**Erratum: LISA for cosmologists: Calculating the signal-to-noise
ratio for stochastic and deterministic sources
[Phys. Rev. D **100**, 104055 (2019)]**

Tristan L. Smith and Robert R. Caldwell

(Received 28 September 2021; published 21 January 2022)

DOI: [10.1103/PhysRevD.105.029902](https://doi.org/10.1103/PhysRevD.105.029902)

In preparation for our paper, a factor of $1/2$ was inconsistently implemented in the definition of the quantity F_{ABC}^P in Eq. (11). This does not change the validity of the derivations that we present, but does impact the examples in which we compute the signal to noise ratio (SNR). The factor of $1/2$ means that Eq. (14) gives the detector response per one-way interferometer path, whereas Eqs. (55) and (56) give the detector noise per round-trip path. A simple way to make the detector noise consistent with the response functions is to reduce the noise spectral densities in Eqs. (53) and (54) by a factor of 4, since both expressions are quadratic in path length:

$$S_I \rightarrow S_I/4 \quad (1)$$

$$S_{II} \rightarrow S_{II}/4. \quad (2)$$

With this change the revised Eqs. (55) and (56) give the detector noise per one-way path. This change brings the numerical values of the SNR into agreement with other computations in the literature (e.g., Refs. [1–3]).

In Sec. VIIIA for the example of a flat stochastic gravitational wave background (SGWB), we obtain $\Omega_{\text{GW}} = 1.18 \times 10^{-13}$ or $\Omega_{\text{GW}} h^2 = 5.30 \times 10^{-14}$ for $T = 4$ yr, and $\Omega_{\text{GW}} = 1.36 \times 10^{-13}$ or $\Omega_{\text{GW}} h^2 = 6.12 \times 10^{-14}$ for $T = 3$ yr.

In Sec. VIIC, for the example of binary inspiral in a GW150914-like system using Eqs. (86) and (87) the SNR is 4.64, 6.72 for 4, 10 yr, whereas using Eqs. (85) and (86) the SNR is 5.33, 7.51.

In Sec. VIID, for the example of a monotone binary, the noise power spectral density is $\Sigma_h(f_m) = 7.2 \times 10^{-41} \text{ Hz}^{-1}$. For $T = 1, 4$ yr observation, the sky- and polarization-averaged SNR is 80, 160.

We list all of the quantities that are affected by this change in Table I. The minimum detectable amplitudes all refer to SNR = 5.

A fully revised manuscript may be found on the arXiv or obtained through a request to the authors.

TABLE I. List of all quantities that are affected by the corrected response function.

Corrected quantitative results		
Quantity	Value quoted in original paper	Corrected value
Minimum detectable SGWB amplitude 100% duty cycle (Ω_{GW})	4.7×10^{-13}	1.18×10^{-13}
Minimum detectable SGWB amplitude 100% duty cycle ($\Omega_{\text{GW}} h^2$)	2.1×10^{-13}	5.30×10^{-14}
Minimum detectable SGWB amplitude 75% duty cycle (Ω_{GW})	5.4×10^{-13}	1.36×10^{-13}
Minimum detectable SGWB amplitude 75% duty cycle ($\Omega_{\text{GW}} h^2$)	2.4×10^{-13}	6.12×10^{-14}
Minimum detectable integrated SGWB amplitude four-year mission (Ω_{GW})	4.7×10^{-13}	1.2×10^{-13}
SNR for GW150914 $T = 4$ yr(10 yr) before merger, Eqs. (86), (87)	2.3(3.4)	4.64(6.72)
SNR for GW150914 $T = 4$ yr(10 yr) before merger, Eqs. (85), (86)	2.7(3.8)	5.33(7.51)
Noise power spectral density [$\Sigma_h(f_m)$] for ZTFJ1539	$1.5 \times 10^{-40} \text{ Hz}^{-1}$	$7.2 \times 10^{-41} \text{ Hz}^{-1}$
SNR for ZTFJ1539 $T = 1$ yr(4 yr)	40(80)	80(160)

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