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Statistical and systematic uncertainties in extracting the source properties of neutron star - black hole binaries with gravitational waves

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Gravitational waves emitted by neutron star black hole mergers encode key properties of neutron stars – such as their size, maximum mass and spins – and black holes. However, it is challenging to generate accurate waveforms from these systems hard with numerical relativity, and not much is known about systematic uncertainties due to waveform modeling. We simulate gravitational waves from neutron star black hole mergers by hybridizing numerical relativity waveforms produced with the SpEC code with a recent numerical relativity surrogate `NRHybSur3dq8Tidal`. These signals are analyzed using a range of available waveform families, and statistical and systematic errors are reported. We find that at a network signal-to-noise ratio (SNR) of 30, statistical uncertainties are usually larger than systematic offsets, while at an SNR of 70 the two become comparable. The individual black hole and neutron star masses, as well as the mass ratios, are typically measured very precisely, though not always accurately at high SNR. At a SNR of 30, the neutron star tidal deformability can only be bound from above, while for louder sources it may be measured and constrained away from zero. All neutron stars in our simulations are non-spinning, but in no case we can constrain the neutron star spin to be smaller than ~ 0.4 (90% credible interval). At lower mass ratios, waveform families whose late inspiral has been tuned specifically for neutron star black hole signals typically yield the most accurate characterization of the source parameters. Their measurements are in tension with those obtained using waveform families tuned against binary neutron stars, even for mass ratios that could be relevant for both binary neutron stars and neutron star black holes mergers. At higher mass ratios, waveforms that account for higher order modes yield the best results.

I. INTRODUCTION

With the detection of the binary neutron star (BNS) merger GW170817, and the associated counterparts across all of the electromagnetic (EM) spectrum (AT2017gfo/GRB170817A) [1–7], the era of multi-messenger astrophysics based on photons and gravitational waves (GWs) has begun, providing new tools to explore the universe. The multi-messenger observation of GW170817 yielded constraints on the neutron star equation of state (EoS) [2] as well as on the heavy metal production in neutron star mergers [4, 7–11], demonstrated that at least a fraction of the short gamma-ray bursts are produced by BNSs [7], and enabled the first-ever measurement of the Hubble constant based on standard sirens [12–14]. Even when an EM counterpart is not found, as was the case for the second BNS detection, GW190425 [15], GWs alone provide precious information, for example, about the masses and spins of neutron stars in binaries. All of the component objects in GW170817

and GW190425 were consistent with having small or no spin, and the total mass of GW190425 was found to be significantly higher than that of any known galactic binary pulsar [15].

Heterogeneous binary systems made of one neutron star and one black hole (NSBH) have yet to be discovered. While they are usually expected to exist, to date no uncontroversial observational evidence has been found, and a 90% upper limit on their merger rate has been set by advanced LIGO [16] and Virgo [17] in the first and the second observing runs to be $610 \text{ Gpc}^{-3} \text{ y}^{-1}$ [18]. The recently announced source GW190814 [19] might have been the first NSBH detected with GWs, but the unusually high mass of the lighter object in the binary, a $2.6 M_{\odot}$ compact object, leaves intact the possibility that GW190814 is in fact a BBH. When detected, NSBHs will come with features that makes them exceptional laboratories for physics, astrophysics and cosmology. A precise localization of the host galaxy and the study of the light-curves associated with the EM emission, which is expected as long as the neutron star is tidally disrupted [20–24], will give precious information about the environment in which NSBHs form.

Since the black hole in NSBHs contributes most of the total spin, these systems can yield precise measurement

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of black hole spins [19, 25]. Furthermore, the potentially large mass ratio¹ will enhance the effect of eventual spin precession [26], also making it easier to measure the black hole spin with good precision [27–30]. Similarly, the impact of higher multipoles is larger for systems with large mass ratios, paving the way to tests of the multipolar structure of general relativity [31, 32].

If the black hole is light enough, or with significant spin [24, 33–41] (otherwise the neutron star will cross the event horizon before it can be significantly disrupted) tidal effects might also be present. Furthermore, the potential presence of significant spin-induced orbital precession would break the degeneracy between luminosity distance and orbital inclination, which could make NSBHs significant contributors to the measurement of the Hubble constants with standard sirens [42].

Key to the interpretation of GW detections and signal analysis is the development of accurate and computationally efficient GW waveforms, which are used to measure the parameters of the signal by matching the model waveform² against the GW data. GW models are usually calibrated against waveforms obtained directly with numerical relativity (NR) codes, which solve Einstein’s equations computationally [43–46]. The presence of matter in combination with a singularity makes NSBHs exciting systems to study, but also very challenging to simulate compared to BBH systems. Due to the high computational cost and significant technical difficulties of numerical relativity NSBH simulations, there are only a handful of high-resolution NSBH waveforms [47–51] publicly available in the Simulating eXtreme Spacetimes (SXS) GW database, and just over one hundred lower-resolution NSBH waveforms generated using the SACRA code, which have been used to calibrate various waveform approximants but are not publicly available [53, 54].

Furthermore, one usually does not directly use NR waveforms to measure the parameters of detected compact binary coalescences (CBCs) (but see [55, 56]), due to their high individual computational cost and sparsity across the parameter space. Instead, surrogate, phenomenological (IMRPhenom) or effective-one-body (EOB) GW models are produced, which are calibrated against NR simulations. To make these waveform models fast enough to be calculated millions of times, as required by stochastic samplers, and in some cases due to limitations in the very NR simulations that the models are calibrated against, only some of the relevant physical features are included (e.g. spin precession but not tidal deformability). Due to the lack of a large NR database, and the fact that all of

the physics that is relevant to describe a CBC can induce *measurable* effects in NSBHs³, these systems are potentially very prone to systematic errors due to waveform modeling.

In this paper we create hybrids from recent NSBH NR simulations to produce full inspiral-merger-ringdown waveforms that are then added to the data stream of a three-detector gravitational-wave network made of the two advanced LIGO [57] and the advanced Virgo [17] detectors. We simulate signals at various mass ratios, signal-to-noise ratios (SNRs) and orbital orientation, and measure their parameters with stochastic samplers, using a suite of phenomenological and effective-one-body models. Our work significantly extends what done by Ref. [50], which explored NSBH waveform systematics by only looking at waveform overlaps, instead of performing a full Markov Chain Monte Carlo (MCMC) measurement of all of the binary parameters. Ref. [58] looked at parameter estimation for NSBH sources, but only used a single waveform model, Lackey_Tidal_2013_SEOBNRv2_ROM [53, 59, 60], and did not measure extrinsic parameters or source-frame masses.

The rest of the paper is organized as follows: in Sec. II we describe the generation of simulated signals and the source characterization algorithm; we report our main findings in Sec. III and then conclude in Sec. IV. Given the large volume of data produced, we have to make drastic cuts in what can be reported in the main body of the paper without disrupting the flow: the appendices provide extensive summary tables for all of the analyses we performed and additional details about their setup.

II. METHOD

Gravitational waves emitted by a binary of compact objects in a quasi-circular orbit can be quantified by 15 parameters, including masses, spin vectors, sky position, luminosity distance and orbital orientation. Each neutron star in the binary adds additional complexity caused by tidal deformation of neutron stars and their potential disruption before merger. Waveforms that include tidal features often parametrize finite size effects through a single additional parameter proportional to the lowest-order correction to the phase evolution of the waveform in a post-Newtonian expansion, the effective tidal deformability of the binary, or using the tidal deformability of each individual neutron star. Other finite size effects, including for example the impact of the disruption of a neutron star by a black hole, are then modeled as a function of the tidal deformability of the neutron star. While there

¹ Note that two conventions exist for the mass ratio. The LIGO-Virgo collaboration usually defines the mass ratio in the range [0, 1]. We will follow the opposite convention (primarily used in the numerical relativity community) and define $q = m_1/m_2$, with $m_1 \geq m_2$.

² In what follows we will use “waveform model” and “waveform approximant” as synonyms, as they are both commonly used in GW literature.

³ For example, higher order modes are formally present in all CBC signals, but are suppressed for systems close to equal mass. This is the reason why waveform models that do not model them perform well with most of the binary black holes detected to date. The same will not necessarily be true for NSBHs.

is no obvious reason for the tidal deformability to be the most relevant parameter for finite-size effects beyond the lowest-order post-Newtonian correction to the waveform, the tidal deformability has worked well enough as a proxy for other finite-size effects within the accuracy of existing waveform models⁴.

As mentioned above, our goal is to verify if current waveform approximants can be used to accurately constrain the unknown parameters of NSBH systems. If not, we wish to check which parameters are more susceptible to biases, and at which SNR these biases become significant compared to the statistical uncertainties. In this section, we describe the generation and construction of inspiral-merger-ringdown NSBH waveforms used in this work and the data analysis approaches to measure their parameters.

A. Simulated signals

One can decompose the GW strain into a sum of spin-weighted spherical harmonics [31] as:

$$h(t, \iota, \psi_0) = \sum_{l=2}^{\infty} \sum_{m=-l}^l h_{lm}(t) {}^{-2}Y_{lm}(\iota, \psi_0) \quad (1)$$

where ι is the angle between the line of sight and the orbital angular momentum⁵, ${}^{-2}Y_{lm}$ are spin -2 weighted spherical harmonics, and ψ_0 is the initial binary phase [61].

In general, the ($l = 2, |m| = 2$) mode is dominant [31, 62–66] and higher order modes (HOMs) are suppressed. This is particularly true for low-mass systems with mass ratio close to unity [65, 67–73]. The impact of HOMs is also reduced in systems which are observed close to “face-on” i.e. with the orbital angular momentum aligned with the line-of-sight, since in this case the angular structure of the spin-weighted spherical harmonics suppresses the magnitude of the higher order terms relative to the dominant ($l = 2, |m| = 2$) mode. This is consistent with the fact that GW190412 [32] and GW190814 [19], the first two CBC detections with visible imprints of HOMs, are also the GW events with the most asymmetric mass ratio reported to date, at around 4 and 9, respectively.

To generate realistic GWs emitted from NSBHs we use NR simulations carried out by the SXS collaboration. We consider 3 different NR simulations [47, 51] produced by the SXS collaboration using the SpEC code [37, 75] at the highest available resolution. In Table I, we report

their corresponding masses, spins, HOMs, as well as the tidal deformability of the neutron star, defined as:

$$\Lambda_{\text{NS}} = \frac{2}{3} k_2 \left(\frac{R_{\text{NS}}}{M_{\text{NS}}} \right)^5 \quad (2)$$

where R_{NS} and M_{NS} are the radius and mass of the neutron star, k_2 is its tidal Love number describing the susceptibility of its shape to changes in response to a tidal potential [76], and we assume $G=c=1$.

The NR waveforms used in this work are for non-spinning neutron stars and black holes⁶. This limitation does not make our analysis less relevant or urgent since most of the black holes and all of the neutron stars discovered with gravitational waves to date are consistent with having small or no spin [18, 77]. In the rest of this paper we will often use the mass ratios of the systems, as reported in the first column of Table I, to refer to the individual NSBH simulations.

The EoS for cold, supranuclear matter in these simulations is such that the resulting tidal deformability is toward the high end of the region still allowed by previous GW observations [2, 15]. Specifically, the q6 and q2 NR waveforms use a simple Γ -law EoS, where the pressure P , density ρ , temperature T , and specific internal energy ϵ are related by $P = \kappa \rho^\Gamma + \rho T$, $\epsilon = P/\rho/(\Gamma - 1)$. Both simulations use $\Gamma = 2$, while $\kappa = 92.12$ for q6 and $\kappa = 101.45$ for q2⁷. This results in the neutron star having a tidal deformability of $\Lambda = 526$ for q6, and $\Lambda = 791$ for q2. Finally, q3 uses a piece-wise polytropic equation of state (H1, defined in [74]). For the $1.35M_\odot$ neutron star considered in the simulation, this EoS leads to a tidal deformability of $\Lambda = 624$.

We note that while q2 and q6 are relatively long waveforms by the standard of hydrodynamic simulations (> 20 cycles), q3 is comparatively short (13.3 cycles). In all cases, these simulations contain only the last few cycles before the two compact objects merge, while we need to simulate the full GW signals starting at the low frequency cut-off of gravitational-wave detectors (i.e. 20 Hz) for the purpose of this study.

We use the hybridization scheme described in [78] to combine the NR simulations with models for the early-inspiral section of the waveform. The late and post-inspiral phases from NR are smoothly attached to the early-inspiral section predicted by the NRHybSur3dq8Tidal⁸ model [79].

⁴ In this work, we will assume that black holes are not tidally deformed, and that all finite-size effects can be accurately modeled as functions of the tidal deformability of neutron stars.

⁵ For a binary with observable spin-precession, the inclination angle ι can vary significantly over time. In this work, we therefore define ι always at a reference gravitational-wave frequency of 100 Hz.

⁶ We will report on spinning NSBHs in a follow-up study, as more NR simulations become available.

⁷ Simulations using a Γ -law EoS can in theory be rescaled to any mass, at constant mass ratio and neutron star compactness. However, we do have to choose a mass scale when injecting the waveform into detector data. In this work, we set the mass of the neutron star to $1.4M_\odot$.

⁸ NRHybSur3dq8Tidal is constructed by adding post-Newtonian tidal effects to the underlying BBH model NRHybSur3dq8 [78]. Therefore it only includes the inspiral part of the waveform.

NR Waveforms	M_{BH}/M_{\odot}	M_{NS}/M_{\odot}	NS equation of state	Λ_{NS}	r_{NS}/km	Modes ($l, m $)
q6 [47]	8.4	1.4	$\Gamma 2(\kappa = 92.12)^a$	526	13.3	(2,2) (3,3) (2,1) (4,4) (5,5)
q3 [51]	4.05	1.35	H1 ^b	624	12.3	(2,2) (3,3) (2,1) (4,4)
q2 [51]	2.8	1.4	$\Gamma 2(\kappa = 101.45)$	791	14.4	(2,2) (3,3)

^a $P = \kappa\rho^2 + \rho T$

^b Piecewise polytrope equation of state defined in [74].

TABLE I: Numerical relativity waveforms used for the post-inspiral part of the simulated signals. Full waveforms are obtained by hybridizing with NRHybSur3dq8Tidal waveforms. See the body for mode details. Note all the neutron stars and black holes are non-spinning.

We use the infrastructure described in [80] to project the hybrid signals into the data streams of a network of 2 aLIGO detectors (Hanford and Livingston) and the Virgo detector. To better isolate biases due to waveform systematics from offsets due to gaussian noise fluctuations, we work with a zero-noise realization of the data [81], i.e. a data stream where the noise is zero at each time or frequency bin (whereas the noise power spectral density (PSD) itself is non-zero).

We probe the effect of the orbital orientation on the results by simulating every source at two different inclinations⁹, representative of a “typical” [82] detection (30°, “face-on”) and of a high-inclination system (70 degrees, close to “edge-on”). As mentioned above, larger inclinations should make the effects of HOMS more visible and, conversely, increase the bias when these effects are not taken into account but where they would have a significant contribution to the overall signal [70, 71]. Finally, all of these systems are put at two distances to give a network SNR of 30 (comparable to the loudest CBC discovered to date) and 70. The masses given in Table I are to be interpreted as detector-frame masses, with the astrophysically relevant source-frame masses being smaller by a factor of $(1+z)$, with z being the redshift of the source. Strictly speaking, the masses reported by NR simulations should be treated as defined in the source-frame, but given the proximity of our sources (Appendix A) these differ by at most a few percent, affecting the mapping between masses and NS radius by an amount much smaller than either statistical or systematic uncertainties. In the rest of the paper we will only report on the measurement of the source-frame masses. Finally, the sky location of all sources is fixed to the (arbitrary) value of 60° for both right ascension and declination.

⁹ Defined in this work as θ_{JN} , the angle between the line-of-sight vector and the total angular momentum.

B. Source characterization and waveform models

With the data stream from all detectors in hand, \mathbf{d} , we want to measure the unknown source parameters, $\boldsymbol{\theta}$. We perform Bayesian inference [83, 84], and calculate the posterior probability density function (PDF) for all source parameters:

$$p(\boldsymbol{\theta}|\mathbf{d}, H) = \frac{p(\boldsymbol{\theta}|H)p(\mathbf{d}|\boldsymbol{\theta}, H)}{p(\mathbf{d}|H)} \quad (3)$$

where $p(\boldsymbol{\theta}|H)$ is the prior probability density of $\boldsymbol{\theta}$, under the model H , while the second term in the numerator is the likelihood:

$$p(\mathbf{d}|\boldsymbol{\theta}, H) \propto \exp\left(-\frac{1}{2}\langle \mathbf{d} - h(\boldsymbol{\theta}) | \mathbf{d} - h(\boldsymbol{\theta}) \rangle\right) \quad (4)$$

where we have defined the noise weighted scalar product:

$$\langle a(\boldsymbol{\theta}, f) | b(\boldsymbol{\theta}, f) \rangle \equiv 4 \int_{f_{\text{low}}}^{f_{\text{high}}} \frac{a(\boldsymbol{\theta}, f)b(\boldsymbol{\theta}, f)^*}{S_n(f)} df$$

with $S_n(f)$ being the PSD of the detector noise. In this work, we use the design sensitivity for aLIGO and Virgo [85, 86].

Finally, the normalization constant $p(\mathbf{d}|H)$ is the evidence for the model H :

$$Z = p(\mathbf{d}|H) = \int d\boldsymbol{\theta} p(\mathbf{d}|\boldsymbol{\theta}, H)p(\boldsymbol{\theta}|H) \quad (5)$$

The choice of the waveform approximants, i.e. the waveform models we use to characterize the hybrid waveforms described in the previous section, enters the analysis through the term $h(\boldsymbol{\theta})$ in the likelihood.

At the time of writing, no waveform that accounts for all of tidal effects, spin precession and higher order modes

was available in the LIGO Algorithm library [87]. We therefore use a range of approximants that have some but not all of these features. These are reported in Table II, together with a list of the physical features that are included and, when relevant, their range of validity (which usually restricts the mass ratio and/or the black hole spin). Details on the priors assumed in the analyses are given in Appendix C.

The waveform approximants used in this study are constructed following either the EOB formalism [60, 67, 88, 99–102], or based on the phenomenological extension to analytical post-Newtonian waveforms (IMRPhenom) [95–98, 103, 104]. Both approaches smoothly extend the inspiral waveforms with models of binary merger-ringdown, calibrated against a set of spin-aligned BBH NR waveforms.

SEOBNRv4_ROM (henceforth: **SEOB**) [88, 89] describes BBH inspiral-merger-ringdown signals with general spins aligned to the orbital angular momentum. For computational efficiency, we evaluate the likelihood (cf. Eq. (4)) using a reduced-order-quadrature rule (ROQ) [105] version of SEOBNRv4 expressed in the frequency-domain. SEOBNRv4_ROM_NRTidal (**SEOBT**) [88–91] builds on the baseline BBH model **SEOB** by adding a correction of the waveform phase through a prescription of tidal effects found in BNS systems, calibrated against a set of BNS NR simulations [90, 91]. Lackey_Tidal_2013_SEOBNRv2_ROM (**LEA+**) [53, 58–60] adds both phase and amplitude corrections specific to NSBH systems, but is constructed from a reduced order model (ROM) of the older SEOBNRv2 BBH baseline waveform model [60]. We also use a ROQ implementation for **LEA+**. SEOBNRv4_ROM_NRTidalv2_NSBH (**SEOBNSBH**) [92] is also based on **SEOB** and dedicated to describing NSBH systems, but adds both a phase correction (through an updated formalism to the one included in **SEOBT** [106]) as well as corrections to the waveform amplitude based on the model of [41].

From the IMRPhenom waveform family we use IMRPhenomNSBH (**IMRNSBH**) [93] constructed from the aligned-spin BBH baseline IMRPhenomC [103] with updated phase [106] and NSBH-specific amplitude [41] corrections similar to **SEOBNSBH**. As a comparison to other GW analyses, we also include a ROQ implementation of IMRPhenomPv2 (**IMRp**) [94–96], which has been used for the majority of CBC analyses in recent years. This waveform is based on the newer aligned-spin IMRPhenomD [95, 96] BBH baseline, extended to also capture spin-induced orbital precession through an effective precession formalism [94]. We also use IMRPhenomPv2_NRTidal (**IMRpT**) [90, 91, 94–96], which further augments **IMRp** by adding a phase correction based on the same BNS tidal description as used in **SEOBT**¹⁰. Fi-

nally, we use IMRPhenomXHM (**IMRXHM**) [97, 98], which is based off the recent IMRPhenomXAS model [104], and describes aligned-spin BBH waveforms including HOMs. Note that **IMRXHM** does not include any phase or amplitude corrections from the presence of a NS in the binary.

Overall, the choice of waveform models used for this study is determined by a compromise between covering a large variety of families and physics, while keeping the computational cost reasonable. This second factor is the reason why we do not include other waveform approximants with HOMs, for example the time-domain **SEOBNRv4HM** [107, 108]¹¹.

We analyze all of the hybrid signals we generated with all of these waveform models (with some exception for the **LEA+** model, used for q3 only due to its limited range of validity on mass ratio). As mentioned above, none of the waveform families account for all of the relevant physical effects. However, given that none of the hybrid waveform we use to simulate the detected sources has precessing spins, we do not expect large biases in the recovered parameters for waveforms that do not model spin precession, while the lack of tides and HOMs might have a visible impact, depending on the mass ratio, inclination angle, and SNR of the source systems.

III. RESULTS

In this section we summarize the main findings of our study, with sections dedicated to the most significant astrophysical parameters that can be inferred from GW observations. As mentioned above, each of the signals is simulated and analyzed at two different values of the inclination angle (30° and 70°) and two values of network SNR (30 and 70). Unless otherwise stated, we quote the 90% credible intervals (CI), either absolute or relative to the true value. Given the very large number of configurations, we will not report all uncertainties in the main body of the paper. The interested reader will find extensive tables in Appendix A.

A. Masses

1. Chirp mass

For low-mass CBCs, the best constrained parameter is the chirp mass \mathcal{M} , defined as

$$\mathcal{M} = \frac{(m_1 m_2)^{3/5}}{(m_1 + m_2)^{1/5}}, \quad (6)$$

¹⁰ Note that the two BNS NRTidal models, **SEOBT** and **IMRpT**, do not include a description for the post-merger section (either a BH ringdown, or a NS remnant oscillation) of the waveform.

¹¹ We note that a frequency-domain ROM of the aligned-spin **SEOBNRv4HM** [107] model was made available as this study reached completion [109].

Waveform Approximant	Tides	Precession	HOMs ($l, m $)	NSBH- amplitude correction	Validity Range
SEOBNRv4_ROM (SEOB) [88, 89]					
SEOBNRv4_NRTidal_ROM (SEOBT) [88–91]	✓				
SEOBNRv4_ROM_NRTidalv2_NSBH (SEOBNSBH) [88, 89, 92]	✓			✓	BH spin=[-0.5,0.8]
Lackey_Tidal_2013_SEOBNRv2_ROM (LEA+) [53, 58–60]	✓			✓	q=[2,5], BH spin=[-0.5,0.5]
IMRPhenomNSBH (IMRNSBH) [93]	✓			✓	BH spin=[-0.5,0.5]
IMRPhenomPv2 (IMRp) [94–96]		✓			
IMRPhenomPv2_NRTidal (IMRpT) [90, 91, 94–96]	✓	✓			
IMRPhenomXHM (IMRXHM) [97, 98]			(2,1),(3,2), (3,3),(4,4)		

TABLE II: Waveform approximants used to characterize the source parameters of the simulated NSBH signals. We report their full name and a short label (in typewriter fonts) used in the body of the paper, whether they support tides, spin precession, HOMs (if yes, which modes), correction to their amplitude tuned for NSBH sources, and eventual restrictions in their parameter space.

where m_1 and m_2 are the component masses¹², as this parameter enters the phase of the GW signal already at the leading-order [110].

\mathcal{M} also appears in the waveform amplitude, together with the luminosity distance D_L and the redshift z :

$$\mathcal{A} \sim ((1+z)\mathcal{M})^{5/3}/D_L. \quad (7)$$

While this would suggest that these two parameters are positively correlated [77], in practice for CBCs with NSs and/or stellar mass BHs, the phasing evolution determines the chirp mass so precisely that it can be treated as known in the amplitude of the signal (which is usually measured much less precisely). Indeed, for NSBH systems like those reported here, there are $\mathcal{O}(1000)$ observable inspiral cycles, leading to a precise \mathcal{M} measurement from the waveform phase alone. For example, for the sources we analyze, the typical fractional uncertainty for \mathcal{M} is $\lesssim 1\%$ whereas the luminosity distance has fractional uncertainties of $\sim 50\%$, due to its correlation with the orbital orientation, see Section III D.

However, we do find a clear anti-correlation between $\mathcal{M}^{\text{source}}$ and D_L , as seen in Fig. 1 – 6, where $\mathcal{M}^{\text{source}}$ is the rest frame (or source-frame) chirp mass of the NSBH binary: when one increases the other decreases. This

behavior can be explained by the fact that what is measured from the GW data is the detector-frame chirp mass, which is larger than the source-frame mass by a factor $(1+z)$. Thus, to convert detector-frame chirp mass to the astrophysically interesting source-frame chirp mass, one must use the measured luminosity distance (and assume a cosmology; we use the Planck 2015 cosmological parameters¹³ [111]). For a given measured detector-frame mass, if the source were a bit farther away (higher z), the source-frame mass would have to be slightly smaller in order to yield the same detector-frame value. This is indeed what we find, and is worth stressing, as the uncertainty on the luminosity distance is often a significant factor in the statistical and systematic uncertainty for the source-frame chirp mass.

For the majority of the systems we analyze, especially for the inclination 30° binaries, the true $\mathcal{M}^{\text{source}}$ value is recovered within the 90% CI and little difference is seen between approximants. The situation is quite different for the systems with inclination equal to 70° : for those the source frame chirp mass is usually underestimated. In turn, this happens because the distance is overestimated (as explained in Sec. III D below). This bias is reduced or even absent when using IMRXHM, at large inclinations and

¹² Following the GW literature, we will use $m_1 \geq m_2$ (cf. footnote 1).

¹³ We use the cosmology defined in the TT+lowP+lensing+ext column of Table 4 from [111]. This corresponds to $\Omega_M = 0.3065$, $\Omega_\Lambda = 0.6935$, $w_0 = -1$ and $H_0 = 67.90 \text{ kms}^{-1}\text{Mpc}^{-1}$.

SNRs for systems with large mass ratios, since in that case HOMs become observable enough to help break the distance-inclination degeneracy, thus yielding unbiased chirp mass estimates (e.g. Fig. 6 and Sec. III D).

On the other hand, there is only marginal difference in the recovery of $\mathcal{M}^{\text{source}}$ between waveform models that do or do not include NS tidal effects. However, the same is not necessarily true for other parameters, as discussed below.

2. Mass ratio

As we will discuss further in Sec. III B, it is largely the binary mass ratio, together with the BH-spin and NS EoS, that determines if and to which extent the finite-size of the NS will leave an observable imprint in the detected GW signal [36, 40, 41]. For the $q = 2$ binaries, for which tidal effects are largest, the waveforms that do not model the tidal disruption of the NS recover a strongly biased posterior of q , with the true value of $q = 2$ only barely included in the tail of the posterior for the SNR 70 sources, Fig. 1b and 2b.

For the $q = 3$ binaries of Fig. 3 and 4, the tidal effects present are less prominent and hence all waveform models show similar performance in recovering the true mass ratio, with the exception of non-tidal IMRPhenom-based waveforms, IMRp and IMRXHM, which are only marginally consistent with the true value.

When $q = 6$, Fig. 5 and 6, the NS is not expected to disrupt before plunging into the BH, and hence tidal effects are unmeasurable (as shown in Fig. 17 and 18). For these sources, it is the two waveform models that are explicitly calibrated against (near-equal-mass) BNS simulations, SEObT and IMRpT, that produce biased mass-ratio posteriors, with IMRpT being farther away from the true value. While at SNR 30, Fig. 5a and 6a, a second peak at more equal mass ratios is already visible, the main peak is still present around the true value of $q = 6$. It is only for the loud signals that the peak at the true value disappears, resulting in a significant bias, especially for IMRpT. For the q6 sources, the two specialized NSBH waveforms – IMRNSBH and SEOBNSBH – and the BBH waveforms measure the mass ratio comparably well, whereas IMRXHM remains overall better suited as it can deliver more precise distance and inclination measurements, Sec. III D, and hence a better measurement of the source-frame masses.

Overall, we find the absolute statistical uncertainties for q of the order of ~ 0.8 at SNR 70 for the NSBH-tuned models without much dependence on the true value of q . While the absolute value of the 90% CI stays roughly constant with q , the relative uncertainty is 3 times smaller for q6 than for q2. For the SNR 30 signals, the absolute uncertainties are naturally higher and fall into the range of 1.2-1.6 for all three systems.

3. Neutron star and black hole masses

One of the most attractive features of NSBH binaries is the potential of a precise measurement of the neutron star mass, including putting constraints on its maximum value, which is still under debate [112–116]. Unfortunately, this is hard to achieve even at high SNRs with BNSs, due to their mass ratio being close to unity [117].

This is particularly true if one follows an agnostic approach, without assuming *a priori* that a compact object lighter than $2 M_{\odot}$ is necessarily a NS, and allows for the object to assume spins larger than what a NS could nominally support [118, 119]. In that case, a known spin-mass ratio degeneracy will significantly increase the uncertainty in both parameters [120]. This was clearly shown with the first BNS source [1], for which the upper value of the 90% CI for the primary mass increases by $\sim 40\%$ ($\sim 18\%$) when the spin magnitude limit is increased from 0.05 to 0.89 for spin-aligned (spin-precessing) waveforms. Similar differences have been reported for GW190425 [15]. For BNSs, the spin prior used will usually determine whether it is possible to set a significant upper bound on NS masses.

We want to verify if NS mass measurement obtained from NSBH sources are more precise, as one would expect from their larger mass ratios, and more accurate. We find that for the SNR 30 binaries, Figures 7a to 12a, all waveform models perform comparatively well in recovering the true binary masses. The exception is the $q = 3$ binaries, Fig. 9a and Fig. 10a: two of the models which allows for NS tidal effects and are dedicated to BNS systems (SEObT and IMRpT) have wider tails towards more equal-mass binaries, and hence heavier NSs (Fig. 3a and 4a). This behavior is also seen for the $q = 6$ binaries, Fig. 11a and 12a, with IMRpT having especially large tails. On the other hand, note that the IMRNSBH and SEOBNSBH analyses are more constraining on the NS and BH masses compared to these “BNS-tuned” waveform models, similar to the discussion in Sec. III A 2.

For the SNR 70 binaries, severe biases are visible, due to two different factors. The true values are outside of the 90% credible regions for the $q = 2$ binaries, Fig. 7b and 8b, when using approximants that do not support NS matter effects. This is to be expected since tidal effects are most visible at small mass ratios, and in light of the fact that tides and mass ratios enter the GW phase in combination [121]. This bias of $\sim 0.1 M_{\odot}$ for the recovered NS mass, though only a small fractional error, could be detrimental when propagated to the inference on the NS EoS, which is very sensitive to changes in NS mass. It is also interesting to note that the HOMs included in the IMRXHM model do not affect the inferred masses for these binaries, and indeed recover the same biased masses as the other waveform models without NS matter effects. A similar behavior is also seen for the $q = 3$ binaries, Fig. 9b and 10b, where again the models without NS matter effects show larger biases in the NS and BH masses, more so for IMRp and IMRXHM than for SEOb, though the effects

are smaller than the those of the $q = 2$ binaries.

For the $q = 6$ binaries, Fig. 11b and 12b, we see even stronger biases, but the reason is now different. As the more unequal mass ratio reduces the observational impact of the tidal effects, the “BBH-like” models can describe the system quite well, and the models tuned to BNS-like (thus light and nearly equal-mass) tidal effects (SEOBT and IMRpT) greatly misestimate the NS and BH masses. This might be due to the conditioning applied to the end of SEOBT and IMRpT waveforms, which would be outside of the most sensitive part of the detector bandwidth for BNS-like systems, but might leave a detectable imprint for NSBH binaries with increased mass ratio and total masses as the binary merger now occurs at frequencies where the detectors are more sensitive.

Analyses of CBCs containing an object whose mass is reasonably consistent with being a neutron star can intuitively be expected to exhibit some form of tidal effects. Waveform models that allow for such effects could therefore be believed to measure the source parameters better, as they nominally contain a more accurate description of all relevant physical effects. Naively following these assumptions for the $q = 6$ binaries would, as shown here, potentially lead to significant errors in the inferred astrophysics. As an example, the IMRpT analysis in Fig. 12b would, if taken in isolation, have a strong impact on the inferred maximum NS mass, a parameter which in turn affects the constraints on the NS EoS [112–116, 122]. It is worth noticing that one can quantitatively assess the relative goodness of fit to the data of two models by computing Bayes factors. We find that the BNS-tuned tidal waveforms are strongly disfavored even when compared to non-tidal waveforms for $q = 6$ and SNR 70, which could be used as a figure of merit to exclude them from parameter estimation for specific candidate events. Further details about Bayes factors are given below, in Sec. III B 3.

B. Matter effects

Together with the mass, the radius is probably the most interesting astrophysical quantity one can infer from GW observations of neutron stars. As seen above, in Eq. (2), this information is encoded in the tidal deformability of neutron stars, which directly enters the phase evolution of GW signals, though at high post-Newtonian orders [91, 123–126]. In this section we will discuss the measurement of both radius and tidal deformability.

1. NS tidal deformability

While GWs carry information about the NS tidal deformability, whether these effects are in practice observable depends heavily on the binary parameters, for a fixed SNR. Specifically, if the mass ratio is too large, the neutron star will cross the event horizon of the black hole before any significant tidal disruption can occur. The

exact value of the mass ratio above which tidal effects are shut off also depends on the black hole spin (as this affects the position of the outer event horizon) and the neutron star compactness or, equivalently, its radius r_{NS} , see App. B [24, 33–41]. Therefore, we do not expect to gain significant information about tides from the $q = 6$ signals. We stress that, if one is agnostic and does not a priori exclude the existence of black holes with masses comparable to neutron stars, measuring the deformability of the secondary object as being non-zero would be the main way to *prove* that it was not a BH (unless EM emission is detected, which would be an even stronger indication that a NS was involved in the merger).

Indeed, at a mass ratio of $q = 6$, NR simulations that nominally include the effects of a tidally disrupted NS are indistinguishable from “pure BBH” simulations, with the tidal signature on the generated waveform being comparable to, or smaller than, the numerical precision of current NR simulations. Thus, we do not expect to be able to constrain the tidal deformability for these types of high mass-ratio NSBH binaries.

On the other hand, as discussed in Sec. II A, we should be able to constrain the tidal deformability better for binaries with less asymmetric mass ratio. In light of the above discussion, we expect the $q = 2, 3$ binaries to be the more favorable configurations in this analysis to measure Λ_{NS} : a low-mass black hole (within the putative mass gap between neutron stars and black holes [127–129]) with a relatively massive neutron star. We stress that black holes with masses in the gap have likely already been discovered: the lighter component of GW190814 was a $\sim 2.6 M_{\odot}$ compact object, making it either the heaviest NS or the lightest BH ever found [19]. Moreover, the total mass of GW190425 was $\sim 3.3 M_{\odot}$ [15, 130, 131]: if the final product of the merger was a BH, which is likely, it would have masses in between the BHs of our q2 and q3 simulations¹⁴. The tidal deformability is indeed best constrained for the most equal-mass system in our study, the $q = 2$ binaries from Fig. 13a and 14a. We note, however, that for these signals, Λ_{NS} is generally underestimated for all waveform models, while still containing the true value within the 90% CI at SNR 30. For the SNR 70 sources, the statistical uncertainty shrinks, while the offsets remain comparable. We find that the true value of Λ_{NS} is outside of the 90% CI for all waveforms but IMRNSBH and IMRpT, including SEOBNSBH that is nominally tuned for this kind of source. It is worth mentioning that SEOBNSBH, IMRpT and SEOBT all roughly agree with each other and underestimate Λ_{NS} by a similar amount, where SEOBNSBH provides the most constrained among the three.

The situation is not too different for the $q = 3$ binaries, for which we can additionally use the LEA+ model, whose range of validity is limited to $q \in [2, 5]$. At SNR 30,

¹⁴ Whether such black hole would have a high probability of merging again, with a NS, is highly dependent on the environment where it formed.

Figures 15a and 16a, the peak of the Λ_{NS} posterior is close to the true value for all approximants, which also agree well with each other, with the exception of LEA+ and IMRNSBH whose posteriors slightly overestimate the Λ_{NS} (while still containing the true value within their very large 90% CIs). A more complex picture emerges when SNR is 70, Figures 15b and 16b. In this case, we observe that posteriors cluster around two values, one larger and one smaller than the true Λ_{NS} , with the true value roughly in between the two sets. IMRpT yields the longest tail in the posteriors among the approximants with its 90% CI only slightly wider than the others. It is worth stressing that the differences we see do not simply align with the underlying base model (IMRPhenom or EOB), as instead happens for, e.g., the luminosity distance, Sec. III D below. We do not have a simple (or complicated) explanation for these features, which could arise from the detailed way each approximant implements and calibrates tidal corrections. We also observe this difference in the results from NSBH-tuned models: while based on IMRNSBH (or LEA+, for q3) we would be able to place $\Lambda_{\text{NS}} = 0$ at a very low confidence level, SEOBNSBH finds a non-negligible amount of posterior support there, and would not allow to rule out that the secondary is in fact a black hole, for which $\Lambda_{\text{NS}} = 0$. However, even at SNR of 70, the systematic uncertainties caused by waveform models on Λ_{NS} are still within the statistical uncertainties. Thus, we can still use any waveforms that include tidal effects, both the ones dedicated to BNSs and the ones to NSBHs, without producing significant bias in the results. We will require specialized NSBH waveforms only when we have detections with significantly higher SNRs.

For the $q = 6$ sources at SNR of 30, (Fig. 17 and 18), the recovered posteriors on Λ_{NS} are not much different from the prior, explicitly showing that for non-spinning systems at such a large mass ratio, there simply is no information about the NS composition, since the NS plunges into the BH horizon before it is significantly deformed. At SNR 70, this general behavior still persists but with a slightly more discernible fall off at high Λ_{NS} (not visible in the plots, due to the range we show in the horizontal axis, but conveyed by the 90% CIs quoted in App. A). The clear exception is IMRpT whose posterior has a visible peak at $\Lambda_{\text{NS}} = 0$, and is significantly different from the prior only the upper bound of the 90% for IMRpT is visible in this case, Fig. 17d and 18d). To a smaller extent, SEOBT shows the same trend. However, as discussed in Sec. III A, these approximants also recover significantly biased mass parameters, and have an unfavorable Bayes factors compared to other approximants at $q = 6$, further discussed in Sec. III B 3.

2. NS radius

Another astrophysically important quantity, capable of constraining the NS EoS through observations with both gravitational and electromagnetic observations [15, 132–

137], is the radius of NSs, r_{NS} . Unlike Λ_{NS} , r_{NS} is not directly encoded in the GW signal, but rather inferred from the measurements of the NS mass and Λ_{NS} using fitting formulae (see App. B).

We report these posteriors in panels (c) and (d) of Figs. 13 – 18. Overall, the radius is not constrained with high precision at SNR 30, with typical 90% CI widths of ~ 7 km (compare with $\lesssim 4$ km for the BNS GW170817 [134], which had a comparable SNR). The fact that the NSBH sources we study do not provide a radius measurement as precise as GW170817 is due to the dependency of the tidal terms on the mass ratio, and the fact that fewer waveform cycles are in band for CBCs with larger chirp masses.

For the $q = 2$ binaries, the inferred r_{NS} distributions show a smaller spread than the respective Λ_{NS} posteriors. For SNR 70 especially (Fig. 13d and 14d), the inferred r_{NS} is underestimated, though the true value is still contained within the 90% CI.

A similar behavior is seen for the $q = 3$ binaries, again with a reduced spread compared to Λ_{NS} . In the SNR 70 analyses (Fig. 15d and 16d), SEOBNSBH and SEOBT return distributions for r_{NS} centered around the true value, whereas IMRNSBH and LEA+ slightly overestimate r_{NS} , while still including the true value at a high confidence level.

For the SNR 70 sources, typical 90% CIs are of $\sim 3 - 4$ km ($4 - 5$ km) for $q = 2$ ($q = 3$).

As with Λ_{NS} , the $q = 6$ analyses recover very broad posteriors for r_{NS} , and only exclude extremely large values of the radius (≥ 20 km), Figs. 17 and 18. It is worth stressing that most of this information does not come from Λ_{NS} , but rather from the measurement of the NS mass (cf. Sec. III A 3). Finally, while IMRpT finds very biased posteriors for Λ_{NS} and NS mass at $q = 6$, the two biases cancel out, giving a derived posterior on r_{NS} not too different from what is obtained with other approximants.

3. Model selection

As mentioned above, the relative goodness of fit of waveform models to the data in hand can be quantified by calculating the Bayes factors between them. If one calls H_1 the model where the approximant A_1 is used to analyze the data, and H_2 the model where the approximant A_2 is used, the Bayes factor can be obtained as the ratio of the models' evidences, Eq (5):

$$\mathcal{B}_{A_2}^{A_1} \equiv \frac{Z(H_1)}{Z(H_2)} \quad (8)$$

with $\mathcal{B}_{A_2}^{A_1} > 0$ if the model H_1 , i.e. if the waveform model A_1 , is preferred.

By comparing the ratio of evidences for competing models, one can quantify the relative belief that a given model represents the true signal in the data [138] in a way that also automatically penalizes models with more degrees of freedom, or larger priors.

In Table III – VI we present the *natural log* Bayes factor between some of the models used in this study.

As one would expect, for the $q = 2$ binaries, there is more support for models that include tidal effects. SEOBNSBH and IMRNSBH match the data equally well, and have large odds ratios relative to all other models, likely due to the fact that they have been tuned specifically for NSBH signals, and that they do not allow for spin in the neutron star and a smaller range for BH spin, see the priors in Table XIV. *Because* the true NS spin is actually zero, they are not penalized by that limitation and their odds ratios are boosted by the smaller prior volume. The same would not necessarily be true if the source contained a spinning NS. SEOBT is significantly favored over SEOB, while there is only a mild support for IMRpT over IMRp even at SNR 70. The results for the $q = 3$ binaries are qualitatively different. Here we observe that IMRXHM is mildly preferred over the NSBH approximants. This suggests that at this mass ratio, not modeling HOMs is penalized more than not modeling tides. The LEA+ model (with further constraints on the mass ratio prior) performs as well as the other two, newer, NSBH approximants. The comparisons between SEOBT and SEOB, as well as between IMRpT and IMRp, are rather inconclusive compared to the results for the $q = 2$ binaries.

For the $q = 6$ binaries, IMRXHM has the highest Bayes factor among all the approximants, since it is the only approximant included here that accounts for higher order modes, which are more significant at higher mass ratios. Moreover, the preference is stronger for higher inclinations where HOMs have a more significant amplitude. The NSBH approximants perform slightly better than non-tidal waveforms, and much better than the two tidal models that are tuned for BNS systems. It is interesting to observe that for $q = 6$ SEOB (IMRp) does better than SEOBT (IMRpT). This is not due to the fact that tides are unmeasurable, and hence “unnecessary” in the model: as we have seen before, no significant constraints can be placed on Λ_{NS} (cf. Sec. III B 1) for these sources, and not much information is gained relative to the prior distribution. In this case, no significant Occam penalty [139] is assigned to the models with tides.

Hence, the fact that the BNS-tuned tidal waveform models are disfavored over their related non-tidal models for the $q = 6$ binaries must be attributed to them failing to properly describe the NSBH waveforms in that mass range.

C. Spins

There are multiple reasons why an accurate measurement of the spin of black holes in NSBH systems is important. First, the BH spin should be measured more precisely in NSBHs than in BBHs, since the potentially large mass ratio of NSBHs enhances the effect of spin precession and spin-orbit coupling, yielding a larger amount of phase and amplitude modulation than what would be

present in an equal mass system with similar spins. NSBHs might very well be the systems that yield the most precise measurement of BH spins in the next few years. It is thus important that accuracy follows. Second, spins are a good tracer of the formation channel of compact binaries [140–145]. A precise and accurate measurement of spins could be key to determine whether the formation pathways of BBH and NSBH systems are different. When the masses of the compact objects in a binary are comparable, GWs provide a good measurement of the effective spin, χ_{eff} , the mass-weighted projection of the total spin along the orbital angular momentum [103, 146–148] (using $G=c=1$)

$$\chi_{\text{eff}} = \frac{\mathbf{S}_1/m_1 + \mathbf{S}_2/m_2}{(m_1 + m_2)} \cdot \frac{\mathbf{L}}{|\mathbf{L}|} \quad (9)$$

where $\mathbf{S}_{1,2}$ are the individual spin vectors and \mathbf{L} the orbital angular momentum of the system. However, the individual spins are poorly constrained [18, 28, 149], but see [150]. As the mass ratio increases, the spin of the primary becomes the leading contribution to χ_{eff} (this is even more true for an NSBH, as NSs are expected to have small spins), and one indeed finds that the spin of the primary becomes measurable [27, 151, 152].

Because our simulations have non-spinning black holes and neutron stars, we will not be able to probe the quality of spin measurement for large spins. However, it is still very interesting to show if waveform systematics affect the measurement of small spins in NSBH because a) most of the BH found to date are consistent with having small or no spin [141, 142, 144, 153] and b) there is correlation between effective spin and mass ratio [120], as well as between mass ratio and tidal parameters [121], hence different waveforms might produce visibly different posteriors. We expect biases to be more visible when the mass ratio is small enough that the NS can acquire significant tidal deformation and disruption before merging with the BH.

Indeed, this is what our simulations show, in Figures 7 and 8 for $q = 2$. While some differences in behavior between tidal and non-tidal approximants are already visible at SNR 30, it is only when the signals are very loud, SNR 70, that the tension becomes significant compared to the statistical uncertainties. For these loud simulations, Fig 7b and 8b, the models that include tides recover the true value of χ_{eff} and q , with the 90% CI of ~ 0.1 (with some small differences depending on inclination angle and the tidal waveform model). As discussed in Sec. III A, biases in the models without tides are also visible for the mass ratio and hence the component masses. For all of these parameters, the true values are marginally included in, or excluded from, the 90% CI. We do not observe significant differences between the NSBH-tuned waveforms and the other tidal waveforms. It is also worth stressing that the IMRXHM waveform does not perform better, or even differently, than the other non-tidal waveforms, showing explicitly that even at this high SNR, the missing tidal terms have a dominant effect on the waveform systematics

over the missing HOMs for small mass ratios.

As the mass ratio increases, the biases in the spin posteriors become less and less apparent, as one would expect given that the effect of tides decreases with more unequal masses. However, for $q = 3$ at SNR 70, Fig. 9b, 10b, we still see a bias for the non-tidal IMR models, whereas the EOB models are consistent with the true values of χ_{eff} and masses. In general, we find that non-tidal IMR models tend to overestimate χ_{eff} and, due to its correlation with q [120], to overestimate the mass ratio. Finally, for $q = 6$, Fig. 11 and 12, the true value of χ_{eff} and masses are within the 90% CIs even at SNR 70 for all non-tidal approximants. This suggests that for $q \simeq 5$, even higher SNRs would be needed for the measurement of χ_{eff} and masses to be limited by waveform systematics. On the other hand, we observe that IMRpT, which has tidal effects, gives significantly biased results at this high mass ratio. While the statistical uncertainties are large enough at SNR 30 that the posterior is still consistent with the true values, the same is no longer true at SNR 70, and for both masses and χ_{eff} the IMRpT posteriors are in significant tension with the true values (even the component spin magnitude is heavily biased, as presented in Tab. XII). In fact, while less pronounced than for IMRpT, one can see that even SEOBT starts diverging from the other approximants at $q = 6$. This can be explained with the fact that SEOBT and IMRpT are constructed with the goal of matching the late inspiral of BNSs, for which the mass ratios are close to 1. Thus, for mass ratios high enough, waveforms without tidal terms actually do better than waveforms with tidal terms tuned to only BNS mergers. As the true mass ratio increases, we are using these two models further and further from their range of validity.

This explanation for the biases is corroborated by the total lack of biases in LEA+ (only used for the q3 analysis), IMRNSBH and SEOBNSBH, all of which have tidal terms that have phase and amplitude corrections tuned against NSBH systems.

For all of the simulations, we find that LEA+, IMRNSBH and SEOBNSBH yield the most precise estimates of χ_{eff} . However, more than representing a true feature of these models, this is merely a consequence of their prior support as they do not allow spin in the NS and thus enable a narrower range of spin for the BH. This reduces correlations in the GW phase, and hence yields a better measurement of the only spin parameter.

To summarize, we find that 90% statistical uncertainties for χ_{eff} are typically around ~ 0.16 for SNR=30 sources (with small variations depending on the mass ratio) and ~ 0.08 for the SNR 70 sources. In fact, the ratio of statistical uncertainties for any given source when measured at SNR 70 and at 30 is close to the ratio of SNRs, as one would expect for loud enough sources for which the Fisher matrix limit is valid [81, 154–156].

Given that the NS spin is expected to be very low, and the mass ratio of these events is far from unity, one might hope to also measure the BH spin, and not only χ_{eff} . In

general, we find that IMRp and IMRpT yield consistently larger uncertainties, followed by spin-aligned waveforms (SEOBT,IMRXHM,SEOB) and by single-spin waveforms (LEA+, IMRNSBH and SEOBNSBH). As in the case of χ_{eff} , these differences can be explained with the reduced parameter space covered by different models. IMRp and IMRpT include a prescription for effective spin-orbit precession, and cover a higher dimensionality than any other waveform in our set. Conversely, LEA+, IMRNSBH and SEOBNSBH only allow the black hole to be spinning (with a smaller maximum amplitude), and only along the orbital angular momentum (see Table II), while setting the neutron star spin to be exactly 0.

If it worth stressing that for none of our configurations can we constrain the magnitude of the neutron star spin (for the waveform models that allow it to vary from 0) to below ~ 0.4 at the 90% confidence level. This suggests that even for NSBH sources, constraining the neutron star spin to values consistent with the range of spins of known pulsars will be challenging and require extremely loud sources. We do not expect this conclusion to depend significantly on the fact that our BHs did not have any spins, as similarly poor constraints on the NS spin in $10-1.4 M_{\odot}$ NSBH with precessing spins were reported by Ref. [27], though they worked with inspiral-only waveforms that did not include tides or higher order modes.

D. Extrinsic parameters

In this section we focus on the measurement of the inclination angle and the luminosity distance, both of great importance for fully exploiting the scientific potential of NSBH sources. At least some of the NSBHs are expected to produce EM radiation as they merge [20–24, 33–41], making an accurate measurement of their luminosity distance crucial for a successful EM follow-up program. Furthermore, the potentially small statistical uncertainty in their luminosity distance results allows NSBHs to be valuable standard sirens in a measurement of the Hubble constant [42]. Measurement of the orbital orientation could be used to distinguish between competing kilonova models [157] and, more generally, to study their detailed emission angular pattern at all wavelengths.

We report the inclination/luminosity distance corner plots for the face-on (i.e. true inclination 30°) systems in Figs. 19, 21, and 23. It is worth underlining a few common features (we remind that the full set of results can be found in Appendix A). First, the only waveform model with HOMs among those we use, IMRXHM, yields both smaller statistical errors and smaller offsets relative to the true value. Smaller statistical errors are not unexpected, since the true signals *do* have HOMs, which are known to help break the distance-inclination degeneracy [158] hence reducing the statistical uncertainty. One might be surprised that systematic errors are smallest for IMRXHM even though it does not allow for tides, even when the mass ratios are small. This can be explained with the fact

that HOMs affect the overall amplitude since they change the angular dependence on the orbital orientation, whilst tides only affect the late inspiral and mostly the phase of the waveform, thus not as directly related to distance and inclination. This also explains why, while IMRXHM performs similarly to other IMR waveforms at small mass ratios for which HOMs are less important, it does significantly better at $q = 6$. For example, at SNR= 30 and inclination 30° , the 90% relative uncertainty for the luminosity distance is 45% for all IMR models when $q = 2$, but decreases to 32% only for IMRXHM when $q = 6$, while staying above 40% for the other IMR approximants. Biases for the luminosity distance usually lie within the 90% credible intervals for the recovered posteriors, with typical offsets of the order of $\sim 5 - 15\%$ of the statistical uncertainty for the runs with SNR 30. As the SNR increases, the statistical uncertainties shrink, making systematic offsets percentually more important, though usually still smaller than the corresponding statistical uncertainty. The only exception is the IMRpT posterior for $q = 6$, Fig. 23b, which is very narrow and only marginally consistent with the true value. As already discussed above, this approximant yields biases for most parameters at $q = 6$, which is quite far from its intended region of validity.

The situation is quite different when the sources are simulated at an inclination angle of 70° . We find that most waveform families severely over-estimate the distance, with the true value barely included in the posterior, Fig. 20, 22, and 24. This results in an orbital orientation measurement closer to face-on/off, and in turn affects the estimation of the source-frame chirp mass, as seen above. This behavior is not unexpected, and can be explained with the strong Bayesian prior in the distance (proportional to D_L^2 , and roughly uniform in comoving volume at the relatively small distances in our simulations), as well as the fact that the waveform approximants without HOMs do not strongly depend on the inclination angle. In a Bayesian framework it is thus often more advantageous to overestimate the distance (which comes with a prior boost) and compensate by measuring an orientation closer to face-on/off. This was explicitly shown for models without HOMs in [42] (see also [159]). It is also consistent with the fact that the only model with HOMs in our set, IMRXHM, usually recovers a posterior closer to the true value, and more and more so as the mass ratio increases, which as discussed above enhances the effect of HOMs. For the $q = 3$ and $q = 6$ runs, the IMRXHM posterior is clearly separated from all approximants, at both SNRs. Overall, the medians for the high inclination runs are offset from the true value by significant fractions of the statistical uncertainty. The smallest offset we observe is $\sim 40\%$ of the statistical uncertainty for IMRXHM when $q = 6$. Typical values are 50% or larger. The relative statistical uncertainties on the distance for the high inclination runs are not significantly smaller than those for the systems closer to face-on. In fact, they can be larger. This is partially an artefact of quoting the 90% credible intervals relative to the true value: as the inclination in-

creases, the true distance of the source must be decreased to keep the same SNR. Since the absolute uncertainty can increase faster with the true distance than decrease with inclination, the relative uncertainties on inclination can get larger (see Fig. 1 of [42]).

It is interesting to compare non-HOM models based on the EOB vs the IMRPhenom formalisms. We see that the EOB-based models usually yield posteriors for the luminosity distance with a tail toward small distances more pronounced than those for the IMRPhenom-based models, and generally peak at similar values. While we do not have a full explanation, we note the EOB-based models we are using do not allow for spin-precession, unlike IMRp and IMRpT. Because spin precession causes amplitude (and phase) modulation that also breaks the distance-inclination degeneracy [26], it is possible that the precessing models yield better constraint posteriors due to the fact that some distance-spin configurations would be excluded when precession is not observed. This interpretation seems to be supported by the behavior of IMRNSBH, which is IMRPhenom-based but does not allow for precession. We see, for example in Fig. 23a how its posteriors follow closely those of the EOB models, rather than those of the other IMRPhenom's.

Overall, our results show that all models broadly agree for the runs with inclinations of 30° . It is only with the high-inclination sources that we start seeing large intra-waveform differences for extrinsic parameters. We see a few instances where two posteriors are nearly disjoint: for $q = 3$ and $q = 6$ at SNR 70, the IMRXHM posterior is in strong tension with IMRp and even more with IMRpT (the tension with the EOB models is milder, since those have longer tails, e.g. Fig. 24). The three NSBH-tuned models do not perform better than the other tidal-models when it comes to the measurement of distance and inclination.

IV. CONCLUSIONS

Observations of neutron star black hole coalescences can lead to significant insights into the nature of neutron stars, for example yielding a precise measurement of their mass and radius, or providing information on their formation channels. However, GWs from NSBHs are very challenging to simulate with current numerical relativity tools. The presence of matter as well as a singularity at the same time, of higher order modes enhanced by the high mass ratio, of the potential of BH spin precession, and the fact that the late inspiral and merger phases will be in a more sensitive frequency band of the detectors than that for BNSs, make it imperative to verify the role of waveform systematics.

In this work, we have created NSBH hybrid waveforms with recent NSBH NR simulations at three mass ratios, $q = 2, 3, 6$. We projected the signal into a three-interferometer network, and ran a full parameter estimation campaign, using most of the relevant waveform families available in the LIGO Algorithm Library [87], in-

cluding three that were especially built for NSBH systems. For each mass ratio, we have considered 4 configurations, where the orbital orientation and the network SNR had all of the pairwise combinations of $\theta_{JN} = 30^\circ, 70^\circ$ and $\text{SNR} = 30, 70$. This gave a total of 88 parameter estimation runs, making this study one of the most extensive analyses of statistical and systematic uncertainties in the analysis of NSBH systems to date.

We found that for signals with a SNR of 30, comparable to the loudest CBC signals detected to date, systematic uncertainties due to waveform modeling are smaller than statistical ones. Some differences are visible, for example in the NS tidal deformability, Λ_{NS} , where in some cases the posterior distributions can cluster in two different groups even at SNR 30. This is more visible for mass ratios of 3, Fig. 15a than 2, Fig. 13a. Significant offsets are also found for the source-frame chirp mass, although they are not due to waveform modeling as much as to a failure to properly measure the source luminosity distance, which is required to convert the detector-frame masses (which are the quantities actually measured from GW data) to the source-frame ones. This is particularly visible for highly inclined sources, Figures 2a, 4a and 6a. The underlying reason, as discussed in Sec. III D, is that the likelihood penalty for measuring an orientation closer to face-on, and hence a larger distance, can be more than compensated for by the fact that the Bayesian prior increases with distance, unless the true inclination angle is within $\sim 15^\circ$ from 90° [42]. This effect will not be seen for a typical detection, as most sources are expected to be detected at small inclination angles (i.e. close to 0° or 180°) [82]. It is also worth stressing that this offset is smaller for the IMRXHM waveform at $q = 6$, since the detectable higher order mode contribution in the true signal allow the IMRXHM model to break the distance-inclination degeneracy. For the WF models tuned against NSBH systems, IMRNSBH, LEA+ and SEOBNSBH, at SNR 30 we obtain 90% statistical credible intervals on the NS source-frame mass of $\sim 0.2 - 0.5 M_\odot$. These uncertainties are comparable to those reported for the BNS GW170817 [2]. This comparison is not entirely fair, as both the dimensionality of the models and the priors used are different. The settings of our IMRpT analyses are more directly comparable to Ref. [2]: for the $q = 2$ analysis and SNR 30 we find a 90% CI uncertainty in the NS mass of $0.5 M_\odot$. While this is nominally less constraining than GW170817's, it must be remembered that the mass posteriors for GW170817 have a hard prior bound (enforcing $m_1 \leq m_2$, Fig. 5 of Ref. [2]), which helps to explain why those posteriors appear narrower.

The situation is starkly different at SNR 70, with biases comparable to or larger than the statistical uncertainties. At $q = 2$, waveform models that do not account for tidal effects yield posterior measurements that do not include the true value in their 90% CIs for the mass ratio, the component masses, and the effective inspiral spin. The overall trend is the same with $q = 3$, but the biases are smaller due to the reduced impact of tides on the GW signal. In

this case, whether the true value is excluded depends on the exact approximant used. At $q = 6$, a configuration for which tidal effects, though formally included in the simulated source, do not play a significant role, the situation is somewhat reversed. Waveform approximants that do not include tides actually perform well, whilst waveforms with post-inspiral evolution tuned against nearly equal mass BNS NR simulations, SEOBT and IMRpT, yield the most severe biases. For those, the recovered masses and spins are entirely different than those from all other waveforms, and systematically offset from the true value, with IMRpT yielding a larger bias than SEOBT. We should stress that we are using these two waveform families in a region of mass ratios quite far from their calibration region, and hence these biases should not be surprising. However, we report them since they clearly show the importance of well-calibrated and faithful waveform models for the systems of interest. While we have not done this test in our study, it would be interesting to show if the opposite is true, and NSBH-tuned waveforms would suffer from similar biases if used to characterize BNS sources.

Bayes factors between pairs of models can be used to reveal whether, and to what extent, some waveform models are inadequate at matching the data. We reported them for a subset of the approximants we used, Sec. III B 3 and show that the BNS-tuned models are clearly disfavored at $q = 6$ even when compared to models that do not include matter effects. These kinds of tests could be used to decide which waveform families should be used for specific analysis, or to combine samples from different waveforms as a way to marginalize over inaccuracies and differences between waveforms [160].

The effective inspiral spin is usually measured accurately and precisely by NSBH-tuned approximants, when applicable, for all configurations. At SNR 70, systematic biases are visible for $q = 3$ from all non-tidal approximants, but are more significant for $q = 2$, where they are larger than statistical uncertainties. It is worth stressing that in none of the configurations and for none of the approximants can we constrain the NS spin to be smaller than ~ 0.4 (the true value was 0). This suggests that even with loud NSBH it might be challenging to set constraints on the NS spin to values comparable to those found in galactic pulsars. This conclusion might need to be checked against NSBH sources in which the BH has a large spin misaligned with the orbital angular momentum, though existing work suggests it might still hold true.

Finally, we found biases in the measurement of the NS tidal deformability, Λ_{NS} . For the $q = 2$ and $q = 3$ sources, the differences in the posteriors are visible even at SNR 30, though much smaller than the statistical uncertainty (which in itself is very large, in excess of 100% of the true value). Perhaps the most interesting of the SNR 30 comparisons is the one shown in Figure 15a, since it shows tension between two approximants that are tuned against NSBH NR simulations, SEOBNSBH and IMRNSBH. While at SNR 30, the offsets are still much smaller than the statistical uncertainties, they are worth stressing as

one would have expected IMRNSBH and SEOBNSBH to perform similarly. It is also worth stressing that LEA+, which belongs to the EOB-baseline family, agrees with IMRNSBH, suggesting the differences we see are not merely due to the underlying difference between EOB or IMRPhenom models, but to the specific technical details such as the way each approximant implements tidal terms and the reference point-particle models. This tension becomes much more visible at SNR 70, Figures 13b, 14b, 15b and 16b, especially for $q = 3$. Here again LEA+ and IMRNSBH roughly agree with each other (and are found to overestimate Λ_{NS}), while IMRpT, SEOBT and SEOBNSBH recover a different, and smaller, value of Λ_{NS} . Since measuring Λ_{NS} away from 0 is perhaps the best way of showing that the secondary object is not a BH when there is no EM counterpart detected, these differences are particularly interesting. Whereas based on IMRNSBH or LEA+, when available, one would exclude for nearly all of the q2 and q3 simulations that $\Lambda_{\text{NS}} = 0$, the SEOBNSBH, IMRpT and SEOBT have a larger support for $\Lambda_{\text{NS}} = 0$. This said, none of the models exclude the true value of Λ_{NS} : for $q = 2$ we find a general tendency to underestimate the tidal deformability, while for $q = 3$ some approximants overestimate and other underestimate it, with the true value found roughly in the middle, e.g. Figure 16b. The most stringent constraints are found for the $q = 2$ sources at SNR 70, with a 90% CI of 500 – 600. For the SNR 30 sources, only an upper bound can be placed. For $q = 6$, the simulated signals do not carry information about tides. We indeed find that nearly all families return a posterior on Λ_{NS} that is very similar to the prior at SNR 30, and only exclude extremely large values at SNR 70, Figures 17b and 18b. IMRpT differs significantly from the other approximants and recovers a Λ_{NS} posterior that peaks at small values. As discussed above, the reason is that waveform approximants tuned for BNS systems are being used far from their calibration range. The mass and Λ_{NS} posteriors can be converted, using phenomenological fits, to a measurement for the NS radius, r_{NS} , appendix B. We find that, at SNR 30, all approximants yield comparable constraints on the radius, with statistical uncertainties of 5 km or larger (which is larger than what was inferred for GW170817 [2]). Interestingly, even for sources where some discrepancy in Λ_{NS} was visible, the posteriors on the radius show a smaller spread. This shows that most of the information comes from the measurement of the NS mass, with Λ_{NS} contributing less to the inference of r_{NS} .

Overall, we find that at least the three approximants that have especially been tuned against NSBH waveforms agree well with each other for most of the parameters, though, critically, they show differences in their measurement of the NS tidal parameters already at SNR 30. These differences, however, are usually smaller or at most comparable to the $1\text{-}\sigma$ statistical uncertainty even at SNRs of 70. While this *might* be enough for most of the sources detected in the next few years, it clearly is insufficient in the next-generation detectors era [161], where typical

SNRs will be 10 times higher. At these SNRs, NSBH signals with mass ratios considered in this work will be distinguishable from BBHs or BNSs at a waveform level (Fig. 2 of [93]), and waveforms calibrated to NSBHs will be crucial for accurately characterize the source properties of the system. This also shows the need for a larger set of numerical relativity simulations, covering a much larger fraction of the relevant parameter space than what is currently available, in order to further calibrate and verify future NSBH waveform models. This will likely require further development of NR simulation codes [162, 163], in order to balance the computational cost and the resolving power necessary to include all significant physical effects [161, 164, 165]. Even before then, residual differences between approximants that are nominally on equal footing might be problematic when performing tests of general relativity with GWs from NSBHs.

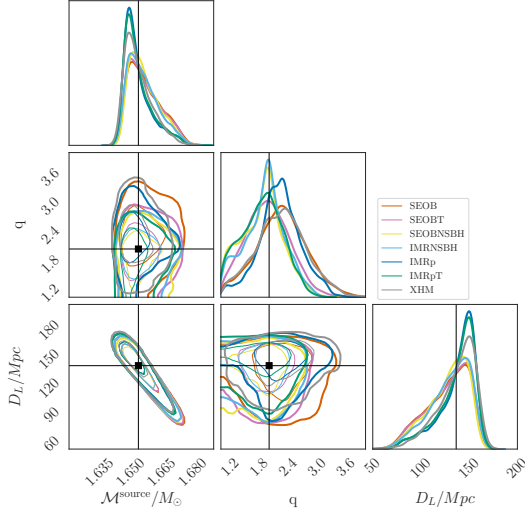
It is worth remembering that all of the simulated signals used in this paper do not have spin. This certainly represents a best-case scenario, though not an unrealistic one since most of the black holes detected to date are consistent with not having spins: it seems likely that even the two NSBH-tuned models we are using would start showing biases if the true signal came from a NSBH source with a large precessing BH spin. Work is ongoing, and will be presented in a forthcoming publication, to consider NSBH NR simulations with spinning BHs [48], though even for those the spin is not precessing.

V. ACKNOWLEDGMENTS

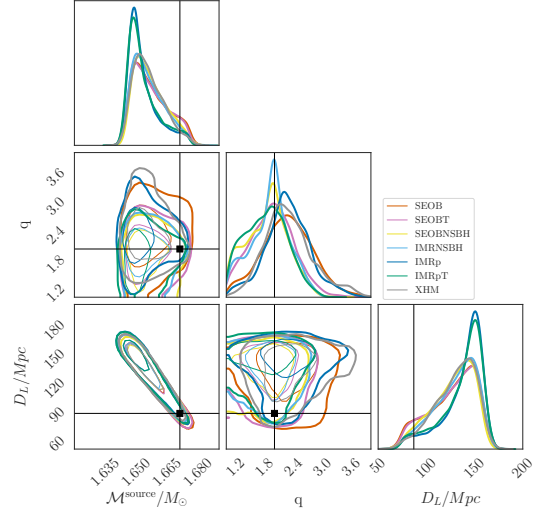
Software: We acknowledge the use of the LIGO Algorithm Library [87], and specifically of the LALInference inference package [83], as released through conda [166]. Plots were produced with matplotlib [167], and corner [168]. We acknowledge use of iPython [169], NumPy [170] and SciPy [171].

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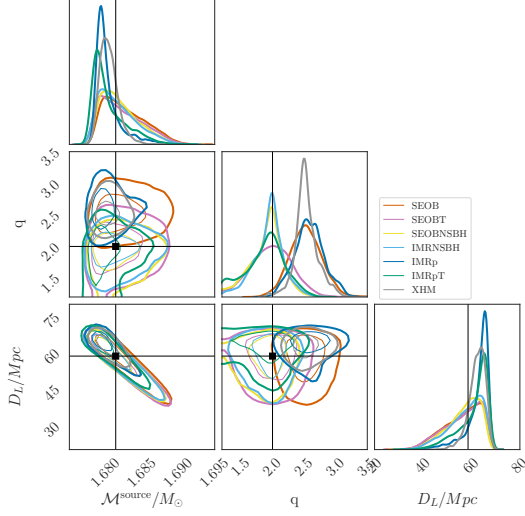
Scientifique (CNRS), the Italian Istituto Nazionale della Fisica Nucleare (INFN) and the Dutch Nikhef, with contributions by Polish and Hungarian institutes. This is LIGO Document Number DCC-P2000176.



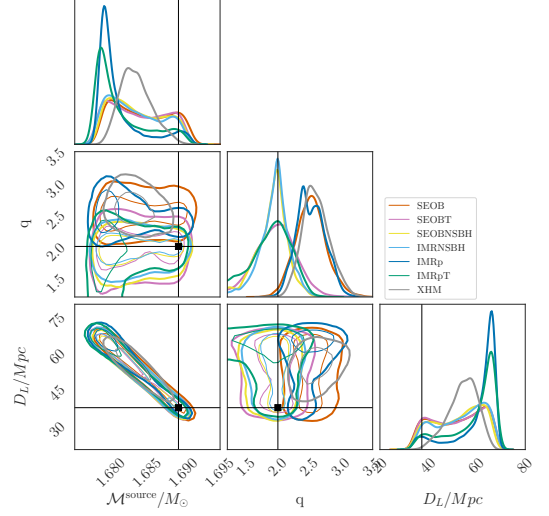
(a) SNR 30.



(a) SNR 30.



(b) SNR 70.



(b) SNR 70.

FIG. 1: Corner plot of posterior distributions for chirp mass $\mathcal{M}^{\text{source}}$, mass ratio q , and luminosity distance D_L , recovered by different approximants for $q = 2$, inclination 30° . The thin (thick) lines mark the 50% (90%) contour, same for all corner plots to follow.

FIG. 2: Corner plot of posterior distributions for chirp mass $\mathcal{M}^{\text{source}}$, mass ratio q , and luminosity distance D_L , recovered by different approximants for $q = 2$, inclination 70° .

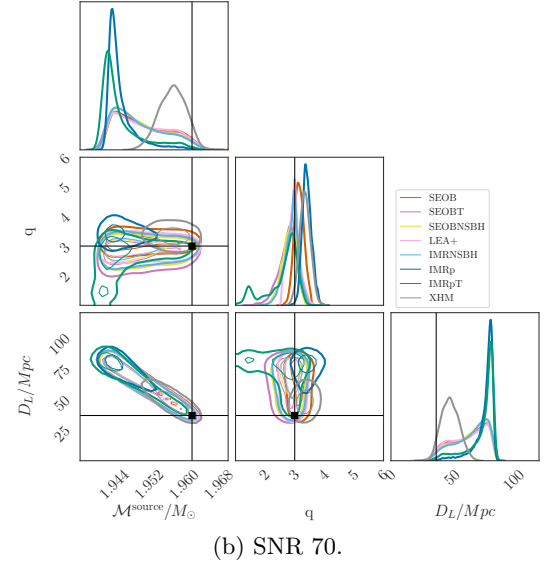
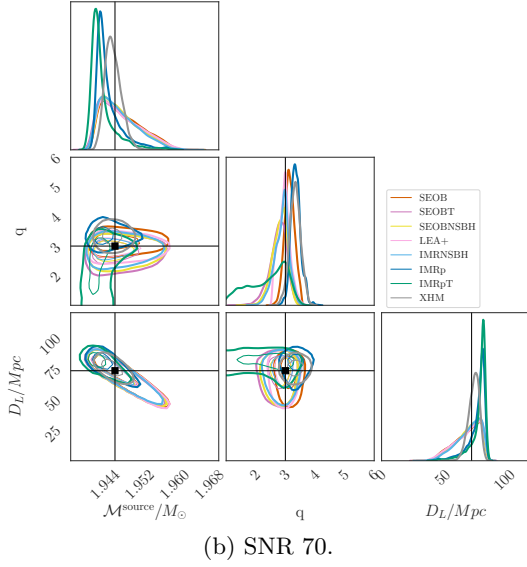
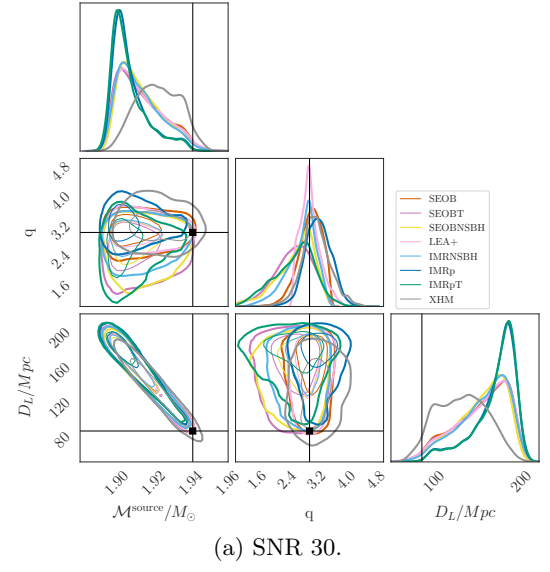
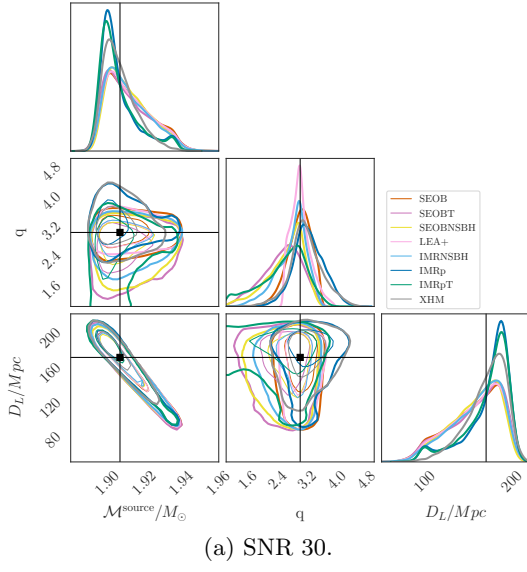


FIG. 3: Corner plot of posterior distributions for chirp mass $\mathcal{M}^{\text{source}}$, mass ratio q , and luminosity distance D_L , recovered by different approximants for $q = 3$, inclination 30° .

FIG. 4: Corner plot of posterior distributions for chirp mass $\mathcal{M}^{\text{source}}$, mass ratio q , and luminosity distance D_L , recovered by different approximants for $q = 3$, inclination 70° .

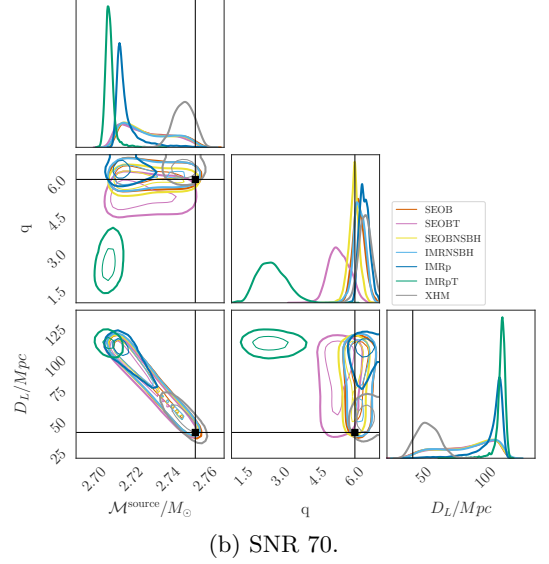
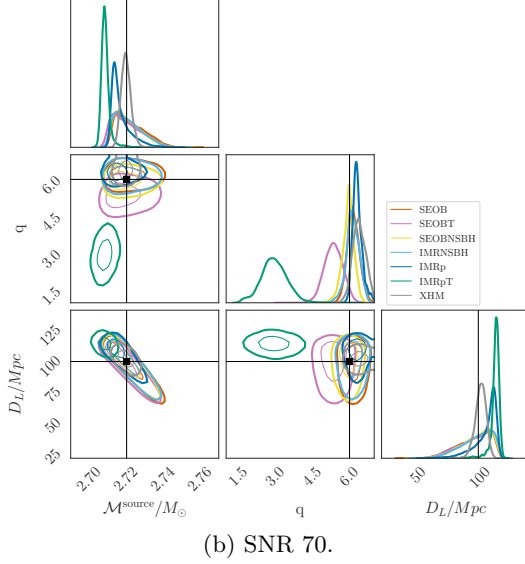
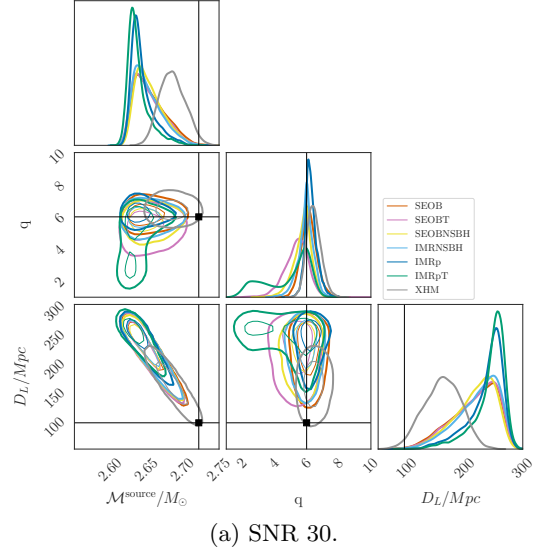
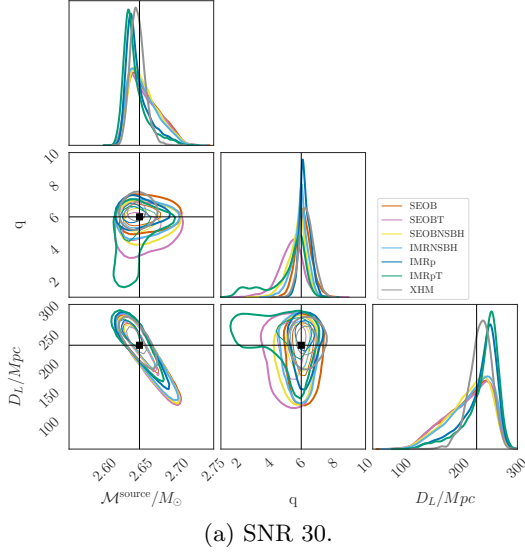
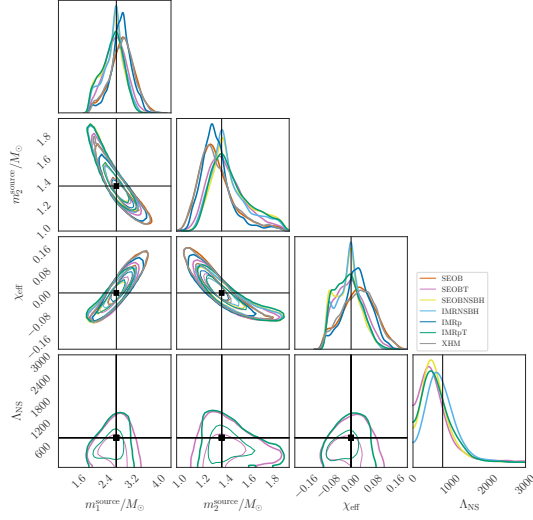
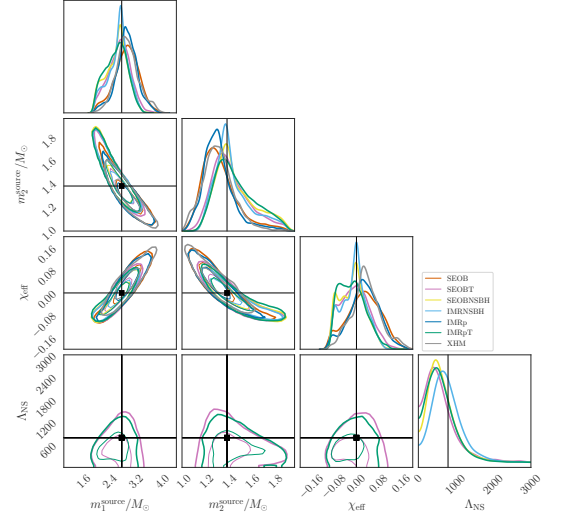


FIG. 5: Corner plot of posterior distributions for chirp mass $\mathcal{M}^{\text{source}}$, mass ratio q , and luminosity distance D_L , recovered by different approximants for $q = 6$, inclination 30° .

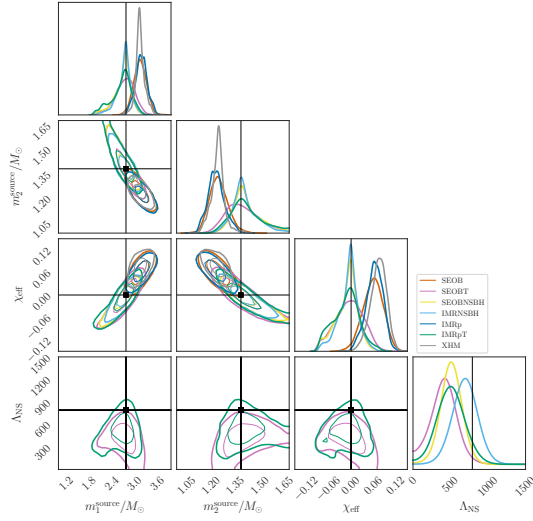
FIG. 6: Corner plot of posterior distributions for chirp mass $\mathcal{M}^{\text{source}}$, mass ratio q , and luminosity distance D_L , recovered by different approximants for $q = 6$, inclination 70° .



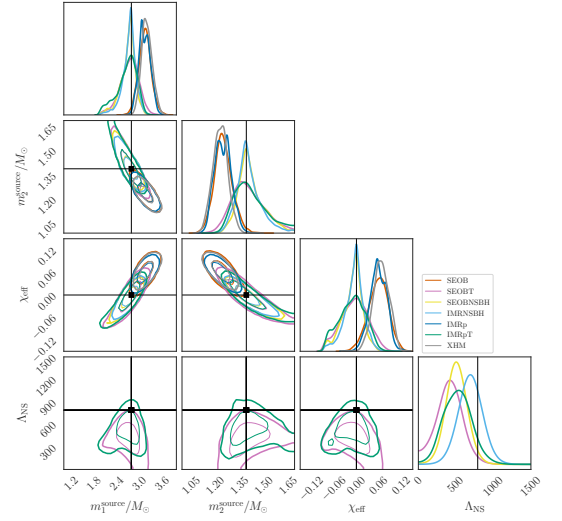
(a) SNR 30.



(a) SNR 30.



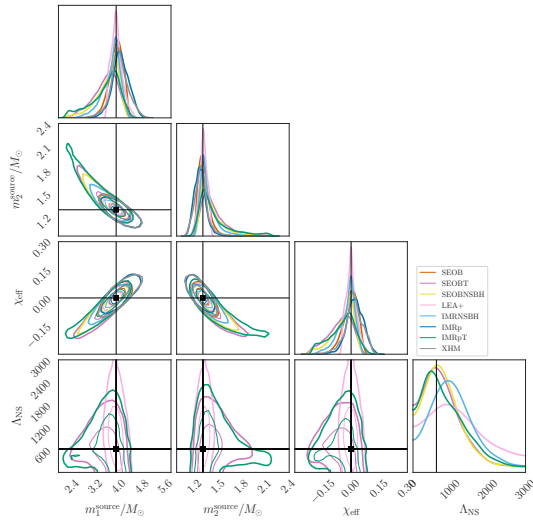
(b) SNR 70.



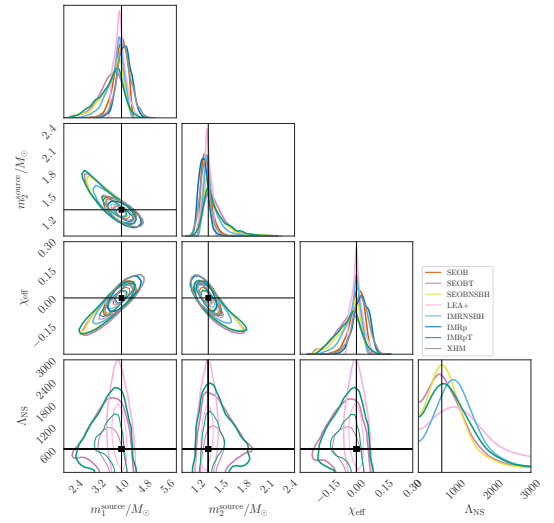
(b) SNR 70.

FIG. 7: Corner plot of posterior distributions for component masses m_1^{source} and m_2^{source} , the effective spin χ_{eff} and the tidal deformability Λ_{NS} recovered by different approximants for $q = 2$, inclination 30° .

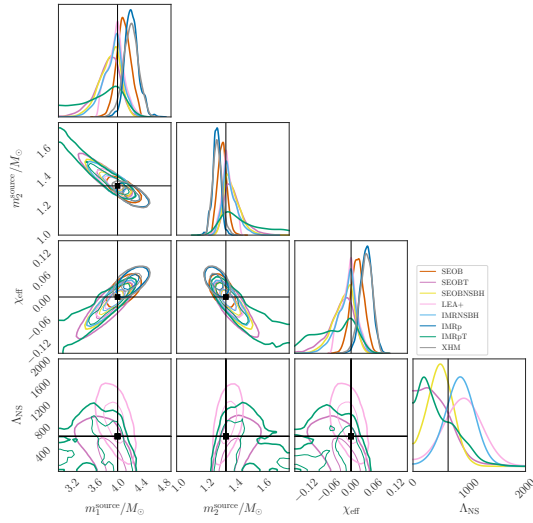
FIG. 8: Corner plot of posterior distributions for component masses m_1^{source} and m_2^{source} , the effective spin χ_{eff} and the tidal deformability Λ_{NS} recovered by different approximants for $q = 2$, inclination 70° .



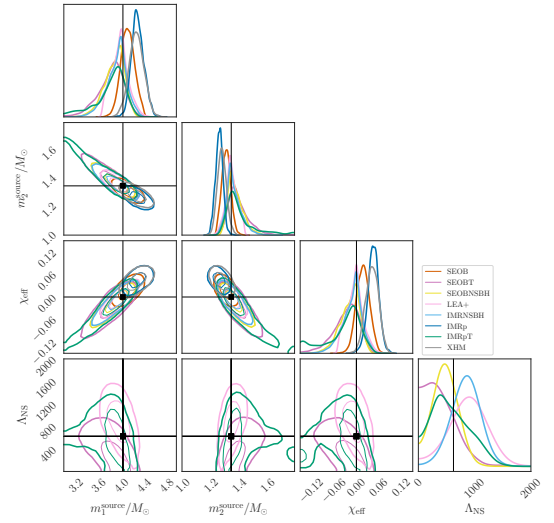
(a) SNR 30.



(a) SNR 30.



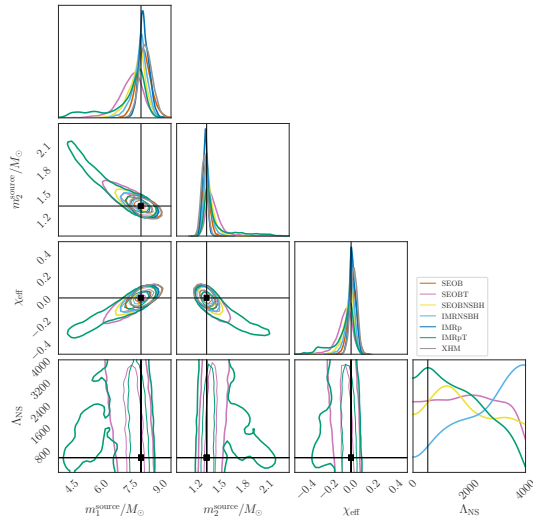
(b) SNR 70.



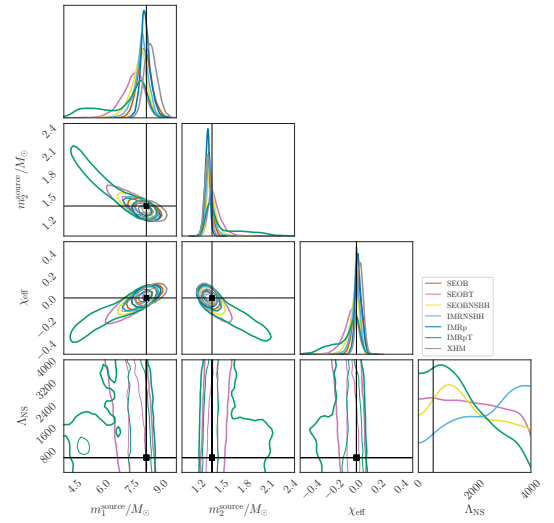
(b) SNR 70.

FIG. 9: Corner plot of posterior distributions for component masses m_1^{source} and m_2^{source} , the effective spin χ_{eff} and the tidal deformability Λ_{NS} recovered by different approximants for $q = 3$, inclination 30° .

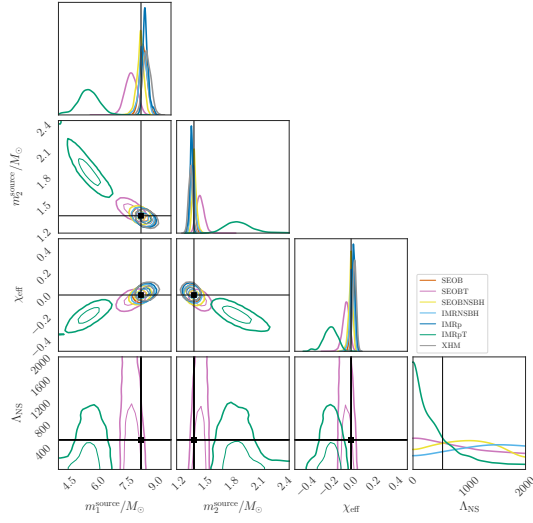
FIG. 10: Corner plot of posterior distributions for component masses m_1^{source} and m_2^{source} , the effective spin χ_{eff} and the tidal deformability Λ_{NS} recovered by different approximants for $q = 3$, inclination 70° .



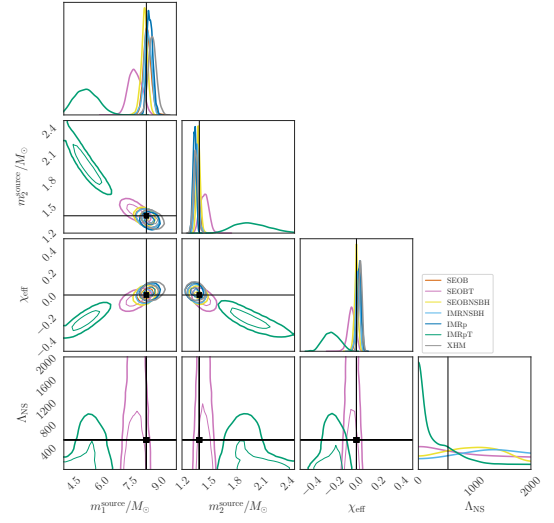
(a) SNR 30.



(a) SNR 30.



(b) SNR 70.



(b) SNR 70.

FIG. 11: Corner plot of posterior distributions for component masses m_1^{source} and m_2^{source} , the effective spin χ_{eff} and the tidal deformability Λ_{NS} recovered by different approximants for $q = 6$, inclination 30° .

FIG. 12: Corner plot of posterior distributions for component masses m_1^{source} and m_2^{source} , the effective spin χ_{eff} and the tidal deformability Λ_{NS} recovered by different approximants for $q = 6$, inclination 70° .

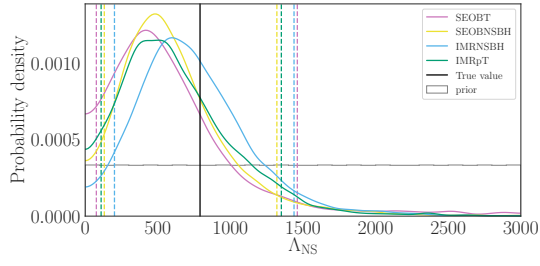
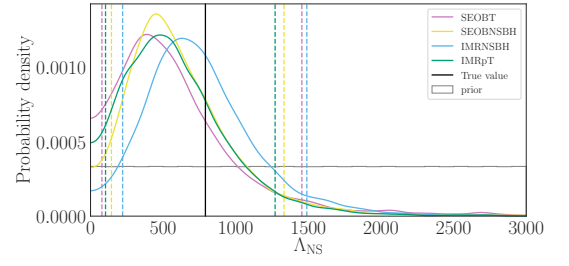
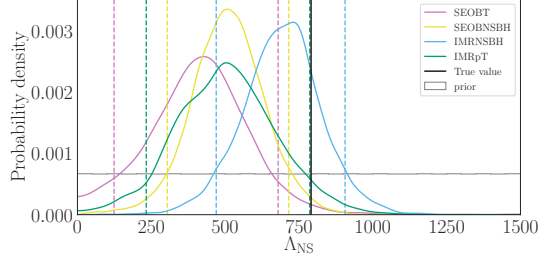
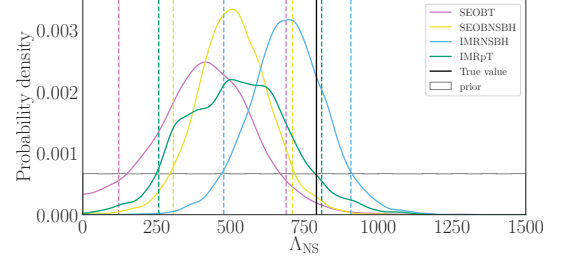
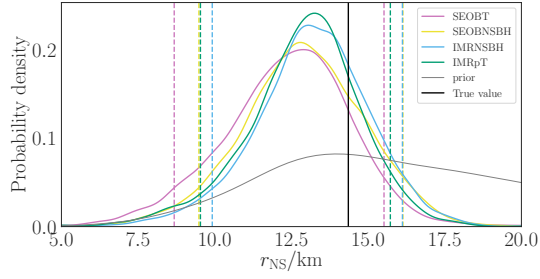
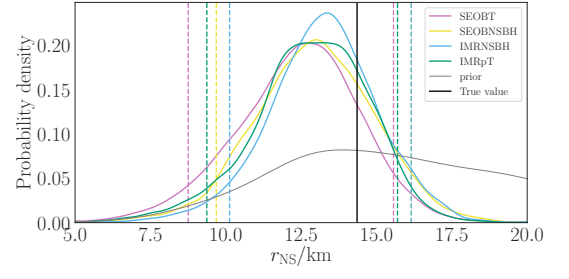
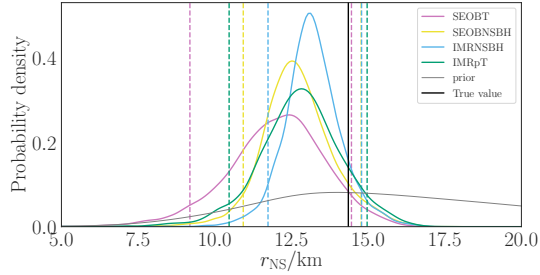
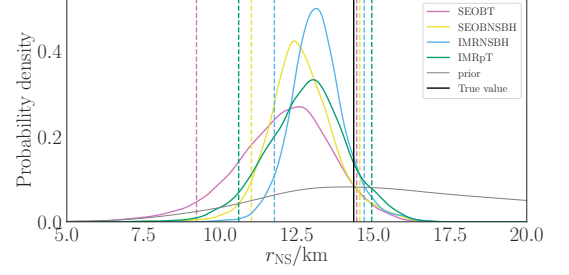
(a) Λ_{NS} , SNR 30.(a) Λ_{NS} , SNR 30.(b) Λ_{NS} , SNR 70.(b) Λ_{NS} , SNR 70.(c) r_{NS} , SNR 30.(c) r_{NS} , SNR 30.(d) r_{NS} , SNR 70.(d) r_{NS} , SNR 70.

FIG. 13: Posterior distributions for Λ_{NS} and r_{NS} recovered by different approximants for $q = 2$, inclination 30° . The dashed lines mark the 90% credible intervals, same for all 1D plots to follow.

FIG. 14: Posterior distributions for Λ_{NS} and r_{NS} recovered by different approximants for $q = 2$, inclination 70° .

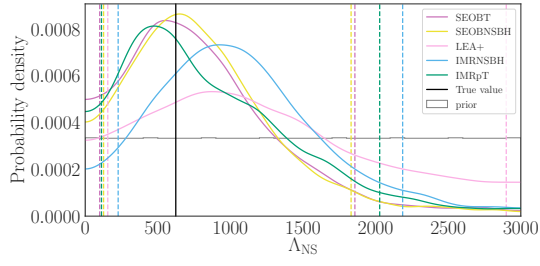
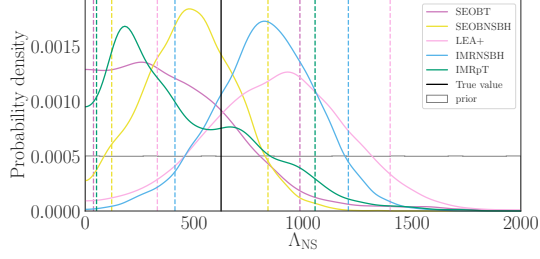
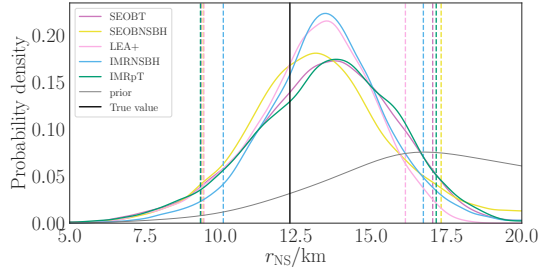
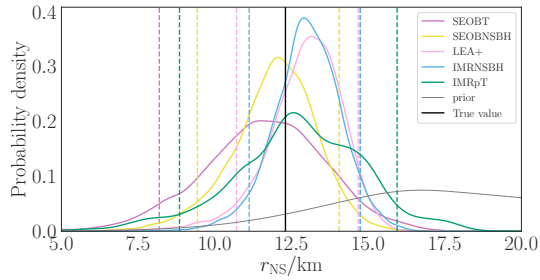
(a) Λ_{NS} , SNR 30.(b) Λ_{NS} , SNR 70.(c) r_{NS} , SNR 30.(d) r_{NS} , SNR 70.

FIG. 15: Posterior distributions for Λ_{NS} and r_{NS} recovered by different approximants for $q = 3$, inclination 30° .

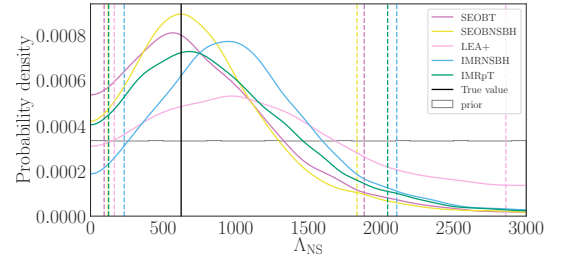
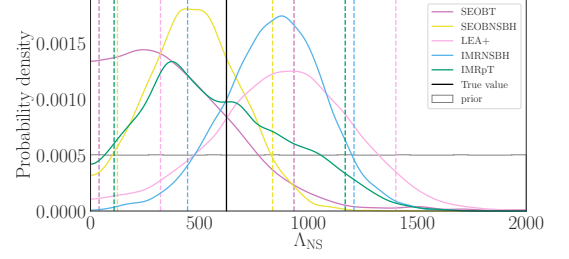
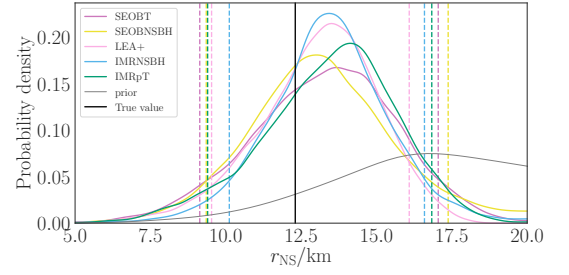
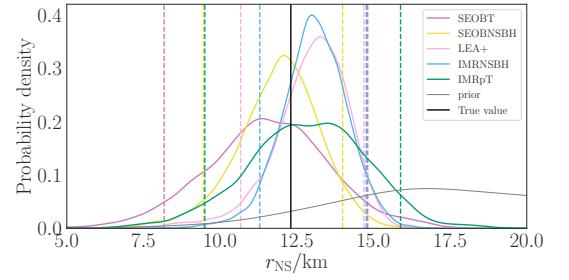
(a) Λ_{NS} , SNR 30.(b) Λ_{NS} , SNR 70.(c) r_{NS} , SNR 30.(d) r_{NS} , SNR 70.

FIG. 16: Posterior distributions for Λ_{NS} and r_{NS} recovered by different approximants for $q = 3$, inclination 70° .

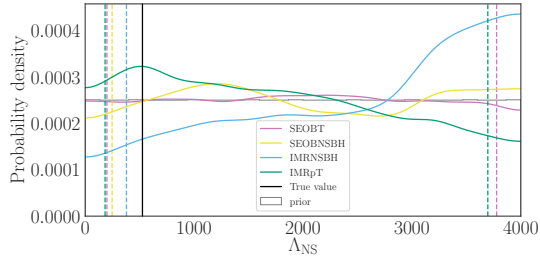
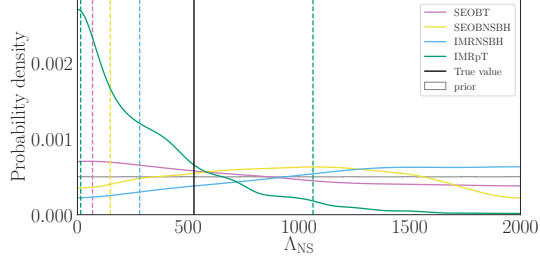
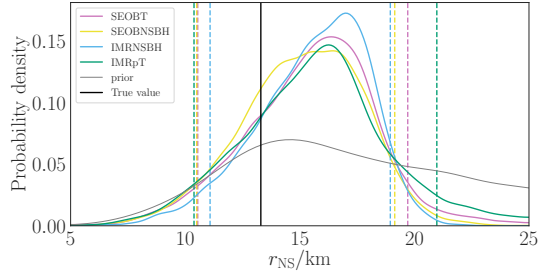
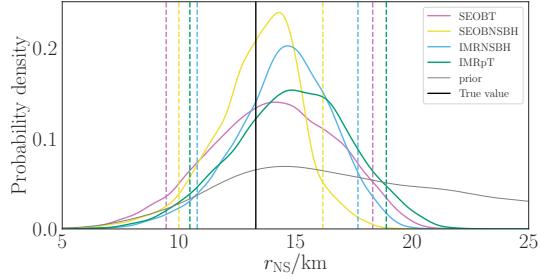
(a) Λ_{NS} , SNR 30.(b) Λ_{NS} , SNR 70.(c) r_{NS} , SNR 30.(d) r_{NS} , SNR 70.

FIG. 17: Posterior distributions for Λ_{NS} and r_{NS} recovered by different approximants for $q = 6$, inclination 30° .

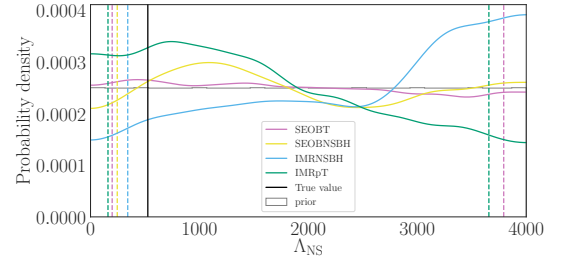
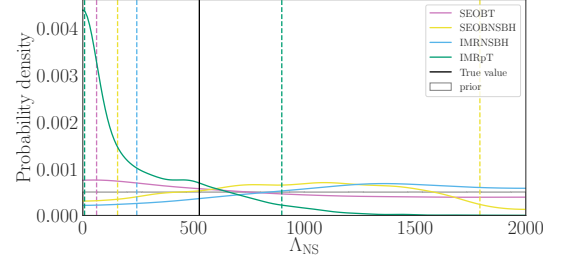
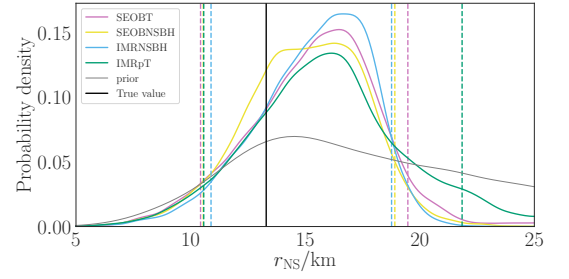
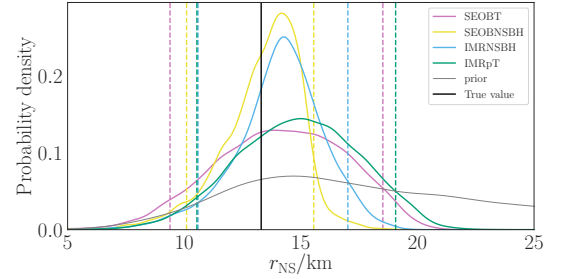
(a) Λ_{NS} , SNR 30.(b) Λ_{NS} , SNR 70.(c) r_{NS} , SNR 30.(d) r_{NS} , SNR 70.

FIG. 18: Posterior distributions for Λ_{NS} and r_{NS} recovered by different approximants for $q = 6$, inclination 70° .

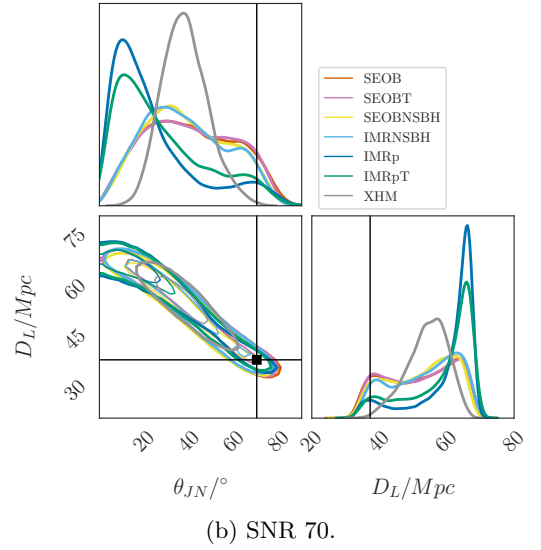
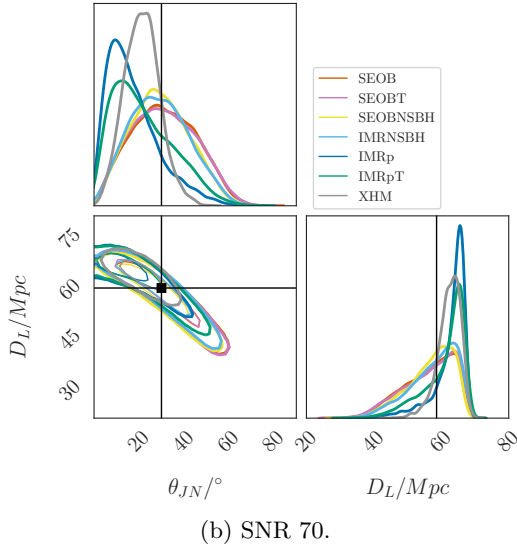
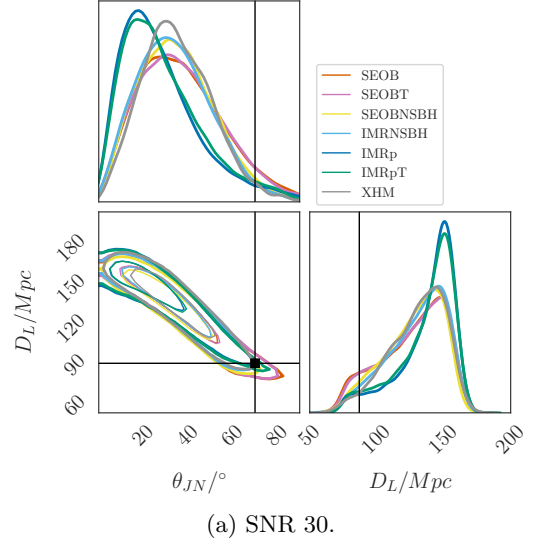
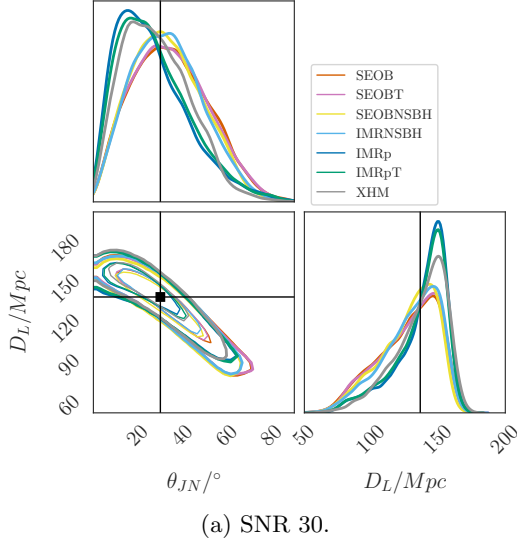


FIG. 19: 2D contour plot of posterior distributions for luminosity distance D_L and the inclination angle θ_{JN} , recovered by different approximants for $q = 2$, inclination 30° .

FIG. 20: 2D contour plot of posterior distributions for luminosity distance D_L and the inclination angle θ_{JN} , recovered by different approximants for $q = 2$, inclination 70° .

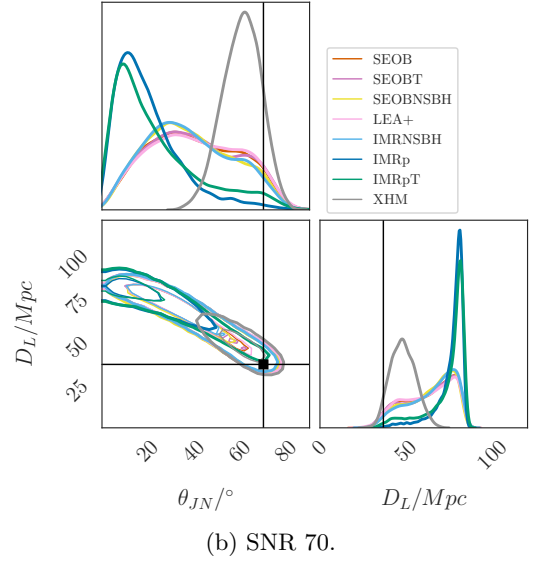
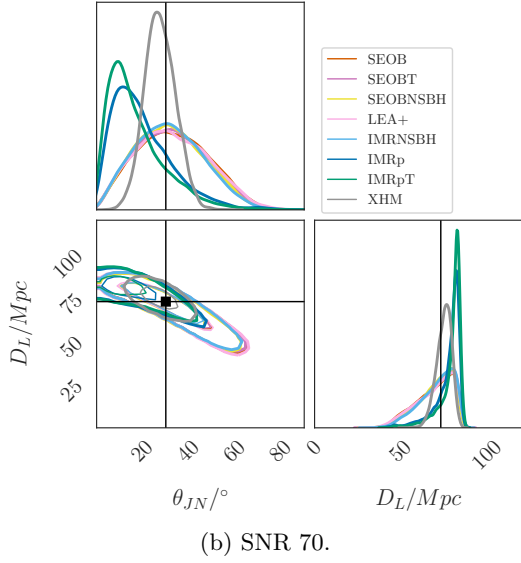
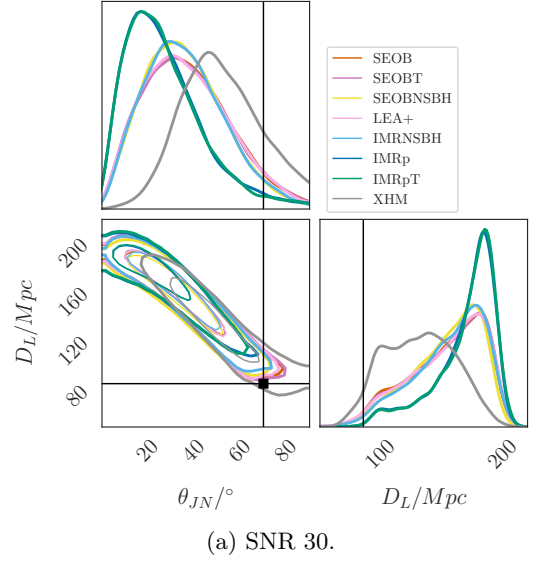
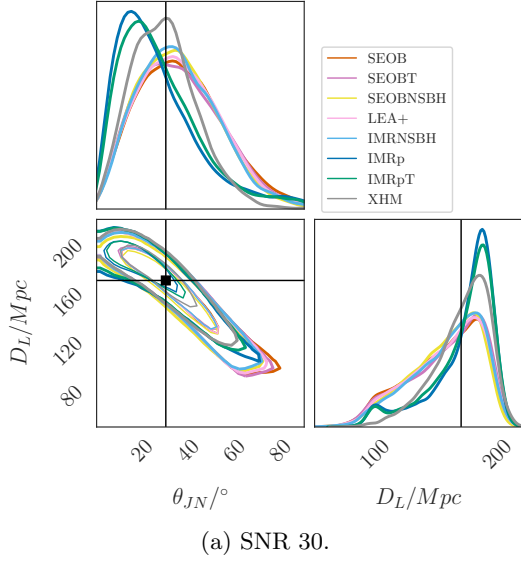


FIG. 21: 2D contour plot of posterior distributions for luminosity distance D_L and the inclination angle θ_{JN} , recovered by different approximants for $q = 3$, inclination 30° .

FIG. 22: 2D contour plot of posterior distributions for luminosity distance D_L and the inclination angle θ_{JN} , recovered by different approximants for $q = 3$, inclination 70° .

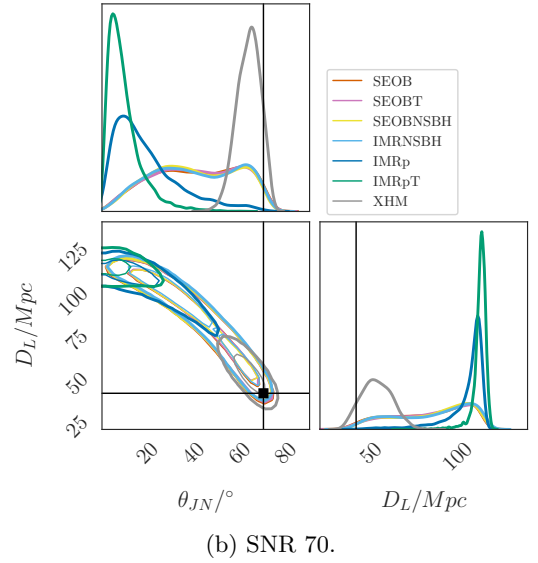
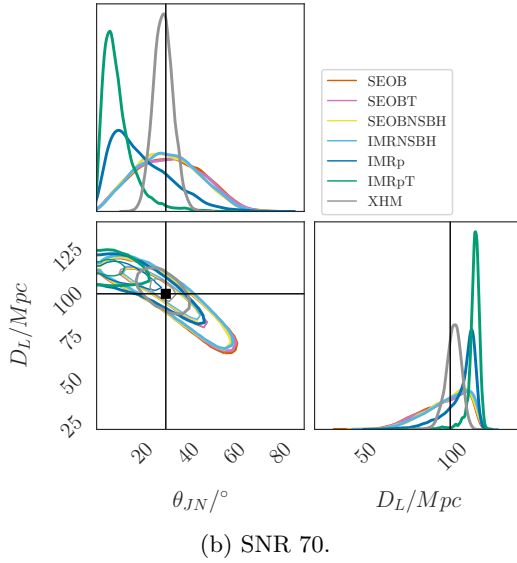
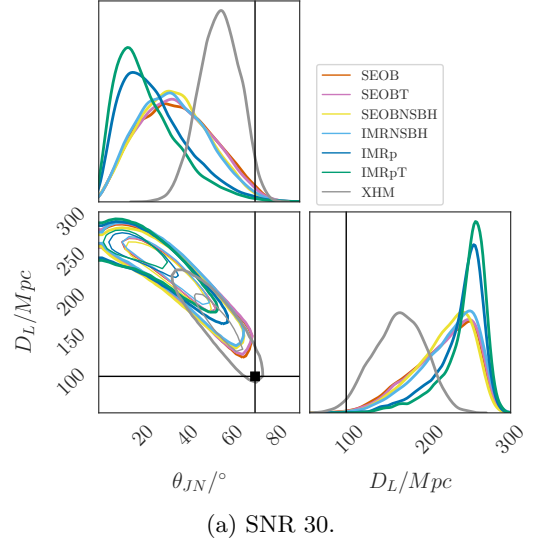
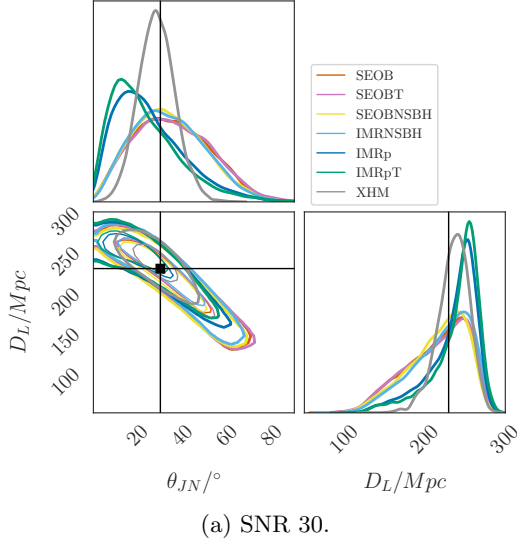


FIG. 23: 2D contour plot of posterior distributions for luminosity distance D_L and the inclination angle θ_{JN} , recovered by different approximants for $q = 6$, inclination 30° .

FIG. 24: 2D contour plot of posterior distributions for luminosity distance D_L and the inclination angle θ_{JN} , recovered by different approximants for $q = 6$, inclination 70° .

SEOBNSBH/Sim.	SEOB	SEOBT	LEA+	IMRNSBH	IMRp	IMRpT	IMRXHM
BHNSq2s0	51.96	41.15	-	-0.35	42.54	42.64	37.74
BHNSq3s0	2.30	2.99	-0.06	-0.27	3.95	3.46	-1.45
BHNSq6s0	2.62	4.81	-	0.24	4.29	6.35	-2.92

TABLE III: $\ln\mathcal{B}$ for different approximants, SNR 30, inclination 30° , reported as the odds ratio to the $\ln\mathcal{B}$ of SEOBNSBH.

SEOBNSBH/Sim.	SEOB	SEOBT	LEA+	IMRNSBH	IMRp	IMRpT	IMRXHM
BHNSq2s0	51.62	40.71	-	-0.29	42.25	42.11	36.50
BHNSq3s0	2.43	2.96	0.06	-0.10	3.58	3.70	-5.69
BHNSq6s0	2.42	5.11	-	0.01	3.90	6.13	-15.67

TABLE IV: $\ln\mathcal{B}$ for different approximants, SNR 30, inclination 70° , reported as the odds ratio to the $\ln\mathcal{B}$ of SEOBNSBH.

SEOBNSBH/Sim.	SEOB	SEOBT	LEA+	IMRNSBH	IMRp	IMRpT	IMRXHM
BHNSq2s0	279.38	214.03	-	-0.27	220.98	218.13	-
BHNSq3s0	3.42	4.97	0.29	-0.27	6.15	6.61	-7.40
BHNSq6s0	2.64	21.20	-	1.55	6.45	17.39	-25.41

TABLE V: $\ln\mathcal{B}$ for different approximants, SNR 70, inclination 30° , reported as the odds ratio to the $\ln\mathcal{B}$ of SEOBNSBH.

SEOBNSBH/Sim.	SEOB	SEOBT	LEA+	IMRNSBH	IMRp	IMRpT	IMRXHM
BHNSq2s0	277.24	212.81	-	-0.18	219.73	216.79	203.87
BHNSq3s0	3.46	5.41	0.40	-0.30	6.89	7.42	36.48
BHNSq6s0	2.67	23.47	-	1.53	6.21	17.87	-102.23

TABLE VI: $\ln\mathcal{B}$ for different approximants, SNR 70, inclination 70° , reported as the odds ratio to the $\ln\mathcal{B}$ of SEOBNSBH.

Appendix A: Full parameter estimation results

We report the parameter estimation results for an extended set of parameters relative to Sec. III. The results are presented as the median of the marginalised 1D posterior distributions for different parameters, and the corresponding symmetric 90% credible intervals.

The $q = 2$ analyses are reported in Table VII for SNR 30 and Table VIII for SNR 70. For $q = 3$, the SNR 30 results are shown in Table IX and the SNR 70 results in Table X. Finally, the $q = 6$ results are reported in Table XI for SNR 30 and Table XII for SNR 70.

Appendix B: Calculating neutron star radii

The compactness of the neutron star is estimated using a fit from [172] (Eq. 78, in Section 4.4.1),

$$C_{\text{NS}} = \sum_{k=0}^2 a_k (\ln \Lambda_{\text{NS}})^k, \quad (\text{B1})$$

with fitting parameters $a_0 = 0.371$, $a_1 = -0.0391$, and $a_2 = 0.001056$ from [173]. As reported in [172], this fit, when compared to a large set of NS EoS models, has the largest deviation of 6.5% that is significantly smaller than the statistical uncertainties reported in Sec. III B 2.

The neutron star radius r_{NS} is in turn related to the compactness through

$$r_{\text{NS}} = \frac{m_{\text{NS}}}{C_{\text{NS}}}, \quad (\text{B2})$$

with m_{NS} being the neutron star mass reported in the rest frame of the NSBH binary. Again, we are here assuming $G = c = 1$.

Appendix C: Prior

We use priors which are routinely used in LVC publications [1, 2, 84, 174–178], in Table XIII, XIV and XV, for $q = 2, 3$, and 6 respectively.

We use uniform priors for *detector-frame* component masses. When using a ROQ, additional prior constraints are imposed on the detector-frame chirp mass and mass ratio, which limit their range. For one system, the mass prior bounds are the same for all the aligned-spin waveforms, while slightly different from those for the precessing-spin waveforms due to different choices by the ROQ basis. Note that the LEA+ ROQ basis is constructed with prior constraints only on the component BH and NS masses.

The black hole spin prior is uniform in the dimensionless spin magnitude in the range $[0, 0.99]$, and isotropic for the orientation for precessing-spin approximants. For

non-precessing waveforms, the prior on the (aligned) spin magnitude is equal to the projection of an isotropic spin vector along the orbital angular momentum.

For waveform models that support tidal deformation of the neutron star, we use a prior uniform over Λ_{NS} within the range of validity.

We choose a prior for sky localization and the orientation of the orbital angular momentum with respect to the line of sight that is uniform over the sphere; a prior for the distance that is proportional to luminosity distance squared; and a uniform prior over the arrival time and phase.

Approximant	$M^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	1.65	2.0	2.72	1.36	4.08	0.00	0.00	0.00	0.00	0.00	0.0	0.0	0.0	0.0	791	14.4	136.5	30
SEOB	$1.66^{+0.02}_{-0.01}$	$2.3^{+0.9}_{-0.9}$	$2.9^{+0.6}_{-0.7}$	$1.3^{+0.4}_{-0.2}$	$4.2^{+0.4}_{-0.3}$	$0.03^{+0.09}_{-0.11}$	$0.07^{+0.17}_{-0.07}$	$0.13^{+0.41}_{-0.12}$	$0.0^{+0.5}_{-0.4}$	-	-	-	-	-	-	-	$131.4^{+25.9}_{-46.0}$	34^{+30}_{-25}
SEOBt	$1.66^{+0.02}_{-0.01}$	$2.0^{+0.8}_{-0.7}$	$2.7^{+0.5}_{-0.6}$	$1.4^{+0.3}_{-0.2}$	$4.1^{+0.3}_{-0.2}$	$0.00^{+0.09}_{-0.07}$	$0.07^{+0.19}_{-0.06}$	$0.12^{+0.41}_{-0.12}$	$0.0^{+0.4}_{-0.4}$	503 ⁺⁹⁵⁷	$12.5^{+3.0}_{-3.8}$	$131.8^{+25.1}_{-45.9}$	34^{+30}_{-24}					
SEOBNSBH	$1.66^{+0.02}_{-0.01}$	$1.9^{+0.7}_{-0.8}$	$2.6^{+0.5}_{-0.6}$	$1.4^{+0.4}_{-0.2}$	$4.0^{+0.3}_{-0.2}$	$0.00^{+0.07}_{-0.07}$	$0.05^{+0.11}_{-0.05}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}_{-0.0}$	547 ⁺⁷⁷³	$12.9^{+3.2}_{-3.4}$	$131.9^{+22.8}_{-44.1}$	32^{+28}_{-23}					
IMRNSBH	$1.65^{+0.02}_{-0.01}$	$1.9^{+0.6}_{-0.8}$	$2.7^{+0.5}_{-0.6}$	$1.4^{+0.4}_{-0.2}$	$4.1^{+0.3}_{-0.2}$	$0.00^{+0.08}_{-0.08}$	$0.04^{+0.11}_{-0.04}$	$0.0^{+0.1}_{-0.1}$	$0.0^{+0.0}_{-0.0}$	694 ⁺⁷⁴⁴	$13.2^{+2.9}_{-3.3}$	$133.4^{+23.5}_{-44.5}$	32^{+28}_{-23}					
IMRp	$1.65^{+0.02}_{-0.01}$	$2.2^{+0.8}_{-0.8}$	$2.9^{+0.5}_{-0.6}$	$1.3^{+0.3}_{-0.2}$	$4.2^{+0.3}_{-0.2}$	$0.02^{+0.08}_{-0.09}$	$0.14^{+0.36}_{-0.13}$	$0.31^{+0.54}_{-0.28}$	$0.0^{+0.4}_{-0.3}$	-	-	-	-	-	-	$143.1^{+16.5}_{-47.3}$	24^{+34}_{-18}	
IMRpT	$1.65^{+0.02}_{-0.01}$	$1.9^{+0.7}_{-0.7}$	$2.6^{+0.5}_{-0.6}$	$1.4^{+0.4}_{-0.2}$	$4.0^{+0.3}_{-0.2}$	$-0.01^{+0.08}_{-0.07}$	$0.13^{+0.40}_{-0.12}$	$0.25^{+0.53}_{-0.23}$	$-0.0^{+0.3}_{-0.2}$	569 ⁺⁷⁸²	$13.0^{+2.7}_{-3.5}$	$142.1^{+17.1}_{-46.9}$	25^{+33}_{-19}					
XHM	$1.65^{+0.01}_{-0.01}$	$2.2^{+0.9}_{-0.9}$	$2.9^{+0.6}_{-0.7}$	$1.3^{+0.4}_{-0.2}$	$4.2^{+0.4}_{-0.3}$	$0.03^{+0.09}_{-0.10}$	$0.07^{+0.15}_{-0.06}$	$0.13^{+0.30}_{-0.12}$	$0.0^{+0.4}_{-0.3}$	-	-	-	-	-	-	-	$141.1^{+20.0}_{-41.6}$	27^{+28}_{-20}

(a) Inclination 30° .

Approximant	$M^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	$M_{\text{Ch}2}^{\text{source}}/M_{\odot}$	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	1.67	2.0	2.75	1.37	4.12	0.00	0.00	0.00	0.00	0.00	0.0	0.0	0.0	0.0	791	14.4	87.0	70
SEOB	$1.66^{+0.02}_{-0.01}$	$2.3^{+0.9}_{-0.9}$	$2.9^{+0.6}_{-0.7}$	$1.3^{+0.4}_{-0.2}$	$4.2^{+0.4}_{-0.3}$	$0.03^{+0.09}_{-0.11}$	$0.07^{+0.18}_{-0.07}$	$0.13^{+0.44}_{-0.13}$	$0.0^{+0.5}_{-0.4}$	-	-	-	-	-	-	-	$131.3^{+27.2}_{-50.2}$	35^{+37}_{-25}
SEOBt	$1.66^{+0.02}_{-0.01}$	$2.0^{+0.8}_{-0.8}$	$2.7^{+0.5}_{-0.6}$	$1.4^{+0.4}_{-0.2}$	$4.1^{+0.3}_{-0.2}$	$0.00^{+0.09}_{-0.07}$	$0.07^{+0.21}_{-0.06}$	$0.12^{+0.43}_{-0.11}$	$0.0^{+0.4}_{-0.4}$	502 ⁺⁹⁵³	$12.5^{+3.1}_{-3.8}$	$131.3^{+27.1}_{-49.1}$	35^{+36}_{-26}					
SEOBNSBH	$1.66^{+0.02}_{-0.01}$	$1.9^{+0.7}_{-0.8}$	$2.6^{+0.5}_{-0.6}$	$1.4^{+0.4}_{-0.2}$	$4.0^{+0.3}_{-0.2}$	$-0.01^{+0.07}_{-0.05}$	$0.05^{+0.11}_{-0.05}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}_{-0.0}$	550 ⁺⁷⁸²	$13.0^{+3.2}_{-3.3}$	$131.0^{+25.2}_{-44.9}$	34^{+31}_{-24}					
IMRNSBH	$1.65^{+0.02}_{-0.01}$	$2.0^{+0.6}_{-0.8}$	$2.7^{+0.4}_{-0.6}$	$1.4^{+0.4}_{-0.2}$	$4.1^{+0.3}_{-0.2}$	$0.00^{+0.06}_{-0.08}$	$0.04^{+0.11}_{-0.04}$	$0.0^{+0.0}_{-0.0}$	$0.0^{+0.0}_{-0.0}$	707 ⁺⁷⁸²	$13.3^{+2.9}_{-3.1}$	$133.8^{+24.9}_{-45.6}$	33^{+30}_{-23}					
IMRp	$1.65^{+0.02}_{-0.01}$	$2.3^{+0.8}_{-0.8}$	$2.9^{+0.5}_{-0.6}$	$1.3^{+0.3}_{-0.2}$	$4.2^{+0.4}_{-0.3}$	$0.03^{+0.08}_{-0.10}$	$0.12^{+0.33}_{-0.11}$	$0.32^{+0.53}_{-0.28}$	$0.0^{+0.4}_{-0.3}$	-	-	-	-	-	-	-	$144.2^{+17.0}_{-54.0}$	24^{+41}_{-18}
IMRpT	$1.65^{+0.02}_{-0.01}$	$1.8^{+0.7}_{-0.7}$	$2.6^{+0.5}_{-0.6}$	$1.4^{+0.4}_{-0.2}$	$4.0^{+0.3}_{-0.2}$	$-0.02^{+0.08}_{-0.06}$	$0.14^{+0.43}_{-0.12}$	$0.24^{+0.49}_{-0.21}$	$-0.0^{+0.2}_{-0.2}$	538 ⁺⁷³²	$12.9^{+2.8}_{-3.6}$	$143.1^{+18.1}_{-53.4}$	25^{+40}_{-19}					
XHM	$1.65^{+0.02}_{-0.01}$	$2.2^{+1.0}_{-0.7}$	$2.9^{+0.7}_{-0.5}$	$1.3^{+0.2}_{-0.2}$	$4.2^{+0.5}_{-0.5}$	$0.03^{+0.10}_{-0.08}$	$0.09^{+0.22}_{-0.09}$	$0.15^{+0.41}_{-0.14}$	$0.1^{+0.5}_{-0.3}$	-	-	-	-	-	-	-	$133.4^{+25.2}_{-43.3}$	34^{+30}_{-22}

(b) Inclination 70° .TABLE VII: Properties of the $q=2$, SNR 30 simulation, as recovered by the listed approximants. We report the 1D median, together with the symmetric the 90% credible interval.

Approximant	$\mathcal{M}^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\odot}gn_2^{\text{source}}/M_{\odot}M_{\text{total}}^{\text{source}}/M_{\odot}$	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	1.68	2.0	2.76	1.38	4.15	0.00	0.0	0.0	791	14.4	58.6	30
SEOB	$1.68^{+0.01}_{-0.00}$	$2.5^{+0.4}_{-0.4}$	$3.1^{+0.3}_{-0.3}$	$1.2^{+0.1}_{-0.1}$	$4.4^{+0.2}_{-0.2}$	$0.06^{+0.04}_{-0.05}$	$0.07^{+0.15}_{-0.07}$	$0.1^{+0.2}_{-0.2}$	$0.14^{+0.40}_{-0.14}$	$0.1^{+0.4}$	$57.8^{+8.5}_{-15.4}$	31^{+24}_{-22}
SEOBt	$1.68^{+0.01}_{-0.00}$	$2.0^{+0.5}_{-0.6}$	$2.8^{+0.4}_{-0.5}$	$1.4^{+0.3}_{-0.1}$	$4.1^{+0.2}_{-0.2}$	$0.00^{+0.06}_{-0.07}$	$0.06^{+0.19}_{-0.06}$	$0.0^{+0.2}_{-0.2}$	$0.11^{+0.41}_{-0.11}$	$-0.0^{+0.4}_{-0.4}$	$12.1^{+2.4}_{-2.9}$	31^{+24}_{-22}
SEOBNSBH	$1.68^{+0.01}_{-0.00}$	$1.9^{+0.4}_{-0.6}$	$2.7^{+0.3}_{-0.5}$	$1.4^{+0.3}_{-0.1}$	$4.1^{+0.2}_{-0.2}$	$0.00^{+0.04}_{-0.06}$	$0.03^{+0.09}_{-0.02}$	$0.0^{+0.1}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}$	$12.7^{+2.1}_{-1.7}$	29^{+23}_{-20}
IMRNSBH	$1.68^{+0.00}_{-0.00}$	$2.0^{+0.4}_{-0.5}$	$2.7^{+0.3}_{-0.4}$	$1.4^{+0.2}_{-0.1}$	$4.1^{+0.2}_{-0.2}$	$0.00^{+0.05}_{-0.05}$	$0.02^{+0.08}_{-0.02}$	$0.0^{+0.1}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}$	$13.2^{+1.6}_{-1.4}$	29^{+23}_{-20}
IMRp	$1.68^{+0.00}_{-0.00}$	$2.6^{+0.4}_{-0.4}$	$3.2^{+0.3}_{-0.3}$	$1.2^{+0.1}_{-0.1}$	$4.4^{+0.2}_{-0.2}$	$0.06^{+0.03}_{-0.04}$	$0.12^{+0.21}_{-0.11}$	$0.1^{+0.1}_{-0.1}$	$0.28^{+0.50}_{-0.24}$	$0.1^{+0.3}$	$64.4^{+3.0}_{-12.6}$	15^{+26}_{-11}
IMRpT	$1.68^{+0.01}_{-0.00}$	$1.9^{+0.5}_{-0.7}$	$2.7^{+0.3}_{-0.6}$	$1.4^{+0.3}_{-0.1}$	$4.1^{+0.2}_{-0.2}$	$0.00^{+0.05}_{-0.06}$	$0.09^{+0.25}_{-0.08}$	$0.0^{+0.1}_{-0.1}$	$0.18^{+0.46}_{-0.16}$	$0.0^{+0.2}$	$63.2^{+4.0}_{-15.2}$	19^{+27}_{-15}
XHM	$1.68^{+0.00}_{-0.00}$	$2.5^{+0.4}_{-0.3}$	$3.1^{+0.2}_{-0.2}$	$1.2^{+0.1}_{-0.1}$	$4.4^{+0.2}_{-0.1}$	$0.07^{+0.03}_{-0.04}$	$0.06^{+0.12}_{-0.05}$	$-0.0^{+0.1}_{-0.2}$	$0.28^{+0.43}_{-0.26}$	$0.3^{+0.4}$	$62.8^{+4.2}_{-6.8}$	21^{+13}_{-12}

(a) Inclination 30° .

Approximant	$\mathcal{M}^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\odot}gn_2^{\text{source}}/M_{\odot}M_{\text{total}}^{\text{source}}/M_{\odot}$	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	1.69	2.0	2.78	1.39	4.17	0.00	0.0	0.0	791	14.4	37.3	70
SEOB	$1.68^{+0.01}_{-0.00}$	$2.5^{+0.4}_{-0.4}$	$3.1^{+0.3}_{-0.3}$	$1.2^{+0.1}_{-0.1}$	$4.4^{+0.2}_{-0.2}$	$0.06^{+0.04}_{-0.05}$	$0.07^{+0.14}_{-0.07}$	$0.1^{+0.1}_{-0.2}$	$0.14^{+0.40}_{-0.13}$	$0.0^{+0.4}$	$53.3^{+13.1}_{-17.5}$	39^{+33}_{-29}
SEOBt	$1.68^{+0.01}_{-0.00}$	$2.0^{+0.5}_{-0.6}$	$2.7^{+0.4}_{-0.5}$	$1.4^{+0.3}_{-0.1}$	$4.1^{+0.2}_{-0.2}$	$0.00^{+0.06}_{-0.06}$	$0.06^{+0.20}_{-0.05}$	$0.0^{+0.2}_{-0.2}$	$0.11^{+0.41}_{-0.10}$	$0.0^{+0.4}$	$12.2^{+2.3}_{-3.0}$	39^{+32}_{-29}
SEOBNSBH	$1.68^{+0.01}_{-0.00}$	$1.9^{+0.3}_{-0.5}$	$2.7^{+0.2}_{-0.4}$	$1.4^{+0.2}_{-0.1}$	$4.1^{+0.1}_{-0.2}$	$0.00^{+0.04}_{-0.05}$	$0.02^{+0.07}_{-0.02}$	$0.0^{+0.1}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}$	$12.6^{+1.9}_{-1.6}$	37^{+33}_{-26}
IMRNSBH	$1.68^{+0.01}_{-0.00}$	$2.0^{+0.3}_{-0.5}$	$2.7^{+0.2}_{-0.4}$	$1.4^{+0.2}_{-0.1}$	$4.1^{+0.1}_{-0.2}$	$0.00^{+0.04}_{-0.05}$	$0.02^{+0.07}_{-0.02}$	$0.0^{+0.1}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}$	$13.2^{+1.5}_{-1.4}$	36^{+33}_{-26}
IMRp	$1.68^{+0.01}_{-0.00}$	$2.5^{+0.4}_{-0.3}$	$3.1^{+0.2}_{-0.2}$	$1.2^{+0.1}_{-0.1}$	$4.4^{+0.2}_{-0.2}$	$0.06^{+0.03}_{-0.04}$	$0.12^{+0.15}_{-0.11}$	$0.0^{+0.1}_{-0.1}$	$0.30^{+0.44}_{-0.26}$	$0.1^{+0.3}$	$64.1^{+3.8}_{-26.4}$	18^{+51}_{-14}
IMRpT	$1.68^{+0.01}_{-0.00}$	$1.9^{+0.4}_{-0.7}$	$2.7^{+0.3}_{-0.6}$	$1.4^{+0.3}_{-0.1}$	$4.1^{+0.2}_{-0.2}$	$0.00^{+0.05}_{-0.06}$	$0.09^{+0.24}_{-0.08}$	$0.0^{+0.1}_{-0.1}$	$0.16^{+0.52}_{-0.15}$	$0.0^{+0.2}$	$62.5^{+5.2}_{-24.8}$	22^{+46}_{-18}
XHM	$1.68^{+0.00}_{-0.00}$	$2.6^{+0.4}_{-0.3}$	$3.2^{+0.3}_{-0.2}$	$1.2^{+0.1}_{-0.1}$	$4.4^{+0.2}_{-0.1}$	$0.07^{+0.03}_{-0.04}$	$0.06^{+0.11}_{-0.06}$	$0.1^{+0.1}_{-0.1}$	$0.12^{+0.31}_{-0.11}$	$0.1^{+0.3}$	$54.8^{+8.2}_{-11.6}$	38^{+18}_{-15}

(b) Inclination 70° .TABLE VIII: Properties of the $q=2$, SNR 70 simulation, as recovered by the listed approximants. We report the 1D median, together with the symmetric the 90% credible interval.

Approximant	$M^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\text{ch}}^{\text{source}}/M_{\text{ch}}^{\text{total}}/M_{\odot}$	$ s_1 $	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	1.95	3.0	3.99	1.33	5.31	0.00	0.0	0.0	0.0	624	12.3	73.0	30
SEOB	$1.95^{+0.01}_{-0.01}$	$3.1^{+0.3}_{-0.3}$	$4.1^{+0.2}_{-0.2}$	$1.3^{+0.1}_{-0.1}$	$5.4^{+0.2}_{-0.1}$	$0.02^{+0.03}_{-0.03}$	$0.04^{+0.12}_{-0.03}$	$0.0^{+0.1}_{-0.1}$	$0.11^{+0.40}_{-0.41}$	-	-	$71.1^{+11.8}_{-21.9}$	33^{+26}_{-27}
SEOBt	$1.95^{+0.01}_{-0.00}$	$2.8^{+0.4}_{-0.6}$	$3.8^{+0.3}_{-0.4}$	$1.4^{+0.1}_{-0.1}$	$5.2^{+0.2}_{-0.3}$	$-0.02^{+0.04}_{-0.07}$	$0.05^{+0.15}_{-0.04}$	$-0.0^{+0.1}_{-0.2}$	$0.0^{+0.4}_{-0.4}$	380^{+606}_{-342}	$11.7^{+3.1}_{-3.5}$	$71.7^{+11.3}_{-21.6}$	32^{+27}_{-23}
SEOBNSBH	$1.95^{+0.01}_{-0.00}$	$2.8^{+0.3}_{-0.5}$	$3.9^{+0.2}_{-0.4}$	$1.4^{+0.1}_{-0.1}$	$5.2^{+0.2}_{-0.2}$	$-0.02^{+0.04}_{-0.05}$	$0.03^{+0.08}_{-0.02}$	$-0.0^{+0.0}_{-0.1}$	$0.0^{+0.0}_{-0.0}$	481^{+358}_{-359}	$12.0^{+2.0}_{-2.6}$	$71.7^{+10.7}_{-21.2}$	31^{+26}_{-22}
LEA+	$1.95^{+0.01}_{-0.00}$	$3.0^{+0.3}_{-0.3}$	$4.0^{+0.2}_{-0.3}$	$1.3^{+0.1}_{-0.1}$	$5.3^{+0.2}_{-0.2}$	$0.00^{+0.04}_{-0.04}$	$0.02^{+0.04}_{-0.02}$	$0.0^{+0.0}_{-0.1}$	-	888^{+512}_{-556}	$13.1^{+1.6}_{-2.4}$	$71.5^{+11.5}_{-22.7}$	32^{+28}_{-23}
IMRNSBH	$1.95^{+0.01}_{-0.00}$	$2.9^{+0.3}_{-0.5}$	$3.9^{+0.2}_{-0.4}$	$1.3^{+0.1}_{-0.1}$	$5.3^{+0.2}_{-0.3}$	$0.00^{+0.03}_{-0.05}$	$0.02^{+0.07}_{-0.02}$	$-0.0^{+0.0}_{-0.1}$	$0.0^{+0.0}_{-0.0}$	831^{+377}_{-419}	$13.0^{+1.7}_{-1.9}$	$72.5^{+10.6}_{-21.3}$	31^{+26}_{-22}
IMRp	$1.94^{+0.01}_{-0.00}$	$3.3^{+0.3}_{-0.7}$	$4.2^{+0.2}_{-0.2}$	$1.3^{+0.0}_{-0.1}$	$5.5^{+0.2}_{-0.1}$	$0.04^{+0.03}_{-0.03}$	$0.07^{+0.14}_{-0.06}$	$0.0^{+0.1}_{-0.1}$	$0.28^{+0.43}_{-0.22}$	-	-	$79.9^{+4.2}_{-18.8}$	17^{+28}_{-13}
IMRpT	$1.94^{+0.01}_{-0.00}$	$2.5^{+0.7}_{-1.3}$	$3.6^{+0.5}_{-1.1}$	$1.4^{+0.6}_{-0.1}$	$5.1^{+0.4}_{-0.6}$	$-0.04^{+0.07}_{-0.11}$	$0.17^{+0.34}_{-0.15}$	$-0.1^{+0.1}_{-0.3}$	$0.34^{+0.52}_{-0.31}$	371^{+685}_{-319}	$12.8^{+3.2}_{-3.9}$	$81.2^{+3.4}_{-16.8}$	14^{+28}_{-11}
XHM	$1.95^{+0.00}_{-0.00}$	$3.3^{+0.3}_{-0.3}$	$4.2^{+0.2}_{-0.2}$	$1.3^{+0.1}_{-0.1}$	$5.5^{+0.2}_{-0.2}$	$0.04^{+0.03}_{-0.03}$	$0.04^{+0.10}_{-0.04}$	$0.0^{+0.1}_{-0.1}$	$0.13^{+0.39}_{-0.13}$	-	-	$75.3^{+6.0}_{-8.3}$	27^{+12}_{-10}

(a) Inclination 30° .

Approximant	$M^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\text{ch}}^{\text{source}}/M_{\text{ch}}^{\text{total}}/M_{\odot}$	$ s_1 $	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	1.96	3.0	4.02	1.34	5.36	0.00	0.0	0.0	0.0	624	12.3	36.6	70
SEOB	$1.95^{+0.01}_{-0.01}$	$3.1^{+0.3}_{-0.3}$	$4.1^{+0.2}_{-0.2}$	$1.3^{+0.1}_{-0.1}$	$5.4^{+0.2}_{-0.1}$	$0.02^{+0.03}_{-0.03}$	$0.04^{+0.11}_{-0.03}$	$0.0^{+0.1}_{-0.1}$	$0.10^{+0.39}_{-0.10}$	-	-	$65.5^{+15.8}_{-27.1}$	39^{+32}_{-28}
SEOBt	$1.95^{+0.01}_{-0.01}$	$2.8^{+0.4}_{-0.6}$	$3.8^{+0.3}_{-0.5}$	$1.4^{+0.2}_{-0.1}$	$5.2^{+0.2}_{-0.3}$	$-0.02^{+0.04}_{-0.07}$	$0.05^{+0.18}_{-0.05}$	$-0.0^{+0.1}_{-0.2}$	$0.13^{+0.45}_{-0.12}$	357^{+577}_{-318}	$11.6^{+3.2}_{-3.4}$	$65.8^{+15.6}_{-27.8}$	39^{+33}_{-28}
SEOBNSBH	$1.95^{+0.01}_{-0.01}$	$2.9^{+0.3}_{-0.5}$	$3.9^{+0.2}_{-0.4}$	$1.4^{+0.1}_{-0.1}$	$5.3^{+0.2}_{-0.2}$	$-0.01^{+0.03}_{-0.05}$	$0.02^{+0.07}_{-0.02}$	$-0.0^{+0.0}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	474^{+362}_{-350}	$12.0^{+2.0}_{-2.5}$	$66.7^{+14.0}_{-27.5}$	37^{+33}_{-26}
LEA+	$1.95^{+0.01}_{-0.01}$	$3.0^{+0.4}_{-0.3}$	$4.0^{+0.2}_{-0.2}$	$1.3^{+0.1}_{-0.1}$	$5.3^{+0.2}_{-0.2}$	$0.00^{+0.04}_{-0.04}$	$0.02^{+0.04}_{-0.02}$	$0.0^{+0.0}_{-0.1}$	-	891^{+510}_{-570}	$13.1^{+1.6}_{-2.4}$	$64.7^{+16.6}_{-26.9}$	40^{+32}_{-29}
IMRNSBH	$1.95^{+0.01}_{-0.01}$	$2.9^{+0.3}_{-0.5}$	$3.9^{+0.2}_{-0.4}$	$1.3^{+0.1}_{-0.1}$	$5.3^{+0.2}_{-0.3}$	$0.00^{+0.03}_{-0.05}$	$0.02^{+0.06}_{-0.02}$	$0.0^{+0.0}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	852^{+357}_{-407}	$13.1^{+1.7}_{-1.8}$	$67.7^{+13.7}_{-28.3}$	36^{+33}_{-26}
IMRp	$1.94^{+0.01}_{-0.00}$	$3.4^{+0.3}_{-0.2}$	$4.3^{+0.2}_{-0.2}$	$1.3^{+0.0}_{-0.0}$	$5.5^{+0.1}_{-0.1}$	$0.04^{+0.02}_{-0.03}$	$0.08^{+0.16}_{-0.06}$	$0.0^{+0.1}_{-0.1}$	$0.27^{+0.51}_{-0.24}$	-	-	$78.7^{+4.1}_{-21.5}$	17^{+32}_{-13}
IMRpT	$1.94^{+0.01}_{-0.00}$	$2.8^{+0.4}_{-0.0}$	$3.8^{+0.3}_{-1.2}$	$1.4^{+0.5}_{-1.4}$	$5.2^{+0.2}_{-0.7}$	$-0.02^{+0.05}_{-0.13}$	$0.07^{+0.30}_{-0.07}$	$-0.0^{+0.1}_{-0.3}$	$0.23^{+0.55}_{-0.21}$	530^{+639}_{-422}	$12.9^{+3.0}_{-3.4}$	$78.7^{+4.5}_{-33.3}$	18^{+47}_{-14}
XHM	$1.96^{+0.00}_{-0.01}$	$3.3^{+0.3}_{-0.3}$	$4.3^{+0.2}_{-0.2}$	$1.3^{+0.1}_{-0.0}$	$5.5^{+0.2}_{-0.2}$	$0.04^{+0.03}_{-0.03}$	$0.04^{+0.09}_{-0.04}$	$0.0^{+0.1}_{-0.1}$	$0.15^{+0.38}_{-0.14}$	-	-	$47.8^{+12.3}_{-10.8}$	61^{+13}_{-15}

(b) Inclination 70° .

TABLE X: Properties of the $q=3$, SNR 70 simulation, as recovered by the listed approximants. We report the 1D median, together with the symmetric the 90% credible interval.

Approximant	$M^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\oplus}n_2^{\text{source}}/M_{\oplus}M_{\text{total}}^{\text{source}}/M_{\oplus}$	$M_{\oplus}^{\text{source}}/M_{\oplus}$	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	2.65	6.0	8.00	1.33	9.33	0.00	0.0	0.0	0.0	526	13.3	229.5	30
SEOB	$2.652^{+0.043}_{-0.023}$	$6.1^{+0.9}_{-0.8}$	$8.1^{+0.7}_{-0.6}$	$1.3^{+0.1}_{-0.1}$	$9.4^{+0.6}_{-0.5}$	$0.01^{+0.06}_{-0.07}$	$0.03^{+0.08}_{-0.03}$	$0.0^{+0.1}_{-0.1}$	$0.15^{+0.45}_{-0.14}$	-	-	$221.4^{+44.0}_{-77.4}$	34^{+30}_{-25}
SEOBt	$2.650^{+0.043}_{-0.024}$	$5.4^{+1.1}_{-1.4}$	$7.5^{+0.8}_{-1.2}$	$1.4^{+0.2}_{-0.1}$	$8.9^{+0.7}_{-1.0}$	$-0.05^{+0.09}_{-0.14}$	$0.06^{+0.20}_{-0.06}$	$-0.1^{+0.1}_{-0.2}$	$0.18^{+0.49}_{-0.17}$	1990^{+1788}_{-1790}	$15.7^{+4.0}_{-5.2}$	$221.5^{+44.5}_{-78.6}$	34^{+31}_{-25}
SEOBNSBH	$2.651^{+0.041}_{-0.022}$	$5.8^{+0.7}_{-1.2}$	$7.9^{+0.6}_{-0.9}$	$1.4^{+0.1}_{-0.1}$	$9.2^{+0.5}_{-0.8}$	$-0.01^{+0.05}_{-0.10}$	$0.03^{+0.11}_{-0.03}$	$-0.0^{+0.1}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	1938^{+2524}_{-1691}	$15.3^{+3.8}_{-4.8}$	$221.9^{+40.9}_{-75.4}$	33^{+29}_{-23}
IMRNSBH	$2.648^{+0.042}_{-0.022}$	$6.0^{+0.7}_{-1.0}$	$8.0^{+0.5}_{-0.8}$	$1.3^{+0.1}_{-0.1}$	$9.3^{+0.5}_{-0.7}$	$0.00^{+0.05}_{-0.08}$	$0.02^{+0.08}_{-0.02}$	$0.0^{+0.1}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	2591^{+2120}_{-2212}	$15.9^{+3.0}_{-4.8}$	$227.3^{+40.8}_{-76.9}$	33^{+29}_{-23}
IMRp	$2.640^{+0.040}_{-0.014}$	$6.2^{+0.6}_{-0.5}$	$8.1^{+0.5}_{-0.4}$	$1.3^{+0.0}_{-0.1}$	$9.5^{+0.4}_{-0.4}$	$0.01^{+0.05}_{-0.04}$	$0.05^{+0.12}_{-0.04}$	$0.0^{+0.1}_{-0.1}$	$0.32^{+0.54}_{-0.28}$	-	-	$243.9^{+25.8}_{-75.2}$	22^{+33}_{-17}
IMRpT	$2.636^{+0.040}_{-0.014}$	$5.6^{+0.8}_{-3.2}$	$7.6^{+0.8}_{-2.9}$	$1.4^{+0.6}_{-0.1}$	$9.0^{+0.6}_{-2.3}$	$-0.04^{+0.07}_{-0.26}$	$0.09^{+0.43}_{-0.08}$	$-0.0^{+0.1}_{-0.4}$	$0.37^{+0.53}_{-0.33}$	1710^{+1986}_{-1527}	$15.7^{+5.3}_{-5.3}$	$248.3^{+23.9}_{-73.6}$	20^{+33}_{-15}
XHM	$2.644^{+0.020}_{-0.014}$	$6.3^{+0.8}_{-0.7}$	$8.2^{+0.6}_{-0.5}$	$1.3^{+0.1}_{-0.1}$	$9.5^{+0.5}_{-0.5}$	$0.02^{+0.05}_{-0.05}$	$0.03^{+0.07}_{-0.03}$	$0.0^{+0.1}_{-0.1}$	$0.13^{+0.34}_{-0.13}$	-	-	$237.1^{+27.3}_{-37.0}$	28^{+14}_{-13}

(a) Inclination 30° .

Approximant	$M^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\oplus}n_2^{\text{source}}/M_{\oplus}M_{\text{total}}^{\text{source}}/M_{\oplus}$	$M_{\oplus}^{\text{source}}/M_{\oplus}$	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	2.72	6.0	8.23	1.37	9.60	0.00	0.0	0.0	0.0	526	13.3	95.4	70
SEOB	$2.651^{+0.045}_{-0.023}$	$6.1^{+0.9}_{-0.8}$	$8.1^{+0.7}_{-0.6}$	$1.3^{+0.1}_{-0.1}$	$9.4^{+0.6}_{-0.5}$	$0.01^{+0.06}_{-0.07}$	$0.03^{+0.08}_{-0.03}$	$0.0^{+0.1}_{-0.1}$	$0.15^{+0.47}_{-0.15}$	-	-	$222.5^{+44.3}_{-80.5}$	34^{+29}_{-25}
SEOBt	$2.648^{+0.043}_{-0.023}$	$5.4^{+1.0}_{-1.4}$	$7.6^{+0.8}_{-1.2}$	$1.4^{+0.2}_{-0.1}$	$9.0^{+0.7}_{-1.0}$	$-0.05^{+0.09}_{-0.14}$	$0.06^{+0.19}_{-0.05}$	$-0.0^{+0.1}_{-0.2}$	$0.16^{+0.48}_{-0.16}$	1936^{+1857}_{-1738}	$15.6^{+3.9}_{-5.2}$	$224.3^{+43.3}_{-79.0}$	34^{+29}_{-24}
SEOBNSBH	$2.650^{+0.040}_{-0.021}$	$5.8^{+0.7}_{-1.1}$	$7.9^{+0.6}_{-0.9}$	$1.4^{+0.1}_{-0.1}$	$9.2^{+0.5}_{-0.8}$	$-0.01^{+0.05}_{-0.10}$	$0.03^{+0.11}_{-0.03}$	$-0.0^{+0.1}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	1864^{+2526}_{-1619}	$15.1^{+3.8}_{-4.6}$	$223.0^{+40.8}_{-73.8}$	33^{+27}_{-23}
IMRNSBH	$2.647^{+0.042}_{-0.021}$	$6.0^{+0.7}_{-0.9}$	$8.0^{+0.5}_{-0.7}$	$1.3^{+0.1}_{-0.1}$	$9.3^{+0.5}_{-0.6}$	$0.00^{+0.05}_{-0.08}$	$0.02^{+0.07}_{-0.02}$	$0.0^{+0.1}_{-0.1}$	$0.00^{+0.00}_{-0.00}$	2449^{+2227}_{-2108}	$15.7^{+3.1}_{-4.8}$	$229.5^{+40.2}_{-77.5}$	32^{+28}_{-23}
IMRp	$2.639^{+0.039}_{-0.014}$	$6.2^{+0.7}_{-0.5}$	$8.2^{+0.5}_{-0.4}$	$1.3^{+0.0}_{-0.1}$	$9.5^{+0.5}_{-0.3}$	$0.01^{+0.05}_{-0.04}$	$0.05^{+0.12}_{-0.04}$	$0.0^{+0.1}_{-0.1}$	$0.30^{+0.56}_{-0.28}$	-	-	$245.0^{+25.9}_{-72.9}$	23^{+31}_{-17}
IMRpT	$2.634^{+0.037}_{-0.013}$	$5.4^{+1.0}_{-3.1}$	$7.5^{+0.8}_{-2.9}$	$1.4^{+0.6}_{-0.1}$	$8.9^{+0.7}_{-2.2}$	$-0.05^{+0.09}_{-0.28}$	$0.11^{+0.45}_{-0.10}$	$-0.0^{+0.1}_{-0.4}$	$0.41^{+0.50}_{-0.36}$	1549^{+2108}_{-1389}	$15.8^{+5.9}_{-5.3}$	$250.8^{+22.8}_{-67.5}$	19^{+31}_{-14}
XHM	$2.685^{+0.031}_{-0.027}$	$6.4^{+0.8}_{-0.6}$	$8.4^{+0.6}_{-0.5}$	$1.3^{+0.1}_{-0.1}$	$9.8^{+0.5}_{-0.5}$	$0.03^{+0.05}_{-0.05}$	$0.04^{+0.08}_{-0.04}$	$0.0^{+0.1}_{-0.1}$	$0.14^{+0.37}_{-0.13}$	-	-	$162.3^{+48.7}_{-34.6}$	54^{+14}_{-15}

(b) Inclination 70° .

TABLE XI: Properties of the $q=6$, SNR 30 simulation, as recovered by the listed approximants. We report the 1D median, together with the symmetric the 90% credible interval.

Approximant	$M^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\oplus}^{\text{source}}/M_{\oplus}^{\text{total}}/M_{\oplus}$	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	2.72	6.0	8.22	1.37	9.59	0.00	0.0	0.0	526	13.3	98.4	30
SEOB	$2.721^{+0.016}_{-0.009}$	$6.1^{+0.4}_{-0.4}$	$8.3^{+0.3}_{-0.3}$	$1.4^{+0.0}_{-0.0}$	$9.7^{+0.3}_{-0.3}$	$0.01^{+0.03}_{-0.03}$	$0.02^{+0.06}_{-0.02}$	$0.0^{+0.1}_{-0.1}$	$0.13^{+0.42}_{-0.13}$	$0.0^{+0.5}_{-0.4}$	$97.2^{+14.9}_{-26.5}$	31^{+24}_{-23}
SEOBt	$2.719^{+0.016}_{-0.009}$	$5.4^{+0.6}_{-0.6}$	$7.7^{+0.5}_{-0.5}$	$1.4^{+0.1}_{-0.1}$	$9.1^{+0.4}_{-0.4}$	$-0.04^{+0.05}_{-0.05}$	$0.08^{+0.13}_{-0.07}$	$-0.1^{+0.1}_{-0.1}$	$0.21^{+0.55}_{-0.20}$	$0.1^{+0.6}_{-0.4}$	921^{+2273}_{-852}	$14.1^{+4.2}_{-4.6}$ $97.8^{+14.6}_{-26.6}$ 31^{+24}_{-22}
SEOBNSBH	$2.720^{+0.015}_{-0.008}$	$6.0^{+0.4}_{-0.4}$	$8.2^{+0.3}_{-0.4}$	$1.4^{+0.0}_{-0.0}$	$9.6^{+0.3}_{-0.3}$	$0.00^{+0.03}_{-0.04}$	$0.01^{+0.03}_{-0.01}$	$0.0^{+0.0}_{-0.0}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}_{-0.0}$	979^{+1225}_{-831}	$13.6^{+2.6}_{-3.6}$ $99.0^{+13.6}_{-25.7}$ 30^{+24}_{-21}
IMRNSBH	$2.720^{+0.015}_{-0.008}$	$6.1^{+0.4}_{-0.5}$	$8.3^{+0.3}_{-0.4}$	$1.4^{+0.0}_{-0.0}$	$9.7^{+0.3}_{-0.3}$	$0.00^{+0.03}_{-0.03}$	$0.02^{+0.03}_{-0.01}$	$0.0^{+0.0}_{-0.0}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}_{-0.0}$	1562^{+2058}_{-1281}	$14.6^{+3.1}_{-3.9}$ $99.5^{+13.9}_{-25.8}$ 30^{+23}_{-21}
IMRp	$2.715^{+0.013}_{-0.003}$	$6.3^{+0.4}_{-0.2}$	$8.4^{+0.3}_{-0.2}$	$1.3^{+0.0}_{-0.0}$	$9.8^{+0.3}_{-0.2}$	$0.02^{+0.03}_{-0.02}$	$0.04^{+0.09}_{-0.03}$	$0.0^{+0.1}_{-0.0}$	$0.24^{+0.51}_{-0.21}$	$0.1^{+0.4}_{-0.3}$	-	$108.1^{+5.6}_{-23.2}$ 17^{+27}_{-13}
IMRpT	$2.708^{+0.004}_{-0.003}$	$2.9^{+0.9}_{-0.8}$	$5.5^{+0.9}_{-0.9}$	$1.9^{+0.3}_{-0.2}$	$7.3^{+0.7}_{-0.6}$	$-0.18^{+0.09}_{-0.11}$	$0.42^{+0.16}_{-0.14}$	$-0.4^{+0.1}_{-0.1}$	$0.66^{+0.29}_{-0.50}$	$0.4^{+0.3}_{-0.3}$	262^{+801}_{-247}	$14.9^{+3.9}_{-4.5}$ $112.8^{+3.5}_{-7.2}$ 8^{+16}_{-6}
XHM	$2.720^{+0.005}_{-0.004}$	$6.4^{+0.5}_{-0.4}$	$8.5^{+0.4}_{-0.3}$	$1.3^{+0.0}_{-0.0}$	$9.9^{+0.3}_{-0.3}$	$0.03^{+0.03}_{-0.03}$	$0.03^{+0.06}_{-0.03}$	$0.0^{+0.1}_{-0.1}$	$0.14^{+0.41}_{-0.14}$	$0.1^{+0.5}_{-0.4}$	-	$100.5^{+6.4}_{-8.0}$ 29^{+7}_{-7}

(a) Inclination 30° .

Approximant	$M^{\text{source}}/M_{\odot}$	q	$m_1^{\text{source}}/M_{\oplus}^{\text{source}}/M_{\oplus}^{\text{total}}/M_{\oplus}$	χ_{eff}	$ s_1 $	$s_{1,z}$	$ s_2 $	$s_{2,z}$	Λ_2	r_{NS}/km	D_L/Mpc	$\theta_{JN}/^{\circ}$
True value	2.75	6.0	8.32	1.39	9.71	0.00	0.0	0.0	526	13.3	40.9	70
SEOB	$2.727^{+0.025}_{-0.014}$	$6.2^{+0.4}_{-0.4}$	$8.4^{+0.3}_{-0.3}$	$1.4^{+0.0}_{-0.0}$	$9.7^{+0.3}_{-0.3}$	$0.01^{+0.03}_{-0.03}$	$0.02^{+0.07}_{-0.02}$	$0.0^{+0.1}_{-0.1}$	$0.14^{+0.43}_{-0.13}$	$0.0^{+0.5}_{-0.4}$	$87.1^{+24.6}_{-42.1}$	41^{+27}_{-30}
SEOBt	$2.725^{+0.025}_{-0.014}$	$5.3^{+0.7}_{-0.6}$	$7.7^{+0.6}_{-0.5}$	$1.4^{+0.1}_{-0.1}$	$9.2^{+0.5}_{-0.5}$	$-0.04^{+0.05}_{-0.05}$	$0.08^{+0.14}_{-0.08}$	$-0.1^{+0.1}_{-0.1}$	$0.24^{+0.56}_{-0.23}$	$0.2^{+0.6}_{-0.5}$	973^{+2407}_{-911}	$14.2^{+4.3}_{-4.8}$ $87.6^{+24.5}_{-42.0}$ 41^{+27}_{-30}
SEOBNSBH	$2.725^{+0.026}_{-0.013}$	$6.0^{+0.3}_{-0.4}$	$8.2^{+0.3}_{-0.3}$	$1.4^{+0.0}_{-0.0}$	$9.6^{+0.2}_{-0.3}$	$0.00^{+0.03}_{-0.03}$	$0.01^{+0.03}_{-0.01}$	$0.0^{+0.0}_{-0.0}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}_{-0.0}$	986^{+807}_{-829}	$13.7^{+1.9}_{-3.5}$ $89.5^{+23.5}_{-43.0}$ 39^{+28}_{-28}
IMRNSBH	$2.726^{+0.026}_{-0.014}$	$6.2^{+0.4}_{-0.4}$	$8.4^{+0.3}_{-0.3}$	$1.4^{+0.0}_{-0.0}$	$9.7^{+0.3}_{-0.3}$	$0.01^{+0.03}_{-0.03}$	$0.02^{+0.03}_{-0.02}$	$0.0^{+0.0}_{-0.0}$	$0.00^{+0.00}_{-0.00}$	$0.0^{+0.0}_{-0.0}$	1350^{+1672}_{-1106}	$14.2^{+2.8}_{-3.6}$ $89.7^{+23.5}_{-43.2}$ 40^{+28}_{-29}
IMRp	$2.715^{+0.018}_{-0.003}$	$6.4^{+0.4}_{-0.3}$	$8.5^{+0.3}_{-0.2}$	$1.3^{+0.0}_{-0.0}$	$9.8^{+0.2}_{-0.2}$	$0.03^{+0.03}_{-0.02}$	$0.04^{+0.08}_{-0.04}$	$0.0^{+0.1}_{-0.0}$	$0.29^{+0.47}_{-0.26}$	$0.1^{+0.3}_{-0.3}$	-	$108.9^{+5.2}_{-32.0}$ 16^{+34}_{-12}
IMRpT	$2.708^{+0.005}_{-0.003}$	$2.6^{+1.0}_{-0.8}$	$5.1^{+1.0}_{-0.9}$	$2.0^{+0.4}_{-0.3}$	$7.1^{+0.8}_{-0.6}$	$-0.22^{+0.12}_{-0.11}$	$0.47^{+0.24}_{-0.16}$	$-0.4^{+0.1}_{-0.2}$	$0.48^{+0.41}_{-0.36}$	$0.3^{+0.3}_{-0.2}$	186^{+713}_{-178}	$14.9^{+4.2}_{-4.4}$ $113.2^{+3.6}_{-7.5}$ 8^{+17}_{-6}
XHM	$2.748^{+0.008}_{-0.009}$	$6.4^{+0.5}_{-0.4}$	$8.6^{+0.4}_{-0.3}$	$1.3^{+0.0}_{-0.0}$	$10.0^{+0.3}_{-0.3}$	$0.03^{+0.02}_{-0.02}$	$0.03^{+0.07}_{-0.03}$	$0.0^{+0.1}_{-0.1}$	$0.16^{+0.41}_{-0.15}$	$0.1^{+0.5}_{-0.4}$	-	$52.5^{+15.4}_{-13.8}$ 64^{+8}_{-9}

(b) Inclination 70° .

TABLE XII: Properties of the $q=6$, SNR 70 simulation, as recovered by the listed approximants. We report the 1D median, together with the symmetric the 90% credible interval.

Approximant	$\mathcal{M}^{\text{det}}/M_{\odot}$	q	$m_1^{\text{det}}/M_{\odot}$	$m_2^{\text{det}}/M_{\odot}$	s_1	s_2	Λ_2	D_L/Mpc
SEOB	[1.480,2.711]	[1.0,8.0]	[1.043, 9.664]	[1.043, 9.664]	[-0.99,0.99]	[-0.99,0.99]	-	[0,500]
SEOBT	[1.480,2.711]	[1.0,8.0]	[1.043, 9.664]	[1.043, 9.664]	[-0.99,0.99]	[-0.99,0.99]	[0,4000]	[0,500]
SEOBNSBH	[1.480,2.711]	[1.0,8.0]	[1.043, 9.664]	[1.043, 3.000]	[-0.5,0.8]	-	[0,5000]	[0,500]
IMRNSBH	[1.480,2.711]	[1.0,8.0]	[1.043, 9.664]	[1.043, 3.000]	[-0.5,0.5]	-	[0,5000]	[0,500]
IMRp	[1.421,2.602]	[1.0,8.0]	[1.001, 9.277]	[1.001, 9.277]	[-0.99,0.99]	[-0.99,0.99]	-	[0,500]
IMRpT	[1.421,2.602]	[1.0,8.0]	[1.001, 9.277]	[1.001, 9.277]	[-0.99,0.99]	[-0.99,0.99]	[0,4000]	[0,500]
XHM	[1.480,2.711]	[1.0,8.0]	[1.043, 9.664]	[1.043, 9.664]	[-0.99,0.99]	[-0.99,0.99]	-	[0,500]

TABLE XIII: Prior bounds for the $q=2$ simulation.

Approximant	$\mathcal{M}^{\text{det}}/M_{\odot}$	q	$m_1^{\text{det}}/M_{\odot}$	$m_2^{\text{det}}/M_{\odot}$	s_1	s_2	Λ_2	D_L/Mpc
SEOB	[1.776,3.253]	[1.0,8.0]	[1.252,11.597]	[1.252,11.597]	[-0.99,0.99]	[-0.99,0.99]	-	[0,500]
SEOBT	[1.776,3.253]	[1.0,8.0]	[1.252,11.597]	[1.252,11.597]	[-0.99,0.99]	[-0.99,0.99]	[0,4000]	[0,500]
SEOBNSBH	[1.776,3.253]	[1.0,8.0]	[1.252,11.597]	[1.252, 3.000]	[-0.5,0.8]	-	[0,5000]	[0,500]
LEA+	-	[2.0,5.0]	[3.0,7.2]	[1.2,1.45]	[-0.5,0.5]	-	[0,4000]	[0,500]
IMRNSBH	[1.776,3.253]	[1.0,8.0]	[1.252,11.597]	[1.252, 3.000]	[-0.5,0.5]	-	[0,5000]	[0,500]
IMRp	[1.421,2.602]	[1.0,8.0]	[1.001, 9.277]	[1.001, 9.277]	[-0.99,0.99]	[-0.99,0.99]	-	[0,500]
IMRpT	[1.421,2.602]	[1.0,8.0]	[1.001, 9.277]	[1.001, 9.277]	[-0.99,0.99]	[-0.99,0.99]	[0,4000]	[0,500]
XHM	[1.776,3.253]	[1.0,8.0]	[1.252,11.597]	[1.252,11.597]	[-0.99,0.99]	[-0.99,0.99]	-	[0,500]

TABLE XIV: Prior bounds for the $q=3$ simulation.

Approximant	$\mathcal{M}^{\text{det}}/M_{\odot}$	q	$m_1^{\text{det}}/M_{\odot}$	$m_2^{\text{det}}/M_{\odot}$	s_1	s_2	Λ_2	D_L/Mpc
SEOB	[2.184,4.016]	[1.000,17.944]	[1.000,22.953]	[1.000,22.953]	[-0.99,0.99]	[-0.99,0.99]	-	[0,500]
SEOBT	[2.184,4.016]	[1.000,17.944]	[1.000,22.953]	[1.000,22.953]	[-0.99,0.99]	[-0.99,0.99]	[0,4000]	[0,500]
SEOBNSBH	[2.184,4.016]	[1.000,17.944]	[1.000,22.953]	[1.000,3.000]	[-0.5,0.8]	-	[0,5000]	[0,500]
IMRNSBH	[2.184,4.016]	[1.000,17.944]	[1.000,22.953]	[1.000,3.000]	[-0.5,0.5]	-	[0,5000]	[0,500]
IMRp	[2.184,4.016]	[1.000,8.000]	[1.001,14.317]	[1.001,14.317]	[-0.99,0.99]	[-0.99,0.99]	-	[0,500]
IMRpT	[2.184,4.016]	[1.000,8.000]	[1.001,14.317]	[1.001,14.317]	[-0.99,0.99]	[-0.99,0.99]	[0,4000]	[0,500]
XHM	[2.184,4.016]	[1.000,17.944]	[1.000,22.953]	[1.000,22.953]	[-0.99,0.99]	[-0.99,0.99]	-	[0,500]

TABLE XV: Prior bounds for the $q=6$ simulation.

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