the supergravity action is

$$\frac{1}{\sqrt{-g}} \hat{\mathcal{L}}_{\text{sg}}^0 = \frac{1}{2} M_P^2 (R + 6A_\mu A^\mu), \quad (3.19)$$

$$\frac{1}{\sqrt{-g}} \hat{\mathcal{L}}_{\text{sg}}^{\text{m}} = -\frac{1}{6} |z^\alpha|^2 R - \delta_{\alpha\bar{\beta}} D^\mu z^\alpha D_\mu \bar{z}^{\bar{\beta}} - \delta^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}} - \frac{1}{2} (\text{Re } f)^{-1 AB} P_A P_B, \quad (3.20)$$

where the  $U(1)$   $\mathcal{R}$  covariant derivative acting on scalars is

$$D_\mu z^\alpha = \partial_\mu z^\alpha - i A_\mu z^\alpha. \quad (3.21)$$

and

$$W_\alpha \equiv \frac{\partial W(z)}{\partial z^\alpha} = d_{\alpha\beta\gamma} z^\beta z^\gamma. \quad (3.22)$$

When the corresponding fermion and vector fields are added to this action, one finds that the total action has an unbroken local Poincaré supersymmetry. The crucial difference with generic case of supergravity theory is that the matter part of action whose scalar-gravity part is given in Eq. (3.20), remains superconformal invariant. In particular, the scalar-gravity action in Eq. (3.20) is invariant under simultaneous local conformal transformation of the metric and scalars, the vector being inert,

$$g'_{\mu\nu} = e^{-2\sigma(x)} g_{\mu\nu}, \quad z' = e^{\sigma(x)} z, \quad \bar{z}' = e^{\sigma(x)} \bar{z}, \quad A'_\mu = A_\mu. \quad (3.23)$$

It is also invariant under local  $U(1)$   $\mathcal{R}$ -symmetry, which is part of the superconformal  $SU(2, 2|1)$  symmetry,

$$g'_{\mu\nu} = g_{\mu\nu}, \quad z' = e^{i\Lambda(x)} z, \quad \bar{z}' = e^{-i\Lambda(x)} \bar{z}, \quad A'_\mu = A_\mu + \partial_\mu \Lambda(x). \quad (3.24)$$

The action (3.19) is nonconformal; it describes the gravitational multiplet, including the auxiliary field  $A_\mu$ , and it is not invariant under local conformal transformations nor under local  $\mathcal{R}$ -symmetry. The total self-coupling of the gravitational multiplet breaks the superconformal symmetry down to super-Poincaré.

---

<sup>5</sup>While the gauge fixing for dilatations agrees with the choice made in [1], we made a different gauge choice for special supersymmetry, which is chosen here in order that the compensating multiplet does not mix with the physical multiplets.

### 3.4 A simple Jordan frame supergravity action with superconformal matter

Let us summarize what we have learned in the previous couple of sections. A spectacular property of the supergravity total action in this class of models is its local supersymmetry, whereas the kinetic terms are canonical and the potential is that of global SUSY. The total action can be split into the action of pure supergravity part and superconformal matter part. For example, the total scalar-gravity part of the supergravity action is

$$\hat{\mathcal{L}}_{\text{sg}}^0 + \hat{\mathcal{L}}_{\text{sc}}^{\text{m}} = \sqrt{-g_J} \left[ \frac{1}{2} M_P^2 (R + 6 A_\mu A^\mu) - \frac{1}{6} |z^\alpha|^2 R - \delta_{\alpha\bar{\beta}} g^{\mu\nu} D_\mu z^\alpha D_\nu \bar{z}^{\bar{\beta}} - \hat{V}_J \right], \quad (3.25)$$

where

$$\hat{V}_J = \delta^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}} + \frac{1}{2} (\text{Re } f)^{-1 AB} \mathcal{P}_A \mathcal{P}_B. \quad (3.26)$$

It consists of a pure supergravity part to which one has to add the global supersymmetry action interacting with supergravity multiplet. If we would like to cut out the supergravity multiplet and get the global SUSY action, we would have to remove from the action the first 3 terms, remove  $\sqrt{-g}$ , replace the curved metric by the flat one and remove the  $U(1)$  field  $A_\mu$  from the covariant derivative. We would get:

$$\hat{\mathcal{L}}_{\text{susy}}^{\text{m}} = -\delta_{\alpha\bar{\beta}} \eta^{\mu\nu} \partial_\mu z^\alpha \partial_\nu \bar{z}^{\bar{\beta}} - V_{\text{susy}}. \quad (3.27)$$

The potential  $V_{\text{susy}}$  is precisely the same as in supergravity and given in Eq. (3.26).

Vice versa, if one would like to promote any scale-free global SUSY theory into supergravity, one starts with the action (3.27) with the potential in Eq. (3.26). First, one has to add a factor  $\sqrt{-g}$  to the global action and replace the flat space-time metric  $\eta^{\mu\nu}$  in the kinetic term by the curved metric  $g^{\mu\nu}$ . The partial derivatives of the scalars have to be made  $\mathcal{R}$ -covariant as in Eq. (3.21) to make the chiral multiplets superconformal:

$$\hat{\mathcal{L}}_{\text{susy}}^{\text{m}} = -\delta_{\alpha\bar{\beta}} \eta^{\mu\nu} \partial_\mu z^\alpha \partial_\nu \bar{z}^{\bar{\beta}} - V_{\text{susy}} \quad \Rightarrow \quad \hat{\mathcal{L}}_{\text{sc}}^{\text{m}} = \sqrt{-g_J} \left( -\frac{1}{6} |z^\alpha|^2 R - \delta_{\alpha\bar{\beta}} g^{\mu\nu} D_\mu z^\alpha D_\nu \bar{z}^{\bar{\beta}} - \hat{V}_J \right), \quad (3.28)$$

where

$$\hat{V}_J = V_{\text{susy}} = \delta^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}} + \frac{1}{2} (\text{Re } f)^{-1 AB} \mathcal{P}_A \mathcal{P}_B, \quad D_\mu z^\alpha = \partial_\mu z^\alpha - i A_\mu z^\alpha. \quad (3.29)$$

It remains to add an action of pure supergravity  $\sqrt{-g_J} \left[ \frac{1}{2} M_P^2 (R + 6 A_\mu A^\mu) \right]$ . The result is in Eq. (3.25). The same for vector multiplets. If we include also fermions, the rules require also to introduce the interaction with gravitino, as shown in Appendix A. The resulting action of the form (3.25), (3.26) has a local super-Poincaré symmetry and the matter action has a superconformal symmetry. When



fermions and vectors are included, the generalization of eqs. (3.25), (3.26) is given in Appendix A. The principle is the same, all chiral and vector multiplets start interacting with the gravitational Weyl supermultiplet, and the pure supergravity action is added. The total action, including fermions and vectors, has local Poincaré supersymmetry.

If we would like to embed the scale-free global SUSY theory into supergravity in the Einstein frame, we would have to use the Kähler potential

$$\hat{\mathcal{K}}(z, \bar{z}) = -3M_P^2 \log \left( -\frac{1}{3M_P^2} \hat{\Phi}(z, \bar{z}) \right) = -3M_P^2 \log \left( 1 - \frac{1}{3M_P^2} \delta_{\alpha\bar{\alpha}} z^\alpha \bar{z}^{\bar{\alpha}} \right). \quad (3.30)$$

The total scalar-gravity part of action will be

$$\hat{\mathcal{L}}_{\text{supergrav}}^E = \sqrt{-g_E} \left( \frac{1}{2} M_P^2 R - \hat{\mathcal{K}}_{\alpha\bar{\beta}} g_E^{\mu\nu} \partial_\mu z^\alpha \partial_\nu \bar{z}^{\bar{\beta}} - V_E \right), \quad (3.31)$$

where  $V_E$  is

$$V_E = e^{\frac{\kappa}{M_P^2}} \left( \nabla_\alpha W g^{\alpha\bar{\beta}} \nabla_{\bar{\beta}} \bar{W} - \frac{3W\bar{W}}{M_P^2} \right) + \frac{1}{2} (\text{Re } f)^{-1 AB} P_A P_B, \quad (3.32)$$

and  $P_A = -3 \frac{\mathcal{P}_A}{\Phi(z, \bar{z})}$ . The action in the Einstein frame is significantly different from the global SUSY action, the kinetic terms are not canonical, the F-term potential is complicated and no part of this action has a conformal or  $\mathcal{R}$ -symmetry. The dependence on  $M_P$  is all over the place. Thus, for scale-free globally supersymmetric models there is an obvious advantage to study their supergravity embedding in the simple Jordan frame with manifest superconformal symmetry of the matter action as shown in eqs. (3.25). Note also that for this class of models the potential in the Einstein frame can be given in the form

$$V_E = 9 \frac{\hat{V}_J}{\Phi^2} = \frac{\delta^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}} + \frac{1}{2} (\text{Re } f)^{-1 AB} \mathcal{P}_A \mathcal{P}_B}{(1 - \delta_{\gamma\bar{\delta}} z^\gamma \bar{z}^{\bar{\delta}} / 3M_P^2)^2}, \quad (3.33)$$

and it is positive semi-definite.

It is instructive to compare the CSS class of models with no-scale supergravity. The review of no-scale supergravity models can be found in Sec. XII in [23]. No-scale models have a positive semi-definite potential in the Einstein frame. This condition is also satisfied by the CSS models. However, the second feature of no-scale models is that at the minimum  $V_{\text{min}} = 0$  they break supersymmetry spontaneously. Meanwhile in the CSS models, the minimum of the potential is at  $V = 0$ , but supersymmetry is not broken there. Therefore such theories may provide a natural starting point for investigation of the models with a low scale of SUSY breaking.

An interesting property of the no-scale supergravity is that the term  $-3|W|^2$  is absent in the expression for the scalar potential. As we already demonstrated, the CSS models share this property,

but, in addition, the expression for the scalar potential in the Jordan frame does not have the overall factor  $e^{\mathcal{K}}$ , the Kähler connection terms,  $K_\alpha W$ , drop and  $K^{\alpha\bar{\beta}}$  is replaced by  $\delta^{\alpha\bar{\beta}}$ . This is a major simplification, reducing the F-term potential to its global SUSY expression. We found this result directly from the superconformal approach to supergravity, but one can also confirm it by direct calculations presented in Appendix B.

The important property of both CSS and no-scale supergravity models is that in order to describe physics, we have to break some of the symmetries of these models. In case of no-scale model an important example is the KKLT stabilization of the string theory Kähler moduli [24] where the breaking of the no-scale property of supergravity is achieved via the instanton corrections/gaugino condensation. In case of superconformal matter coupling we will introduce in the next section a mechanism of breaking of superconformal symmetry which is useful for inflation as well as for a possible solution of the  $\mu$ -problem in the NMSSM. This mechanism is geometric: the moduli space of chiral fields including the compensator field is not flat anymore, but no dimensional parameters are introduced.

### 3.5 Breaking of superconformal symmetry via $\chi$ -terms: the real part of the holomorphic function in scalar-curvature coupling

An interesting possibility to break the superconformal symmetry of the matter multiplets in the supergravity action without introducing dimensional parameters into the underlying superconformal action (3.1) is to modify the real function  $\mathcal{N}(X, \bar{X})$  as follows

$$\mathcal{N}(X, \bar{X}) = -|X^0|^2 + |X^\alpha|^2 - \chi \left( a_{\alpha\beta} \frac{X^\alpha X^\beta \bar{X}^{\bar{0}}}{X^0} + h.c. \right). \quad (3.34)$$

Here  $\chi$  is a dimensionless parameter and  $a_{\alpha\beta}$  is a numerical matrix. The function  $\mathcal{N}(X, \bar{X})$  has the correct dilatation weight in each  $X$  and  $\bar{X}$  direction. This means that the new Kähler manifold for all  $n + 1$  chiral multiplets  $X^I$ , including the compensator field  $X^0$ , is not flat anymore. The metric  $G_{I\bar{J}} = \frac{\partial \mathcal{N}(X, \bar{X})}{\partial X^I \partial \bar{X}^{\bar{J}}}$  is not flat and the curvature  $R_{I\bar{K}J\bar{L}}$  is proportional to  $\chi$ . We keep a cubic,  $X^0$  independent superpotential and a flat vector moduli space and an  $X^0, \bar{X}^{\bar{0}}$  independent momentum map  $\mathcal{P}_A$ , as above. The gauge-fixing of this class of models with  $\mathcal{N}(X, \bar{X}) = \Phi(z, \bar{z})$  and  $X^0 = \bar{X}^{\bar{0}} = \sqrt{3}M_P$  leads to a Jordan frame supergravity, described in general case in [1]. The resulting supergravity action in which the matter multiplet is not superconformal due to  $\chi$  terms is given by

$$\frac{1}{\sqrt{-g_J}} \mathcal{L}_{\text{sg}}^J = \frac{1}{2} M_P^2 (R + 6 A_\mu A^\mu) - \frac{1}{6} \left( |z^\alpha|^2 - \chi (a_{\alpha\beta} z^\alpha z^\beta + h.c.) \right) R - \delta_{\alpha\bar{\beta}} D_\mu z^\alpha D^\mu \bar{z}^{\bar{\beta}} - V_J, \quad (3.35)$$

where

$$V_J = G^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}} + \frac{1}{2} (\text{Re } f)^{-1 AB} \mathcal{P}_A \mathcal{P}_B . \quad (3.36)$$

Here  $G^{\alpha\bar{\beta}}$  is the matter part of the inverse metric  $G^{IJ}$  of the enlarged space including the compensator. We compute it in Appendix C. The action corresponds to a Jordan frame supergravity with the frame function given in Eq. (2.9) where the holomorphic function is  $J(z) = -\chi a_{\alpha\beta} z^\alpha z^\beta$ . Note that the inverse metric  $G^{\alpha\bar{\beta}}$  in the potential (3.36) is not flat anymore, it depends on moduli. However, the kinetic term for scalars is canonical since  $\Phi_{\alpha\bar{\beta}} = \delta_{\alpha\bar{\beta}}$ . An additional simplification is also observed in the potential (3.36) : it has the form rather close to the global supersymmetry potential. The difference comes from the non-flat inverse metric  $G^{\alpha\bar{\beta}}$ . In particular, *certain directions may still keep a flat metric and the corresponding part of the potential remains superconformal*. As we will see later, this property is useful for the studies of inflation in the Jordan frame.

### 3.6 A simple example

A simple example is the case of two scalars, the field  $S$ , which is not included in the  $\chi$ -term, and the field  $H$ , which is included in the  $\chi$ -term. We start with the superconformal theory (3.34)

$$\mathcal{N}(X, \bar{X}) = -|X^0|^2 + |S|^2 + |H|^2 - \frac{3}{4} \chi \left( \frac{H^2 \bar{X}^0}{X^0} + h.c. \right). \quad (3.37)$$

After gauge-fixing  $X^0 = \bar{X}^0 = \sqrt{3} M_P$  we find a Jordan frame supergravity with the frame function

$$\begin{aligned} \Phi &= -3M_P^2 + |S|^2 + |H|^2 - \frac{3}{4} \chi (H^2 + \bar{H}^2) \\ &= -3M_P^2 + |S|^2 - \frac{1}{4} \left( 1 + \frac{3}{2} \chi \right) (H - \bar{H})^2 + \frac{1}{4} \left( 1 - \frac{3}{2} \chi \right) (H + \bar{H})^2. \end{aligned} \quad (3.38)$$

The action has the following curvature dependent terms:

$$\left[ \frac{1}{2} M_P^2 - \frac{1}{6} |S|^2 - \frac{1}{6} |H|^2 + \frac{1}{8} \chi (H^2 + \bar{H}^2) \right] R. \quad (3.39)$$

If  $S = 0$  and the field  $H$  is real, so that  $H = \bar{H} = \frac{h}{\sqrt{2}}$ , we find the following action:

$$\frac{1}{2} \left[ M_P^2 + \left( -\frac{1}{6} + \frac{1}{4} \chi \right) h^2 \right] R. \quad (3.40)$$

This will explain a particular relation between the standard model action [2] and the NMSSM action during inflation, as well as the relation  $\xi = -\frac{1}{6} + \frac{1}{4} \chi$ , see (7.18) in Sect. 7.1.

We may also rewrite the curvature-dependent terms of the action (3.39) in the following form:

$$\left[ \frac{1}{2} M_P^2 - \frac{1}{6} \left( |S|^2 - \frac{1}{4} \left( 1 + \frac{3}{2} \chi \right) (H - \bar{H})^2 + \frac{1}{4} \left( 1 - \frac{3}{2} \chi \right) (H + \bar{H})^2 \right) \right] R. \quad (3.41)$$

### 3.7 Shift symmetric models

In the main part of the paper we will be interested in the regime where  $\chi \gg 1$ . However, there are two other special cases which may be equally interesting,  $\chi = \pm \frac{2}{3}$ :

1)  $\chi = -\frac{2}{3}$ , in which case the frame function is given by

$$\Phi_{\chi=-2/3} = -3M_P^2 + |S|^2 + \frac{1}{2}(H + \bar{H})^2. \quad (3.42)$$

In this case the field  $S$  remains conformally coupled, but the imaginary part of the field  $H$ , which is given by  $(H - \bar{H})/2i$ , decouples from the curvature scalar in (3.41), i.e. this field becomes minimally coupled.

2)  $\chi = \frac{2}{3}$ ,

$$\Phi_{\chi=2/3} = -3M_P^2 + |S|^2 - \frac{1}{2}(H - \bar{H})^2. \quad (3.43)$$

In this case  $S$  remains conformally coupled, but the real part of the field  $H$ , which is given by  $(H + \bar{H})/2$ , decouples from the curvature scalar in (3.41), i.e. it becomes minimally coupled.

Consider a particular class of superconformal models with the superconformal symmetry broken by the real part of the holomorphic function, as shown in (3.34). For superconformal models we are interested in the relation between the frame function and the Kähler potential of the form

$$\mathcal{K} = -3M_P^2 \log \left( -\frac{1}{3M_P^2} \Phi(z, \bar{z}) \right). \quad (3.44)$$

If we break the superconformal symmetry of matter preserving this relation between the frame function and the Kähler potential, we are led to a class of models where the Kähler potential has a shift symmetry:

1)  $\chi = -\frac{2}{3}$ ,

$$\mathcal{K}(z, \bar{z})_{\chi=-2/3} = -3M_P^2 \log \left( 1 - \frac{1}{3M_P^2} \left( |S|^2 + \frac{1}{2}(H + \bar{H})^2 \right) \right). \quad (3.45)$$

This Kähler potential has a shift symmetry with respect to  $H - \bar{H}$ .

2)  $\chi = \frac{2}{3}$ ,

$$\mathcal{K}(z, \bar{z})_{\chi=2/3} = -3M_P^2 \log \left( 1 - \frac{1}{3M_P^2} \left( |S|^2 - \frac{1}{2}(H - \bar{H})^2 \right) \right). \quad (3.46)$$

This Kähler potential has a shift symmetry with respect to  $H + \bar{H}$ .

Thus, a new class of models with the shift symmetric Kähler potential was derived here from the superconformal approach to supergravity. These models provide a natural basis for a broad class of

new models of chaotic inflation in supergravity, with a functional freedom of choice of the inflaton potential [21].

### 3.8 Stabilization of moduli at the origin of the moduli space

We may be interested for applications in a method of breaking superconformal symmetry which enforces some scalars to be fixed at the origin of the moduli space. The method to achieve the moduli stabilization at the origin of the moduli space due to quartic corrections to the Kähler potential,  $\mathcal{K}(S, \bar{S}) = S\bar{S} - \frac{(S\bar{S})^2}{\Lambda^2}$ , was studied in [25]. It was argued there that the quartic term originates from the loop corrections representing the effective potential from the massive fields which have been integrated out. The sign of the second term in the Kähler potential is negative. In such case the supergravity potential was shown in [25] to stabilize at  $S = 0$ , at the origin of moduli space. Such quartic terms may be generated by radiative corrections, or they may even be present in the Kähler potential from the very beginning.

Here we will show how these terms may emerge from superconformal coupling of matter if one introduces an additional coupling of matter to the compensator. We split the  $n + 1$  scalars into a group where  $X^0$  is the compensator field,  $X^a = \{X^1, \dots, X^{n-1}\}$  are matter scalars and  $X^n$  is the field which we would like to stabilize at the minimum of the potential at  $X^n = 0$ . We consider a superconformal theory where the  $X^n$  direction is not present in the  $a_{\alpha\beta}$ -matrix:

$$\mathcal{N}(X, \bar{X}) = -|X^0|^2 + |X^a|^2 + |X^n|^2 - \chi \left( a_{ab} \frac{X^a X^b \bar{X}^{\bar{0}}}{X^0} + h.c. \right) - 3\zeta \frac{|X^n \bar{X}^{\bar{n}}|^2}{X^0 \bar{X}^{\bar{0}}}. \quad (3.47)$$

After gauge-fixing at  $X^0 = \bar{X}^{\bar{0}} = \sqrt{3} M_P$  and using notation  $X^a = z^a$ ,  $X^n = S$ , we find the frame function

$$\Phi(z^a, \bar{z}^{\bar{b}}; S, \bar{S}) = -3e^{-\frac{1}{3}\mathcal{K}(z^a, \bar{z}^{\bar{b}}; S, \bar{S})} = -3M_P^2 + |z^a|^2 + |S|^2 - \chi \left( a_{ab} z^a \bar{z}^{\bar{b}} + h.c. \right) - \zeta \frac{|S\bar{S}|^2}{M_P^2}. \quad (3.48)$$

As we will find later, in agreement with the proposal in [4] and previous work in [25], the term  $-\zeta \frac{|S\bar{S}|^2}{M_P^2}$  will allow us to stabilize the inflationary trajectory in the NMSSM at  $S = 0$ . An interesting feature of this mechanism is that the term  $-\zeta \frac{|S\bar{S}|^2}{M_P^2}$  vanishes on the inflationary trajectory when the moduli stabilization is achieved.

## 4 Supergravity embedding of the scale-free NMSSM

### 4.1 Superconformal embedding of the NMSSM into supergravity

The original motivation for the NMSSM model which, in addition to two charged Higgs doublets  $H_u, H_d$ , has a gauge singlet Higgs field  $S$ , was the hope for the elegant solution of the  $\mu$ -problem. In MSSM there is a problem to explain a small value of the  $\mu$ -term in the quadratic part of the superpotential  $W = \mu H_u \cdot H_d$ . This term is required for the phenomenological reasons. In presence of the gauge singlet  $S$  one can start with the cubic superpotential  $\lambda S H_u \cdot H_d$  and hope to find a way to produce small vev of  $S$  so that  $\mu_{\text{eff}} = \lambda \langle S \rangle$  will produce the effective desired value of  $\mu$ -term.

From the superconformal approach we have a totally different motivation for the gauge singlet field. The scale-free NMSSM has cubic potential. Without the gauge singlet the term  $\lambda S H_u \cdot H_d$  would not be possible. So from our perspective the motivation for the gauge singlet  $S$  is the requirement of a scale invariance of a globally supersymmetric theory, which permits a simple promotion to local supersymmetry with the superconformal matter-supergravity coupling.

We start with the scale-free NMSSM model reviewed most recently in [14]. The Higgs field sector of the NMSSM gauge theory has one gauge singlet and two gauge doublet chiral superfields  $z_H = \{S, H_u, H_d\}$ .

$$S, \quad H_u = \begin{pmatrix} H_u^+ \\ H_u^0 \end{pmatrix}, \quad H_d = \begin{pmatrix} H_d^0 \\ H_d^- \end{pmatrix}. \quad (4.1)$$

and  $H_u \cdot H_d \equiv -H_u^0 H_d^0 + H_u^+ H_d^-$ . The Higgs part of the model depends on five chiral superfields. The superpotential is

$$W_{\text{Higgs}} = -\lambda S H_u \cdot H_d + \frac{\rho}{3} S^3. \quad (4.2)$$

The quarks and leptons  $z_{QL} = \{Q, U_R, D_R, L, E_R\}$  are introduced via Yukawa cubic superpotential  $W_{\text{Yukawa}}$  so that the total superpotential for all superfields  $z^\alpha = (z_H, z_{QL})$  is cubic

$$W_{\text{total}} = W_{\text{Yukawa}} + W_{\text{Higgs}} = \frac{1}{3} d_{\alpha\beta\gamma} z^\alpha z^\beta z^\gamma. \quad (4.3)$$

The D- and F-term potentials of the general form (3.26) for the NMSSM are given explicitly in eqs. (9) and (10) of [26] where also the complete set of Feynman rules is presented. All kinetic terms are canonical, both for chiral as well as vector superfields. Also the Yukawa and vector part of the action as well as interaction between the chiral and vector multiplets are given explicitly. We do not add

the soft breaking terms to the NMSSM at this point since we would like first to embed the globally supersymmetric action into supergravity.

We have shown above that the scale-invariant version of the NMSSM has all conditions satisfied so that the simplest possible embedding of the scale invariant version of it into supergravity is possible. One should take the globally supersymmetric action of the form (3.27) with details in [26] and follow the rules explained around eqs. (3.28),(3.29). This gives the promotion to supergravity of the scale-invariant globally supersymmetric NMSSM.

The full supergravity action corresponds to the choice of the frame function

$$\hat{\Phi}(z, \bar{z}) = -3M_P^2 + (S\bar{S} + H_u H_u^\dagger + H_d H_d^\dagger), \quad (4.4)$$

which corresponds to the underlying superconformal theory (3.3) with an extra compensator field  $X^0$

$$\mathcal{N}(X, \bar{X}) = -X^0 \bar{X}^0 + (S\bar{S} + H_u H_u^\dagger + H_d H_d^\dagger). \quad (4.5)$$

The supergravity potential in the Jordan frame is the same as the global one given in eqs. (9) and (10) of [26]. For example, the Higgs-gravity part of the supergravity action consists of the supergravity part, given in Eq. (3.19) and the matter part of supergravity action, which is superconformal, when interacting with the Weyl multiplet:

$$\hat{\mathcal{L}}_{\text{sc}}^{\text{m}} = \sqrt{-g_J} \left[ -\frac{R}{6} \left( S\bar{S} + H_u H_u^\dagger + H_d H_d^\dagger \right) - D_\mu H^u D^\mu H_u^\dagger - D_\mu H^d D^\mu H_d^\dagger - D_\mu S D^\mu S^\dagger - \hat{V}_J \right], \quad (4.6)$$

where

$$\hat{V}_J = \left| \frac{\partial W}{\partial S} \right|^2 + \left| \frac{\partial W}{\partial H_u} \right|^2 + \left| \frac{\partial W}{\partial H_d} \right|^2 + \frac{g'^2}{8} (|H_u^0|^2 - |H_d^0|^2)^2 + \frac{g^2}{8} (H_u^\dagger \vec{\tau} H_u + H_d^\dagger \vec{\tau} H_d)^2. \quad (4.7)$$

Here  $D_\mu$  acting on scalars includes the gauge field of the  $U(1)$   $\mathcal{R}$  symmetry  $A_\mu$ , which is an auxiliary field of supergravity, see Eq. (3.21). It can be replaced by its on shell value as the function of scalars and fermions, see e.g. Eq. (5.13) of [1]. For example, its bosonic part  $\mathcal{A}_\mu(z, \bar{z})$  is given in Eq. (2.3).

## 4.2 Breaking superconformal symmetry of matter in the NMSSM supergravity

Here we consider a possibility to break the superconformal symmetry of the matter multiplets in supergravity action geometrically, without introducing dimensional parameters into the underlying superconformal action. One of the possibilities was studied in Sec. 3.5. It corresponds to the choice of the frame function

$$\Phi_\chi(z, \bar{z}) = -3M_P^2 + |S|^2 + |H_u|^2 + |H_d|^2 + \frac{3}{2}\chi(H_u \cdot H_d + \text{h.c.}). \quad (4.8)$$

This choice was proposed in [3] for the purpose of the Higgs-type inflation in the NMSSM. The underlying superconformal action is defined by the function  $\mathcal{N}_\chi(X, \bar{X})$  which is homogeneous of first degree in both  $X$  and  $\bar{X}$ :

$$\mathcal{N}(X, \bar{X})_\chi = -|X^0|^2 + |S|^2 + |H_u|^2 + |H_d|^2 + \frac{3}{2}\chi \left( H_u \cdot H_d \frac{\bar{X}^0}{X^0} + h.c. \right). \quad (4.9)$$

Note that there is no superconformal symmetry breaking in  $S$  direction of the moduli space, namely, the metric  $G_{S\bar{S}} = 1 = G^{S\bar{S}}$  remains flat, decoupled from the compensator sector and from the  $H_u$  and  $H_d$  sectors. Meanwhile, the moduli space of the Higgs doublets,  $H_u$  and  $H_d$ , is mixed with the compensator field  $X^0, \bar{X}^0$ , and it is non-flat, with  $\chi$ -dependent curvature.

The Kähler function of the enlarged space (4.9) after the gauge-fixing with  $X^0 = \bar{X}^0 = \sqrt{3}M_P$  corresponds to the frame function (4.8). The bosonic part of the supergravity action is as before  $\sqrt{-g_J} [\frac{1}{2}M_P^2(R + 6A_\mu A^\mu)]$  and the matter part of supergravity action, which is superconformal (up to terms with  $\chi$ ), is

$$\sqrt{-g_J} \left[ -\frac{1}{6} (|S|^2 + |H_u|^2 + |H_d|^2) - \frac{1}{4}\chi(H_u \cdot H_d + h.c.) \right] R - |D_\mu H_u|^2 - |D_\mu H_d|^2 - |D_\mu S|^2 - V_J \quad (4.10)$$

In this case, as shown in Sec. 3.5, the D-term potential in the Jordan frame is the same as in the case  $\chi = 0$ ; however, the  $F$ -term potential, as given by (3.36) has a specific deviation from the quartic superconformal potential, since the metric  $G^{\alpha\bar{\beta}}$  is not flat at  $\chi \neq 0$ :

$$V_J = G^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}} + \frac{g'^2}{8} (|H_u^0|^2 - |H_d^0|^2)^2 + \frac{g^2}{8} (H_u^\dagger \bar{\tau} H_u + H_d^\dagger \bar{\tau} H_d)^2. \quad (4.11)$$

The metric  $G^{\alpha\bar{\beta}}$  is the part of the inverse  $G^{I\bar{J}}$  to the  $G_{I\bar{J}} = \frac{\partial \mathcal{N}(X, \bar{X})}{\partial X^I \partial \bar{X}^{\bar{J}}}$  metric. It is easy to compute using Eq. (4.9). One may notice, using  $W_{\text{Higgs}} = -\lambda S H_u \cdot H_d + \frac{\rho}{3} S^3$  that at  $S = 0$  the only contribution to the F-term potential comes from the term

$$(V_J^F)|_{S=0} = \frac{\partial W}{\partial S} G^{S\bar{S}} \frac{\partial \bar{W}}{\partial \bar{S}} = \lambda^2 G^{S\bar{S}} |H_u \cdot H_d|^2. \quad (4.12)$$

Since the field  $S$  does not enter in the  $\chi$ -term, one finds that  $G^{S\bar{S}} = 1$  and therefore even after this breaking of superconformal symmetry the specific part of the potential remains quartic. This plays an important role for inflation where the inflationary trajectory is at  $S = 0$ .

To embed the NMSSM gauge theory into the Einstein frame supergravity with the superconformal symmetry breaking explained above, we have to use the Kähler potential

$$\mathcal{K}_\chi(z, \bar{z}) = -3 \log \left[ 1 - \frac{1}{3} \left( S\bar{S} + H_u H_u^\dagger + H_d H_d^\dagger \right) - \frac{1}{2}\chi(H_u \cdot H_d + h.c.) \right]. \quad (4.13)$$



## 5 Phenomenological aspects of the NMSSM

Here we start with the current point of view on the NMSSM, and its problems, following [14], where the globally supersymmetric model is studied in presence of the terms breaking supersymmetry softly, which originate from a hidden sector of the theory. We afterwards discuss the issues of the NMSSM from the superconformal symmetry approach that we find useful both for the particle physics phenomenology as well as for cosmology.

One of the reasons to augment the MSSM by the gauge singlet Higgs field  $S$  and study the NMSSM was that the superpotential of the MSSM contained the term  $\mu H_u \cdot H_d$ . It is difficult to explain the required smallness of this term. In the NMSSM, one may generate the  $\mu$ -term as  $-\lambda \langle S \rangle H_u \cdot H_d$  from the superpotential  $W_\lambda = -\lambda S H_u \cdot H_d$ . The problem, however, is to explain why one cannot add the term  $\mu H_u \cdot H_d$  to the NMSSM. To address this problem, one may assume that the superpotential of the NMSSM must be scale invariant. This requirement forbids term such as  $\mu H_u \cdot H_d$ , as well as the tadpole term  $\sim S$  and the term  $\sim S^2$ , and allows only the cubic superpotential  $W_{\text{Higgs}} = \lambda S H_u \cdot H_d + \frac{\rho}{3} S^3$ .

Scale-invariance of the NMSSM superpotential allows its consistent embedding into the CSS. From the top-down perspective, this scale invariance can be interpreted as a consequence of the original superconformal symmetry, protected by the decoupling of the light fields from the conformal compensator. However, scale-invariance of the of the NMSSM superpotential may result in the cosmological domain problem, which we are going to analyze now.

At low energies one usually considers adding to the global SUSY potential the soft SUSY breaking terms. The soft terms are of two types. There are mass terms for each Higgs,

$$V_{\text{soft}}^m = m_{H_u}^2 |H_u|^2 + m_{H_d}^2 |H_d|^2 + m_S^2 |S|^2. \quad (5.1)$$

There are also terms related to a superpotential contribution to the potential: i.e. there is a coupling  $A_\lambda, A_\rho$  for each cubic term in the superpotential, times the real part of the superpotential. In case of the NMSSM with the cubic superpotential they are

$$V_{\text{soft}}^W(\text{NMSSM}) = A_\lambda \lambda S H_u \cdot H_d + A_\rho \frac{1}{3} \rho S^3 + h.c. \quad (5.2)$$

A continuous global  $\mathcal{R}$ -symmetry of the total potential, when each scalar transforms as  $z' = z^{i\Lambda}, \bar{z} = e^{-i\Lambda} \bar{z}$ , is broken down to a discrete one due to the  $V_{\text{soft}}^W(\text{NMSSM})$  term. Namely, (5.2) is invariant under  $\mathbb{Z}_3$  symmetry:

$$S' = e^{\frac{2\pi i n}{3}} S, \quad H'_u = e^{\frac{2\pi i n}{3}} H_u, \quad H'_d = e^{\frac{2\pi i n}{3}} H_d, \quad (5.3)$$

where  $n \in \mathbb{Z}$  and we assume that  $A_\lambda$  and  $A_\rho$  are real. In such case, the theory has domain walls created once the  $\mathbb{Z}_3$  symmetry is spontaneously broken after a restoration of a symmetric phase in the hot early universe. This creates large anisotropies of the CMB and contradicts a successful nucleosynthesis.

An interesting role is played here by the local  $U(1)$   $\mathcal{R}$ -symmetry, which is part of the superconformal  $SU(2, 2|1)$  symmetry (and it is not included into the super-Poincaré symmetry). As explained in [1], the  $\chi$ -term required for inflation in the NMSSM must be a sum of holomorphic and anti-holomorphic terms to keep the Jordan frame kinetic terms canonical,  $\Phi(z, \bar{z}) = -3 M_P^2 + \delta_{\alpha\bar{\beta}} z^\alpha \bar{z}^{\bar{\beta}} + J(z) + \bar{J}(\bar{z})$ . These  $J(z) + \bar{J}(\bar{z})$  terms in the frame function and in the Kähler potential not only break the continuous  $\mathcal{R}$ -symmetry, but also break the discrete  $\mathbb{Z}_3$  symmetry (5.3). A study of the  $\mathbb{Z}_3$  symmetry breaking terms in the supergravity Einstein frame potential shows that the symmetry breaking term is an order six operator  $\sim \chi \frac{\lambda^2 h^6}{M_P^2}$ . According to [27], this amount of  $\mathbb{Z}_3$  symmetry breaking may not be sufficient to make the domain walls disappear before the nucleosynthesis. However, we have to take into account that the  $J(z) + \bar{J}(\bar{z})$  terms in the Kähler potential may change the soft breaking SUSY terms in the potential, in presence of a hidden sector [28–30]. This possibility was proposed in [4]. Here we will present a more detailed investigation of this scenario.

Suppose that chiral superfields  $z^\alpha = \{\phi^a, \varphi^i\}$  are split into an observable sector  $\phi^a$  and the hidden sector  $\varphi^i$ . Whereas the observable fields have weak scale vev's  $\sim 10^{-16} M_P$ , the hidden sector scalars have a much larger scale, but they are still much smaller than  $M_P$ . Therefore one may expect that at present  $W = W_{\text{obs}} + W_{\text{hid}} \approx W_{\text{hid}}$ , and  $e^{\frac{\kappa}{2M_P^2}} \approx 1$ . In what follows, we will assume that  $W_{\text{obs}}$  is cubic in  $\phi^a$ , but we will not specify the superpotential of the hidden sector. Up to an irrelevant complex phase, the gravitino mass is given by

$$m_{3/2} = e^{\frac{\kappa}{2M_P^2}} \frac{\langle W \rangle}{M_P^2} \approx \frac{\langle W_{\text{hid}} \rangle}{M_P^2}. \quad (5.4)$$

In what follows, the discussion will proceed in the Jordan frame supergravity since it makes the conceptual points very clear. We will write the  $\chi$ -term, which we added to the Kähler potential, as the real part of the holomorphic function  $J(\phi)$ , quadratic in fields from the observable sector,  $J(\phi) = -\chi C_{ab} \phi^a \phi^b$ . This allows us to keep the Jordan frame kinetic terms for the observable sector canonical and to have only a dimensionless superconformal symmetry breaking parameter  $\chi$ :

$$\mathcal{K}(z, \bar{z}) = -3M_P^2 \log \left[ 1 - \frac{\phi^a \bar{\phi}_a}{3M_P^2} - \frac{J(\phi)}{3M_P^2} - \frac{\bar{J}(\bar{\phi})}{3M_P^2} - \dots \right]. \quad (5.5)$$

Here  $\dots$  stands for the terms depending on the hidden sector superfields. For the values of the fields

much smaller than the Planck scale we may expand the logarithm in Eq. (5.5):

$$\mathcal{K}(z, \bar{z}) = \phi^a \bar{\phi}_a + J(\phi) + \bar{J}(\bar{\phi}) + \dots \quad (5.6)$$

Now we can use the Kähler invariance to switch to a different Kähler potential and superpotential,

$$\mathcal{K}_{\text{eff}}(z, \bar{z}) = \mathcal{K}(z, \bar{z}) - J(\phi) - \bar{J}(\bar{\phi}) + \dots, \quad W_{\text{eff}} = W e^{\frac{J(\phi)}{M_P^2}}. \quad (5.7)$$

The new Kähler potential is canonical, but the superpotential has a correction,

$$\mathcal{K}_{\text{eff}}(\phi, \bar{\phi}) = \phi^a \bar{\phi}_a, \quad W_{\text{eff}} \rightarrow W e^{\frac{J(\phi)}{M_P^2}} \approx W + \frac{\langle W_{\text{hid}} \rangle}{M_P^2} J(\phi) \approx W + m_{3/2} J(\phi). \quad (5.8)$$

Here we took into account (5.4). In the specific case of the NMSSM, where  $J = \frac{3}{2} \chi H_u \cdot H_d$ , one finds [4]

$$W_{\text{eff}} = -\lambda S H_u \cdot H_d + \frac{\rho}{3} S^3 + \frac{3}{2} \chi m_{3/2} H_u \cdot H_d. \quad (5.9)$$

Thus, the mere existence of the real part of the holomorphic quadratic correction to the frame function for observable Higgs fields, breaking the superconformal symmetry in a way required for inflation, is responsible also for the specific contribution  $\frac{3}{2} \chi m_{3/2} H_u \cdot H_d$  to the  $\mu$ -term in the effective superpotential for small fields,

$$\mu_{\text{eff}} = \frac{3}{2} \chi m_{3/2} - \lambda \langle S \rangle. \quad (5.10)$$

This is a specific realization of the Giudice-Masiero mechanism [29]. Note that the term  $\frac{3}{2} \chi m_{3/2} H_u \cdot H_d$  breaks  $\mathbb{Z}_3$  symmetry of the real part of the scale-invariant superpotential. To evaluate the significance of this effect, one may estimate the correction to the soft breaking part of the potential originating from the term  $\frac{3}{2} \chi m_{3/2} H_u \cdot H_d$ :

$$V_{\text{soft}}^{W_{\text{eff}}} = A_\lambda \lambda S H_u \cdot H_d + A_\rho \frac{\rho}{3} S^3 + B_\mu \mu_{\text{eff}} H_u \cdot H_d + h.c.. \quad (5.11)$$

This term contains the  $\mathbb{Z}_3$ -noninvariant term

$$\Delta V = \frac{3}{2} B_\mu \chi m_{3/2} (H_u \cdot H_d + h.c.). \quad (5.12)$$

According to [27],  $\mathbb{Z}_3$  symmetry does not lead to the cosmological domain wall problem if the difference in vacuum energy between the different vacua separated by the domain walls is greater than  $10^{-7} \frac{v}{M_P} v^4 \sim 10^{-25} v^4$ . We may now compare the potential energy for two vacua, which are degenerate for  $\chi = 0$ . Consider  $\frac{z}{|z|} = e^{\frac{2\pi i}{3} n}$  and take one vacuum with  $n = 0$  and another one with  $n = 1$ . For  $B_\mu \sim \chi m_{3/2} \sim v$ , the energy difference is  $\sim \frac{3}{2} B_\mu \chi m_{3/2} v^2 \sim v^4$ , which is many orders of magnitude greater than the energy separation required for the absence of domain walls.

One may wonder, whether all of these nice properties will be spoiled by the tadpole problem? Indeed, in generic models interactions with heavy particles may induce large terms linear in  $S$  in the superpotential, see e.g. [27, 31–33]. Fortunately, this problem can be solved under certain conditions, as explained in [14]. In particular, in the theories with  $R$  symmetry [34] a solution to the tadpole problem was suggested. We believe that this solution applies to our model. Some other proposals how to stabilize the singlets in supergravity and avoid domain walls can be found in [35].

Let us summarize our approach to the NMSSM phenomenology.

1) There are several different versions of the NMSSM, and many inequivalent ways to incorporate each of these versions into supergravity. We propose to incorporate the NMSSM into a canonical superconformal supergravity (CSS). This singles out the scale-invariant version of the NMSSM. In general, the embedding of a global SUSY model can be quite complicated, but the embedding of the NMSSM into the CSS is a trivial exercise in the Jordan frame: one simply replaces usual derivatives by covariant derivatives. The resulting theory has superconformal symmetry, and all kinetic terms are canonical. This is a unique property of the CSS approach, not shared by other methods of embedding of the NMSSM into supergravity.

2) After the embedding, all fields in the NMSSM are massless. Then one introduces masses due to gravitational effects and interaction with hidden sector. This explains the smallness of all masses in the NMSSM as a consequence of the underlying superconformal symmetry.

3) Adding the  $\chi$ -term to the Kähler potential is equivalent to adding a nonminimal coupling of the Higgs field to gravity, which is consistent with our ideology of breaking the superconformal symmetry by gravitational effects. Whereas the  $\chi$ -term was added in order to realize Higgs inflation, it plays an additional role: it leads to a specific realization of the Giudice-Masiero mechanism of generation of the  $\mu$ -term in the NMSSM [29]. This mechanism breaks  $\mathbb{Z}_3$  symmetry and resolves the domain wall problem in the NMSSM, whereas the tadpole problem may be solved due to  $\mathcal{R}$  symmetry of our construction.

## 6 On Higgs-type inflation with nonminimal coupling in standard model

### 6.1 Basic model

Here we review the Higgs-type inflation with nonminimal scalar-curvature  $\xi$ -coupling studied in [2]. We will focus on three different ranges of the Higgs field vev's, at the beginning of the last 60 e-foldings, at the exit from inflation, and at the present values of the SM Higgs. In [2] the SM potential with canonical kinetic term for the Higgs field  $h$  is coupled to a gravitational field in the Jordan frame:

$$\mathcal{L}_J = \sqrt{-g_J} \left[ \frac{M^2 + \xi h^2}{2} R(g_J) - \frac{1}{2} \partial_\mu h \partial_\nu h g_J^{\mu\nu} - \frac{\lambda}{4} (h^2 - v^2)^2 \right]. \quad (6.1)$$

At present,  $h = v \sim 10^{-16} M_P$ , and  $M_P^2 = M^2 + \xi v^2$ . Since  $v$  is extremely small, we will ignore it in our investigation, and take  $M = M_P = 1$ . The frame function for the action (6.1) in this approximation is  $\Phi = -3(1 + \xi h^2)$  and the rescaling of the metric function  $\Omega^2 = 1 + \xi h^2$  and the action can be rewritten as

$$\mathcal{L}_J = \sqrt{-g_J} \left[ \frac{1 + \xi h^2}{2} R(g_J) - \frac{1}{2} \partial_\mu h \partial_\nu h g_J^{\mu\nu} - \frac{\lambda}{4} h^4 \right]. \quad (6.2)$$

In the Einstein frame the action is

$$\mathcal{L}_E = \sqrt{-g_E} \left( \frac{1}{2} R(g_E) - \frac{1}{2} \partial_\mu \psi \partial_\nu \psi g_E^{\mu\nu} - U(\psi) \right), \quad (6.3)$$

where

$$U(\psi) = \frac{\lambda}{4} \left( \frac{h^2(\psi) - v^2}{1 + \xi h(\psi)^2} \right)^2, \quad (6.4)$$

and  $\psi$  is a canonically normalized scalar in the Einstein frame, defined by

$$d\psi \equiv dh \sqrt{\frac{\Omega^2 + 6\xi^2 h^2}{\Omega^4}}. \quad (6.5)$$

A solution of this equation is

$$\psi = \sqrt{1 + 6\xi^{-1}} \text{ArcSinh}(\sqrt{\xi + 6\xi^2} h) - \sqrt{6} \text{ArcTanh} \left( \frac{\sqrt{6}\xi h}{\sqrt{1 + \xi h^2 + 6\xi^2 h^2}} \right). \quad (6.6)$$

It is useful to present this solution in a simpler, asymptotic form for three different ranges of  $h$ .

1) In the interval  $0 < h \ll \frac{1}{\xi}$  one has

$$\psi \approx h, \quad U(\psi) \approx \frac{\lambda}{4} \psi^4. \quad (6.7)$$

2) In the interval  $\frac{1}{\xi} \ll h \ll \frac{1}{\sqrt{\xi}}$  one has

$$\psi \approx \sqrt{\frac{3}{2}} \xi h^2, \quad U(\psi) \approx \frac{\lambda}{6\xi^2} \left( \frac{\psi}{1 + \sqrt{\frac{2}{3}}\psi} \right)^2. \quad (6.8)$$

At the upper part of this interval one has  $\psi = O(1)$ . The existence of this intermediate range was not taken into account in many recent papers on Higgs inflation. It will play an important role in our discussion of the unitarity bound in the next subsection.

3) Finally, for  $h \gg \frac{1}{\sqrt{\xi}}$  (or, equivalently,  $\psi \gg 1$ ) one has

$$\psi \approx \sqrt{\frac{3}{2}} \ln(\xi h^2), \quad U(\psi) \approx \frac{\lambda}{4\xi^2} \left( 1 + e^{-\frac{2\psi}{\sqrt{6}}} \right)^{-2}. \quad (6.9)$$

In this regime, the potential in the Einstein frame is very flat, which leads to inflation. As one can see from (6.9), the constant ( $\psi$ -independent) term in the potential  $U(\psi)$  is  $\frac{\lambda}{4\xi^2}$ , so nothing would work without the nonminimal scalar-curvature coupling proportional to  $\xi$ .

The slow-roll parameters, for  $\xi h^2 \gg 1$ , are,

$$\epsilon \simeq \frac{4}{3\xi^2 h^4}, \quad (6.10)$$

$$\eta \simeq -\frac{4}{3\xi h^2}. \quad (6.11)$$

Slow roll ends when  $\epsilon \simeq 1$ , so the field value at the end of inflation is  $h_{\text{end}} \simeq (4/3)^{1/4}/\sqrt{\xi} \simeq 1/\sqrt{\xi}$ . The number of e-foldings  $N \gg 1$  during the slow roll of the field  $h$  from its initial value  $h_0$  is given by

$$N \simeq \frac{3}{4} \xi h_0^2. \quad (6.12)$$

For a particular case  $N \sim 60$ , the amplitude of scalar perturbations of metric corresponds to the COBE normalization for

$$\frac{\xi}{\sqrt{\lambda}} \simeq 5 \times 10^4. \quad (6.13)$$

The Hubble constant during inflation in this model is  $H \approx \sqrt{\frac{\lambda}{3}} \frac{1}{2\xi}$ .

## 6.2 The unitarity bound?

Recently several authors argued that one cannot rely on the description of various processes in the Higgs inflation model on energy scale exceeding the unitarity bound  $\Lambda \sim 1/\xi$  [15–18]. For the non-supersymmetric standard model described above, with  $\lambda = O(1)$ , this bound is dangerously close to

the Hubble constant during inflation  $H \approx \sqrt{\frac{\lambda}{3}} \frac{1}{2\xi}$ . In the NMSSM one may consider the regime with  $\lambda \ll 1$ , where the concerns about the unitarity bound do not seem to appear [4]. This can be done by using the rescaling of several parameters of the model.

Indeed, one can easily check that all observational consequences of the inflationary model described above, including the value of the potential, the Hubble constant, the slow-roll parameters, the number of e-folds of inflation, the amplitude of scalar perturbations of metric, the spectral index  $n_s$ , and the ratio of tensor perturbations to scalar perturbations  $r$ , depend only on two combinations of parameters:  $\xi h^2$  and  $\frac{\lambda}{\xi^2}$ . Therefore all observational consequences of this model are invariant with respect to the simultaneous rescaling  $\lambda \rightarrow c^2 \lambda$ ,  $\xi \rightarrow c \xi$  and  $h \rightarrow h/\sqrt{c}$ . This means that one can study the inflationary regime for  $\lambda = 1$ ,  $\xi \simeq 5 \times 10^4$ , and then rescale it to smaller values of  $\lambda$  to avoid the problems with the unitarity bound.

It is good to know that we have this possibility. However, whereas it is possible to use small  $\lambda$  in the NMSSM, one cannot do it in the original nonsupersymmetric model of [2]. Therefore it would be interesting to double-check whether one should worry about the unitarity bound in general.

Most of the arguments suggesting the existence of this bound are based on the investigation of the theory in the small field approximation  $\psi \approx h$ , where one can use an expansion  $\psi = h(1 + \xi^2 h^2 + \dots)$ . For example, Ref. [16] considers the potential (6.4) at small values of the field  $\psi$  where the potential can be expanded in powers of  $\psi$  as

$$U(\psi)_{\psi \rightarrow 0} \Rightarrow \frac{\lambda}{4} \psi^4 (1 - 4\xi^2 \psi^2 + O((\xi^2 \psi^2)^2)) + \dots \quad (6.14)$$

One may consider the term  $-\lambda \xi^2 \psi^6$ , take two of the fields  $\psi$ , form a loop and integrate. This will produce a term proportional to  $\lambda \xi^2 \Lambda^2 \psi^4$ , where  $\Lambda$  is a cutoff. Repeating this step for all higher order terms, one may come to a conclusion that quantum corrections to  $\frac{\lambda}{4} \psi^4$  become uncontrollable if  $\Lambda > 1/\xi$ .

However, it was suggested in [36] that “the apparent generation of the new physics is an artifact of considering only two terms of the expansion when all terms are important”. For example, one-loop quantum corrections to the scalar potential involve knowledge of the scalar propagator in an external classical field  $\psi$ , which is equivalent to a resummation of diagrams with arbitrary number of external lines of the scalar field. One-loop corrections to the potential are proportional to  $(U''(\psi))^2 \ln |U''(\psi)|$ . Therefore these corrections during inflation are suppressed by an extra power of  $\frac{\lambda}{\xi^2}$ , as well as by the asymptotic flatness of the potential (6.9). Here we would like to take another look at this issue, and

give an independent argument, which can be applied not only to the scalar potential, but also for kinetic terms and scattering amplitudes.

The key observation used in the derivation of the unitarity bound was that for  $h \ll 1/\xi$ , the expansion of the potential contains powers of  $(\xi^2 \psi^2)^{2n}$ . Replacing the operators  $\psi^2$  by  $\Lambda^2$  results in quantum corrections containing powers of  $\xi^2 \Lambda^2$ , and, consequently, to the estimate for the energy cutoff  $\Lambda \sim 1/\xi$ . However, this is true only if one is interested in quantum effects at very small values of the Higgs field,  $h \ll 1/\xi$ , which is very far from the inflationary region  $h \gtrsim 1/\sqrt{\xi}$ .

As we already mentioned, for  $h \gg 1/\xi$  the expansion of the potential in powers of  $\psi$  is dramatically different. Indeed, expansion of the potential  $U(\psi)$  (6.8) in powers of  $\psi$  in the intermediate range  $\frac{1}{\xi} \ll h \ll \frac{1}{\sqrt{\xi}}$  does not contain the dangerous factors  $(\xi^2 \psi^2)^{2n}$  :

$$U(\psi) \approx \frac{\lambda}{6\xi^2} \left( \frac{\psi}{1 + \sqrt{\frac{2}{3}}\psi} \right)^2 = \frac{\lambda}{6\xi^2} \left[ \psi^2 - 2\sqrt{\frac{2}{3}}\psi^3 + \dots \right]. \quad (6.15)$$

The dependence on  $\xi$  in this expression is extracted into a single overall coefficient  $\frac{\lambda}{6\xi^2}$ , and all terms in the expansion are proportional to  $\psi^n$ . This is very much different from the small field regime, where the higher order terms were proportional to  $(\xi^2 \psi^2)^{2n}$ . To estimate how vulnerable Eq. (6.15) could be with respect to quantum corrections, one may again replace some of the operators  $\psi^2$  in this expansion by  $\Lambda^2$ . One can easily see that the higher order corrections will remain small for  $\Lambda \ll \psi$ . At the lower boundary of the range  $\frac{1}{\xi} \ll h \ll \frac{1}{\sqrt{\xi}}$ , this leads to the same bound as before:  $\Lambda \sim 1/\xi$ . However, at the upper boundary one has  $\psi = O(1)$ , which means that quantum corrections are not expected to be important until one reaches super-Planckian energies, which are well above the energy scale of inflation.

One can reach similar conclusions for quantum corrections during inflation, when  $\psi > 1$  and the potential is given by Eq. (6.9). This means that the typical energy scale of inflation,  $H \sim \sqrt{\lambda}/\xi$ , is many orders of magnitude below the UV cutoff during this process. Of course, for the processes which occur long after inflation, when  $h \approx \psi < 1/\xi$ , the unitarity bound will be much smaller,  $\Lambda \sim 1/\xi$  [15–17], but this does not affect our ability to describe physical processes during inflation.

An attempt to derive the unitarity bound without using the small field approximation was made in [18]. The authors considered interaction of the inflaton field with gravity in the Jordan frame and argued that the scattering amplitude  $2h \rightarrow 2h$  exceeds the unitarity bound at energy  $E > 1/\xi$ . However, the estimates made in [18] ignored the non-diagonal kinetic terms mixing the scalar field



with gravity in the Jordan frame. These terms disappear in the Einstein frame, and the estimate of the corresponding  $2h \rightarrow 2h$  scattering amplitude shows that it does not violate the unitarity bound at sub-Planckian energies.

In [15,17] it was argued that investigation of scattering of scalar particles on other scalar and vector particles also gives rise to the unitarity bound  $\Lambda \sim 1/\xi$ . Once again, the calculations in [15,17] are based on the expansion  $\psi = h(1 + \xi^2 h^2 + \dots)$ , which is valid only for  $h < 1/\xi$ . In the most interesting interval of the values of the Higgs field  $h \gg 1/\xi$ , one can repeat the arguments given above and again come to the conclusion that the higher order corrections are suppressed for  $\Lambda < \psi$ . The authors of Ref. [19] mentioned a possibility of the unitarity cutoff  $\Lambda \sim 1/\sqrt{\xi}$ , but for  $\xi \sim 10^4$  this cutoff is two orders of magnitude higher than the Hubble constant during inflation, so it is harmless.

In conclusion, we do not think that one should worry too much about the unitarity bound during inflation in the Higgs inflation model of Ref. [2]. However, those who want to feel even better protected against this problem may either try to find a consistent UV completion of this model [20], or switch to the NMSSM and study the model with  $\lambda \ll 1$ , where the presumed unitarity bound  $\Lambda \sim 1/\xi$  is well above the typical energy scale of inflation  $H \approx \sqrt{\frac{\lambda}{3}} \frac{1}{2\xi}$  [4]. That is what we are going to do now.

## 7 Inflation in the NMSSM

Here we will start with the Jordan frame supergravity (we set  $M_P = 1$  throughout this section) with the following frame function, as outlined in Secs. 3.5, 3.8 :

$$\Phi(z, \bar{z}) = -3 + (S\bar{S} + H_u H_u^\dagger + H_d H_d^\dagger) + \frac{3}{2}\chi(H_u \cdot H_d + \text{h.c.}) - \zeta(S\bar{S})^2. \quad (7.1)$$

Here the term  $(S\bar{S} + H_u H_u^\dagger + H_d H_d^\dagger)$  corresponds to the superconformal coupling of the chiral multiplets. The term  $+\frac{3}{2}\chi(H_u \cdot H_d + \text{h.c.})$  is the real part of the holomorphic quadratic function in the curvature-scalar coupling, it breaks the superconformal symmetry of the chiral multiplet coupling. This term reflects the need of the superconformal symmetry breaking to provide a realistic Higgs-type inflationary model proposed in [3] and developed in [1,4]. Finally, the term  $\zeta(S\bar{S})^2$  is added to provide the stability of the origin of the moduli space,  $S = 0$ , as proposed in [4] and used in earlier models in [25].

To embed the NMSSM gauge theory into the Einstein frame supergravity we will use the Kähler

potential and the superpotential

$$\mathcal{K}(z, \bar{z}) = -3 \log\left(-\frac{1}{3}\Phi\right) = -3 \log\left[1 - \frac{1}{3}\left((S\bar{S} + H_u H_u^\dagger + H_d H_d^\dagger)\right) - \frac{\chi}{2}(H_u \cdot H_d + \text{h.c.}) + \frac{\zeta}{3}(S\bar{S})^2\right]. \quad (7.2)$$

$$W = -\lambda S H_u \cdot H_d + \frac{\rho}{3} S^3 \quad (7.3)$$

where the Higgs doublets are defined in (4.1). Note that

$$H_u \cdot H_d \equiv -H_u^0 H_d^0 + H_u^+ H_d^- . \quad (7.4)$$

$$H_u H_u^\dagger + H_d H_d^\dagger = H_u^0 (H_u^0)^\dagger + H_d^0 (H_d^0)^\dagger + H_u^+ (H_u^+)^\dagger + H_d^- (H_d^-)^\dagger \quad (7.5)$$

As in [1], we use a consistent truncation in which the charged superfields  $H_u^+$  and  $H_d^-$  are absent. We will present later in Appendix E the condition for the stability of the inflationary trajectory with regard to the vanishing charged fields. We will use the fact that the dependence on the neutral and charged Higgs fields in  $H_u H_u^\dagger + H_d H_d^\dagger$  is symmetric, whereas the one in  $H_u \cdot H_d$  is anti-symmetric.

Below we use a simplified action of the NMSSM, containing only three superfields:  $S$ ,  $H_u^0$  and  $H_d^0$ , such that:

$$H_1 = \begin{pmatrix} 0 \\ H_u^0 \end{pmatrix}, \quad H_2 = \begin{pmatrix} H_d^0 \\ 0 \end{pmatrix}. \quad (7.6)$$

With this truncation, the frame function, the Kähler potential and the superpotential are:

$$\Phi(z, \bar{z}) = -3 + (|S|^2 + |H_u^0|^2 + |H_d^0|^2) - \frac{3}{2}\chi(H_u^0 H_d^0 + \overline{H_u^0 H_d^0}) - \zeta |S|^4, \quad (7.7)$$

$$\mathcal{K}(z, \bar{z}) = -3 \log\left[1 - \frac{1}{3}\left(|S|^2 + |H_u^0|^2 + |H_d^0|^2\right) + \frac{1}{2}\chi\left(H_u^0 H_d^0 + \overline{H_u^0 H_d^0}\right) + \frac{1}{3}\zeta |S|^4\right], \quad (7.8)$$

$$W = \lambda S H_u^0 H_d^0 + \frac{\rho}{3} S^3. \quad (7.9)$$

The D-term potential in the Jordan frame remains simple

$$V_J^D = \frac{g'^2}{8}(|H_u^0|^2 - |H_d^0|^2)^2 + \frac{g^2}{8}(H_u^\dagger \vec{\tau} H_u + H_d^\dagger \vec{\tau} H_d)^2. \quad (7.10)$$

The  $S$ -dependent terms in the F-term potential even in the Jordan frame are complicated due to  $\zeta$ -corrections. However, we will establish that the stabilization of some scalars takes place and only one real scalar remains light during inflation. We will find out that during inflation all complicated corrections to the potential drop and we can explain the inflationary dynamics regime using the simple features of the superconformal matter coupling and its particular breaking.

## 7.1 Basic features of inflation in the NMSSM

For a numerical investigation of inflation in the NMSSM model with three chiral multiplets and truncated charged Higgs fields we use the Mathematica code [37] designed to compute the Einstein frame potentials and scalar kinetic terms for any number of moduli with generic Kähler potential  $\mathcal{K}(z, \bar{z})$  and generic superpotential  $W(z)$ .

The potential in the NMSSM depends on three complex superfields:

$$S = se^{i\alpha}/\sqrt{2}, \quad H_u^0 = h_1 e^{i\alpha_1}/\sqrt{2}, \quad H_d^0 = h_2 e^{i\alpha_2}/\sqrt{2}. \quad (7.11)$$

Note that here we slightly deviate from notation of our previous paper [1]: We divided all fields by  $\sqrt{2}$ . The main reason to do it is to keep the fields  $h$  canonically normalized in the Jordan frame. It will simplify the comparison of inflation in the NMSSM with inflation in the nonsupersymmetric standard model [2].

The standard mixing of the Higgs fields is defined as:

$$h_1 \equiv h \cos \beta, \quad h_2 \equiv h \sin \beta, \quad (7.12)$$

which leaves us with two real fields,  $h$  and  $\beta$ , instead of  $h_1$  and  $h_2$ . The  $D$ -flat direction, defined by  $V_J^D = 0$  requires that

$$\beta = \pi/4; \quad h_1^2 = h_2^2 = h^2/2. \quad (7.13)$$

In this subsection we will consider the simplest inflationary solution with  $\beta = \pi/4$ ,  $\alpha_i = 0$  and  $s = 0$ . In the next subsections we will investigate the conditions required for stability of this solution with respect to the  $\beta$ ,  $\alpha_i$  and  $s$ .

We find that in the Jordan frame the total supergravity action for the field  $h$ , under the condition that  $\beta = \pi/4$ ,  $\alpha_i = 0$  and  $s = 0$ , is reduced to

$$\mathcal{L}_J(h, g_J; \lambda) = \sqrt{-g_J} \left[ \frac{1}{2} \left( 1 - \frac{1}{6} h^2 + \frac{1}{4} \chi h^2 \right) R(g_J) - \frac{1}{2} (\partial_\mu h)^2 - \frac{\lambda^2}{16} h^4 \right]. \quad (7.14)$$

An interesting question to ask here is: *why the complete supergravity action of the NMSSM model with the frame function in (7.7) and superpotential in (7.9) at the inflationary trajectory with all fields real and  $s = 0$  is so simple in the Jordan frame?*

The answer to this question consists of several parts. 1) The first term appears directly from our expression for the frame function (7.7). 2) The kinetic term for scalars at  $S = 0$  is canonical due to

our choice of geometric breaking of superconformal symmetry, which does not affect this important property. 3) The value of the auxiliary field  $\mathcal{A}_\mu(z, \bar{z})$  vanishes for real scalars. 4) The potential in the Jordan frame (3.36) in the D-flat direction with cubic superpotential is

$$V_J = G^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}}. \quad (7.15)$$

The term  $\chi(H_u^0 H_d^0 + \overline{H_u^0 H_d^0})$  in the frame function signals the deviation from the superconformal theory. This deviation, however, is controllable. Namely, with  $W_{\text{Higgs}} = -\lambda S H_u \cdot H_d + \frac{\rho}{3} S^3$  if we succeed to stabilize the theory at  $S = 0$ , the only contribution to the potential at  $S = 0$  comes from the term

$$V_J|_{S=0} = \frac{\partial W}{\partial S} G^{S\bar{S}} \frac{\partial \bar{W}}{\partial \bar{S}} = \lambda^2 G^{S\bar{S}} |H_u \cdot H_d|^2 = \lambda^2 |H_u \cdot H_d|^2 = \frac{\lambda^2}{16} h^4. \quad (7.16)$$

The metric  $G^{S\bar{S}} = 1$  since the field  $S$  does not enter in the superconformal symmetry breaking  $\chi$ -term, and therefore even after this breaking of superconformal symmetry the  $h$ -dependent part of the potential remains quartic: in the D-flat direction for real fields it is equal to  $\frac{\lambda^2}{16} h^4$ .

It is easy to compare the supergravity action on inflationary trajectory (7.14) with the nonsuper-symmetric Jordan frame action (6.2), which we reproduce here again to simplify the comparison:

$$\mathcal{L}_J^{\text{SM}} = \sqrt{-g_J} \left[ \frac{1 + \xi h^2}{2} R(g_J) - \frac{1}{2} (\partial_\mu h)^2 - \frac{\lambda}{4} h^4 \right]. \quad (7.17)$$

These two actions coincide after the following identification of the parameters:

$$\xi \longleftrightarrow -\frac{1}{6} + \frac{1}{4}\chi, \quad \lambda \longleftrightarrow \frac{\lambda^2}{4}. \quad (7.18)$$

On the left hand side of each equation in (7.18) we have parameters of the standard model as in Eq. (7.17). On the right hand side of each equation above we have parameters of the NMSSM inflation model as in Eq. (7.14).

After the identification (7.18), all features of inflation in the NMSSM can be deduced from the results of Ref. [2] presented in Sec. 3. In particular, the slow-roll parameters are

$$\epsilon \simeq \frac{64}{3\chi^2 h^4}, \quad (7.19)$$

$$\eta \simeq -\frac{16}{3\chi h^2}. \quad (7.20)$$

Slow roll ends when  $\epsilon, \eta \simeq 1$ , so the field value at the end of inflation is  $h_{\text{end}} \simeq 2.2/\sqrt{\chi}$ . The number of e-foldings during the slow roll of the field  $h$  from its initial value  $h_0$ , for  $h_0 \gg h_{\text{end}}$ , is given by

$$N \simeq \frac{3}{16} \xi h_0^2. \quad (7.21)$$

For  $N \sim 60$ , the amplitude of scalar perturbations of metric corresponds to the COBE normalization for

$$\chi \simeq 10^5 \lambda. \quad (7.22)$$

The asymptotic value of the Einstein frame potential  $V_E$  at large  $h$  is  $\frac{\lambda^2}{\chi^2}$ , and the Hubble constant during inflation in this model is  $H \approx \frac{1}{\sqrt{3}} \frac{\lambda}{\chi}$ .

To give a particular example, let us take  $\lambda = 10^{-2}$ . In this case one should have  $\chi = 10^3$ . Inflation ends at  $h_{\text{end}} \sim 0.07$ . The last 60 e-folds of inflation begin at  $h_0 \simeq 0.37$ . All observational consequences are the same as in the nonsupersymmetric model [2]. In particular, the spectral index is  $n_s \sim 0.97$ , and the tensor to scalar ratio is  $r \approx 0.0033$ . These results are valid for  $\chi \gg 1$ . They are invariant with respect to the simultaneous rescaling  $\lambda \rightarrow c \lambda$ ,  $\chi \rightarrow c \chi$  and  $h_0 \rightarrow h_0/\sqrt{c}$ . For a complete investigation of inflation in this model one would also need to study quantum corrections in supersymmetric theory as it was done for the standard model case in [2].

## 7.2 Stabilization of the non-inflaton directions in the moduli space

We would like to split all 6 components of the 3 complex scalars  $S, H_u^0, H_d^0$  in (7.11) into heavy and light ones. First of all, we impose a unitary gauge, when one combination of the neutral components of  $H_u^0$  and  $H_d^0$  is the Goldstone boson and is absent in the unitary gauge. We take a condition  $\alpha_1 = \alpha_2$ .

We study stabilization of angles  $\alpha, \beta, \gamma \equiv \alpha_1 + \alpha_2$  and of the field  $s$  using the complete and explicit expressions for the kinetic terms and the potential in the Einstein frame derived using the Mathematica code [37] for the Kähler potential in (7.8) and superpotential in (7.9). We present some details of the action in the Jordan frame and the Einstein frame for the real fields  $h$  and  $s$  in Appendix B.

### 7.2.1 Stabilization of angles

Now we must check the stability of the inflationary solution with respect to the fields  $\beta = \pi/4$ ,  $\alpha = 0$ ,  $\gamma \equiv \alpha_1 + \alpha_2 = 0$  and  $s = 0$ . We already checked in [1] that during inflation the  $CP$ -invariant solution in which  $S, H_u^0$  and  $H_d^0$  are real, is stable with respect to the field  $\beta$ . The degree of stability is described by the mass squared of the field  $\beta$ . During inflation, in the limit  $\chi h^2 \gg 1$ , one has the kinetic term  $\frac{2}{\chi}(\partial\beta)^2$  and the second derivative of the potential over  $\beta$  is  $V_{\beta,\beta}(\beta = \pi/4) = \frac{4(g^2 + g'^2)}{\chi^2}$ . This means that the effective mass

$$m_\beta^2 = \frac{g^2 + g'^2}{\chi} = \frac{g^2 + g'^2}{\lambda^2} 3\chi H^2, \quad (7.23)$$

is greater than  $H^2 = \frac{1}{3} \frac{\lambda^2}{\chi^2}$  in the most natural case  $\lambda^2 < 3\chi(g^2 + g'^2)$ , one has  $m_\beta^2 \gg H^2$ . Thus, there is no slow-roll regime with respect to the change of  $\beta$  during inflation, because the mass squared of perturbations of the angle  $\beta$  is much greater than  $H^2 = \frac{1}{3} \frac{\lambda^2}{\chi^2}$ . During inflation the field  $\beta$  rapidly approaches  $\pi/4$  and stays there. For  $\lambda^2 \ll g^2, g'^2$ , the regime with  $\beta = \pi/4$ ,  $h_1^2 = h_2^2 = h^2/2$  remains stable even long after inflation, until the soft supersymmetry breaking terms become important and change the final value of  $\beta$  [1].

Now we should study the dependence of the potential on angles  $\alpha$  and  $\alpha_1 = \alpha_2$  near the inflationary trajectory  $s = 0$ ,  $\beta = \pi/4$ . The potential at  $s = 0$  does not depend on  $\alpha$ . Therefore instead of investigation of excitations of  $\alpha$  one should study stability of the potential with respect to the field  $s$  for different  $\alpha$ . For small  $s$  and  $\lambda\rho < 0$ , the minimum of the potential with respect to  $\alpha$  occurs at  $\alpha = 0$  [1]. As we see later, stability in this direction is achieved by adding the term  $\frac{\zeta}{3}(S\bar{S})^2$  in the Kähler potential, following the suggestion made in [4] and ideas developed in [25].

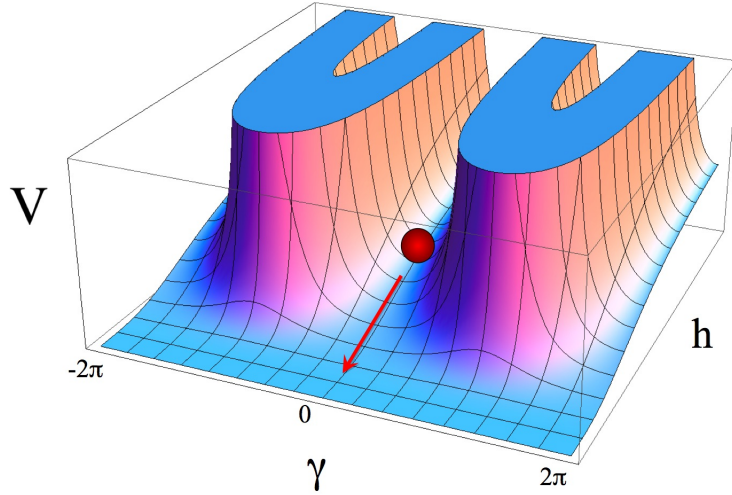


Figure 1: Stabilization of the angle  $\gamma = \alpha_1 + \alpha_2 = 0$  near the inflationary trajectory. The infinitely high horseshoe barriers correspond to the singularity of the Kähler geometry. These barriers separate the admissible range of variables from the forbidden part of the landscape (inside the horseshoes), where the argument of the logarithm in the expression for the Kähler potential becomes negative.

As explained above, the combination  $\alpha_1 - \alpha_2$  describes a Goldstone boson, which is replaced by a longitudinal component of the vector field. In the unitary gauge  $\alpha_1 - \alpha_2 = 0$ . The remaining combination  $\gamma = \alpha_1 + \alpha_2$  corresponds to a scalar field with mass which during inflation is

$$m_\gamma^2 \approx 4H^2. \quad (7.24)$$

To see it we should analyze the potential along the inflationary direction  $s = 0$ ,  $\beta = \pi/4$ .

$$V(h, \gamma) = \frac{9\lambda^2 h^4}{(12 - 2h^2 + 3\chi h^2 \cos \gamma)^2}, \quad (7.25)$$

where  $\gamma = \alpha_1 + \alpha_2$ . At  $\chi h^2 \gg 1$   $V(h, \gamma) = \frac{\lambda^2}{\chi^2 \cos^2 \gamma}$ . Therefore its second derivative at  $\gamma = 0$  is given by  $V_{\gamma, \gamma}(\gamma = 0) = \frac{2\lambda^2}{\chi^2}$ . The matrix of kinetic terms for the fields  $H_i$  in the limit  $\chi h^2 \gg 1$  at  $S = 0$  in the unitary gauge at  $\beta = \pi/4$  simplifies to

$$\mathcal{L}_{kin} \Rightarrow \frac{3}{h^2}(\partial h)^2 + \frac{3}{4}(\partial \gamma)^2. \quad (7.26)$$

Therefore the mass of the canonical field

$$m_\gamma^2 = \frac{2}{3}V_{\gamma, \gamma}(\gamma = 0) = \frac{4\lambda^2}{3\chi^2} = 4V/3 = 4H^2. \quad (7.27)$$

Since  $m_\gamma^2$  is of the same order as  $H^2$ , during inflation the field  $\gamma$  rapidly rolls towards  $\gamma = 0$  and stays there. Therefore during inflation we have  $\gamma = 0$ , or, equivalently,  $\alpha_1 = \alpha_2 = 0$ . During inflation,  $\mathbb{Z}_3$  symmetry is broken by the term  $\frac{3}{2}\chi(H_u \cdot H_d + \text{h.c.})$  in the Kähler potential. The potential has a minimum with respect to  $\gamma$  only at  $\gamma = 0$  (or, more exactly, at  $\gamma = 2\pi n$ ), see Fig. 1. Therefore inflation naturally singles out only one of the three possible minima related to each other by  $\mathbb{Z}_3$  symmetry. However, long after inflation, when soft supersymmetry breaking terms become important at small  $h$  a new strong mechanism of breaking  $\mathbb{Z}_3$  symmetry takes place as we have shown in Sec. 5. It originates from the real part of the quadratic holomorphic term  $\frac{3}{2}\chi(H_u \cdot H_d + \text{h.c.})$  in the Kähler potential, however, it removes domain walls via the induced soft term in the potential,  $V_{\text{soft}} \sim \chi(H_u \cdot H_d + \text{h.c.})$ .

### 7.2.2 Stabilization of the field $S$

As we have shown in [1], the original version of inflation in the NMSSM model [3] suffered from the tachyonic instability with respect to the field  $S$ . However, one may circumvent this problem by taking into account interactions of the field  $S$  with superheavy fields that one may add to the model, or simply by adding a term  $-\frac{1}{3}\zeta(S\bar{S})^2$  to the frame function and Kähler potential [4, 25]. This term helps to stabilize the inflationary trajectory in the toy model considered in [4]. Here we will check what happens in the NMSSM model. An investigation required an extensive use of the Mathematica program [37]; we will present here only the main results, some details are given in the Appendix B.

During the inflationary regime, the leading behavior of the F-term potential in this model for  $\xi h^2 \gg 1$  is given by

$$V_E \sim \frac{\lambda^2}{\chi^2} - \left( |\lambda\rho| + \frac{\lambda^2}{6\chi}(2 - 3\zeta\chi h^2) \right) \frac{4s^2}{\chi^2 h^2} + O(s^4). \quad (7.28)$$

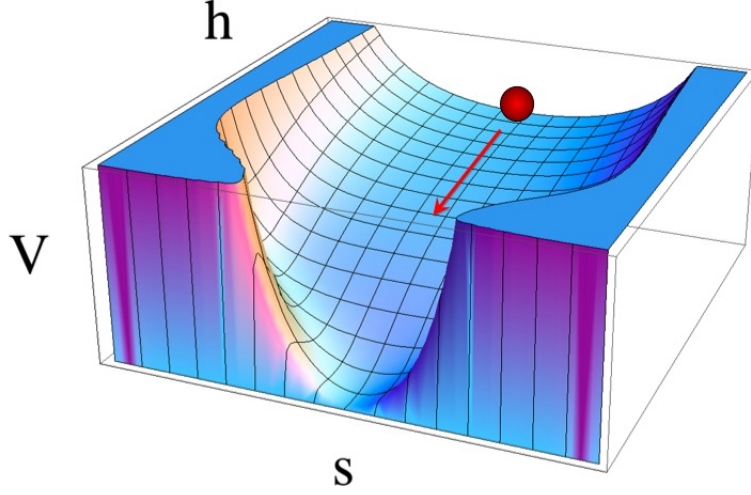


Figure 2: Stabilization of the field  $s$  near the inflationary trajectory  $s = 0$  for  $\zeta > \frac{2|\lambda\rho|}{\lambda^2 h^2} + 0.0327$ .

To find the effective mass of the  $s$  field, attention must be paid to the nonminimal normalization of the field  $S = se^{i\alpha}/\sqrt{2}$ . At constant  $\alpha$ , the kinetic term of field  $S$  is given by

$$g_{S\bar{S}}\partial S\partial\bar{S} = \frac{4}{\chi h^2}\partial S\partial\bar{S} = \frac{2}{\chi h^2}(\partial s)^2. \quad (7.29)$$

Here, as we already explained before, the  $\zeta$ -correction to the kinetic term of the  $s$ -field is always small comparative to other terms, we neglected it. For small  $s$ , in the vicinity of the inflationary trajectory, the Lagrangian of the field  $s$  for  $\chi h^2 \gg 1$  is

$$\mathcal{L}_E \approx -\frac{2}{\chi h^2}(\partial s)^2 - \frac{\lambda^2}{\chi^2} + \left(|\lambda\rho| + \frac{\lambda^2}{6\chi}(2 - 3\zeta\chi h^2)\right) \frac{4s^2}{\chi^2 h^2}. \quad (7.30)$$

Therefore the mass of the canonical field  $s$  is:

$$m_s^2 \sim 2 \left( \frac{\lambda^2}{6\chi^2}(3\zeta\chi h^2 - 2) - \frac{|\lambda\rho|}{\chi} \right). \quad (7.31)$$

Thus the condition of stability of the inflationary trajectory at  $s = 0$  is

$$\zeta > \frac{2|\lambda\rho|}{\lambda^2 h^2} + \frac{2}{3\chi h^2}. \quad (7.32)$$

However, this simple analytic form of the bound can be used only at  $\sqrt{\chi}h \gg 1$ , i.e. well before the end of inflation. Meanwhile, the greatest danger of instability occurs at the very end of inflation. A more accurate condition, which is valid even for  $\chi h^2 < 1$ , is

$$\zeta > \frac{2|\lambda\rho|}{\lambda^2 h^2} + \frac{2(y^2 - 32)}{3y^2(y + 4)}, \quad (7.33)$$



where  $y = \chi h^2$ . The function  $\frac{2(y^2-32)}{3y^2(y+4)}$  takes its maximal value 0.0327 at  $y = \chi h^2 \approx 10.9$ . This point corresponds to the moment of maximal vulnerability with respect to the tachyonic instability. Therefore the trajectory  $s = 0$  remains stable for all  $h$  if

$$\zeta > \frac{2|\lambda\rho|}{\lambda^2 h^2} + 0.0327. \quad (7.34)$$

This result is illustrated by Fig. 2, which shows the potential for  $\rho = 0$  and  $\zeta = 0.04$ .

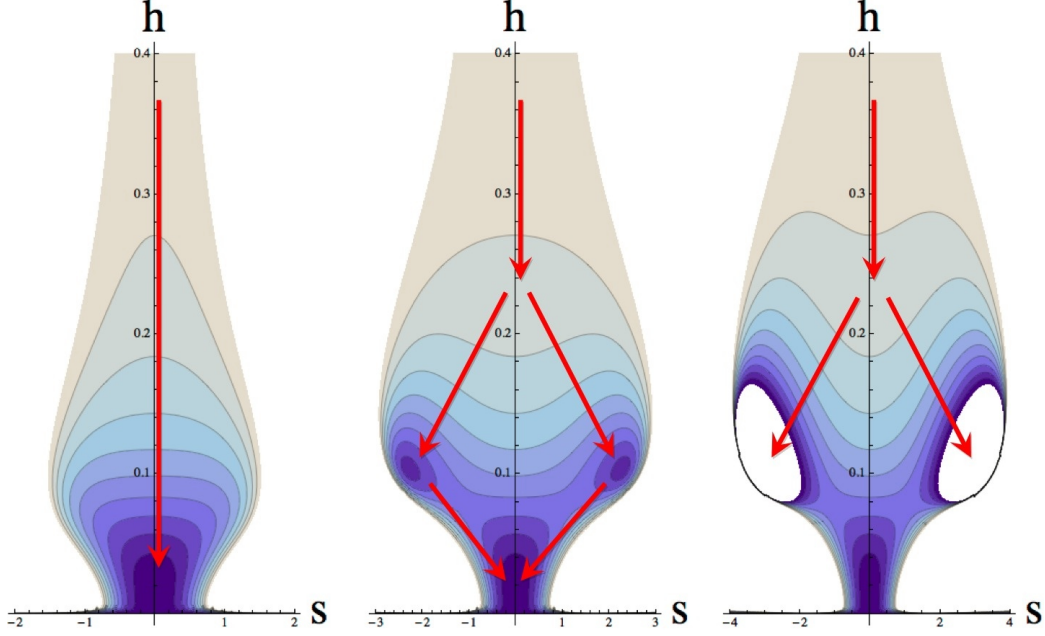


Figure 3: Contour plots of the potential for the fields  $s$  and  $h$  during the last 60 e-folds of inflation for  $\rho = 0$ ,  $\lambda = 10^{-2}$  and  $\xi = 5 \times 10^2$ , for three different values of  $\zeta$ . Darker parts correspond to smaller positive density, white ovals on the right panel correspond to negative values of the potential. Red arrows show the evolution of the fields during inflation.

To illustrate the general situation in a more complete way, we show the contour plot of the potential of the fields  $h$  and  $s$  for  $\rho = 0$  and various values of the coupling constant  $\zeta$  in Fig. 3. The first panel corresponds to the potential shown in Fig. 2, with  $\zeta = 0.04$ . The field  $s$  is stabilized during inflation and after it. For  $\zeta$  slightly smaller than 0.0327, the tachyonic instability of the field  $s$  at the very end of inflation may force the field to deviate from the straight path, which results in tachyonic preheating. This possibility is exotic but not dangerous. Further decrease of  $\zeta$  may result in formation of two local minima of the potential, as shown at the second panel for  $\zeta = 0.01$ . The field may roll to one of the two metastable minima and stuck there until it tunnels to the global minimum at  $h, s = 0$ . This leads to large inhomogeneities, as in the old inflation scenario. Finally, the third panel shows the potential

for  $\zeta = 0.003$ . The field rolls down to one of the two minima with negative values of the potential, looking like two white eyes of an alien, and the universe collapses.

Thus, for  $\zeta$  significantly below  $\frac{2|\lambda\rho|}{\lambda^2 h^2} + 0.0327$  one may encounter production of gross inhomogeneities and even a collapse of the universe. Fortunately, however, one can have a successful inflationary scenario if  $\zeta$  is greater than  $\frac{2|\lambda\rho|}{\lambda^2 h^2} + 0.0327$ , and even if  $\zeta$  is slightly below this limit.

### 7.3 Higher order corrections

In Sect. 6.2 we argued that the unitarity bound discussed in [15–17] is not expected to pose any problems for the description of inflation in this class of models. However, for the full investigation of these models one should also investigate running of the coupling constants, along the lines of Ref. [2] where it was done for the standard model case. For our models the quantum corrections have to be studied with account of supersymmetry.

Independently of the issue of quantum field theory type corrections, one may wonder how stable are our conclusions with respect to modifications of various ingredients of these models. According to (7.22), our model does require small  $\lambda$  and large  $\chi$ , so that  $\lambda/\chi \sim 10^{-5}$ . This, by itself, does not look much better than the standard requirement  $\lambda \sim 10^{-6}$  in the theory  $\lambda^2 \phi^4/16$ , as in Eq. (7.14). However, the choice of the large  $\chi$  may provide an additional robustness of the model with respect to higher terms in the expression for the Kähler potential.

Indeed, inflation in our model occurs at  $h \sim 1/\sqrt{\chi}$ . Suppose that, in addition to the term breaking superconformal symmetry,  $-\frac{1}{2}\chi(H_u \cdot H_d + \text{h.c.})$ , there is also a higher order term  $c(H_u \cdot H_d + \text{h.c.})^2$  in the Kähler potential. For  $c = O(1)$ , this term remains smaller than the first one for  $h < \sqrt{\chi}$ . Therefore if one takes  $\chi \gg 1$ , then for  $h \lesssim \sqrt{\chi}$  one should not be concerned about higher order terms described above. According to (7.21), the total number of e-folds in this regime is  $N \sim \chi^2$ , which is incredibly large for the values of  $\chi$  considered in our paper. The total number of e-folds may be even greater if there is some symmetry which protects the original structure of the Kähler potential at large values of the inflaton field.

## 7.4 Gravitino problem and inflation beyond the NMSSM

The possibility to have an inflationary regime in the NMSSM does not necessarily mean that the cosmological theory based on this scenario is fully consistent. Supergravity is plagued by the cosmological moduli problem and by the gravitino problem. Inflation helps to solve the gravitino problem, but only if the reheating temperature  $T_r$  after inflation is sufficiently small. The bounds on  $T_r$  depend on the gravitino mass and other parameters, but typically it should be smaller than  $10^8$  GeV, see [38–43] for a more detailed discussion of this issue.

One way to avoid this problem is to assume that the energy scale of inflation is very low, which leads to a small reheating temperature. However, inflation in the NMSSM occurs at the energy density  $\frac{\lambda^2}{\chi^2} \sim 10^{-10}$ , in Planck units. If reheating happens instantly, this energy is converted to thermal energy  $T_r^4 \sim 10^{-10}$ . This gives an estimate  $T_r \sim 10^{15}$  GeV.

One may have a much smaller reheating temperature if the inflaton field extremely weakly couples to matter, which leads to a delay in thermalization. During this delay, the energy of the inflaton field decreases, and the reheating temperature becomes smaller. For example, one may consider inflationary models where the inflaton belongs to hidden sector, and its decay to observable particles is suppressed by the small gravitational coupling. But the Higgs fields belong to the observable sector and they couple to matter quite strongly. An investigation of reheating in the Higgs inflation [44] suggested that the reheating temperature is about  $10^{13}$  GeV. A more detailed investigation performed in [45] demonstrated that the process of reheating in this theory is quite complex, being a combination of the perturbative reheating [46], parametric resonance [47], instant preheating [48] and tachyonic preheating [49]. The authors argued that the full investigation of this complicated process should be done by lattice simulations [50]. However, there is no obvious reason to expect that this investigation will yield the reheating temperature 5 orders of magnitude smaller than the estimate  $T_r \sim 10^{13}$  GeV made in [44].

There are several possible ways to address this problem, even if the future investigation confirms that  $T_r \gg 10^8$  GeV. First of all, the gravitino problem disappears if the gravitino mass is below keV, or if it is several orders of magnitude above the TeV scale, see e.g. [41–43] for a recent discussion and more precise bounds on the gravitino mass. Both of these possibilities are realistic. For example, in the model of conformal gauge mediation one can have gravitino with mass below 10 eV [51]. Superheavy gravitino have been discussed in [52–54].

Another solution is to have a second stage of inflation after the NMSSM inflation. This is a realistic possibility since the energy scale of the NMSSM inflation is very high, so it is quite possible to have a second stage of inflation at a much smaller energy scale after the NMSSM inflation. If this stage is short, as in the thermal inflation scenario [55], then it may solve the gravitino problem, and all observational predictions on the NMSSM inflation will remain intact. On the other hand, if the second stage of inflation is sufficiently long, then it will determine all properties of the observable part of the universe. In this respect, it is quite encouraging that one can develop a large class of new models of chaotic inflation based on the ideas discussed in this paper, but without necessarily identifying the inflaton field with the Higgs field of the standard model [21].

## 8 Conclusion

Supergravity phenomenology was mostly developed in the Einstein frame where there is no scalar-curvature coupling. In this paper we propose a superconformal approach to supergravity phenomenology and cosmology. One can start with the  $SU(2,2|1)$  superconformal theory of chiral and vector multiplets interacting with supergravity Weyl multiplet. This theory contains a conformal compensator, which can be gauged away, giving rise to the Planck mass. In this paper we identified a special class of supergravity models: If chiral and vector multiplets of the superconformal theory are decoupled from the conformal compensator, the part of the action describing matter fields in the Jordan frame remains superconformal invariant. This action is unusually simple: kinetic terms are canonical, supergravity potential coincides with the global theory potential, scalars are conformally coupled to gravity. The potential is quartic, the theory has no mass terms, no nonrenormalizable terms, no cosmological constant.

Theories of this type may form a convenient starting point for constructing phenomenological models. In such models, one may attribute smallness of all mass parameters to the effects of breaking of the superconformal symmetry, which can be achieved e.g. due to gravitational effects. In particular, these theories may provide a natural supergravity embedding for the NMSSM. Superconformal symmetry breaking is introduced by the real part of the holomorphic quadratic nonminimal scalar-curvature coupling, by terms designed to stabilize some fields at the origin of moduli space, and by interactions with a hidden sector. This approach to supergravity phenomenology from the underlying superconformal theory allows to address the  $\mu$ -problem and the domain wall problem, and to obtain an inflationary regime in the NMSSM. Efficient reheating after inflation in the NMSSM may lead to the cosmological

gravitino problem. This problem can be solved if one considers models with superlight or superheavy gravitino, or if one postulates a secondary stage of inflation after the NMSSM inflation. Fortunately, the general methods developed during the investigation of the canonical superconformal supergravity and inflation in the NMSSM can be used for construction of a new broad class of models of chaotic inflation in supergravity with a functional freedom of choice of the inflaton potential [21].

## Acknowledgments

We are grateful to S. A. Abel, S. Dimopoulos, M. Einhorn, U. Ellwanger, J. Garcia-Bellido, C. Germani, P. Graham, A. Guth, S. Kachru, B. Kyae, H. M. Lee, P. Nilles, K. Olive, F. Quevedo, S. Shenker, L. Susskind and A. Westphal for useful discussions. The work of RK and AL is supported by NSF Grant No. 0756174. The work of SF is supported by ERC Advanced Grant No. 226455, *Supersymmetry, Quantum Gravity and Gauge Fields (Superfields)*, and in part by INFN, sez. L.N.F. The work of AM is supported by INFN, as a visitor to Stanford University, Stanford, CA, USA. The work of AVP is supported in part by the FWO - Vlaanderen, Project No. G.0235.05, and in part by the Federal Office for Scientific, Technical and Cultural Affairs through the “Interuniversity Attraction Poles Programme – Belgian Science Policy” P6/11-P.

## A Complete CSS action

We present here the full action corresponding to the 4 assumptions given in Sec. 3.2. We will eliminate the scalar auxiliary fields of supergravity, but leave the auxiliary vector  $A_\mu$  as an independent field. We will make use of the gauge conditions mentioned in Sec. 3.3.

Due to the separation of  $X^0$  from the other fields, the action can be split and we can write

$$\begin{aligned} S &= \int d^4x \left\{ [-|X^0|^2 + |X^\alpha|^2]_D + [\tfrac{1}{3}d_{\alpha\beta\gamma}X^\alpha X^\beta X^\gamma]_F + [f_{AB}\bar{\lambda}^A P_L \lambda^B]_F \right\} \\ &= \int d^4x [\mathcal{L}_{\text{SG}} + \mathcal{L}_{\text{conf}}] , \end{aligned} \tag{A.1}$$

where  $S_{\text{SG}}$  is the action of pure supergravity, which is produced from the  $|X^0|^2$ -term, and  $S_{\text{conf}}$  is the conformal action of all the other physical fields. The former is given by

$$\mathcal{L}_{\text{SG}} = e \tfrac{1}{2} M_P^2 \left[ R(\omega(e, \psi)) - \bar{\psi}_\mu \gamma^{\mu\nu\rho} \left( \partial_\nu + \tfrac{1}{4} \omega_\nu^{ab}(e, \psi) \gamma_{ab} \right) \psi_\rho + 6 A^\mu A_\mu \right] , \tag{A.2}$$

where we already replaced  $X^0$  by its gauge-fixed value  $\sqrt{3} M_P$ . The conformal  $D$ -terms are

$$\begin{aligned}
[|X^\alpha|^2]_D = & e \delta_{\alpha\bar{\beta}} \left\{ -\mathcal{D}_\mu X^\alpha \mathcal{D}^\mu \bar{X}^{\bar{\beta}} - \frac{1}{2} \bar{\Omega}^\alpha \not{D} \Omega^{\bar{\beta}} - \frac{1}{2} \bar{\Omega}^{\bar{\beta}} \not{D} \Omega^\alpha + F^\alpha \bar{F}^{\bar{\beta}} \right. \\
& - i \bar{X}^{\bar{\beta}} k_A{}^\alpha D^A - \sqrt{2} \bar{\lambda}^A \left( \Omega^\alpha k_A{}^{\bar{\beta}} + \Omega^{\bar{\beta}} k_A{}^\alpha \right) \\
& + \left[ \frac{1}{2\sqrt{2}} \bar{\psi} \cdot \gamma \left( F^\alpha \Omega^{\bar{\beta}} - \not{D} \bar{X}^{\bar{\beta}} \Omega^\alpha + X^{\bar{\beta}} P_R \lambda^A k_A{}^\alpha \right) - \frac{1}{6\sqrt{2}} \bar{X}^{\bar{\beta}} \bar{\Omega}^\alpha \gamma^{\mu\nu} R'_{\mu\nu}(Q) \right. \\
& + \frac{1}{8} i \varepsilon^{\mu\nu\rho\sigma} \bar{\psi}_\mu \gamma_\nu \psi_\rho \left( \bar{X}^{\bar{\beta}} \mathcal{D}_\sigma X^\alpha + \frac{1}{2} \bar{\Omega}^\alpha \gamma_\sigma \Omega^{\bar{\beta}} + \frac{1}{\sqrt{2}} \bar{X}^{\bar{\beta}} \bar{\psi}_\sigma \Omega^\alpha \right) + \text{h.c.} \left. \right] \\
& \left. + \frac{1}{6} X^\alpha \bar{X}^{\bar{\beta}} \left( -R(\omega(e, \psi)) + \frac{1}{2} \bar{\psi}_\mu \gamma^{\mu\nu\rho} R'_{\nu\rho}(Q) \right) \right\}. \tag{A.3}
\end{aligned}$$

The fermions are chiral:

$$P_L \Omega^\alpha = \Omega^\alpha, \quad P_R \Omega^{\bar{\beta}} = \Omega^{\bar{\beta}}. \tag{A.4}$$

The covariant curvature and covariant derivatives in this equation are

$$\begin{aligned}
\mathcal{D}_\mu X^\alpha &= (\partial_\mu - i A_\mu) X^\alpha - \frac{1}{\sqrt{2}} \bar{\psi}_\mu \Omega^\alpha - A_\mu^A k_A{}^\alpha, \\
\mathcal{D}_\mu \Omega^\alpha &= \left( \partial_\mu + \frac{1}{4} \omega_\mu{}^{ab}(e, \psi) \gamma_{ab} + \frac{1}{2} i A_\mu \right) \Omega^\alpha - \frac{1}{\sqrt{2}} (\not{D} X^\alpha + F^\alpha) \psi_\mu - \sqrt{2} X^\alpha \phi_\mu - A_\mu^A (m_A)^\alpha{}_\beta \Omega^\beta, \\
R'_{\mu\nu}(Q) &= 2 \left( \partial_{[\mu} - \frac{3}{2} i A_{[\mu} \gamma_* + \frac{1}{4} \omega_{[\mu}{}^{ab}(e, \psi) \gamma_{ab} \right) \psi_{\nu]}. \tag{A.5}
\end{aligned}$$

The spin connection  $\omega_\mu{}^{ab}$  in this equation contains  $\psi$ -torsion. The field  $\phi_\mu$  is the composite gauge field of special supersymmetry:

$$\phi_\mu = -\frac{1}{2} \gamma^\nu R'_{\mu\nu}(Q) + \frac{1}{12} \gamma_\mu \gamma^{ab} R'_{ab}(Q). \tag{A.6}$$

The superpotential term is

$$\begin{aligned}
\left[ \frac{1}{3} d_{\alpha\beta\gamma} X^\alpha X^\beta X^\gamma \right]_F e^{-1} &= d_{\alpha\beta\gamma} X^\alpha X^\beta F^\gamma - d_{\alpha\beta\gamma} X^\alpha \bar{\Omega}^\beta \Omega^\gamma + d_{\alpha\beta\gamma} X^\alpha X^\beta \bar{\psi} \cdot \gamma \Omega^\gamma \\
&+ \frac{1}{6} d_{\alpha\beta\gamma} X^\alpha X^\beta X^\gamma \bar{\psi}_\mu P_R \gamma^{\mu\nu} \psi_\nu + \text{h.c.} . \tag{A.7}
\end{aligned}$$

The superconformal-invariant kinetic terms for the gauge multiplets are

$$\begin{aligned}
[f_{AB} \bar{\lambda}^A P_L \lambda^B]_F e^{-1} &= -\frac{1}{4} f_{AB} \left[ 2 \bar{\lambda}^A P_L \not{D} \lambda^B + \hat{F}_{\mu\nu}^{-A} \hat{F}^{\mu\nu -B} - D^A D^B \right. \\
&+ \frac{1}{4} \bar{\psi} \cdot \gamma P_L \left( \frac{1}{2} \gamma^{\mu\nu} \hat{F}_{\mu\nu}^{-A} - i D^A \right) \lambda^B - \frac{1}{8} \bar{\psi}_\mu \gamma^{\mu\nu} P_R \psi_\nu \bar{\lambda}^A P_L \lambda^B \left. \right] + \text{h.c.} . \tag{A.8}
\end{aligned}$$

Here

$$\begin{aligned}
\hat{F}_{\mu\nu}{}^A &= 2 \partial_{[\mu} A_{\nu]}^A + g f_{BC}{}^A A_\mu{}^B A_\nu{}^C + \bar{\psi}_{[\mu} \gamma_{\nu]} \lambda^A, \\
\mathcal{D}_\mu \lambda^A &\equiv \left( \partial_\mu + \frac{1}{4} \omega_\mu{}^{ab}(e, \psi) \gamma_{ab} - \frac{3}{2} i \gamma_* A_\mu \right) \lambda^A + \lambda^C A_\mu{}^B f_{BC}{}^A - \left[ \frac{1}{4} \gamma^{ab} \hat{F}_{ab}{}^A + \frac{1}{2} i \gamma_* D^A \right] \psi_\mu. \tag{A.9}
\end{aligned}$$

Many cancelations occur in terms with gravitinos when the various covariantizations are written in detail, and the torsion terms are extracted from the spin connection. The supergravity action is then

$$\begin{aligned}\mathcal{L}_{\text{SG}} &= e \frac{1}{2} M_P^2 \left[ R(\omega(e)) - \bar{\psi}_\mu \gamma^{\mu\nu\rho} \left( \partial_\nu + \frac{1}{4} \omega_\nu^{ab}(e) \gamma_{ab} \right) \psi_\rho + 6 A^\mu A_\mu + \mathcal{L}_{\text{SG,torsion}} \right], \\ \mathcal{L}_{\text{SG,torsion}} &= -\frac{1}{16} \left[ (\bar{\psi}^\rho \gamma^\mu \psi^\nu) (\bar{\psi}_\rho \gamma_\mu \psi_\nu + 2 \bar{\psi}_\rho \gamma_\nu \psi_\mu) - 4 (\bar{\psi}_\mu \gamma \cdot \psi) (\bar{\psi}^\mu \gamma \cdot \psi) \right].\end{aligned}\quad (\text{A.10})$$

We now choose the physical scalars and fermions

$$z^\alpha = X^\alpha, \quad \chi^\alpha = \Omega^\alpha. \quad (\text{A.11})$$

After elimination of the auxiliary fields  $F^\alpha$  and  $D^A$ , the conformal part of the action becomes

$$\begin{aligned}e^{-1} \mathcal{L}_{\text{conf}} &= \delta_{\alpha\bar{\beta}} \left\{ -D_\mu z^\alpha D^\mu \bar{z}^{\bar{\beta}} - \frac{1}{2} \bar{\chi}^\alpha \not{D} \chi^{\bar{\beta}} - \frac{1}{2} \bar{\chi}^{\bar{\beta}} \not{D} \chi^\alpha - F^\alpha \bar{F}^{\bar{\beta}} \right. \\ &\quad + \frac{1}{6} z^\alpha \bar{z}^{\bar{\beta}} \left[ -R(\omega(e)) + \bar{\psi}_\mu R^\mu + e^{-1} \partial_\mu (e \bar{\psi} \cdot \gamma \psi^\mu) - \mathcal{L}_{\text{SG,torsion}} \right] \\ &\quad + \left[ \frac{1}{8} i \varepsilon^{\mu\nu\rho\sigma} \bar{\psi}_\mu \gamma_\nu \psi_\rho \bar{z}^{\bar{\beta}} D_\sigma z^\alpha \right. \\ &\quad + \frac{1}{\sqrt{2}} \bar{\psi}_\mu \not{D} z^{\bar{\beta}} \gamma^\mu \chi^\alpha - \frac{2}{3\sqrt{2}} \bar{z}^{\bar{\beta}} \bar{\chi}^\alpha \gamma^{\mu\nu} \left( \partial_\mu + \frac{1}{4} \omega_\mu^{ab}(e, \psi) \gamma_{ab} - \frac{3}{2} i A_\mu \gamma_* \right) \psi_\nu \\ &\quad + \left( -\sqrt{2} \bar{\chi}^{\bar{\beta}} - \frac{1}{2} \bar{\psi} \cdot \gamma P_L z^{\bar{\beta}} \right) \lambda^A (m_A)^\alpha{}_\gamma z^\gamma + \text{h.c.} \left. \right\} \\ &\quad + \frac{1}{16} i e^{-1} \varepsilon^{\mu\nu\rho\sigma} \bar{\psi}_\mu \gamma_\nu \psi_\rho \bar{\chi}^{\bar{\beta}} \gamma_\sigma \chi^\alpha - \frac{1}{2} \bar{\psi}_\mu \chi^{\bar{\beta}} \bar{\psi}_\mu \chi^\alpha \left. \right\} \\ + (\text{Re } f_{AB}) &\left\{ -\frac{1}{4} F_{\mu\nu}^A F^{\mu\nu B} - \frac{1}{2} \bar{\lambda}^A \not{D} \lambda^B - \frac{1}{2} D^A D^B \right. \\ &\quad + \frac{1}{8} \bar{\psi}_\mu \gamma^{ab} \left( F_{ab}^A + \hat{F}_{ab}^A \right) \gamma^\mu \lambda^B + \frac{1}{32} i e^{-1} \varepsilon^{\mu\nu\rho\sigma} \bar{\psi}_\mu \gamma_\nu \psi_\rho \bar{\lambda}^A \gamma_* \gamma_\sigma \lambda^B \left. \right\} \\ + \{ d_{\alpha\beta\gamma} &\left[ -z^\alpha \bar{\chi}^\beta \chi^\gamma + z^\alpha z^\beta \bar{\psi} \cdot \gamma \chi^\gamma + \frac{1}{6} z^\alpha z^\beta z^\gamma \bar{\psi}_\mu P_R \gamma^{\mu\nu} \psi_\nu \right] + \text{h.c.} \}.\end{aligned}\quad (\text{A.12})$$

The covariant derivatives  $D_\mu$  have no torsion in the spin-connection, neither supersymmetric covariantization. The same is true for  $F_{ab}$  and  $R^\mu$ :

$$\begin{aligned}D_\mu z^\alpha &= (\partial_\mu - i A_\mu) z^\alpha - A_\mu^A m_A{}^\alpha{}_\beta z^\beta, \\ D_\mu \chi^\alpha &= \left( \partial_\mu + \frac{1}{4} \omega_\mu^{ab}(e) \gamma_{ab} + \frac{1}{2} i A_\mu \right) \chi^\alpha - A_\mu^A m_A{}^\alpha{}_\beta \chi^\beta, \\ D_\mu \lambda^A &= \left( \partial_\mu + \frac{1}{4} \omega_\mu^{ab}(e) \gamma_{ab} - \frac{3}{2} i A_\mu \gamma_* \right) \lambda^A - A_\mu^C \lambda^B f_{BC}^A, \\ F_{\mu\nu}^A &= 2 \partial_{[\mu} A_{\nu]}^A + g f_{BC}^A A_\mu^B A_\nu^C, \\ R^\mu &\equiv \gamma^{\mu\rho\sigma} \left( \partial_\rho + \frac{1}{4} \omega_\rho^{ab}(e) \gamma_{ab} - \frac{3}{2} i A_\rho \gamma_* \right) \psi_\sigma.\end{aligned}\quad (\text{A.13})$$

The auxiliary fields  $A_\mu$  in (A.12) are to be considered as independent fields, which should still be solved for by their field equations. The latter will mix the supergravity part and the superconformal part of the action. The fields  $F^\alpha$  and  $D^A$  on the other hand are to be considered as their expressions

in terms of the other fields:

$$\begin{aligned}\bar{F}^{\bar{\beta}} &= -\delta^{\alpha\bar{\beta}} d_{\alpha\beta\gamma} z^{\beta} z^{\gamma}, \\ D^A &= (\text{Re } f)^{-1AB} \mathcal{P}_A = (\text{Re } f)^{-1AB} i\delta_{\alpha\bar{\beta}} z^{\bar{\beta}} (m_A)^{\alpha}{}_{\gamma} z^{\gamma}.\end{aligned}\tag{A.14}$$

This does not mix the supergravity and the superconformal part. Thus, equations (A.10)-(A.14) provide the generalization of eqs. (3.25), (3.26) when all fermions and vectors are included.

## B Why is the supergravity potential in CSS Jordan frame the same as in global SUSY?

Starting from the superconformal theory potential in (3.4) we have already derived the potential of the CSS in (3.26). The main reason from that point of view is that the modifications to the global SUSY potential originate from the compensating multiplet, containing the scalar  $(y\bar{y})^3 = e^{\mathcal{K}}$ , and the auxiliary field  $F^0$  producing the term  $-3|W|^2$ . This compensating multiplet has been decoupled in CSS. The fact that for the CSS models the supergravity potential is the same as in globally supersymmetric models with canonical kinetic terms is somewhat surprising, from the point of view of the complicated Einstein frame F-term potential in generic supergravity theory

$$V_E = e^{\mathcal{K}} \left( \nabla_{\alpha} W g^{\alpha\bar{\beta}} \nabla_{\bar{\beta}} \bar{W} - 3W\bar{W} \right), \tag{B.1}$$

where

$$\nabla_{\alpha} W \equiv W_{\alpha} + K_{\alpha} W. \tag{B.2}$$

It is therefore instructive to see directly how the cancelation of various terms in the F-term potential takes place, leading to a simple CSS Jordan frame potential.

We define the Jordan frame for the CSS via the frame function  $\Phi(z, \bar{z}) = -3\Omega^2$  related to the Kähler potential  $\mathcal{K}(z, \bar{z}) = -3\log \Omega^2$ . The metric in the Einstein frame is related to the metric in the Jordan frame as  $g_{\mu\nu}^E = \Omega^2 g_{\mu\nu}^J$  and  $\sqrt{g^E} = \Omega^4 \sqrt{g^J}$ . The F-term potential in the Jordan frame specified by the frame function (B.4) is related to the Einstein frame potential as

$$V_J = \Omega^4 V_E = \Omega^4 e^{\mathcal{K}} \left( \nabla_{\alpha} W g^{\alpha\bar{\beta}} \nabla_{\bar{\beta}} \bar{W} - 3W\bar{W} \right). \tag{B.3}$$

We take into account that in CSS

$$\Omega^4 e^{\mathcal{K}} = \Omega^{-2}, \tag{B.4}$$



which means that

$$V_J = \Omega^{-2} \left( \nabla_\alpha W g^{\alpha\bar{\beta}} \nabla_{\bar{\beta}} \bar{W} - 3W\bar{W} \right). \quad (\text{B.5})$$

In these models we have the following Kähler potential and generic cubic superpotential:

$$\mathcal{K}(z, \bar{z}) = -3 \log \left( 1 - \frac{1}{3} \delta_{\alpha\bar{\beta}} z^\alpha \bar{z}^\beta \right), \quad W(z) = \frac{1}{3} d_{\alpha\beta\gamma} z^\alpha z^\beta z^\gamma, \quad (\text{B.6})$$

and

$$\Omega^2 = 1 - \frac{1}{3} \delta_{\alpha\bar{\beta}} z^\alpha \bar{z}^\beta. \quad (\text{B.7})$$

It follows that the Kähler geometry with  $g_{\alpha\bar{\beta}} g^{\alpha\bar{\gamma}} = \delta_{\bar{\beta}}^{\bar{\gamma}}$  has the following properties:

$$\mathcal{K}_\alpha = e^{\frac{1}{3}\mathcal{K}} \delta_{\alpha\bar{\beta}} \bar{z}^\beta, \quad g^{\alpha\bar{\beta}} = e^{-\frac{1}{3}\mathcal{K}} \left( \delta^{\alpha\bar{\beta}} - \frac{1}{3} z^\alpha \bar{z}^\beta \right) = \Omega^2 \left( \delta^{\alpha\bar{\beta}} - \frac{1}{3} z^\alpha \bar{z}^\beta \right), \quad (\text{B.8})$$

so that

$$V_J = \Omega^{-2} g^{\alpha\bar{\beta}} (W_\alpha + W \mathcal{K}_\alpha) \left( \bar{W}_{\bar{\beta}} + \bar{W} \mathcal{K}_{\bar{\beta}} \right) - 3\Omega^{-2} |W|^2. \quad (\text{B.9})$$

With an account of the CSS Kähler geometry properties we may rewrite the potential as follows:

$$V_J = \delta^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}} - 3|W|^2 + \bar{W} W_\alpha z^\alpha + W \bar{W}_{\bar{\alpha}} \bar{z}^\alpha - \frac{1}{3} |W_\alpha z^\alpha|^2. \quad (\text{B.10})$$

Note that all  $\Omega^{-2}$  factors in (B.10) have cancelled. It remains to take into account that the CSS superpotential  $W$  is homogeneous of third degree in  $z^\alpha$ 's, it follows that

$$W_\alpha z^\alpha = 3W, \quad \bar{W}_{\bar{\alpha}} \bar{z}^\alpha = 3\bar{W}. \quad (\text{B.11})$$

This allows to bring the F-term potential to the final form

$$V_J = \delta^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}}, \quad (\text{B.12})$$

where it is clear that it coincides with the global supersymmetric F-term potential.

## C The moduli space geometry in CSS models with symmetry breaking $\chi$ -terms

In CSS models the moduli space geometry is flat as shown in (3.9). When superconformal symmetry is broken by the  $\chi$ -terms of the form given in (3.34), the Jordan frame potential depends on  $G^{\alpha\bar{\beta}}$  according to Eq. (3.36). Here we study the non-flat geometry for the models in (3.34) and, in particular, we compute  $G^{\alpha\bar{\beta}}$ . We start from Eq. (3.34) which we repeat here for convenience

$$\mathcal{N}(X, \bar{X}) = -|X^0|^2 + |X^\alpha|^2 - \chi \left( a_{\alpha\beta} \frac{X^\alpha X^\beta \bar{X}^{\bar{0}}}{X^0} + \bar{a}_{\bar{\alpha}\bar{\beta}} \frac{\bar{X}^{\bar{\alpha}} \bar{X}^{\bar{\beta}} X^0}{\bar{X}^{\bar{0}}} \right). \quad (\text{C.1})$$

The metric

$$G_{I\bar{J}} \equiv \frac{\partial^2 \mathcal{N}}{\partial X^I \partial \bar{X}^{\bar{J}}} \quad (\text{C.2})$$

can be computed to read

$$G_{0\bar{0}} = -1 + \chi \left[ a_{\alpha\beta} \frac{X^\alpha X^\beta}{(X^0)^2} + \bar{a}_{\bar{\alpha}\bar{\beta}} \frac{\bar{X}^{\bar{\alpha}} \bar{X}^{\bar{\beta}}}{(\bar{X}^0)^2} \right]; \quad (\text{C.3})$$

$$G_{0\bar{\beta}} = -2\chi \bar{a}_{\bar{\beta}\bar{\gamma}} \frac{\bar{X}^{\bar{\gamma}}}{\bar{X}^0}; \quad (\text{C.4})$$

$$G_{\alpha\bar{0}} = -2\chi a_{\alpha\gamma} \frac{X^\gamma}{X^0}; \quad (\text{C.5})$$

$$G_{\alpha\bar{\beta}} = \delta_{\alpha\bar{\beta}}. \quad (\text{C.6})$$

The components of the inverse metric  $G^{I\bar{J}}$  (such that  $G^{I\bar{J}} G_{I\bar{K}} = \delta_{\bar{K}}^{\bar{J}}$ ) can be computed to read

$$G^{0\bar{0}} = - \frac{(X^0 \bar{X}^0)^2}{\left[ (X^0 \bar{X}^0)^2 - \chi a_{\gamma\eta} X^\gamma X^\eta (\bar{X}^0)^2 - \chi \bar{a}_{\bar{\gamma}\bar{\eta}} \bar{X}^{\bar{\gamma}} \bar{X}^{\bar{\eta}} (X^0)^2 + 4\chi^2 X^0 \bar{X}^0 \delta^{\gamma\bar{\eta}} a_{\gamma\theta} \bar{a}_{\bar{\eta}\bar{\rho}} X^\theta \bar{X}^{\bar{\rho}} \right]}; \quad (\text{C.7})$$

$$G^{0\bar{\beta}} = - \frac{2\chi \delta^{\lambda\bar{\beta}} a_{\lambda\xi} X^\xi (\bar{X}^0)^2 X^0}{\left[ (X^0 \bar{X}^0)^2 - \chi a_{\gamma\eta} X^\gamma X^\eta (\bar{X}^0)^2 - \chi \bar{a}_{\bar{\gamma}\bar{\eta}} \bar{X}^{\bar{\gamma}} \bar{X}^{\bar{\eta}} (X^0)^2 + 4\chi^2 X^0 \bar{X}^0 \delta^{\gamma\bar{\eta}} a_{\gamma\theta} \bar{a}_{\bar{\eta}\bar{\rho}} X^\theta \bar{X}^{\bar{\rho}} \right]}; \quad (\text{C.8})$$

$$G^{\alpha\bar{0}} = - \frac{2\chi \delta^{\alpha\bar{\lambda}} \bar{a}_{\bar{\lambda}\bar{\xi}} \bar{X}^{\bar{\xi}} (X^0)^2 \bar{X}^0}{\left[ (X^0 \bar{X}^0)^2 - \chi a_{\gamma\eta} X^\gamma X^\eta (\bar{X}^0)^2 - \chi \bar{a}_{\bar{\gamma}\bar{\eta}} \bar{X}^{\bar{\gamma}} \bar{X}^{\bar{\eta}} (X^0)^2 + 4\chi^2 X^0 \bar{X}^0 \delta^{\gamma\bar{\eta}} a_{\gamma\theta} \bar{a}_{\bar{\eta}\bar{\rho}} X^\theta \bar{X}^{\bar{\rho}} \right]}; \quad (\text{C.9})$$

$$G^{\alpha\bar{\beta}} = \delta^{\alpha\bar{\beta}} - \frac{4\chi^2 X^0 \bar{X}^0 \delta^{\alpha\bar{\lambda}} \delta^{\sigma\bar{\beta}} a_{\sigma\zeta} \bar{a}_{\bar{\lambda}\bar{\xi}} X^\zeta \bar{X}^{\bar{\xi}}}{\left[ (X^0 \bar{X}^0)^2 - \chi a_{\gamma\eta} X^\gamma X^\eta (\bar{X}^0)^2 - \chi \bar{a}_{\bar{\gamma}\bar{\eta}} \bar{X}^{\bar{\gamma}} \bar{X}^{\bar{\eta}} (X^0)^2 + 4\chi^2 X^0 \bar{X}^0 \delta^{\gamma\bar{\eta}} a_{\gamma\theta} \bar{a}_{\bar{\eta}\bar{\rho}} X^\theta \bar{X}^{\bar{\rho}} \right]}. \quad (\text{C.10})$$

By performing the gauge-fixing

$$X^0 = \bar{X}^0 = \sqrt{3} M_P; \quad (\text{C.11})$$

$$X^\alpha = y Z^\alpha(z); \quad (\text{C.12})$$

$$y = \bar{y} = 1; \quad (\text{C.13})$$

$$Z^\alpha = z^\alpha, \quad (\text{C.14})$$

one then respectively obtains:

$$G_{0\bar{0}} = -1 + \frac{\chi}{3M_P^2} \left( a_{\alpha\beta} z^\alpha z^\beta + \bar{a}_{\bar{\alpha}\bar{\beta}} \bar{z}^{\bar{\alpha}} \bar{z}^{\bar{\beta}} \right); \quad (\text{C.15})$$

$$G_{0\bar{\beta}} = -\frac{2}{\sqrt{3}} \frac{\chi}{M_P} \bar{a}_{\bar{\beta}\bar{\gamma}} \bar{z}^{\bar{\gamma}}; \quad (\text{C.16})$$

$$G_{\alpha\bar{0}} = -\frac{2}{\sqrt{3}} \frac{\chi}{M_P} a_{\alpha\gamma} z^\gamma; \quad (\text{C.17})$$

$$G_{\alpha\bar{\beta}} = \delta_{\alpha\bar{\beta}}; \quad (\text{C.18})$$

$$G^{0\bar{0}} = -\frac{1}{\left[ 1 - \frac{\chi}{3M_P^2} (a_{\gamma\eta} z^\gamma z^\eta + \bar{a}_{\bar{\gamma}\bar{\eta}} \bar{z}^{\bar{\gamma}} \bar{z}^{\bar{\eta}}) + \frac{4}{3} \frac{\chi^2}{M_P^2} \delta^{\gamma\bar{\eta}} a_{\gamma\theta} \bar{a}_{\bar{\eta}\bar{\rho}} z^\theta \bar{z}^{\bar{\rho}} \right]}; \quad (\text{C.19})$$

$$G^{0\bar{\beta}} = -\frac{2\sqrt{3}M_P\chi\delta^{\lambda\bar{\beta}}a_{\lambda\xi}z^\xi}{\left[ 3M_P^2 - \chi (a_{\gamma\eta} z^\gamma z^\eta + \bar{a}_{\bar{\gamma}\bar{\eta}} \bar{z}^{\bar{\gamma}} \bar{z}^{\bar{\eta}}) + 4\chi^2 \delta^{\gamma\bar{\eta}} a_{\gamma\theta} \bar{a}_{\bar{\eta}\bar{\rho}} z^\theta \bar{z}^{\bar{\rho}} \right]}; \quad (\text{C.20})$$

$$G^{\alpha\bar{0}} = -\frac{2\sqrt{3}M_P\chi\delta^{\alpha\bar{\lambda}}\bar{a}_{\bar{\lambda}\bar{\xi}}\bar{z}^{\bar{\xi}}}{\left[ 3M_P^2 - \chi (a_{\gamma\eta} z^\gamma z^\eta + \bar{a}_{\bar{\gamma}\bar{\eta}} \bar{z}^{\bar{\gamma}} \bar{z}^{\bar{\eta}}) + 4\chi^2 \delta^{\gamma\bar{\eta}} a_{\gamma\theta} \bar{a}_{\bar{\eta}\bar{\rho}} z^\theta \bar{z}^{\bar{\rho}} \right]}; \quad (\text{C.21})$$

$$G^{\alpha\bar{\beta}} = \delta^{\alpha\bar{\beta}} - \frac{4\chi^2 \delta^{\alpha\bar{\lambda}} \delta^{\sigma\bar{\beta}} a_{\sigma\zeta} \bar{a}_{\bar{\lambda}\bar{\xi}} z^\zeta \bar{z}^{\bar{\xi}}}{\left[ 3M_P^2 - \chi (a_{\gamma\eta} z^\gamma z^\eta + \bar{a}_{\bar{\gamma}\bar{\eta}} \bar{z}^{\bar{\gamma}} \bar{z}^{\bar{\eta}}) + 4\chi^2 \delta^{\gamma\bar{\eta}} a_{\gamma\theta} \bar{a}_{\bar{\eta}\bar{\rho}} z^\theta \bar{z}^{\bar{\rho}} \right]}. \quad (\text{C.22})$$

In particular, the metric  $G^{\alpha\bar{\beta}}$  given by Eq. (C.22) is the metric appearing in Eq. (3.36). Notice that clearly  $G^{\alpha\bar{\beta}}$  given by (C.22) is not the inverse of  $G_{\alpha\bar{\beta}}$  given by (C.18), because what really holds is

$$G^{0\bar{\beta}} G_{0\bar{\gamma}} + G^{\alpha\bar{\beta}} G_{\alpha\bar{\gamma}} = \delta_{\bar{\gamma}}^{\bar{\beta}}. \quad (\text{C.23})$$

## D The $h$ and $s$ part of the NMSSM potential in the Jordan and the Einstein frame

Here we present some details of the scalar-gravity part of the supergravity action given in Eq. (2.8). We apply it to the frame function (7.7) and we consider  $\beta = \pi/4, \alpha = \alpha_1 = \alpha_2 = 0$ . The kinetic scalar terms in the Jordan frame are canonical, except for the contribution to the gauge singlet one due to  $\zeta(S\bar{S})^2$  corrections to the NMSSM frame function:

$$\mathcal{L}_J^{\text{kinetic}} = -\frac{\sqrt{-g_J}}{2} \left[ (1 - 2\zeta s^2)(\partial_\mu s)^2 + (\partial_\mu h_1)^2 + (\partial_\mu h_2)^2 \right], \quad (\text{D.1})$$

where  $h_1 \equiv h \cos \beta, h_2 \equiv h \sin \beta$ . Note that the  $\zeta$ -correction to the kinetic term of the  $s$ -field is always small comparative to 1 and can be safely neglected. Along the  $D$ -flat direction with  $\sin(2\beta) = 1$  the

curvature term in the action for real fields  $h$  and  $s$  is

$$\mathcal{L}_J^{\text{curv}} = \frac{\sqrt{-g_J}}{2} \left[ 1 - \frac{1}{6} (s^2 + h^2) + \frac{1}{4} \chi h^2 + \frac{1}{12} \zeta s^4 \right] R(g_J). \quad (\text{D.2})$$

The potential in the Jordan frame for the non-vanishing  $\chi$  and  $\zeta$  and non-vanishing field  $s$  is complicated:

$$V_J(s, h; \chi, \lambda, \rho, \zeta) = \frac{A_0 + s^2 A_2 + s^4 A_4 + s^6 A_6 + s^8 A_8}{1 - 2\zeta s^2 + \frac{1}{3}\zeta G s^4}. \quad (\text{D.3})$$

where we introduced the following notation

$$\begin{aligned} A_0 &\equiv \frac{\lambda^2}{16} h^4, \quad A_2 \equiv -\frac{h^2}{4} (|\lambda\rho| + G\lambda^2(\chi h^2 - 4)), \quad A_4 \equiv \frac{\rho^2}{4} - \frac{\zeta G h^2 \lambda^2}{8} (32 - (12\chi + 1/3) h^2), \\ A_6 &\equiv \frac{1}{12} \zeta G (\lambda^2 - |\lambda\rho| - 6\chi |\lambda\rho|) h^2, \quad A_8 \equiv \frac{1}{12} \zeta G \rho^2, \quad G \equiv \frac{2}{8 + (3\chi - 2) \chi h^2}. \end{aligned} \quad (\text{D.4})$$

At  $\zeta = 0$  the potential simplifies to the form given by  $V_J = G^{\alpha\bar{\beta}} W_\alpha \bar{W}_{\bar{\beta}}$  when only  $\chi$ -term breaks superconformal symmetry of the matter

$$V_J(s, h; \chi, \lambda, \rho, \zeta = 0) = \frac{\lambda^2}{4} h^4 - \frac{h^2}{4} (|\lambda\rho| + 2G\lambda^2(\chi h^2 - 2)) s^2 + \frac{\rho^2}{4} s^4. \quad (\text{D.5})$$

At  $s = 0$  the Jordan potential  $V_J|_{s=0} = G^{S\bar{S}} W_S \bar{W}_{\bar{S}}$  restores the superconformal form at any values of  $\chi, \lambda, \rho, \zeta$ .

$$V_J(s = 0, h; \chi, \lambda, \rho, \zeta) = \frac{\lambda^2}{4} h^4. \quad (\text{D.6})$$

In the Einstein frame for the real fields, the expression of the F-term potential is

$$V_E(s, h; \chi, \lambda, \rho, \zeta) = \frac{9}{\Phi^2} V_J = \frac{A_0 + s^2 A_2 + s^4 A_4 + s^6 A_6 + s^8 A_8}{\left[ 1 + \frac{1}{4} \chi h^2 - \frac{1}{6} (s^2 + h^2) + \frac{1}{12} \zeta s^4 \right]^2 (1 - 2\zeta s^2 + \frac{1}{3} \zeta G s^4)}, \quad (\text{D.7})$$

where all notations are given above. For  $\zeta = 0$  the potential in (E.2) reduces to the one studied in [1]

$$V_E(s, h; \chi, \lambda, \rho, \zeta = 0) = \frac{\frac{\lambda^2}{4} h^4 - |\lambda\rho| s^2 h^2 - \frac{2\lambda^2 s^2 h^2 (\chi h^2 - 4)}{8 + 3\chi^2 h^2 - 2\chi h^2} + \rho^2 |s|^4}{4 \left[ 1 - \frac{1}{6} (s^2 + h^2) + \frac{1}{4} \chi h^2 \right]^2}. \quad (\text{D.8})$$

From this equation it is obvious that at large  $\chi h^2$  the mass squared of the  $s$  field is negative at  $\zeta = 0$ , as explained in detail in [1]. The crucial positive contribution to the gauge singlet mass term comes from the negative term  $-2\zeta s^2$  in the second bracket in the denominator of (E.2), where  $\frac{1}{3}\zeta(S\bar{S})^2$  is the quartic correction to the Kähler potential suggested in [4] and also studied before in [25].

## E The mass of the charged Higgs field

The mass of the physical charged Higgs field in globally supersymmetric NMSSM is *e.g.* given in [14] in Eq. (2.29). In the absence of the soft breaking terms the F-term and the D-term contribution to mass (in our notation) is

$$(m_{\pm}^2)_{susy} = \frac{1}{8}(-\lambda^2 h^2 + \frac{g_2^2}{2} h^2) . \quad (\text{E.1})$$

For  $g_2^2 > 2\lambda^2$ , the charged Higgs field is not tachyonic, i.e. it is stabilized at  $h_{\pm} = 0$ .

In supergravity we look first at the  $s = 0$  expression of the  $F$ -term potential given by Eq. (D.8):

$$V_E^F(s = 0, h; \chi, \lambda) = \frac{\frac{\lambda^2}{16} h^4}{\left[1 + \frac{1}{4}\chi h^2 - \frac{1}{6}h^2\right]^2} . \quad (\text{E.2})$$

As explained in Sec. 7, we can extend this potential using the  $SU(2)$  symmetry to include the charged Higgs, so that  $h^2 \rightarrow h^2 - h_{\pm}^2$  in the part of the potential associated with the holomorphic functions, *i.e.* in the terms originating from the superpotential and from the  $\chi$ -terms in the Kähler potential. However, one has to replace  $h^2 \rightarrow h^2 + h_{\pm}^2$  in the  $\chi$ -independent part of the Kähler potential. This leads to the following potential

$$V_E^F(s = 0, h; h_{\pm}, \chi, \lambda) = \frac{\frac{\lambda^2}{16} (h^2 - h_{\pm}^2)^2}{\left[1 + \frac{1}{4}\chi(h^2 - h_{\pm}^2) - \frac{1}{6}(h^2 + h_{\pm}^2)\right]^2} . \quad (\text{E.3})$$

On the other hand, the D-term potential depends on charged Higgs field as follows

$$V_E^D(h; h_{\pm}, g, \chi) = \frac{g^2 h^2 h_{\pm}^2}{16 \left[1 + \frac{1}{4}\chi(h^2 - h_{\pm}^2) - \frac{1}{6}(h^2 + h_{\pm}^2)\right]^2} . \quad (\text{E.4})$$

One can easily see that *after* inflation, for  $\chi h^2 \ll 1$ , the mass squared of the charged Higgs field  $m_{\pm}^2$  coincides with its value in the globally supersymmetric case, and therefore the stability condition requires that  $g_2^2 > 2\lambda^2$ . *During* inflation, one has  $h^2 \ll 1 \ll \chi h^2$ , and the second derivatives of (E.3) and (E.4) respectively read

$$V_{\pm}''^F \sim -\frac{16\lambda^2}{\chi^3 h^4} , \quad V_{\pm}''^D \sim \frac{g_2^2}{\chi^2 h^2} . \quad (\text{E.5})$$

In this regime, for  $g_2^2 > 2\lambda^2$ , the D-term contribution to  $m_{\pm}^2$  is much greater than the F-term contribution. In order to calculate  $m_{\pm}^2$ , one should also take into account that the kinetic terms for the fields  $h_{\pm}$  are not canonical. By so doing, one finds that during inflation

$$m_{\pm}^2 \sim \frac{g_2^2}{2\chi} \gg H^2 = \frac{\lambda^2}{3\chi^2} . \quad (\text{E.6})$$

This means that this field is strongly stabilized at  $h_{\pm} = 0$ .

In other words, under the condition  $g_2^2 > 2\lambda^2$  the charged Higgs field vanishes *during* and *after* inflation.

## References

- [1] S. Ferrara, R. Kallosh, A. Linde, A. Marrani and A. Van Proeyen, “Jordan Frame Supergravity and Inflation in NMSSM,” *Phys. Rev. D* **82**, 045003 (2010) [arXiv:1004.0712 [hep-th]].
- [2] F. L. Bezrukov and M. Shaposhnikov, “The Standard Model Higgs boson as the inflaton,” *Phys. Lett. B* **659**, 703 (2008) [arXiv:0710.3755 [hep-th]]. F. Bezrukov, D. Gorbunov and M. Shaposhnikov, “On initial conditions for the Hot Big Bang,” *JCAP* **0906**, 029 (2009) [arXiv:0812.3622 [hep-ph]]. J. Garcia-Bellido, D. G. Figueroa and J. Rubio, “Preheating in the Standard Model with the Higgs-Inflaton coupled to gravity,” *Phys. Rev. D* **79**, 063531 (2009) [arXiv:0812.4624 [hep-ph]]. A. De Simone, M. P. Hertzberg and F. Wilczek, “Running Inflation in the Standard Model,” *Phys. Lett. B* **678**, 1 (2009) [arXiv:0812.4946 [hep-ph]]. F. Bezrukov and M. Shaposhnikov, “Standard Model Higgs boson mass from inflation: two loop analysis,” *JHEP* **0907**, 089 (2009) [arXiv:0904.1537 [hep-ph]]. A. O. Barvinsky, A. Y. Kamenshchik, C. Kiefer, A. A. Starobinsky and C. Steinwachs, “Asymptotic freedom in inflationary cosmology with a non-minimally coupled Higgs field,” *JCAP* **0912**, 003 (2009) [arXiv:0904.1698 [hep-ph]].
- [3] M. B. Einhorn and D. R. T. Jones, “Inflation with Non-minimal Gravitational Couplings in Supergravity,” *JHEP* **1003**, 026 (2010) [arXiv:0912.2718 [hep-ph]].
- [4] H. M. Lee, “Chaotic inflation in Jordan frame supergravity,” *JCAP* **1008**, 003 (2010) [arXiv:1005.2735 [hep-ph]].
- [5] N. D. Birrell and P. C. W. Davies, *Quantum Fields in Curved Space* (Cambridge University Press, 1982); A.A. Grib, S.G. Mamayev and V.M. Mostepanenko, *Vacuum quantum effects in strong fields* (Moscow, Energoatomizdat, 1988 (in Russian) St. Petersburg, Friedmann Lab. Publ., 1994 (in English)).
- [6] D. S. Salopek, J. R. Bond and J. M. Bardeen, “Designing Density Fluctuation Spectra in Inflation,” *Phys. Rev. D* **40**, 1753 (1989). T. Futamase and K. i. Maeda, “Chaotic inflationary scenario in models having nonminimal coupling with curvature,” *Phys. Rev. D* **39**, 399 (1989).
- [7] E. Cremmer, B. Julia, J. Scherk, S. Ferrara, L. Girardello and P. van Nieuwenhuizen, “Spontaneous Symmetry Breaking And Higgs Effect In Supergravity Without Cosmological Constant,” *Nucl. Phys. B* **147**, 105 (1979).
- [8] R. Barbieri, S. Ferrara, D. V. Nanopoulos and K. S. Stelle, “Supergravity, R Invariance And Spontaneous Supersymmetry Breaking,” *Phys. Lett. B* **113**, 219 (1982).

- [9] E. Cremmer, S. Ferrara, L. Girardello and A. Van Proeyen, “Yang-Mills Theories With Local Supersymmetry: Lagrangian, Transformation Laws And Superhiggs Effect,” Nucl. Phys. B **212**, 413 (1983).
- [10] G. Girardi, R. Grimm, M. Muller and J. Wess, “Superspace Geometry And The Minimal, Non-minimal, And New Minimal Supergravity Multiplets,” Z. Phys. C **26**, 123 (1984).
- [11] J. Wess and J. Bagger, *Supersymmetry and supergravity*, (Princeton University Press, Princeton, 1992).
- [12] S. Weinberg, *The quantum theory of fields. Vol. 3: Supersymmetry*, (Cambridge, UK: Univ. Pr., 2000).
- [13] R. Kallosh, L. Kofman, A. D. Linde and A. Van Proeyen, “Superconformal symmetry, supergravity and cosmology,” Class. Quant. Grav. **17**, 4269 (2000) [Erratum-ibid. **21**, 5017 (2004)] [arXiv:hep-th/0006179].
- [14] U. Ellwanger, C. Hugonie and A. M. Teixeira, “The Next-to-Minimal Supersymmetric Standard Model,” arXiv:0910.1785 [hep-ph].
- [15] C. P. Burgess, H. M. Lee and M. Trott, “Power-counting and the Validity of the Classical Approximation During Inflation,” JHEP **0909**, 103 (2009) [arXiv:0902.4465 [hep-ph]]. C. P. Burgess, H. M. Lee and M. Trott, “Comment on Higgs Inflation and Naturalness,” arXiv:1002.2730 [hep-ph].
- [16] J. L. F. Barbon and J. R. Espinosa, “On the Naturalness of Higgs Inflation,” Phys. Rev. D **79**, 081302 (2009) [arXiv:0903.0355 [hep-ph]].
- [17] M. P. Hertzberg, “On Inflation with nonminimal Coupling,” arXiv:1002.2995 [hep-ph].
- [18] C. Germani and A. Kehagias, “New Model of Inflation with Non-minimal Derivative Coupling of Standard Model Higgs Boson to Gravity,” Phys. Rev. Lett. **105**, 011302 (2010) [arXiv:1003.2635 [hep-ph]].
- [19] F. Bezrukov, A. Magnin, M. Shaposhnikov and S. Sibiryakov, “Higgs inflation: consistency and generalisations,” arXiv:1008.5157 [hep-ph].
- [20] G. F. Giudice and H. M. Lee, “Unitarizing Higgs Inflation,” arXiv:1010.1417 [hep-ph].
- [21] R. Kallosh and A. Linde, “New models of chaotic inflation in supergravity,” arXiv:1008.3375 [hep-th].

- [22] M. Kaku, P. K. Townsend and P. van Nieuwenhuizen, “Gauge Theory Of The Conformal And Superconformal Group,” *Phys. Lett. B* **69**, 304 (1977); S. Ferrara, M. Kaku, P. K. Townsend and P. van Nieuwenhuizen, “Gauging The Graded Conformal Group With Unitary Internal Symmetries,” *Nucl. Phys. B* **129**, 125 (1977); M. Kaku and P. K. Townsend, “Poincaré Supergravity As Broken Superconformal Gravity,” *Phys. Lett. B* **76**, 54 (1978). M. Kaku, P. K. Townsend and P. van Nieuwenhuizen, “Properties Of Conformal Supergravity,” *Phys. Rev. D* **17**, 3179 (1978).
- [23] S. Ferrara, “Supersymmetry,” North-Holland, World Scientific (1987) 1343p.
- [24] S. Kachru, R. Kallosh, A. D. Linde and S. P. Trivedi, “De Sitter vacua in string theory,” *Phys. Rev. D* **68**, 046005 (2003) [arXiv:hep-th/0301240].
- [25] M. Yamaguchi and J. Yokoyama, “New inflation in supergravity with a chaotic initial condition,” *Phys. Rev. D* **63**, 043506 (2001) [arXiv:hep-ph/0007021]. R. Kitano, “Gravitational gauge mediation,” *Phys. Lett. B* **641**, 203 (2006) [arXiv:hep-ph/0607090]. R. Kallosh and A. D. Linde, “O’KKLT,” *JHEP* **0702**, 002 (2007) [arXiv:hep-th/0611183].
- [26] F. Franke and H. Fraas, “Neutralinos and Higgs bosons in the next-to-minimal supersymmetric standard model,” *Int. J. Mod. Phys. A* **12**, 479 (1997) [arXiv:hep-ph/9512366].
- [27] S. A. Abel, S. Sarkar and P. L. White, “On the Cosmological Domain Wall Problem for the Minimally Extended Supersymmetric Standard Model,” *Nucl. Phys. B* **454**, 663 (1995) [arXiv:hep-ph/9506359].
- [28] R. Barbieri, S. Ferrara and C. A. Savoy, “Gauge Models With Spontaneously Broken Local Supersymmetry,” *Phys. Lett. B* **119**, 343 (1982).
- [29] G. F. Giudice and A. Masiero, “A Natural Solution to the  $\mu$  Problem in Supergravity Theories,” *Phys. Lett. B* **206**, 480 (1988).
- [30] S. Shirai, F. Takahashi, T. T. Yanagida and K. Yonekura, “Gauge Mediation with Sequestered Supersymmetry Breaking,” *Phys. Rev. D* **78**, 075003 (2008) [arXiv:0808.0848 [hep-ph]].
- [31] H. P. Nilles, M. Srednicki and D. Wyler, “Constraints On The Stability Of Mass Hierarchies In Supergravity,” *Phys. Lett. B* **124**, 337 (1983).
- [32] A. B. Lahanas, “Light Singlet, Gauge Hierarchy And Supergravity,” *Phys. Lett. B* **124**, 341 (1983).
- [33] S. Ferrara, D. V. Nanopoulos and C. A. Savoy, “Hierarchical Supergravity Induced  $SU(2) \times U(1)$  Breaking In  $SU(5)$  Guts,” *Phys. Lett. B* **123**, 214 (1983).



- [34] S. A. Abel, “Destabilising divergences in the NMSSM,” Nucl. Phys. B **480**, 55 (1996) [arXiv:hep-ph/9609323].
- [35] C. F. Kolda, S. Pokorski and N. Polonsky, “Stabilized singlets in supergravity as a source of the mu-parameter,” Phys. Rev. Lett. **80**, 5263 (1998) [arXiv:hep-ph/9803310]. C. Panagiotakopoulos and K. Tamvakis, “Stabilized NMSSM without domain walls,” Phys. Lett. B **446**, 224 (1999) [arXiv:hep-ph/9809475].
- [36] R. N. Lerner and J. McDonald, “Higgs Inflation and Naturalness,” JCAP **1004**, 015 (2010) [arXiv:0912.5463 [hep-ph]].
- [37] R. Kallosh and S. Prokushkin, “Supercosmology,” arXiv:hep-th/0403060.
- [38] J. R. Ellis, A. D. Linde and D. V. Nanopoulos, “Inflation Can Save The Gravitino,” Phys. Lett. B **118**, 59 (1982).
- [39] M. Y. Khlopov and A. D. Linde, “Is It Easy To Save The Gravitino?,” Phys. Lett. B **138**, 265 (1984).
- [40] T. Moroi, “Effects of the gravitino on the inflationary universe,” arXiv:hep-ph/9503210.
- [41] M. Kawasaki, K. Kohri, T. Moroi and A. Yotsuyanagi, “Big-Bang Nucleosynthesis and Gravitino,” Phys. Rev. D **78**, 065011 (2008) [arXiv:0804.3745 [hep-ph]].
- [42] A. Ferrantelli, “Gravitino phenomenology and cosmological implications of supergravity,” arXiv:1002.2835 [hep-ph].
- [43] J. Hasenkamp and J. Kersten, “Leptogenesis, Gravitino Dark Matter and Entropy Production,” arXiv:1008.1740 [hep-ph].
- [44] F. Bezrukov, D. Gorbunov and M. Shaposhnikov, “On initial conditions for the Hot Big Bang,” JCAP **0906**, 029 (2009) [arXiv:0812.3622 [hep-ph]].
- [45] J. Garcia-Bellido, D. G. Figueroa and J. Rubio, “Preheating in the Standard Model with the Higgs-Inflaton coupled to gravity,” Phys. Rev. D **79**, 063531 (2009) [arXiv:0812.4624 [hep-ph]].
- [46] A. D. Dolgov and A. D. Linde, “Baryon Asymmetry In Inflationary Universe,” Phys. Lett. B **116**, 329 (1982); L. F. Abbott, E. Farhi and M. B. Wise, “Particle Production In The New Inflationary Cosmology,” Phys. Lett. B **117**, 29 (1982)

- [47] L. Kofman, A. D. Linde and A. A. Starobinsky, “Reheating after inflation,” *Phys. Rev. Lett.* **73**, 3195 (1994) [arXiv:hep-th/9405187]; L. Kofman, A. D. Linde and A. A. Starobinsky, “Towards the theory of reheating after inflation,” *Phys. Rev. D* **56**, 3258 (1997) [arXiv:hep-ph/9704452].
- [48] G. N. Felder, L. Kofman and A. D. Linde, “Instant preheating,” *Phys. Rev. D* **59**, 123523 (1999) [arXiv:hep-ph/9812289].
- [49] G. N. Felder, J. Garcia-Bellido, P. B. Greene, L. Kofman, A. D. Linde and I. Tkachev, “Dynamics of symmetry breaking and tachyonic preheating,” *Phys. Rev. Lett.* **87**, 011601 (2001) [arXiv:hep-ph/0012142]; G. N. Felder, L. Kofman and A. D. Linde, “Tachyonic instability and dynamics of spontaneous symmetry breaking,” *Phys. Rev. D* **64** (2001) 123517 [arXiv:hep-th/0106179].
- [50] S. Y. Khlebnikov and I. I. Tkachev, “Classical decay of inflaton,” *Phys. Rev. Lett.* **77**, 219 (1996) [arXiv:hep-ph/9603378]. G. N. Felder and I. Tkachev, “LATTICEEASY: A program for lattice simulations of scalar fields in an expanding universe,” [arXiv:hep-ph/0011159]. A. V. Frolov, “DEFROST: A New Code for Simulating Preheating after Inflation,” *JCAP* **0811**, 009 (2008) [arXiv:0809.4904 [hep-ph]].
- [51] M. Ibe, Y. Nakayama and T. T. Yanagida, “Conformal Gauge Mediation and Light Gravitino of Mass  $m_{3/2} < O(10)eV$ ,” *Phys. Lett. B* **671**, 378 (2009) [arXiv:0804.0636 [hep-ph]].
- [52] O. DeWolfe and S. B. Giddings, “Scales and hierarchies in warped compactifications and brane worlds,” *Phys. Rev. D* **67**, 066008 (2003) [arXiv:hep-th/0208123].
- [53] N. Arkani-Hamed and S. Dimopoulos, “Supersymmetric unification without low energy supersymmetry and signatures for fine-tuning at the LHC,” *JHEP* **0506**, 073 (2005) [arXiv:hep-th/0405159]; N. Arkani-Hamed, S. Dimopoulos, G. F. Giudice and A. Romanino, “Aspects of split supersymmetry,” *Nucl. Phys. B* **709**, 3 (2005) [arXiv:hep-ph/0409232].
- [54] R. Blumenhagen, J. P. Conlon, S. Krippendorff, S. Moster and F. Quevedo, “SUSY Breaking in Local String/F-Theory Models,” *JHEP* **0909**, 007 (2009) [arXiv:0906.3297 [hep-th]].
- [55] D. H. Lyth and E. D. Stewart, “Thermal Inflation And The Moduli Problem,” *Phys. Rev. D* **53** (1996) 1784 [arXiv:hep-ph/9510204].